



An Episodic Wide-angle Outflow in HH 46/47

Yichen Zhang¹ , Héctor G. Arce² , Diego Mardones^{3,4} , Sylvie Cabrit^{5,6} , Michael M. Dunham⁷ , Guido Garay³, Alberto Noriega-Crespo⁸, Stella S. R. Offner⁹ , Alejandro C. Raga¹⁰ , and Stuartt A. Corder¹¹

¹ Star and Planet Formation Laboratory, RIKEN Cluster for Pioneering Research, Wako, Saitama 351-0198, Japan; yichen.zhang@riken.jp

² Astronomy Department, Yale University, P.O. Box 208101, New Haven, CT 06520, USA

³ Departamento de Astronomía, Universidad de Chile, Casilla 36-D, Santiago, Chile

⁴ Centre for Astrochemical Studies, Max-Planck-Institute for Extraterrestrial Physics, Giessenbachstrasse 1, D-85748 Garching, Germany

⁵ LERMA, Observatoire de Paris, UMR 8112 du CNRS, ENS, UPMC, UCP, 61 Av. de l'Observatoire, F-75014 Paris, France

⁶ Institut de Planétologie et d'Astrophysique de Grenoble (IPAG) UMR 5274, F-38041 Grenoble, France

⁷ Department of Physics, State University of New York at Fredonia, 280 Central Avenue, Fredonia, NY 14063, USA

⁸ Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA

⁹ Department of Astronomy, The University of Texas at Austin, Austin, TX 78712, USA

¹⁰ Instituto de Ciencias Nucleares, UNAM, Ap. 70-543, 04510 D.F., Mexico

¹¹ Joint ALMA Observatory, Av. Alonso de Córdova 3107, Vitacura, Santiago, Chile

Received 2019 February 11; revised 2019 July 31; accepted 2019 August 1; published 2019 September 17

Abstract

During star formation, the accretion disk drives fast MHD winds, which usually contain two components, a collimated jet and a radially distributed wide-angle wind. These winds entrain the surrounding ambient gas producing molecular outflows. We report a recent observation of ^{12}CO (2–1) emission of the HH 46/47 molecular outflow by the Atacama Large Millimeter/submillimeter Array, in which we identify multiple wide-angle outflowing shell structures in both the blueshifted and redshifted outflow lobes. These shells are highly coherent in position–position–velocity space, extending to $\gtrsim 40\text{--}50\text{ km s}^{-1}$ in velocity and 10^4 au in space, with well-defined morphology and kinematics. We suggest these outflowing shells are the result of the entrainment of ambient gas by a series of outbursts from an intermittent wide-angle wind. Episodic outbursts in collimated jets are commonly observed, yet detection of a similar behavior in wide-angle winds has been elusive. Here we show clear evidence that the wide-angle component of the HH 46/47 protostellar outflows experiences variability similar to that seen in the collimated component.

Key words: Herbig–Haro objects – ISM: clouds – ISM: individual objects (HH 46, HH 47) – ISM: jets and outflows – stars: formation

1. Introduction

Outflows play an important role in star formation and the evolution of molecular clouds and cores, as they remove angular momentum from the accretion disk (e.g., Bjerkeli et al. 2016; Hirota et al. 2017; Lee et al. 2017a; Zhang et al. 2018), carve out cavities in their parent cores (e.g., Arce & Sargent 2006), and inject energy and momentum into the star-forming environments (e.g., Arce et al. 2010; Plunkett et al. 2013). During star formation, the accreting circumstellar disk drives bipolar magneto-centrifugal winds (e.g., Konigl & Pudritz 2000; Shu et al. 2000). Models predict that these protostellar winds have both collimated and wide-angle components (e.g., Kwan & Takamaru 1995; Shang et al. 1998; Matt et al. 2003). The collimated portion of the wind, which is usually referred to as a jet, is typically traced by optical line emission in later-stage exposed sources (e.g., Reipurth & Bally 2001), or sometimes in molecular emissions in early-stage embedded sources (e.g., Tafalla et al. 2010; Lee et al. 2017a). The wide-angle component (presumably arising from a larger stellocentric radius in the disk) is thought to be slower than the collimated component, and does not produce the striking features seen in jets. In young embedded sources, the wide-angle component of a disk wind may be detected with

high-resolution molecular line observations (e.g., Tabone et al. 2017; Louvet et al. 2018). In more evolved pre-main sequence stars this component has been observed with optical atomic emission lines (e.g., Bacciotti et al. 2000).

