

Very high energy particle acceleration powered by the jets of the microquasar SS 433

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SS 433 is a binary system containing a supergiant star that is overflowing its Roche lobe with matter accreting onto a compact object (either a black hole or neutron star) ¹⁻³. Two jets of ionized matter with a bulk velocity of $\sim 0.26c$ extend from the binary, perpendicular to the line of sight, and terminate inside W50, a supernova remnant that is being distorted by the jets ^{2,4-8}. SS 433 differs from other microquasars in that the accretion is believed to be super-Eddington ⁹⁻¹¹, and the luminosity of the system is $\sim 10^{40} \text{ erg s}^{-1}$ ^{2,9,12,13}. The lobes of W50 in which the jets terminate, about 40 pc from the central source, are expected to accelerate charged particles, and indeed radio and X-ray emission consistent with electron synchrotron emission in a magnetic field have been observed ¹⁴⁻¹⁶. At higher energies ($> 100 \text{ GeV}$), the particle fluxes of γ rays from X-ray hotspots around SS 433 have been reported as flux upper limits ^{6,17-20}. In this energy regime, it has been unclear whether the emission is dominated by electrons that are interacting with photons from the cosmic microwave background through inverse-Compton scattering or by protons interacting with the ambient gas. Here we report TeV γ -ray observations of the SS 433/W50 system where the lobes are spatially resolved. The TeV emission is localized to structures in the lobes, far from the center of the system where the jets are formed. We have measured photon energies of at least 25 TeV, and these are certainly not Doppler boosted, because of the viewing geometry. We conclude that the emission from radio to TeV energies is consistent with a single population of electrons with energies extending to at least hundreds of TeV in a magnetic field of ~ 16 micro-Gauss.

In the SS 433/W50 complex, several regions located west of the central binary (w1 and w2) and east (e1, e2, e3) are observed to emit hard X-rays ⁶. Previous searches for very high-energy (VHE) γ -ray emission from the hotspots between roughly 100 GeV and 10 TeV have produced null results ¹⁷⁻²⁰, though an excess observed at $\sim 800 \text{ MeV}$ may be associated with SS 433 and W50 ²¹. The High Altitude Water Cherenkov (HAWC) observatory is a wide field-of-view VHE γ -ray observatory surveying the Northern sky above 1 TeV, and is optimized for photon detection above 10 TeV ²². SS 433 transits 15° from the zenith of the HAWC detector each day, and has been observed with $> 90\%$ uptime since the start of detector operations in 2015.

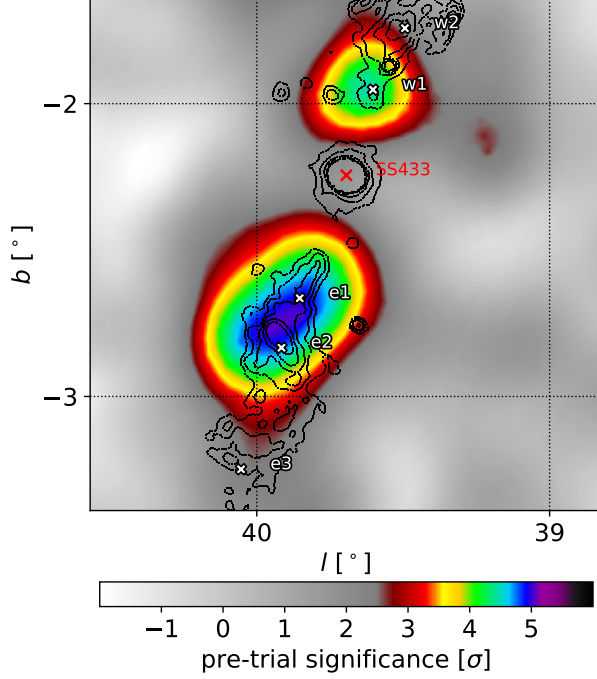


Figure 1: **VHE γ -ray image of the SS 433/W50 region.** The color scale indicates the statistical significance of the excess counts above the background of nearly isotropic cosmic rays before accounting for statistical trials. The figure shows the γ -ray excess measured after the fitting and subtraction of γ -rays from the spatially extended source MGRO J1908+06. The jet termination regions e1, e2, e3, w1, and w2 observed in the X-ray data are indicated, as well as the location of the central binary. The solid contours show the X-ray emission observed from this system.

In 1017 days of measurements with HAWC, an excess of γ rays with a post-trials significance of 5.4σ has been observed in a joint fit of the eastern and western interaction regions of the jets of SS 433. The emission is plotted in galactic coordinates in Fig. 1, which includes an overlay of the X-ray observations of the jets and the central binary. The γ -ray emission is spatially coincident with the X-ray hotspots w1 and e1; no significant emission is observed at the location of the central binary where the jets are produced.

Spatial and spectral fits to SS 433 are performed in a semicircular region of interest (RoI) designed to mask out diffuse emission from the Galactic Plane. The RoI also removes significant spatially extended emission from the nearby γ -ray source MGRO J1908+06. The spatial distribution and spectrum of γ rays from MGRO J1908+06 are fit using an electron diffusion model²³, and point-like sources centered on e1 and w1 are fit on top of this extended emission. As a systematic check, the regions are also fit using X-ray spatial templates and extended Gaussian functions. Neither improves the statistical significance of the fits. Upper limits on the angular size of the emission regions are 0.25° for the east hotspot and 0.35° for the west hotspot at 90% confidence.

Given the distance to the source of 5.5 kpc, this corresponds to a physical size of 24 pc and 34 pc, respectively. The constraint is tighter on the eastern hotspot due to its higher statistical significance.

The VHE γ -ray flux is consistent with a hard E^{-2} spectrum, though current data from HAWC are not of sufficient significance to constrain the spectral index. Therefore we report the flux of both hotspots at 20 TeV, where systematic uncertainties due to the choice of spectral model are minimized and the sensitivity of HAWC is maximized. At e1, the VHE flux is $2.4^{+0.6}_{-0.5}(\text{stat.})^{+1.3}_{-1.3}(\text{syst.}) \times 10^{-16} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$, and at w1 the flux is $2.1^{+0.6}_{-0.5}(\text{stat.})^{+1.2}_{-1.2}(\text{syst.}) \times 10^{-16} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$. HAWC detects γ rays from the interaction regions up to at least 25 TeV. The energies of these γ rays are a factor of three to ten higher than previous measurements from microquasars^{24,25}. Since most γ -ray telescopes are optimized for measurements below 10 TeV, this may explain why these photons were not observed in previous observational campaigns.

