



# VERY HIGH ENERGY OBSERVATIONS OF THE BINARIES V 404 CYG AND 4U 0115+634 DURING GIANT X-RAY OUTBURSTS

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## ABSTRACT

Transient X-ray binaries produce major outbursts in which the X-ray flux can increase over the quiescent level by factors as large as  $10^7$ . The low-mass X-ray binary V 404 Cyg and the high-mass system 4U 0115+634 underwent such major outbursts in 2015 June and October, respectively. We present here observations at energies above hundreds of GeV with the VERITAS observatory taken during some of the brightest X-ray activity ever observed from these systems. No gamma-ray emission has been detected by VERITAS in 2.5 hr of observations of the microquasar V 404 Cyg from 2015, June 20–21. The upper flux limits derived from these observations on the gamma-ray flux above 200 GeV of  $F < 4.4 \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$  correspond to a tiny fraction (about  $10^{-6}$ ) of the Eddington luminosity of the system, in stark contrast to that seen in the X-ray band. No gamma-rays have been detected during observations of 4U 0115+634 in the period of major X-ray activity in 2015 October. The flux upper limit derived from our observations is  $F < 2.1 \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$  for gamma-rays above 300 GeV, setting an upper limit on the ratio of gamma-ray to X-ray luminosity of less than 4%.

**Key words:** acceleration of particles – binaries: general – gamma rays: general – stars: individual (V 404 Cyg, 4U 0115+634)

## 1. INTRODUCTION

Variable very high energy gamma-ray emission ( $>100 \text{ GeV}$ , VHE) has been observed from several X-ray binary systems (LS I +61 303, LS 5039, HESS J0632+057, 1FGL J1018.65856, PSR 1259–63, and tentatively Cygnus X-1; see Dubus 2013, 2015 for recent reviews). All are high-mass X-ray binaries with an O or Be star as stellar companions. The observed gamma-ray emission mostly exhibits modulation with

the orbital period (e.g., Aharonian et al. 2006), but short intensive flares (e.g., Archambault et al. 2016) have also been observed from some of the gamma-ray binaries. The origin of the high-energy emission and the acceleration mechanism for the underlying population of relativistic particles are not well understood. They depend, among other features, on the nature of the compact object and the massive star, the geometry of the orbit and, if applicable, the state of the accretion disk. Two

major scenarios can be found among the wealth of theoretical work (see Dubus 2013 for a review): in the microquasar model, charged particles are accelerated in an accretion-driven relativistic jet, similar to the processes observed in active galactic nuclei (AGNs). Alternatively, the high-energy emission may be driven by shock interaction between a rotation-powered pulsar wind and the strong wind of the massive stellar companion.

Binaries can exhibit huge outbursts during which the X-ray fluxes increase by factors of  $10\text{--}10^7$  compared to the flux observed during the quiescent stage (Remillard & McClintock 2006). The systems sometimes reach or even exceed their Eddington luminosity and temporarily become the brightest objects in the X-ray sky. In general, the outbursts show no preferred orbital phase and might last for several orbital periods. VHE emission from X-ray binaries during outbursts has been predicted by a few authors, e.g., Orellana et al. (2007). Therein, a process is proposed which is based on a mechanism first presented by Cheng & Ruderman (1989): a hadronic beam is accelerated in the pulsar magnetosphere and impacts on the transient accretion disk. Neutral pions produced in this interaction decay into a possibly observable flux of gamma-rays.

The known VHE-emitting binaries show strong variability in X-rays. The correlation of the VHE emission with the X-ray flux is unclear: some binaries exhibit clear correlations (e.g. HESS J0632+057 Aliu et al. 2014a), while in others the correlation is not observed during all periods (e.g. in LS I +61 303, Anderhub et al. 2009; Acciari et al. 2011b). LS I +61 303 stands out among the VHE binaries, with strong and short X-ray flares observed with flux doubling times of a few seconds (Smith et al. 2009). For low-mass X-ray binaries, no VHE gamma-ray emission has been discovered during giant outbursts. However, the sheer amount of energy released in the X-ray range during the giant flares is a strong motivation to observe these types of object at gamma-ray energies (Levinson & Blandford 1996). In this paper, observations above a few hundred GeV with the VERITAS observatory of the two binaries V 404 Cyg and 4U 0115+634 during major outbursts in 2015 are presented.