Both jets and wide-angle winds can interact with the ambient molecular gas and entrain material to form slower, but much more massive outflows, which are typically observed in CO and other molecules and are generally referred to as molecular outflows. The entrainment process is not yet fully understood. Models include entrainment through jet–bow shocks (internal and/or leading) (e.g., Raga & Cabrit 1993) and wide-angle winds (e.g., Li & Shu 1996). In the jet–bow-shock entrainment model, a jet propagates into the surrounding cloud, and forms bow shocks which push and accelerate the ambient gas producing outflow shells surrounding the jet (e.g., Tafalla et al. 2016). In the wide-angle wind entrainment model, a radial wind blows into the ambient material, forming an expanding outflow shell (e.g., Lee et al. 2000). These two mechanisms may act simultaneously, as jet and wide-angle wind may coexist.

The accretion of material from a circumstellar disk onto a forming star is believed to be episodic (e.g., Dunham & Vorobyov 2012). The variation in the accretion rate may arise from various instabilities in the accretion disk (e.g., Zhu et al. 2010; Vorobyov & Basu 2015). In protostars, which are embedded in their parent gas cores (i.e., the so-called Class 0 and Class I sources), the most significant evidence of episodic accretion comes from jets that show a series of knots (which



Original content from this work may be used under the terms of the [Creative Commons Attribution 3.0 licence](https://creativecommons.org/licenses/by/3.0/). Any further distribution of this work must maintain attribution to the author(s) and the title of the work, journal citation and DOI.

sometimes are evenly spaced) along their axes (e.g., Lee et al. 2007; Plunkett et al. 2015). These knots often trace bow shocks that are formed by variations in the mass-loss rate or jet velocity, which in turn may be caused by variation in the accretion rate. However, such variability has not yet been seen in wide-angle outflows, which in principle should experience the same variations as jets.

Here we report recent ^{12}CO (2–1) observations of the HH 46/47 molecular outflow using the Atacama Large Millimeter/submillimeter Array (ALMA) that reveal multiple wide-angle outflowing shells, which we argue were formed by an episodic wide-angle wind. The HH 46/47 outflow is driven by a low-mass early Class I source (HH 47 IRS, which is also known as HH 46 IRS 1, IRAS 08242-5050) with a bolometric luminosity of $L_{\text{bol}} = 12 L_{\odot}$ that resides in the Bok globule ESO 216-6A, located on the outskirts of the Gum Nebula at a distance of 450 pc (Schwartz 1977; Reipurth 2000; Noriega-Crespo et al. 2004). Previous ALMA ^{12}CO (1–0) observations (Arce et al. 2013; Zhang et al. 2016, referred to as Papers I and II hereafter) showed a highly asymmetric CO outflow (with the redshifted lobe extending a factor of four more than the blueshifted lobe), as the driving source lies very close to the edge of the globule. In addition to the wide molecular outflow, collimated jets are also optically seen in the blueshifted lobe (Reipurth & Heathcote 1991; Eisloffel & Mundt 1994; Hartigan et al. 2005) and in the infrared in the redshifted lobe (Micono et al. 1998; Noriega-Crespo et al. 2004; Velusamy et al. 2007). Detailed analysis of the morphology and kinematics of the molecular outflow showed evidence of wide-angle wind entrainment for the blueshifted outflow lobe and jet–bow-shock entrainment for the redshifted lobe (see Papers I and II). The difference between the two molecular outflow lobes is likely due to the fact that the blueshifted jet is mostly outside of the globule where the outflow cavity has little or no molecular gas inside, while the redshifted jet is pushing through the core, surrounded by dense gas. However, even in the redshifted side, the energy distribution shows that more energy is injected by the outflow at the base of the outflow cavity, which is consistent with a wide-angle wind entrainment scenario, rather than at the jet–bow-shock heads, as a jet-entrainment scenario would suggest.

