

Nuclear mass measurements map the structure of atomic nuclei and accreting neutron stars

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We present mass excesses (ME) of neutron-rich isotopes of Ar through Fe, obtained via time of flight $B\rho$ mass spectrometry at the National Superconducting Cyclotron Laboratory. Our new results have significantly reduced systematic uncertainties relative to a prior analysis, enabling the first determination of ME for $^{58,59}\text{Ti}$, ^{62}V , ^{65}Cr , $^{67,68}\text{Mn}$, and $^{69,70}\text{Fe}$. Our results show the $N = 34$ subshell weaken at Sc and vanish at Ti, along with the absence of an $N = 40$ subshell at Mn. This leads to a cooler accreted neutron star crust, highlighting the connection between the structure of nuclei and neutron stars.

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The rest mass m is a basic property of an atomic nucleus, essential for calculating astrophysical processes such as x-ray burst light curves and r -process nucleosynthesis, and key to mapping the evolution of nuclear structure across the nuclear landscape [1–3]. While nuclear masses nearly follow the whole-number rule, $m \approx Am_u$, where A is the mass number and $m_u = 931.49$ MeV is the atomic mass unit, the $\lesssim 1\%$ deviation from this relationship due to nuclear binding is notoriously difficult to predict. State-of-the-art mass models (e.g., [4–7]) often disagree in their predictions of the atomic mass excess (ME), $\text{ME}(Z, A) = m - (Z + N)m_u$, where Z is the proton number and N is the neutron number, by more than an MeV. Similar discrepancies are present when comparing predictions to experimentally measured masses. As such, experiments mapping the evolution of the nuclear mass surface across the nuclear landscape are essential.

For neutron-rich nuclides, mass measurements have revealed the emergence and disappearance of the magic

numbers that indicate enhanced nuclear binding (e.g., [8–15]). For instance, $N = 34$ semimagicity emerges for neutron-rich calcium isotopes [14], whereas there are signatures that the $N = 40$ subshell gap disappears for neutron-rich manganese [16]. While the evolution of these subshells were mapped by spectroscopy experiments that often long-preceded the corresponding mass measurements (e.g., [17–24]), nuclear masses provided the first model-independent confirmation of this spectroscopic evidence via fundamental ground state properties.

These evolutions in nuclear structure are closely linked to the thermal structure of accreting neutron stars. Nuclei produced by surface burning processes are buried by subsequent accretion, resulting in a number of nuclear reactions in the neutron star crust that drive it from thermal equilibrium [25]. Electron-capture (EC) reactions near closed shells result in relatively large EC heating due to the large change in the EC Q value $Q_{\text{EC}} = \text{ME}(Z, A) - \text{ME}(Z - 1, A)$ [26]. EC in regions of deformation between closed shells occur on isobars with a small odd-even staggering in Q_{EC} as well as low-lying excited states, which results in the EC- β^- -decay cycling process

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known as urca cooling [27]. Therefore, whether EC heating or cooling occurs, and the strength of the heat source or heat sink, strongly depend on nuclear masses [28–30].

To simultaneously map the evolution of the $N = 34$ and $N = 40$ subshells and constrain the thermal structure of accreting neutron stars, we performed mass measurements of neutron-rich isotopes of Ar through Fe ($Z = 18$ –26). First results from these measurements have been reported in Refs. [10,29,30]. The present work is a re-evaluation of the original data, incorporating recently published high-precision Penning trap mass data [12,13] as additional calibration nuclides, which greatly reduces the systematic uncertainty present in measurement results and greatly expands the number of nuclides for which masses are obtained. Our results extend the known nuclear mass surface, provide model-independent confirmation of the emergence of $N = 34$ and disappearance of $N = 40$ semimagicity, and significantly update predictions for urca cooling in accreted neutron star crusts.