V 404 Cyg (Ginga 2023+338) is a low-mass X-ray binary system (LMXB) comprised of an evolved K-type star, almost filling its Roche lobe, and a black hole (Casares et al. 1992). The orbital period of the system is 6.47 days, and the binary mass function is  $f(M) = 6.07 \pm 0.05$ , with typical estimates of the black-hole mass around  $9 M_{\odot}$  or larger (Casares & Jonker 2014). Its distance of  $2.39 \pm 0.14$  kpc is very precisely known through parallax measurements (Miller-Jones et al. 2009). V 404 Cyg is among the brightest LMXBs during quiescence. The source underwent transient flaring periods in the optical in 1938, 1956 and 1989. X-ray emission was first observed during an outburst in 1989 by the Ginga X-ray observatory (Makino et al. 1989). The intense, correlated, rapid variability across the spectrum during outbursts is believed to be related to the accretion disk or to jet-like ejection events in the system (Munoz-Darias et al. 2016). Similarities with the non-thermal emission in the XRBs Cyg X-1 and GRS 1915+105 have been identified by Marti et al. (2016) and Roques et al. (2015). Variable X-ray emission in V 404 Cyg is also observed during quiescence, the mechanism of which is not well understood (Bernardini & Cackett 2014).

On 2015 June 15, at 18:32 UT (MJD 57188.772), the start of a new outburst from V 404 Cyg was observed by the X-ray observatory *Swift* (GCN#17929). In the hard X-ray band, the source reached unprecedented flux levels, up to 50 times that of the Crab Nebula. Segreto et al. (2015) estimate the 1–500 keV luminosity during the flare peak to be  $1.6 \times 10^{39} \text{ erg s}^{-1}$ , consistent with the Eddington luminosity of a  $12 M_{\odot}$  black hole system. The Gamma-ray Burst Monitor on board the *Fermi* satellite triggered 96 times in 10 days, with the flux reaching 3.7 Crab in the 100–300 keV band on June 19 (Jenke et al. 2015). The microquasar was not detected by the Large Area Telescope (LAT) on board the *Fermi* satellite during the flare period. Siebert et al. (2016) report flux upper limits for the period MJD 57199.616 to 57200.261 of  $8 \times 10^{-7} \text{ cm}^{-2} \text{ s}^{-1}$  for gamma rays with energies between 100 MeV and 1 GeV, and  $3 \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$  in the 1–10 GeV interval.

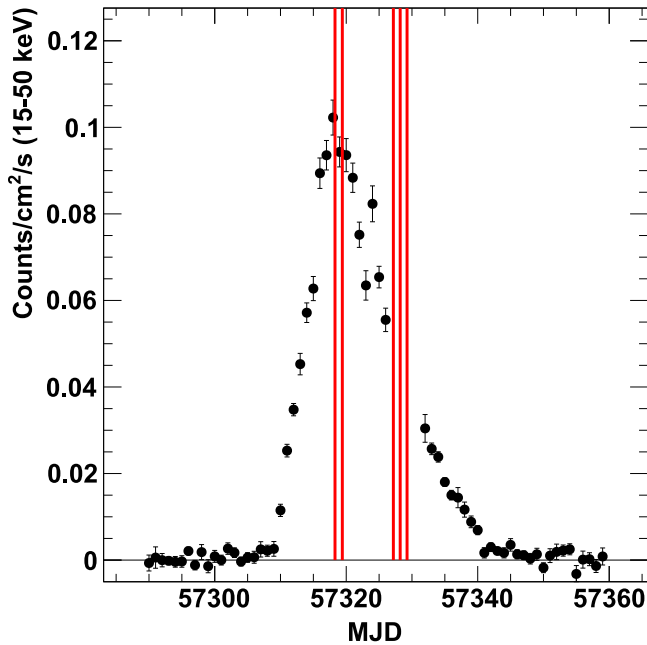
4U 0115+634 is a binary system consisting of an X-ray pulsar and the B0.2Ve star V635 Cas (Reig & Fabregat 2015) in an eccentric ( $e = 0.34$ ) orbit of 24.3 days (Rappaport et al. 1978). The period of the pulsar is 3.6 s. The distance to the system is estimated to be  $6.0 \pm 1.5$  kpc (Reig & Fabregat 2015). 4U 0115+634 was discovered during an outburst in 1970 by the *Uhuru* satellite and has since been observed to have major outbursts approximately every 4–5 years. The strength and frequency of the outburst are thought to be connected to the timescale of formation and dissipation of the disk of the Be star (Martin et al. 2014). The outbursts usually last  $\approx 1$  month, with luminosities up to  $\approx 10^{38} \text{ erg s}^{-1}$  (see, e.g., Tsygankov et al. 2007). The accreting neutron star exhibits several cyclotron resonant-scattering features, which lead to an estimate of the magnetic field strength of  $B > 10^{13} \text{ G}$  (Santangelo et al. 1999). The Whipple collaboration observed this system in quiescence for 124 hr between 1985 and 1988, and reported a 95% confidence upper limit of  $1 \times 10^{-11} \text{ cm}^{-2} \text{ s}^{-1}$  above 0.7 TeV (Macomb et al. 1991). The Gas Slit Camera of the Monitor of All-sky X-ray Image instrument on board the *International Space Station* reported on 2015 October 15 the onset of a new major outburst, reaching a peak of 0.5 Crab flux units in the 15–40 keV band about seven days after the initial alert (Nakajima et al. 2015).