The γ rays detected by HAWC are produced by radiative or decay processes from much higher energy particles. The detection yields important information about the mechanisms and sites of particle acceleration, the types of particles accelerated (e.g., protons or electrons), and the radiative processes which produce the spectrum of emission from radio to VHE γ rays. Two scenarios for explaining the HAWC observations of the e1 and w1 regions can be tested. The first is that protons are primarily responsible for the observed γ rays. Protons must have an energy of at least 250 TeV to produce 25 TeV γ rays through hadronic collisions with ambient gas. Proton-proton collisions yield π^0 particles which decay to VHE γ rays, and $\pi^\pm$ particles which decay to the secondary electrons and positrons responsible for radio to X-ray emission via synchrotron radiation. This scenario is of particular interest because there is spectroscopic evidence for ionized nuclei in the inner jets of SS 433^{8,26}. The alternative scenario requires electrons of at least 130 TeV to up-scatter the low-energy photons from the cosmic microwave background (CMB) to 25 TeV γ rays. In this case, the radio to X-ray emission is dominated by synchrotron radiation from this same population of electrons in the magnetized plasma of the jets and lobes.

The fact that the VHE emission is detected along a line of sight nearly orthogonal to the jet axis means that charged particle trajectories become isotropic before they interact to produce the γ rays. The embedded magnetic fields in the VHE regions can easily deflect the accelerated particles because their typical gyroradii are much smaller than the size of the emission regions, ~ 30 pc. The jets are only mildly relativistic, so the emission from the interaction regions will have a negligible Doppler beaming effect and remain nearly isotropic.

The flux of VHE γ rays observed by HAWC makes the proton scenario for SS 433 unlikely, because the total energy required to produce the highly relativistic protons is too high. The jets of SS 433 are known to be radiatively inefficient, with most of the jet energy transformed into the thermal energy of W50^{16,27} rather than particle acceleration. We model the primary proton spectrum as a power law with an exponential cutoff, $dN/dE_p \propto E_p^{-2} \exp(-E_p/1 \text{ PeV})$. If we assume 10% of the jet kinetic energy converts into accelerated protons, and the ambient gas density is 0.05 cm^{-3} ^{16,27}, then the resulting flux of γ rays from proton-proton collisions is much less than the observed γ -ray flux, as shown in the dash-dotted line of Fig. 2. In fact, for a target proton density as large as 0.1 cm^{-3} in the e1 region^{16,27}, the total energy of the proton population needs to

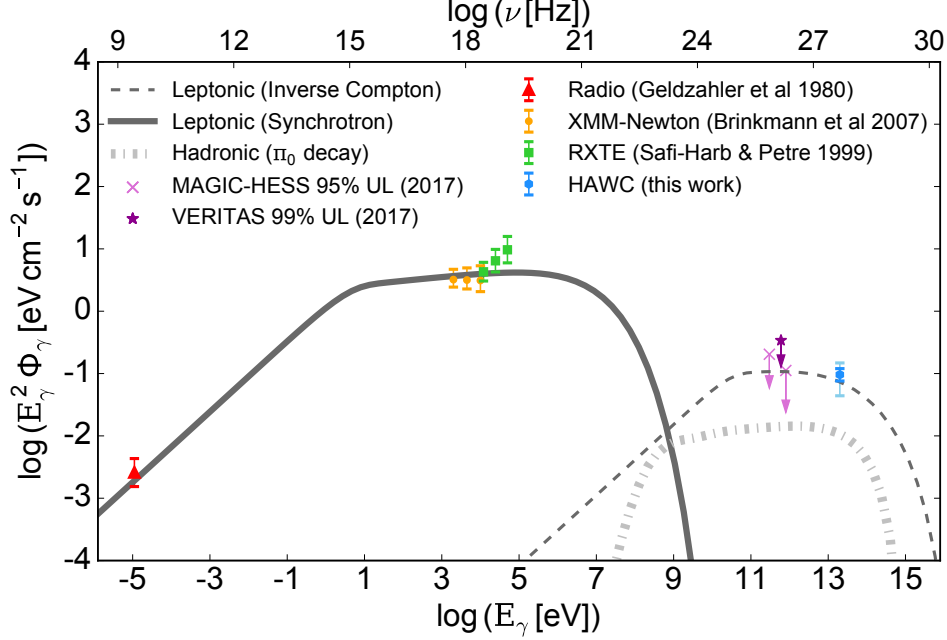


Figure 2: **Broadband spectral energy distribution of the eastern emission region.** The data include radio ¹⁴, soft X-ray ¹⁵, hard X-ray ¹⁶, and VHE γ -ray upper limits ^{19,20}, and HAWC observations of e1. Error bars indicate 1σ uncertainties, with the thick (thin) errors on the HAWC flux indicating statistical (systematic) uncertainties and arrows indicating flux upper limits. The multiwavelength spectrum produced by electrons assumes a single electron population following a power-law spectrum with an exponential cutoff. The electrons produce radio to X-ray photons through synchrotron emission in a magnetic field (thick solid line) and TeV γ rays through inverse Compton scattering of the cosmic microwave background (thin dashed line). The dash-dotted line represents the radiation produced by protons, assuming 10% of the jet kinetic energy converts into protons.

be $\sim 3 \times 10^{50}$ erg to explain the observed γ rays, assuming an E_γ^{-2} spectrum. This is comparable to the total jet energy available during the presumed 30,000-year lifetime of SS 433 ². Furthermore, because the synchrotron emission from secondary electrons from charged pion decay is always lower than the γ -ray flux from π^0 decay, and the observed X-ray flux is higher than the γ -ray flux, the X-rays cannot originate solely from secondary electrons. Finally, the proton scenario requires that the protons remain trapped in the region observed by HAWC for the lifetime of SS 433 ². This means the protons must diffuse very slowly, with a diffusion coefficient $\sim 1/1000$ of the typical value of the interstellar medium (ISM) $D_{\text{ISM}} \approx 3 \times 10^{28} (E/3 \text{ GeV})^{1/3} \text{ cm}^2 \text{ s}^{-1}$ ²⁸. This value, comparable to the theoretical Bohm limit, is very small but not impossible. Given the uncertainties in the historical jet flux, the ambient particle density, and the radiative efficiency, we cannot exclude the possibility that some fraction of the γ -ray flux is produced by protons. However, we do rule out the possibility that the VHE γ rays are entirely produced by protons.