Magnetic-rigidity corrected time-of-flight (TOF- $B\rho$) mass measurements were performed at the National Superconducting Cyclotron Laboratory. The measurement technique and measurements are described in detail in Refs. [10,29–32] and are only briefly summarized here. A 140 MeV/nucleon beam of ^{82}Se accelerated by the coupled cyclotrons impinged on a Be target and the resulting fully stripped (charge $q = Z$) fragments were transmitted through the A1900 fragment separator [33], momentum-analyzed at the target location of the S800 spectrograph [34], and stopped in the focal plane of the S800 [35]. Particle identification was performed event-by-event using the TOF- ΔE method, where TOF was provided by fast-timing scintillators separated by a flight path $L_{\text{path}} = 60.6$ m and energy loss ΔE determined using an ionization chamber. A relative measurement of $B\rho$, which is the momentum over q , was obtained via a position measurement using a microchannel plate detector located at the dispersive focus of the S800 [36,37].

Nominally, $m = (\text{TOF}/L_{\text{path}})(qB\rho/\gamma)$, where γ is the Lorentz factor. However, determining L_{path} and $B\rho$ to sufficient precision is not practicable. Instead, an empirical relationship between m/q and TOF is determined from a fit to nuclides of known m which are simultaneously measured alongside nuclides of interest. This work improves on prior results [10,29,30] by including seven additional calibration nuclides, bringing the total to 27. High-precision ME for $^{59-63}\text{Cr}$ [13] and $^{54,55}\text{Ti}$ [12] substantially improve constraints on the $m/q(\text{TOF})$ relationship, whose ambiguity previously provided one of the dominant contributions to our ME uncertainties [30].

Mass fits were performed as described in Ref. [30]. Several fit functions were explored, of the form

$$\frac{m}{q}(\tau) = a_0 + a_1\tau + a_2z + a_3\tau^2 + a_4z^2 + a_5z\tau + f(z, \tau),$$

where $\tau = \text{TOF} - \langle \text{TOF} \rangle$, $z = Z - \langle Z \rangle$, and $f(z, \tau)$ is a function of higher-order in z and/or τ . The addition of the new Ti and Cr reference nuclides resolved the previously existing ambiguity in the Z dependence, while the Cr masses additionally clarified that a higher-order TOF component was needed

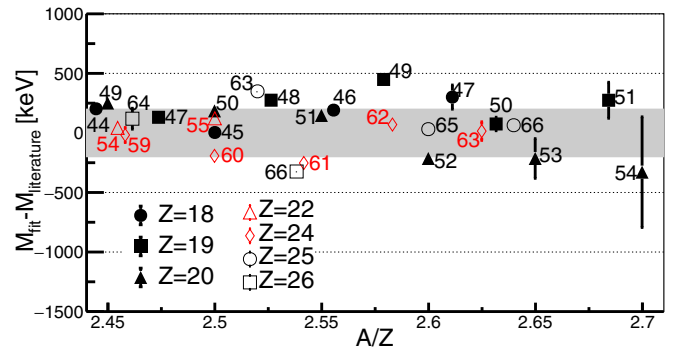


FIG. 1. Residuals of the $m/q(\text{TOF})$ fit to calibration nuclides, where the isotope of an element is indicated by the number next to the data point. Red data are new calibration nuclides in this re-evaluation. The gray band represents the average systematic mass uncertainty from the χ^2 normalization.

to adequately minimize fit residuals, which are shown for the best-fit in Fig. 1. The best-fit function has $f(z, \tau) = a_6z^3 + a_7\tau^4$, whereas a fit of slightly lower quality was obtained with $f(z, \tau) = a_6z^3 + a_7\tau^3$. This set of acceptable functions was determined by the following criterion. (1) The fit residuals must lack systematic trends. (2) The fit residuals must be robust to the arbitrary removal of reference nuclides. (3) The difference between χ^2 for a fit function and the best-fit $\Delta\chi_i^2 = \chi_i^2 - \chi_{\text{min}}^2$ must be within three standard deviations of the best-fit ($\Delta\chi_i^2 \lesssim 37$), which is a valid metric based on the Gaussian distribution of the fit parameters after repeated fits varying ME for reference nuclides in a Monte Carlo procedure [30,38].