## 2. VERITAS OBSERVATORY

The Very Energetic Radiation Imaging Telescope Array System VERITAS<sup>27</sup> is a ground-based gamma-ray observatory built to observe high-energy photons in the energy range from 85 GeV to  $>30$  TeV. The VERITAS observatory is located at the Fred Lawrence Whipple observatory in southern Arizona, USA (1.3 km above sea level, N31°40', W 110°57'). It consists of four 12 m diameter imaging atmospheric Cherenkov telescopes which observe simultaneously the faint Cherenkov light that is emitted when particle showers initiated by the impact of high-energy gamma rays pass through the atmosphere.

The flux sensitivity of the instrument reaches 1% of the flux of the Crab Nebula in 25 hr in the elevation range relevant for this paper (assuming a detection significance of five standard deviations). The sensitivity of VERITAS to short flares of high-energy emission exceeds significantly the sensitivity of satellite experiments at these energies due to an effective area of  $>10^4 \text{ m}^2$  above 100 GeV ( $>10^5 \text{ m}^2$  above 1 TeV). The angular

<sup>27</sup> <http://veritas.sao.arizona.edu/>



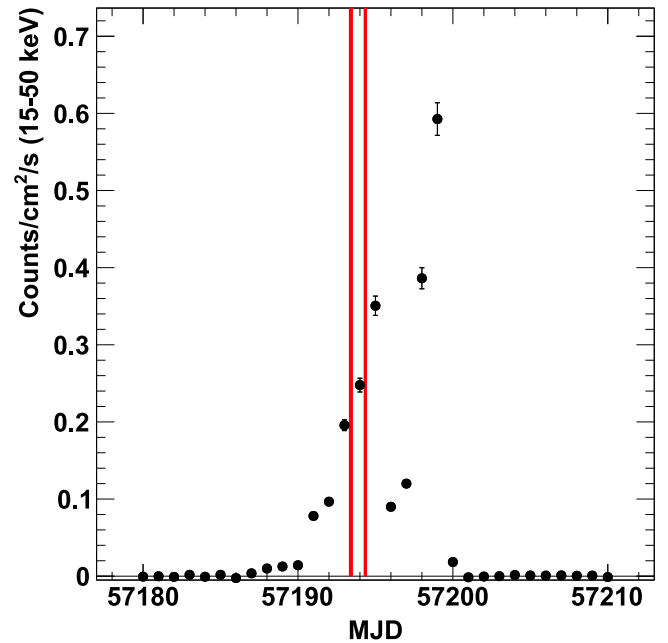
**Figure 1.** Daily average light curve (15–50 keV) of 4U 0115+634 for the major outburst in 2015 October obtained with the Burst Alert Telescope (BAT) on board the *Swift* satellite. The red lines indicate the time and length of the observing periods with VERITAS.

resolution is better than  $0.1^\circ$  at 1 TeV (Holder 2011). Observations with VERITAS are only possible during dark nights and moderate moonlight conditions, which constrains follow-up observations of flaring objects like 4U 0115+634 and V 404 Cyg. Some of the observations of 4U 0115+634 were taken during moonlight conditions, for which the higher background light levels lead to a lower sensitivity at the energy threshold (Griffin 2015). The analysis results for 4U 0115+634 are therefore given at a slightly higher energy threshold (300 GeV) than the results on V 404 Cyg (200 GeV).