Highly relativistic electrons, on the other hand, can produce γ rays much more efficiently, primarily via inverse Compton scattering of CMB photons to γ rays. The inverse Compton losses due to up-scattering of infrared and optical photons are suppressed due to the Klein-Nishina effect and are thus dominated by scattering of CMB photons ²⁹. In this scenario, the ratio of the VHE

γ -ray to X-ray fluxes constrains the energy density in the magnetic field compared to the energy density in CMB photons. We have modeled the broadband spectral energy distribution (SED) of the eastern emission region 15' to 33' from the center of SS 433. The solid and dashed lines in Fig. 2 show the SED of a leptonic model for e1 produced by an injected flux of relativistic electrons with an energy spectrum $dN/dE \propto E^{-\alpha} \exp(-E/E_{\max})$ in a magnetic field of strength B . We use the parameters $\alpha = 1.9$, $E_{\max} = 3.5$ PeV, and $B = 16 \mu\text{G}$ (see the methods section). The estimate of the magnetic field strength is consistent with the equipartition of energy between the relativistic electrons and magnetic fields, which is common in astrophysical systems¹⁶. The required total energy budget for relativistic electrons is three orders of magnitude lower than the total jet energy.

The maximum electron energy of ~ 1 PeV has significant implications for electron acceleration sites and acceleration mechanisms in SS 433. SS 433 is distinguished from other binary systems with relativistic objects because it realizes a supercritical accretion of gas onto the central engine². Powerful accretion flows and the inner jets near the compact object have therefore been proposed as possible acceleration sites of relativistic particles²⁶. However, the observation from HAWC suggests that ultra-relativistic electrons are not accelerated near the center of the binary. If electrons were accelerated in the central region, they would have cooled by the time they reached the sites of observed VHE emission. Due to their small gyroradii, high-energy electrons may transport in a magnetized medium via diffusion or advection. The distance traveled via diffusion within the cooling time of an electron of energy E is $r_d = 2\sqrt{Dt_{\text{cool}}} \approx 36 \text{ pc } (E/1 \text{ PeV})^{-1/3} (B/16 \mu\text{G})^{-1}$, using the diffusion coefficient typical of the ISM²⁸. This distance would be even smaller for diffusion coefficients lower than the typical ISM value. Similarly, the distance traveled by electrons being advected with the jet flow is $r_{\text{adv}} = 0.26c \cdot t_{\text{cool}} \approx 4 \text{ pc } (E/1 \text{ PeV})^{-1} (B/16 \mu\text{G})^{-2}$ for a jet velocity of $0.26c$. Both distance scales are smaller than the 40 pc distance between the binary and e1, indicating the electrons are not accelerated near the center of the system.

Instead, the highly energetic electrons in SS 433 are likely accelerated in the jets and near the VHE γ -ray emission regions. This presents a challenge to current acceleration models. For example, particle acceleration may be driven by the dissipation of the magnetic fields in the jets, but above several hundred TeV the electron acceleration time exceeds the electron cooling time, assuming a $16 \mu\text{G}$ magnetic field. Thus the system does not appear to have sufficient acceleration power, unless there are very concentrated magnetic fields along the jets. If instead particle acceleration is driven by standing shocks produced by the bulk flow of the jets, it is possible to reach PeV energies if the size of the acceleration region is larger than the electrons' gyroradii. However, shocks in the interaction regions are not currently resolved by X-ray or γ -ray measurements.

Studies of microquasars such as SS 433 provide valuable probes of the particle acceleration mechanisms in jets, since these objects are believed to be scale models of the much larger and more powerful jets in active galactic nuclei (AGN)³⁰. AGN are the most prevalent VHE extragalactic sources and are believed to be the sources of the highest energy hadronic cosmic rays. While AGN are not spatially resolved at VHE energies, with this observation we have identified a VHE source in which we can image particle acceleration powered by jets. Future high-resolution VHE obser-

uations of SS 433 are possible with pointed atmospheric Cherenkov telescopes to better localize the emission sites, and further high-energy measurements with HAWC will record the spectrum at high energies and better constrain the maximum energy of accelerated particles.

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Author Contributions C. D. Rho (crho2@ur.rochester.edu) and H. Zhou (hao@lanl.gov) analyzed the data and performed the maximum likelihood analysis. Modeling of the leptonic and hadronic emission has been carried out by K. Fang (kefang@umd.edu) and H. Zhang (zhan3038@purdue.edu). S. BenZvi (sybenzvi@pas.rochester.edu) and B. Dingus (dingus@lanl.gov) helped prepare the manuscript. The full HAWC Collaboration has contributed through the construction, calibration, and operation of the detector, the development and maintenance of reconstruction and analysis software, and vetting of the analysis presented in this manuscript. All authors have reviewed, discussed, and commented on the results and the manuscript.

Competing Interests The authors declare no competing interests.

Data Availability The datasets analyzed during this study are available at a public data repository maintained by the HAWC Collaboration: <https://data.hawc-observatory.org/>.

Code Availability The study was carried out using the Analysis and Event Reconstruction Integrated Environment Likelihood Fitting Framework (AERIE-LiFF) and the Multi-Mission Maximum Likelihood (3ML) software. The code is open-source and publicly available on Github: <https://github.com/rjlauer/aerie-liff> and <https://github.com/giacomov/3ML>. The software includes instructions on installation and usage.

Methods

Data Reduction and Maximum Likelihood Analysis This analysis uses 1017 days of data from the High Altitude Water Cherenkov Observatory collected between November 26, 2014 and December 20, 2017. HAWC is an array of 300 tightly packed identical water Cherenkov detectors deployed 4100 m above sea level on the slope of volcano Sierra Negra, Mexico ³¹. Each detector is a cylindrical water tank standing 5 m tall and 7.3 m in diameter, filled with 180 000 L of purified water. At the bottom of each tank are four photomultipliers (PMTs) facing upward. The PMTs record the Cherenkov photons created by the relativistic secondary particles produced when primary cosmic rays and γ rays interact at the top of the atmosphere. The HAWC array covers 22,000 m². Its construction ended in December 2014, and the full array was commissioned in March 2015.