The smaller number and closer similarity of the set of acceptable functions substantially reduce the fit-function uncertainty relative to the previous evaluation [30]. Additional uncertainty contributions come from the 9.1 keV/ q systematic uncertainty added to all nuclides to normalize the reduced χ^2 to one for the best-fit and the uncertainty in a_i due to TOF uncertainties in the reference nuclides. See Ref. [30] for details.

Our resultant ME are reported in Tables I and II with comparisons to literature values and results from our previous evaluation, where Fig. 2 shows the two neutron separation energy $S_{2n}(Z, A) = 2\text{ME}(0, 1) + \text{ME}(Z, A - 2) - \text{ME}(Z, A)$. All but one of our updated ME are within 1 standard deviation σ of our previous ME and all are within 2σ , while the majority of uncertainties have been reduced by a factor of two. We report ME for $^{58,59}\text{Ti}$, ^{62}V , ^{65}Cr , $^{67,68}\text{Mn}$, and $^{69,70}\text{Fe}$ for the first time.

The new trend in S_{2n} for V is largely due to $^{58,59}\text{V}$, which are much less bound in our work compared to the 2016 AME evaluated result, but in agreement with the privately communicated results Ref. [39] refers to as 1998Ba.A that were included in that evaluation. The abnormal behavior for Fe in S_{2n} shown in Fig. 2 is difficult to understand in terms of nuclear structure effects, since a new single-particle orbital is not expected to be filled. Additionally, anomalous behavior of $m/q(\text{TOF})$ is an unlikely explanation, since the function is smooth in that region and a similar feature is not seen

TABLE I. Atomic mass excesses (in keV) of nuclides determined in this work compared to the previous evaluation [10,29,30], the 2016 Atomic Mass Evaluation (AME) [39], and literature published after the 2016 AME. An asterisk * near the AME value indicates this is an extrapolation and not directly based on experimental data. The *I* following an isotope indicates a known or suspected long-lived (> 100 ns) isomer component. For instance, for ^{67}Fe the known isomer at 387 keV excitation energy is responsible for the additional asymmetric error bar, while for ^{69}Fe our results should be interpreted as an upper bound.

Isotope	This work	Previous evaluation	AME 2016	Literature
^{48}Ar	-22 390 (260)	-22 280 (310)	-22 280 (310)	-22 330 (120) [14]
^{49}Ar	-16 300 (1100)	-17 820 (1100)	-17 190* (400*)	...
^{52}Sc	-40 620 (230)	-40 300 (520)	-40 443 (82)	-40 525 (65) [15]
^{53}Sc	-38 400 (210)	-38 170 (570)	-38 906 (94)	-38 910 (80) [15]
^{54}Sc	-34 050 (240)	-33 750 (630)	-33 890 (273)	-34 485 (360) [15]
^{55}Sc	-31 090 (220)	-30 520 (580)	-30 159 (454)	...
$^{56}\text{Sc } I$	-25 380 (260) $^{(+0)}_{(-540)}$	-24 850 (590) $^{(+0)}_{(-540)}$	-24 852 (587)	...
^{57}Sc	-20 180 (880)	-21 010 (1320)	-20 996 (1304)	...
^{56}Ti	-39 480 (240)	...	-39 320 (121)	-39 810 (190) [15]
^{64}Cr	-33 640 (300)	-33 480 (440)	-33 480 (440)	...

for nearby Z . ^{67}Fe is known to have an isomeric state at 387 keV [42], which, along with measurement uncertainties, may explain the ~ 1 MeV deviation in S_{2n} from a smooth trend. We suspect ^{69}Fe may also have a long-lived isomer based on the presence of such states in odd- A isotopes of Fe due to intrusion from the $\nu g_{9/2}$ orbital [43]. The kink at $N = 39$ for Cr agrees with the trend calculated using the LNPS' Hamiltonian [13,30,44], but the absolute S_{2n} are discrepant.