### 3. OBSERVATIONS AND RESULTS

VERITAS observed 4U 0115+634 between 2015 October 23 and 2015 November 3 for about 5.5 hr, shortly after the first reports of the X-ray outbursts on 2015 October 15 (Nakajima et al. 2015). Figure 1 shows the X-ray light curve from *Swift* BAT and the VERITAS observing periods. These are also listed in detail in Table 1. All observing times are given after the application of quality-selection criteria, which remove data affected by technical problems. The observations with VERITAS were interrupted by the full-moon period (between 2015 October 25 and October 31), during which the background light levels are too high for the operation of the sensitive photomultiplier cameras of the observatory.

The search region from the putative gamma-ray source is defined by a circle with radius  $0.09^\circ$ . The background in the search region has been estimated from the same observations using the reflected region model (Fomin et al. 1994). No evidence for VHE gamma-ray emission has been found in the VERITAS data. The total number of on-source counts  $N_{\text{on}} = 82$  and off-source counts  $N_{\text{off}} = 567$  (with a ratio of off-to-on area of 6) result in a significance of  $-1.2\sigma$ , following the method for significance calculation in (Li & Ma 1983, Equation (17)). The flux upper limit at the 99% confidence level (Rolke & Lopez 2001) assuming a power-law-like source



**Figure 2.** Daily average *Swift* BAT light curve (15–50 keV) of V 404 Cyg for the major outburst in 2015 June. The red lines indicate the time and length of the VERITAS observing periods with good weather conditions.

spectrum with a spectral index of  $\Gamma = -2.5$  is  $F < 2.1 \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$  at energies above 300 GeV (corresponding to about 1.6% of the flux of the Crab Nebula above the same energy). Flux upper limits for each observing night can be found in Table 1.

VERITAS observed V 404 Cyg for about 3.3 hr during one of the main outbursts reported at lower energies on the nights of 2015 June 20, 21, and 24. Figure 2 indicates the times of the VERITAS observing periods during the X-ray outburst; all details of the covered time ranges are given in Table 2. The data taken on 2015 June 24 were taken under inferior weather conditions and therefore excluded from the analysis. The VERITAS observations were triggered by GCN notification (GCN#17929) obtained through the rapid-follow up system set in place for observations of gamma-ray bursts (Acciari et al. 2011c). The data analysis steps and parameters are exactly the same as for the analysis of 4U 0115+634. No evidence for photon emission above an energy of 200 GeV could be found. The observed number of events in the signal region is 21, consistent with the expected number of background events of 16.6 derived from six background regions located in the same field of view of these observations. The significance for an excess of gamma-rays is  $1.0\sigma$ . These numbers correspond to a flux upper limit to the gamma-ray flux above 200 GeV at 99% confidence level of  $F < 4.4 \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$  (corresponding to about 1.9% of the flux of the Crab Nebula above the same energy). As V 404 Cyg shows strong variability in X-rays on timescales much shorter than a day (Rodriguez et al. 2015; Kimura et al. 2016), nightly flux upper limits above the same energy for the two observing nights of VERITAS are listed in Table 2.

### 4. CONCLUSIONS

The flux upper limits for V 404 Cyg derived from the VERITAS observations and presented in this paper are equivalent to a luminosity of less than  $4.0 \times 10^{33} \text{ erg s}^{-1}$

**Table 1**  
VERITAS Observation Log for 4U 0115+634

Date	MJD (Start of Observations)	Observing Time (minutes)	Elevation Range	Observing Conditions	Flux Upper Flux Limit (99% Confidence) ( $10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ )
2015 Oct 23	57318.28	60	56°–57°	bright moonlight	3.4
2015 Oct 23	57318.38	60	45°–49°	dark	8.7
2015 Oct 24	57319.25	35	45°–44°	bright moonlight	20.7
2015 Nov 01	57327.16	60	45°–52°	dark	3.8
2015 Nov 02	57328.20	60	56°–57°	dark	3.1
2015 Nov 03	57329.24	60	56°–57°	dark	3.1

**Note.** The nightly flux limits are calculated assuming a power-law source spectrum with spectral index of  $\Gamma = -2.5$  and for energies above 300 GeV.

**Table 2**  
VERITAS Observation Log for V 404 Cyg

Date	MJD (Start of Observations)	Observing Time (minutes)	Elevation Range	Observing Conditions	Flux Upper Limit (99% Confidence) ( $10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ )
2015 Jun 20	57193.38	95	78°–82°	dark	5.1
2015 Jun 21	57194.35	60	71°–77°	dark	7.0
2015 Jun 24	57197.34	30	82°–83°	bad weather, excluded	...