Using the relative arrival time of photoelectrons (“hits”) detected by the PMTs, the arrival direction of primary γ rays can be reconstructed with an accuracy of $\sim 1^\circ$ below 1 TeV to $< 0.2^\circ$ above 10 TeV ³². The accuracy of the reconstruction determines the point spread function (PSF) of the detector, and is a function of the energy, zenith angle, and composition of the primary particle. Air showers from γ rays are discriminated from the cosmic-ray background by filtering out “clumpy” patterns of hits, which are characteristic of the energy deposited by hadronic cosmic rays. The cosmic-ray background rejection efficiency improves rapidly as a function of energy, increasing from 90% at 1 TeV to 99.9% at 10 TeV ³².

To compute the statistical significance of γ -ray emission observed with HAWC, a maximum likelihood fit using parametric spatial and spectral models is applied to the data ^{33,34}. The models are forward-folded through the detector response to produce expected counts of γ -ray signal events and cosmic-ray background events. The expectation is then compared to the observed counts N_{obs} . To calculate the expected counts as a function of position on the sky, the events are binned in a fine mesh using the HEALPix pixelization of the unit sphere ³⁵. The pixelization is chosen to be 0.1° , roughly two to five times smaller than the radius of the instrument PSF. To apply models of the energy spectrum of a source, the data are binned according to the fraction of PMTs in the detector triggered by an air shower ³². This measure of shower “size” is used as a coarse proxy for the energy of the primary particle; a total of 9 size bins B is used.

Given a model with $\vec{\theta}$ spatial and spectral parameters, the maximum likelihood of the model having produced the data is

$$\ln \mathcal{L}(N_{\text{obs}}|\vec{\theta}) = \sum_{B=1}^9 \sum_{j=1}^m \ln P(N_{\text{obs}}^{j,B}|\vec{\theta}), \quad (1)$$

where the sum runs over the size bins B and the HEALPix pixels j in the region of interest (RoI) of the fit. P is the Poisson probability of detecting $N_{\text{obs}}^{j,B}$ events in pixel j and size bin B given the model parameters $\vec{\theta}$.

Within the RoI around SS 433 defined in Extended Data Fig. 1, two fits are performed to

maximize the likelihood: a fit which only accounts for the emission from MGRO J1908+06 (null hypothesis), and the combined emission from MGRO J1908+06 and the SS 433 region (alternative hypothesis). The ratio of the maximum likelihood defines a test statistic

$$TS = 2 \left(\ln \mathcal{L}(N_{\text{obs}} | \vec{\theta}_{\text{alt}}) - \ln \mathcal{L}(N_{\text{obs}} | \vec{\theta}_0) \right), \quad (2)$$

where $\vec{\theta}_0$ and $\vec{\theta}_{\text{alt}}$ represent the spatial and spectral parameters of the null and alternative hypotheses, respectively. TS is then converted to a p -value to estimate the statistical significance of emission from SS 433. As discussed in the main text, the alternative hypothesis assumes two point sources with power law spectra $dN/dE = f_0 \cdot (E/20 \text{ TeV})^{-2}$, where the flux normalization f_0 is the free parameter of the spectral model.

Wilks' Theorem is used to convert TS to a p -value³⁶. In the joint likelihood maximization, there are 2 degrees of freedom for the two separately fit flux normalizations of the hotspots at w1 and e1. Therefore, we calculate the one-tailed p -value $\text{pr}(TS > \chi^2 = 41.2 | \text{dof} = 2) = 1.13 \times 10^{-9}$. Since the positions of the point source fits at w1 and e1 were chosen after looking into the data, and because we are searching for other microquasars in the field of view of HAWC, we must apply *a posteriori* corrections to the p -value to account for multiple-comparison effects.

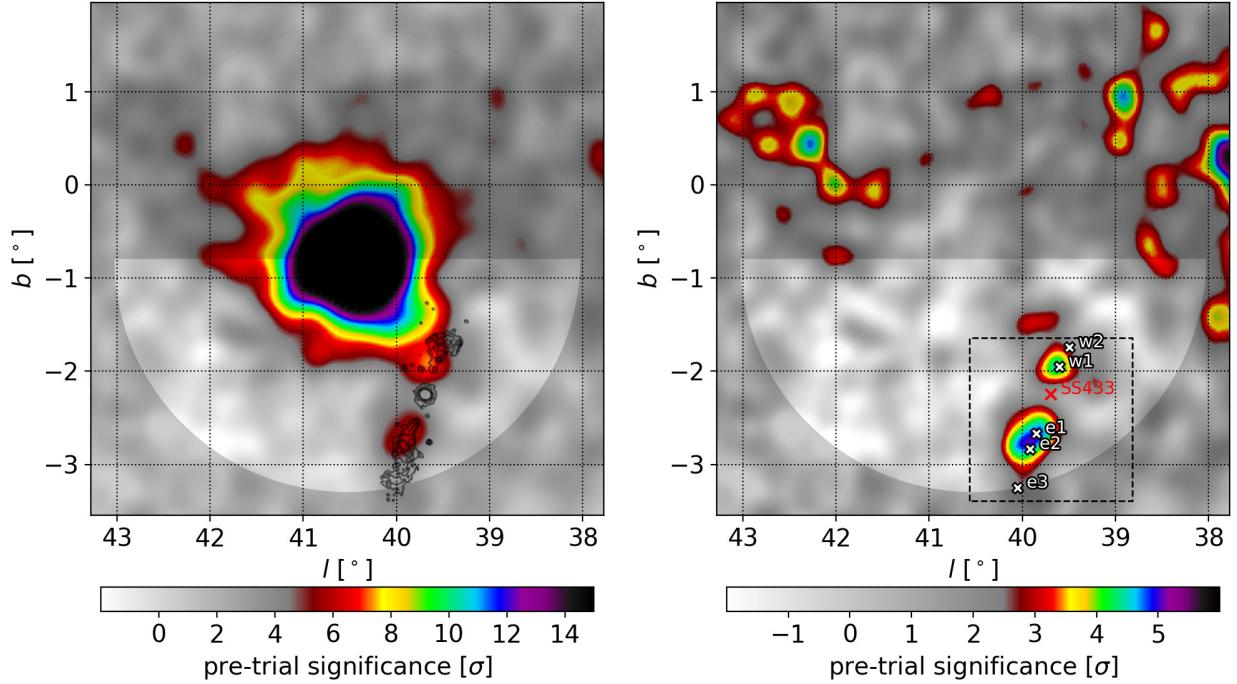
The X-ray interaction regions w1, w2, e1, e2, and e3 are *a priori* candidates for the locations of the maxima, as is the center of the binary system, for a total of six potential hotspots. Given the angular resolution of HAWC, it would not be possible to spatially resolve all six hotspots; at best three regions (east, west, and center) can be separately fit with confidence. There are 23 possible combinations of the six *a priori* locations which can be used to fit one, two, or three hotspots in the eastern, central, and western regions of the source. We add an additional 12 trials to account for the known microquasars in the field of view of HAWC^{9,13,37}. This trial factor is conservative given that several galactic microquasars are already known TeV sources^{38,39}.