ME for $N = 36$ isotopes of Sc, Ti, and V can be used to deduce the evolution of $N = 34$ in this region. Reference [14] demonstrated the presence of $N = 34$ semimagicity for Ca, whereas spectroscopy has indicated this subshell closure is absent for Ti and likely weakened for Sc, with the caveat that $E(2^+_1)$ energies can provide ambiguous constraints on shell gaps [45]. Our data reveal a continuous slope in S_{2n} for Ti following $N = 34$, while Sc trends slightly more negative

beyond this point. This indicates that a weak $N = 34$ subshell gap is present at Sc.

Further insight is provided by the trend in $D_n(Z, A) = (-1)^{N+1}[S_n(Z, A+1) - S_n(Z, A)]$, which is related to the empirical pairing gap [46]. D_n is proportional to the number of angular momentum projection states $(2j+1)$ participating in pairing, providing a signature of gaps in single particle levels. Figure 3(a) shows the trend in D_n near $N = 34$ for isotopes of Ca through V. While the dip after $N = 34$ for the Sc isotopes might initially appear to be the signature of a significant shell gap, shell model calculations suggest this is not the case. Calculations using the GX1A Hamiltonian [47], whose results are shown in Fig. 3(b) to be in qualitative agreement with experiment, indicate that the particularly low D_n for ^{56}Sc is due to the large splitting of levels created by the residual interaction between the $\pi f_{7/2}$ and $\nu f_{5/2}$ orbitals. In particular, the minimum in D_n for Sc is due to the low-lying $J = 1$ level created by this interaction. S_{2n} is not sensitive to

TABLE II. Table I continued, for cases without a previous evaluation or literature value.

Isotope	This work	AME 2016
^{57}Ti	-34 500 (240)	-33 916 (256)
^{58}Ti	-30 890 (250)	-31 110* (200*)
^{59}Ti	-25 220 (270)	-25 510* (200*)
^{57}V	-44 650 (260)	-44 413 (80)
^{58}V	-39 720 (230)	-40 402 (89)
^{59}V	-37 040 (260)	-37 832 (162)
$^{60}\text{V } I$	-32 810 (230) $^{(+0)}_{(-202)}$	-33 242 (220)
^{61}V	-30 380 (280)	-30 506 (894)
^{62}V	-25 340 (420)	-25 476* (298*)
^{65}Cr	-27 280 (780)	-28 220* (300*)
^{67}Mn	-33 960 (330)	-33 460* (300*)
^{68}Mn	-27 710 (1310)	-28 380* (400*)
$^{67}\text{Fe } I$	-45 560 (320) $^{(+0)}_{(-387)}$	-45 610 (270)
^{68}Fe	-44 360 (320)	-43 487 (365)
$^{69}\text{Fe } I$	-40 270 (400) $^{(+0)}_{(-?)}$	-39 030* (400*)
^{70}Fe	-37 710 (490)	-36 510* (400*)

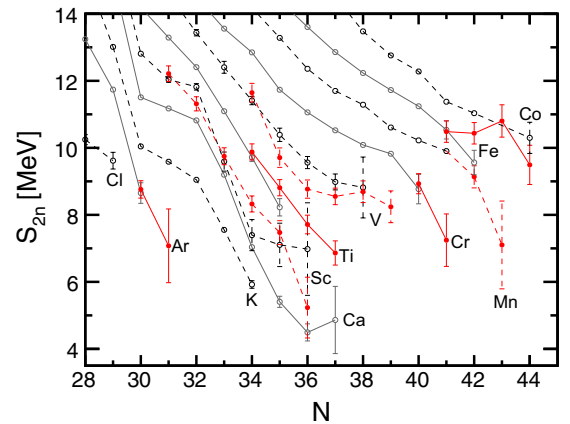


FIG. 2. S_{2n} for isotopes of Cl (left-most trend line) through Co (right-most), where the black and gray open circles correspond to values using the 2016 AME [39] or higher-precision ME from the subsequent literature [12,13,40,41], dashed lines are odd- Z , and solid lines are even- Z . Red filled circles are from this work.