**Note.** The flux limits are calculated assuming a power-law source spectrum with spectral index of  $\Gamma = -2.5$  and for energies above 200 GeV.

(above 200 GeV), assuming a distance to V 404 Cyg of 2.39 kpc and isotropic emission. This corresponds to a tiny fraction (about  $10^{-6}$ ) of the Eddington luminosity of the system, in stark contrast with e.g. the 15–200 keV luminosity of  $4.6 \times 10^{38} \text{ erg s}^{-1}$  derived from *Swift* XRT measurements during the flare peak (Segreto et al. 2015). It is important to note that the X-ray outburst of V 404 Cyg is characterized by large flare episodes followed by low-flux states, with fluctuations in the X-ray luminosity by more than an order of magnitude. Luminosities derived from observations by the *INTEGRAL* Imager on board the *INTEGRAL* satellite (IBIS) range from roughly  $9 \times 10^{36}$  to  $4 \times 10^{38} \text{ erg s}^{-1}$  during the observation periods by VERITAS (Kimura et al. 2016).

V 404 Cyg shows several similarities to an AGN: an accreting black hole and indications of jet-like features. Most X-rays produced during outbursts in binaries are expected to originate in the accretion disk, while in AGNs the synchrotron emission produced by high-energy electrons generally dominates the X-ray flux. However, hard X-rays (above 10 keV) are generally thought to be an indication of high-energy electrons energized in either a corona or a jet region of binaries. Evidence for an electron–positron plasma, expected to be present in jets of microquasars, has been observed through the detection of positron annihilation signatures associated with the outburst of V 404 Cyg on 2015 June 21 contemporaneously with the VERITAS observations (Siebert et al. 2016). The gamma-ray (>200 GeV) to hard X-ray (15–50 keV) flux ratio of  $<10^{-6}$  derived from the VERITAS measurements of V 404 Cyg is orders of magnitude lower than the typical range for these ratios of 0.1–2 observed in gamma-ray bright AGNs (e.g., Aliu et al. 2014b; Aleksić et al. 2015). This could mean that no efficient particle acceleration is taking place in these systems during flares, or that essentially all gamma-rays are absorbed via pair production with the large number of low-energy photons (see e.g., Bednarek & Giovannelli 2007; Dubus 2013). The latter process is especially important if particle acceleration takes place close to the base of the jet, i.e., deeply embedded in

the system. Orellana et al. (2007) suggested that X-ray and gamma-ray luminosities are anticorrelated during outbursts, which means that binary systems are most luminous at VHE gamma-rays at the beginning or ending of major X-ray outbursts.

4U 0115+634 is a pulsar binary system with similarities to the bright gamma-ray binaries LS I+61 303 and HESS J0632+057 (all consisting of Be stars orbited by compact objects on eccentric orbits). The flux upper limits derived from the VERITAS observations for 4U 0115+634 are equivalent to a luminosity of  $<6.4 \times 10^{33} \text{ erg s}^{-1}$  (above 300 GeV), assuming a distance to 4U 0115+634 of 6 kpc. The flux upper limit is not particularly constraining due to the large distance to 4U 0115+634: it corresponds to about 50% of the observed gamma-ray luminosity of LS I+61 303 during outbursts near apastron, and is a factor of 3 above the typical luminosity during the gamma-ray bright phases of HESS J0632+057. It is also less constraining than the upper flux limit of  $<0.5\text{--}1.5 \times 10^{33} \text{ erg s}^{-1}$  (above 300 GeV) obtained for observations during a giant X-ray flare of the nearby Be/pulsar binary 1 A0535+262 (Acciari et al. 2011a). However, the observation of 4U 0115+634 sets a strong upper limit on the ratio of gamma-ray to X-ray luminosity of less than 0.04. This is significantly lower than the typical ratio of 0.5–1.2 found in the known gamma-ray binaries during their peak emission phase (see Table 2 in Dubus 2013). Given these results, and the strong predicted dependency of the gamma-ray emission in Be/pulsar binaries on the properties of the system (e.g., on the inclination angle, size and structure of both the pulsar wind and the stellar wind), 4U 0115+634 may not be a gamma-ray source comparable to the known gamma-ray binaries.

Future observations with the planned Cherenkov Telescope Array, with an order of magnitude better sensitivity than the current gamma-ray instrument, will very likely be necessary to further improve our understanding of the non-thermal processes in these kinds of flaring binary systems.



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