Given 35 total trials, the corrected p -value is 3.96×10^{-8} , which corresponds to a statistical significance of 5.4σ .

Modeling of the Nearby Extended Source MGRO J1908+06 A bright extended source, MGRO J1908+06, is detected with more than 30σ in this dataset and is located less than 2° from the γ -ray hotspots of SS 433 (Extended Data Fig. 1). The region of MGRO J1908+06 contains a pulsar and a supernova remnant, but it is not clear if the observed TeV γ -ray emission is from either or even both of them. A detailed discussion of MGRO J1908+06 is beyond the scope of this paper. However, the morphology of MGRO J1908+06 must be carefully studied in order to minimize the contamination of the emission due to MGRO J1908+06 on the fluxes of the lobes.

A maximum likelihood analysis is performed that simultaneously fits the emission from MGRO J1908+06 and the hotspots at w1 and e1. An electron diffusion model appropriate for older pulsar wind nebulae^{23,40} is used to describe the spatial morphology of MGRO J1908+06. Given the uncertainty of the nature of MGRO J1908+06, two other spatial models with Gaussian and power law radial profiles are also tested in the simultaneous fit. The choice of spatial model

affects the best-fit fluxes from e1 and w1 at the level of $\pm 20\%$. We adopt this value as a systematic uncertainty on the flux from w1 and e1 due to VHE emission from the nearby extended source.



Extended Data Figure 1: **VHE γ rays from MGRO J1908+06 and SS 433/W50.** The color scale indicates the statistical significance of the excess counts above background before accounting for statistical trials. *Left panel:* the bright extended γ -ray source MGRO J1908+06 is shown at the center of the image and SS 433/W50 at the bottom. The dark contours show X-ray emission from SS 433 and its jets ⁴¹. The semicircular area indicates the region of interest used to fit the γ -ray observations. *Right panel:* The figure shows the γ -ray excess measured after the fitting and subtraction of γ -rays from the spatially extended source MGRO J1908+06. The dashed box indicates the region shown in Fig. 1. The jet termination regions e1, e2, e3, w1, and w2 observed in the X-ray data are indicated, as well as the location of the central binary. The dark lines show the contours of X-ray emission observed from this system.

Contamination from Galactic Diffuse Emission The e1 and w1 regions are located roughly 2° from the Galactic plane, so the contamination from the Galactic diffuse emission (GDE) is negligible. However, MGRO J1908+06 has a Galactic latitude of $\sim 1^\circ$. Since the three spatial models used to fit MGRO J1908+06 are radially symmetric, the presence of GDE has the potential to produce an overestimate in the flux from MGRO J1908+06, which could result in an underestimate in the flux measured from w1 and e1. To minimize the effect of the GDE on the fit, the RoI is defined to be a semicircular region centered on the position of MGRO J1908+06 (Extended Data Fig. 1). The RoI is designed to reduce the effect of GDE by excluding the half of the source closest to the Galactic plane.

To estimate the systematic uncertainties associated with the choice of RoI and possible contamination from GDE, a second maximum likelihood fit is performed using a full-disk RoI which includes GDE, emission from MGRO J1908+06, and w1 and e1. The spatial distribution of the GDE is modeled with a Gaussian profile of three different widths of 0.5° , 1° , and 2° in Galactic

latitude, and is treated as constant in Galactic longitude over the width of the RoI. Comparing the results to the fit with a semicircular RoI indicates that contamination from GDE is less than 10% for e1 and less than 20% for w1.

Fit Results The fit results are reported in Extended Data Table 1. Fitting the emission from w1 and e1 simultaneously, we calculate $TS = 41.2$, which corresponds to a 5.4σ observation ($p = 3.96 \times 10^{-8}$) after accounting for *a posteriori* statistical trial factors. To check the consistency of the results, we also fit w1 and e1 separately, re-defining the alternative model to include MGRO J1908+06 and only e1 or w1. The significance of the VHE excess from these locations is below 5σ in the fits, but the estimated fluxes are consistent with the simultaneous fit.

An additional check is made on the effect of fixing versus floating the best-fit positions of the emission from the east and west hotspots. In the original alternative hypothesis, the point sources are centered on e1 and w1. Here, the positions of the point sources are made additional free parameters in the point source fit. We find that allowing the positions of the TeV hotspots to vary does not affect the flux estimates, which are consistent with the fixed-position fits. Moreover, the best-fit positions of the east and west TeV emission regions are consistent with e1 and w1 within statistical uncertainties.