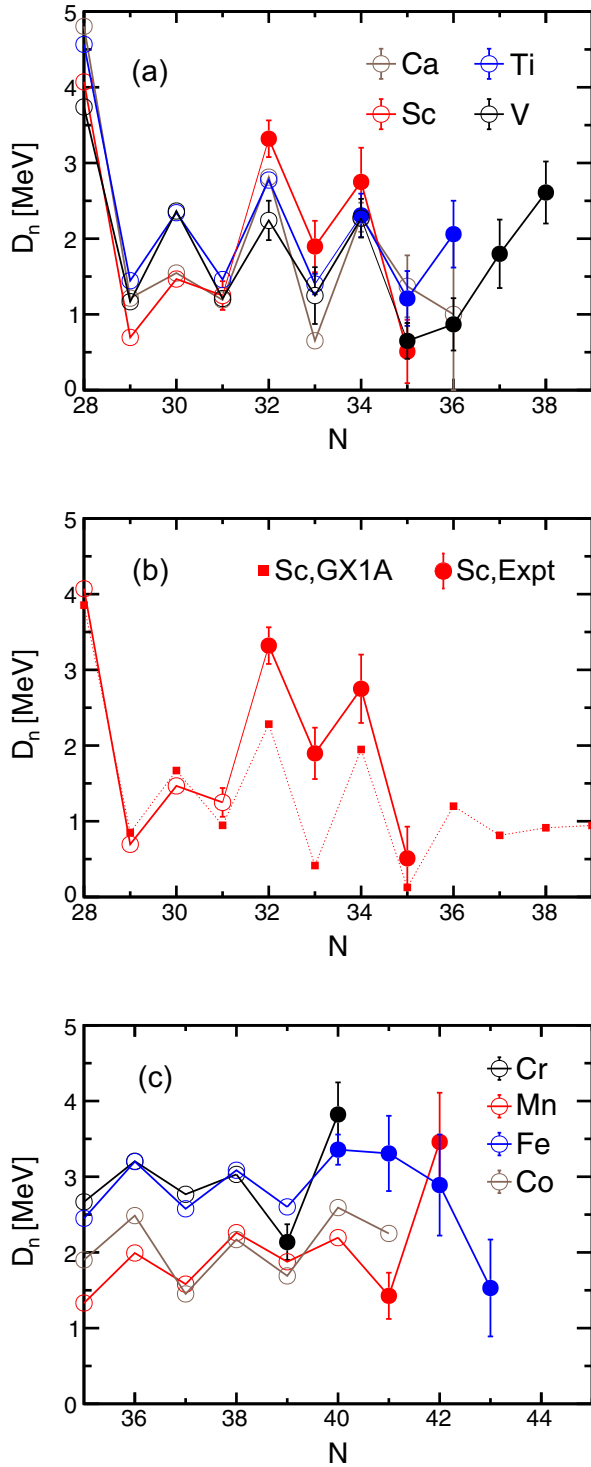


FIG. 3. D_n using our ME (filled circles) and ME from the literature (open circles) (a) near $N = 34$ and (c) near $N = 40$. A comparison to shell model calculations with the GX1A Hamiltonian (filled squares) is shown for Sc in (b).

this effect as it reflects the energetics of neutron shells and not the strength of proton-neutron pairing [3,16]. These results confirm previous indications from spectroscopy of weak

$N = 34$ magicity for Sc, while removing the ambiguity inherent to spectroscopic interpretations of shell structure [18,20,45].