Extended Data Table 1 Fits to the TeV emission from SS 433 using nested point source models				
Lobe	Position (RA, Dec)	dN/dE at 20 TeV [$10^{-16} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$]	TS	Significance (post trials)
Simultaneous fit to E+W hotspots, positions fixed.				
E	19:13:37 04°55'48"	$2.4^{+0.6+1.3}_{-0.5-1.3}$	41.2	5.4σ
W	19:10:37 05°02'13"	$2.1^{+0.6+1.2}_{-0.5-1.2}$		
Separate fit to E+W hotspots, positions fixed.				
E	19:13:37 04°55'48"	$2.5^{+0.7+1.4}_{-0.5-1.4}$	24.3	4.6σ
W	19:10:37 05°02'13"	$2.3^{+0.7+1.3}_{-0.5-1.3}$	20.4	4.2σ
Separate fit to E+W hotspots, positions floated.				
E	19:14:11 $^{+20}_{-39}$ s 04°59'10" $^{+03'30''}_{-06'18''}$	$2.6^{+0.6+1.4}_{-0.5-1.4}$	26.9	4.4σ
W	19:10:40 $^{+17}_{-17}$ s 05°03'40" $^{+03'32''}_{-04'55''}$	$2.4^{+0.6+1.3}_{-0.5-1.3}$	23.4	4.0σ

The choice of spectral model also affects the estimated γ -ray flux at 20 TeV. Extended Data Table 2 shows the dependence of the best-fit VHE flux from e1 and w1 on the assumed spectral models, including statistical uncertainties on the flux normalization at 20 TeV. Two spectral models were tested: a simple power law $dN/dE_\gamma \propto E_\gamma^{-\alpha}$, and a power law with an exponential cutoff $dN/dE_\gamma \propto E_\gamma^{-\alpha} \exp(-E_\gamma/E_{\text{cut}})$. The choice of spectral model can alter the flux normalization by almost a factor of two compared to the default E_γ^{-2} model.

Summary of Systematic Uncertainties The systematic uncertainties in the estimated fluxes from the TeV hotspots of SS 433 include the following contributions: detector systematic effects, modeling ambiguities in MGRO J1908+06, and contamination from Galactic diffuse emission. The systematic uncertainties due to the modeling of MGRO J1908+06 and the contamination from

Extended Data Table 2 Dependence of Flux at 20 TeV on Spectral Assumption						
dN/dE at 20 TeV [$\times 10^{-16}$ TeV $^{-1}$ cm $^{-2}$ s $^{-1}$]						
	Index: -1.5		Index: -2.0		Index: -2.5	
Cutoff Energy	East Lobe	West Lobe	East Lobe	West Lobe	East Lobe	West Lobe
No Cutoff	$1.0^{+0.3}_{-0.2}$	$0.9^{+0.3}_{-0.2}$	$2.4^{+0.6}_{-0.5}$	$2.1^{+0.6}_{-0.5}$	$3.3^{+0.9}_{-0.7}$	$2.4^{+0.9}_{-0.6}$
50 TeV	$4.7^{+1.1}_{-0.9}$	$4.2^{+1.1}_{-0.9}$	$5.0^{+1.2}_{-1.0}$	$4.1^{+1.3}_{-0.9}$	$3.2^{+0.9}_{-0.7}$	$1.7^{+1.1}_{-0.7}$
300 TeV	$1.7^{+0.5}_{-0.4}$	$1.6^{+0.5}_{-0.4}$	$3.3^{+0.8}_{-0.7}$	$2.9^{+0.8}_{-0.7}$	$3.6^{+0.9}_{-0.7}$	$2.4^{+0.9}_{-0.7}$

Fit results with different spectral models for the γ -ray emission.

Galactic diffuse emission are $< \pm 20\%$ and -10% (-20%) for the east (west) hotspot, respectively, and are discussed in previous sections.

The detector response is estimated using Monte Carlo simulations and then optimized using observations of the Crab Nebula³², which appears point-like in the HAWC data. Systematic uncertainties which potentially affect the result presented here include the charge resolution and relative quantum efficiency of the PMTs, the absolute quantum efficiency of the PMTs, changes to the detector layout as construction proceeded, uncertainties in the PSF, and systematic differences in the distribution of arrival times of photoelectrons between data and simulation. The total systematic uncertainty on the flux normalization from detector effects is $\pm 50\%$.

All the components of the systematic uncertainties are summarized in Extended Data Table 3 and combined in quadrature to estimate the total systematic uncertainty on the VHE flux from the w1 and e1. We note that since the systematic uncertainties due to MGRO J1908+06 and GDE are anti-correlated, the quadrature sum overestimates the total systematic uncertainty. However, the effect is not particularly important, since the detector systematic effects are the dominant source of uncertainty.

Extended Data Table 3 Systematic uncertainties on the flux from SS 433

Systematic	East Lobe	West Lobe
Detector Systematic Effects	$\pm 50\%$	
MGRO J1908+06 Modeling	$< \pm 20\%$	
Galactic diffuse contamination	-10%	-20%
Total	$\pm 55\%$	$\pm 55\%$

Systematic 1σ error budget for the VHE γ -ray fits.

X-ray Template Fit and Upper Limit on the Extent of the Emission Regions We performed several maximum likelihood fits modeling the hotspots as spatially extended sources. In the first fit, we generated spatial templates for the eastern and western regions based on the X-ray contours published by ROSAT⁴¹ and then performed a joint likelihood fit with the two γ -ray hotspots and MGRO J1908+06. This produces no improvement in TS over a point-source fit.

In order to constrain the size of the γ -ray emission regions, likelihood fits are applied using a Gaussian morphology convolved with the point spread function of HAWC. To reduce the number of free parameters, we first fit MGRO J1908+06 using an RoI with SS 433 and its hotspots excluded. The extended fit from MGRO J1908+06 is then subtracted from the data, and the residual γ -ray emission from the γ -ray hotspots is fit using two Gaussian functions. The centers of the

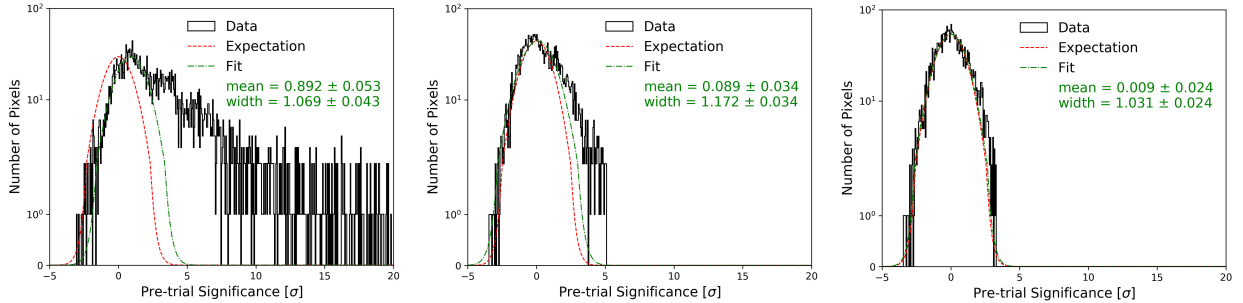
Gaussians are fixed to e1 and w1, and their angular widths are estimated in a simultaneous fit to both the eastern and western regions.

The maximum likelihood fit yields an angular width of $0.14^\circ \pm 0.06$ for the east hotspot and $0.08^\circ_{-0.05}^{+0.14}$ for the west hotspot. We estimate the 90% confidence region on the extent as the value of Gaussian width which produces a decrease $\Delta\text{TS} = -2.71$ from the maximum likelihood value. The resulting 90% upper limits are 0.25° for the east region and 0.35° for the west region.

Upper Limit on Emission from the Central Binary In the present data set, no statistically significant emission is observed from the center of SS 433. Using the method of Feldman and Cousins⁴², we estimate the 90% upper limit on the flux at 20 TeV to be $5.3 \times 10^{-17} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ after fitting MGRO J1908+06 and the emission at e1 and w1.

Upper Limit on Detected γ -ray Energy The binning of γ -ray events into size bins B causes us to lose information about the energies of the γ rays observed from SS 433. To determine the upper energy bound on the flux we observe, we scan over the maximum energy $E_{\gamma,\text{max}}$ used in the forward-folding analysis. Starting at $E_{\gamma,\text{max}} = 15 \text{ TeV}$, we find that TS increases monotonically until $E_{\gamma,\text{max}} = 25 \text{ TeV}$. Increasing $E_{\gamma,\text{max}}$ above this value causes TS to plateau (for e1) or decrease slightly (w1). We infer that the current measurement of e1 and w1 implies a minimum $E_{\gamma,\text{max}} = 25 \text{ TeV}$, and report this as a conservative estimate of the highest energy observed by HAWC.

Study of Residual Emission in the Region of Interest As a final check of the quality of the maximum likelihood fits, we plot the distribution of the significance values in each HEALPix pixel in the RoI around SS 433 in Extended Data Fig. 2. The significance values are plotted in units of Gaussian σ . If only random background fluctuations are present, the significance values follow a standard normal distribution, shown by the dashed line in the figure.



Extended Data Figure 2: Distribution of pixel significance in the region of interest of the fit, defined as deviations from the background expectation, in the HAWC sky map (left), after fitting and subtraction of emission from MGRO J1908+06 (center), and after fitting and removal of emission from MGRO J1908+06 and the γ rays from w1 and e1 (right).

Prior to the maximum likelihood fit, the significance distribution in the RoI is considerably skewed toward positive values due to excess counts above background from MGRO J1908+06 and γ rays from w1 and e1 (left panel of Extended Data Fig. 2). After the subtraction of the maximum likelihood fit to emission from MGRO J1908+06 (center panel), the skew in the distribution is

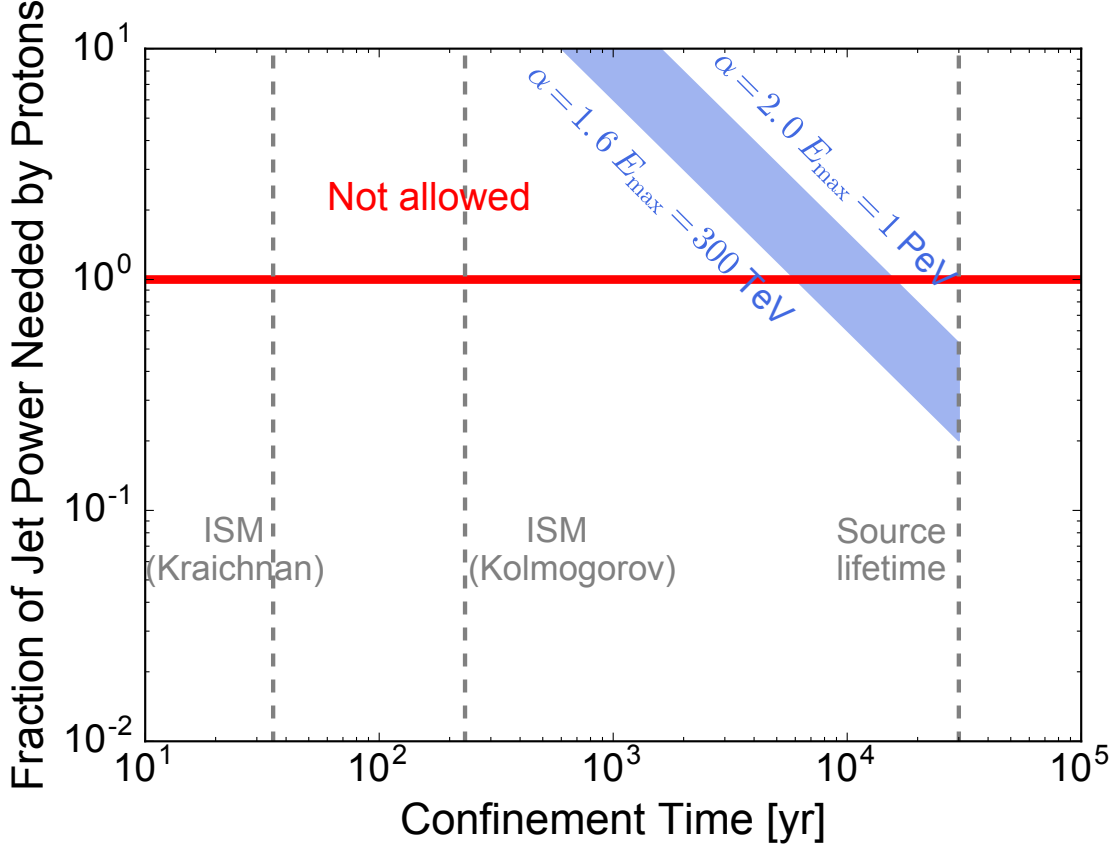
considerably reduced, though still visible due to excess counts from the interaction regions near SS 433. Finally, subtraction of the maximum likelihood flux from w1 and e1 produces a distribution that, within statistical uncertainties, is equivalent to background fluctuations (right panel).