Figure 2 demonstrates a continuous slope in S_{2n} through $N = 40$ for Mn, strengthening the conclusions of Ref. [16] that the $N = 40$ subshell is absent for this element. This is bolstered by the trend in D_n shown in Fig. 3(c). Our data are ambiguous regarding the $N = 40$ subshell at Cr, where the mass of ^{66}Cr is needed to confirm prior spectroscopic and Coulomb-excitation evidence (e.g., [23,48,49]).

The evolution in nuclear structure presented here is directly linked to the thermal structure of accreting neutron stars. The neutrino luminosity from urca cooling $L_\nu \propto X(A)(ft)^{-1}|Q_{\text{EC}}|^5$, where ft is the comparative half-life and $X(A)$ is the mass fraction, and is therefore very sensitive to changes in ME [50,51]. This process operates in the accreted neutron star crust with consequential L_ν for odd- A nuclides with $X(A) \gtrsim 0.5\%$, $ft \lesssim 5$, and $8 \lesssim |Q_{\text{EC}}| \lesssim 15$ MeV [27,52], where the upper limit on Q_{EC} is due to competing EC reaction channels [53]. Two EC parents predicted to produce some of the largest L_ν , with potentially observable consequences [51,52], are ^{55}Sc and ^{65}Mn . For ^{65}Mn , ft will be uncertain by orders of magnitude until measurements are enabled by next-generation rare isotope beam facilities, but estimates from QRPA calculations [27] and using the Moszkowski nomographs [54] result in $\log(ft) \approx 5$, allowing significant L_ν . ^{55}Sc by contrast has more consistent predictions, with $\log(ft) \approx 5$ using QRPA methods [27], the Moszkowski nomographs, and empirical systematics [55], currently making it the highest predicted L_ν urca cooling layer [52,56]. Furthermore, a measurement of ft for this transition has recently taken place [57]. For both $A = 55$ and 65 , $X(A) > 0.5\%$ and are remarkably consistent for a wide variety of assumptions for nuclear burning on the accreting neutron star surface [58]. Therefore Q_{EC} are the final piece of the nuclear physics puzzle for these urca coolers.

Using the newly determined $\text{ME}(^{55}\text{Sc})$, $|Q_{\text{EC}}(^{55}\text{Sc})| = 12.44(0.27)$ MeV, to be compared to the prior [14,39] value $11.51(0.48)$ MeV. Our reported $\text{ME}(^{65}\text{Cr})$ results in $|Q_{\text{EC}}(^{65}\text{Mn})| = 13.69(0.78)$ MeV, to be compared to the prior [16,39] value $12.75(0.30)$ MeV (though the latter uncertainty relies on the AME extrapolation, which assumes an essentially featureless nuclear mass surface and may be underestimated). These results increase L_ν by 50% for EC on ^{55}Sc with half the prior uncertainty and provide the first experimental determination of L_ν for EC on ^{65}Mn . While in agreement with previous predictions, our central value for $Q_{\text{EC}}(^{65}\text{Mn})$ leads to 40% larger L_ν . Therefore the accreted neutron star crust is cooler than previously thought, with an improved precision on the description of the neutron star thermal structure.

In summary, this work highlights the intriguing connection between evolution in nuclear structure and the thermal structure of accreting neutron stars. We find model-independent evidence for the onset of the $N = 34$ subshell for Sc and the likely absence of $N = 40$ magicity for Cr, each of which result in a larger mass difference for transitioning from

odd- Z to odd- N in EC. This is ultimately connected to the strength of the interaction between a nuclear core and an unpaired proton as opposed to an unpaired neutron [16,59], and leads to increasing the phase-space available for the weak transitions involved in urca cooling, which in turn results in a cooler neutron star crust. Our measurements leave $\text{ME}(^{63}\text{V})$ as the final nuclear mass important for urca cooling in the accreted neutron star crust that relies on theoretical mass estimates.

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