VHE Emission due to Leptonic Interactions In the leptonic scenario considered in this paper, relativistic electrons scatter photons from the cosmic microwave background photons to TeV energies via the inverse Compton process, and produce X-ray and radio emissions by the synchrotron radiation. Although a far infrared background in the lobe may contribute to the production of sub-TeV photons, no appreciable infrared emission has been reported near the γ -ray emission region^{2,43}. The electron spectrum is obtained by solving the continuity equation, considering radiative cooling⁴⁴. The best-fit values of the parameters of the injection spectrum, including the flux, spectral index, and maximum energy of the electrons, and the magnetic field strength in the source region, are obtained through Markov Chain Monte Carlo⁴⁵ sampling of their likelihood distributions when fitting to the multi-wavelength data. The radio and soft X-ray data points correspond to the Effelsberg Telescope¹⁴ and the XMM-Newton (the Mos1 detector)¹⁵ observations of a 6' circle centered on e1. A 30% uncertainty attributed to the unknown shape of the HAWC source is added to the statistical and systematic errors of the observational data, though we find the uncertainty has a negligible impact on the fit. The hard X-ray data points and the sub-TeV upper limits are set by the RXTE, MAGIC, HESS, and VERITAS observations of the e1 region^{16,19,20}.

The VHE flux is determined using the flux from e1 at 20 TeV reported in Extended Data Table 1, where separate fits were made to eastern and western hotspots, and the positions were fixed to e1 and w1. The best-fit values of the injection spectrum and magnetic field in the emission region are $\alpha = 1.87^{+0.04}_{-0.07}$, $\log(E_{\text{max}}/\text{PeV}) = 3.53^{+0.31}_{-0.38}$, and $B = 16.04^{+2.60}_{-2.23} \mu\text{G}$. Taking the distance to the source to be 5.5 kpc, the fit suggests a total electron energy of 2.9×10^{47} erg. This is a small fraction of the total energy deposited by the jets of SS 433 over their lifetime, which is $\sim 9 \times 10^{50}$ erg assuming a kinetic jet luminosity of $10^{39} \text{ erg s}^{-1}$ ². Future multi-wavelength observations dedicated to the VHE γ -ray emission region will better constrain the magnetic field strength and the properties of the electron population.

We note that the presence of multi-hundred TeV to PeV electrons would challenge the current particle acceleration mechanisms. Successful acceleration requires that the acceleration rate $\dot{\gamma}_{\text{acc}} \approx eBv/m_e c^2$ based on heuristic considerations (where v is the velocity associated with the notional electromotive force) exceed the cooling rate $\dot{\gamma}_{\text{cool}} \approx 4c\sigma_T\gamma^2(B^2/8\pi)/3m_e c^2$, assuming that synchrotron radiation dominates the cooling processes in the lobes of SS 433. This leads to a maximum electron energy $E_{e,\text{max}} = 271 \text{ TeV } (v/100 \text{ km s}^{-1})^{1/2} (B/16 \mu\text{G})^{-1/2}$. For reference, the Alfvén speed in the lobes is $v_A = 160 \text{ km/s } (n_b/0.05 \text{ cm}^{-3})^{-1/2} (B/16 \mu\text{G})$. A higher Alfvén speed could be achieved if the acceleration takes place in the central spine of the jet, where the mass loading due to black hole accretion is smaller and the magnetic field is stronger. Depending on the exact electron acceleration mechanisms, v could be associated with the jet flow velocity, or with the Alfvén speed. In both cases, using these estimates, it is possible that the maximum electron energy can exceed 1 PeV. However, the timescale of acceleration mechanisms such as second-order Fermi acceleration is proportional to $(v/c)^2$, making the production of multi-hundred TeV electrons less

efficient. Future VHE γ -ray and hard X-ray observations can better constrain the electron cutoff energy, and diagnose the *in situ* particle acceleration mechanism.



Extended Data Figure 3: **Fraction of jet power needed to produce the observed VHE gamma rays in the hadronic scenario.** The blue shaded region shows the energy injection rate of protons, in units of the kinetic luminosity of the jet, in order to produce the observed VHE γ rays by interacting with ambient gas, as a function of the proton confinement time. A gas density of 0.05 cm^{-3} is adopted for the source vicinity^{16,27}. Most hadronic models require $> 100\%$ jet power (above the red solid line) and are thus not allowed. Even when the diffusion coefficient is extremely small (for reference, the dashed grey lines show the source age and the confinement time of 200 TeV protons in a 30 pc region in the ISM with Kraichnan and Kolmogorov-type diffusion) and when the spectral index is much harder than 2, the hadronic scenario still requires a significant energy input from the jet.

VHE Emission due to Hadronic Interactions In the hadronic scenario, high-energy protons interact with the ambient gas in the source, and produce γ rays via the decay $\pi^0 \rightarrow \gamma\gamma$. Extended Data Fig. 3 shows the fraction of jet power needed to be converted to protons to produce the observed γ -ray flux. We assume a proton spectrum $dN/dE \propto E^{-\alpha} \exp(-E/E_{\text{max}})$, and adopt a proton-proton interaction cross section of $\sim 50 \text{ mb}$ ⁴⁶, and a baryon density of $0.01 - 0.1 \text{ cm}^{-3}$ ^{16,27}. The total proton energy is obtained by integrating this spectrum normalized to the VHE γ -ray flux.

If the diffusion coefficient in the source is comparable to that in the ISM, no hadronic mod-

els would be allowed, as they would require a proton injection rate that exceeds the total kinetic luminosity of the jets of SS 433. Even in extreme circumstances, e.g., where the diffusion coefficient is extremely small (possibly due to scattering by turbulence generated from the streaming cosmic rays^{47,48}), particles could remain in the jet as long as the jet lifetime of 10^4 yr². Assuming that protons follow a hard spectrum with $\alpha < 2$, the hadronic scenario would still require that at least 30% of the jet power goes to protons. While a hadronic origin to the VHE flux is possible, it requires rather extreme source parameters and is therefore disfavored.

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