2. Observations

The observations were carried out using ALMA band 6 on 2016 January 6 with the C36-2 configuration and on 2016 June 21, 30 and July 6 with the C36-4 configuration (as part of observations for project 2015.1.01068.S). In the C36-2 configuration observation, 36 antennas were used and the baselines ranged from 15 to 310 m. The total on-source integration time was 75 minutes. J1107-4449 and J0538-4405 were used as bandpass and flux calibrators, and J0811-4929 and J0904-5735 were used as phase calibrators. In the C36-4 configuration observations, 36 antennas were used and the baselines ranged from 15 to 704 m. The total integration time was 150 minutes. J1107-4449, J0538-4405, and J1107-4449 were used as bandpass and flux calibrators, and J0811-4929 was used as a phase calibrator. The observations included only one pointing centered at $8^{\text{h}}25^{\text{m}}43^{\text{s}}.8$, $-51^{\circ}00'36''0$ (J2000), which is the 3 mm continuum peak obtained from the Cycle 1 observations (Paper II). The primary beam size (half power beamwidth) is about $23''$ at Band 6.

The ^{12}CO (2–1) emission at 230.54 GHz was observed with a velocity resolution of about 0.09 km s^{-1} . The center of the ^{12}CO spectral window, which has a bandwidth of 117 MHz ($\sim 150 \text{ km s}^{-1}$), is shifted from the ^{12}CO (2–1) line central frequency by 18 MHz ($\sim 23 \text{ km s}^{-1}$) in order to observe both ^{12}CO and ^{13}CO lines in one spectral setup. As a result, our ^{12}CO observation covers emission from $v_{\text{lsr}} = -94$ to $+56 \text{ km s}^{-1}$. The ^{13}CO (2–1), C^{18}O (2–1), H_2CO ($3_{0,3}-2_{0,2}$), and CH_3OH ($4_{2,2}-3_{1,2}$) lines were observed simultaneously in the same spectral setup. In addition, a spectral window with a bandwidth of 1875 MHz was used to map the 1.3 mm continuum. In this paper we focus on the ^{12}CO (2–1) and continuum data. We defer the discussions of other molecular lines to a future paper.

The data were calibrated and imaged in CASA (McMullin et al. 2007; version 4.5.3). Self-calibration was applied using the continuum data after normal calibration. The task CLEAN was used to image the data. For the spectral data we defined a different clean region for each channel. Robust weighting with a robust parameter of 0.5 was used in the CLEAN process. The resulting synthesized beam is $0''.65 \times 0''.47$ (P.A. = $87^\circ.8$) for the continuum data, and $0''.67 \times 0''.48$ (P.A. = $88^\circ.2$) for the ^{12}CO data. Throughout the paper we define the outflow velocity v_{out} as the LSR velocity of the emission minus the cloud LSR velocity, which is 5.3 km s^{-1} (van Kempen et al. 2009).

3. Results

Figure 1 shows the integrated intensity maps of the blueshifted and redshifted ^{12}CO (2–1) emission. Unlike the previous ^{12}CO (1–0) observations (shown in panel a), our ^{12}CO (2–1) single ALMA pointing observations only allow us to detect the outflow emission up to about $20''$ away from the protostar. Both lobes show conical morphologies of similar size, in contrast to what is seen when the full extent of the two lobes is observed. The ^{12}CO (2–1) emission is also more symmetric with respect to the outflow axis than the $J = 1-0$ emission in which the northern side of the blueshifted outflow is much brighter than its southern side. In Figure 1(b), the red lobe appears to be composed of different shell structures. At a distance of $7''$ from the central source, the inner shells delineate a U-like structure with a width of about $6''$ inside a cone-like shell that is about $9''$ wide. Although multiple shells are not clearly seen in the integrated image of the blue lobe, they are seen in the channel maps (see Figure 3). The 1.3 mm continuum emission shows an elongated structure perpendicular to the outflow axis, which is consistent with previous observations of the 3 mm continuum. The extended continuum emission appears to be shaped by the outflow, as it approximately follows the shape of the outflow cavity.

The shell structures are best seen in the position–position–velocity (PPV) diagrams (Figure 2), where they appear to be highly coherent in space and velocity. At least two shells can be identified in each lobe: Sb1 and Sb2 in the blueshifted lobe (panel a); and Sr1 and Sr2 in the redshifted lobe (panel b). There is possibly a third shell in the blueshifted outflow (Sb3) that appears to have a more complex structure (i.e., less coherent structure) than the other shells. Each of these shells shows a cone-like shape in the PPV space (best seen in shell Sr1 and Sb2), with a high-velocity side and a low-velocity side (also see Figure 5). In the redshifted lobe, both high-velocity and low-velocity sides of Sr1 and Sr2 are distinguishable (Figure 2(b)). However, in the blueshifted lobe, while the

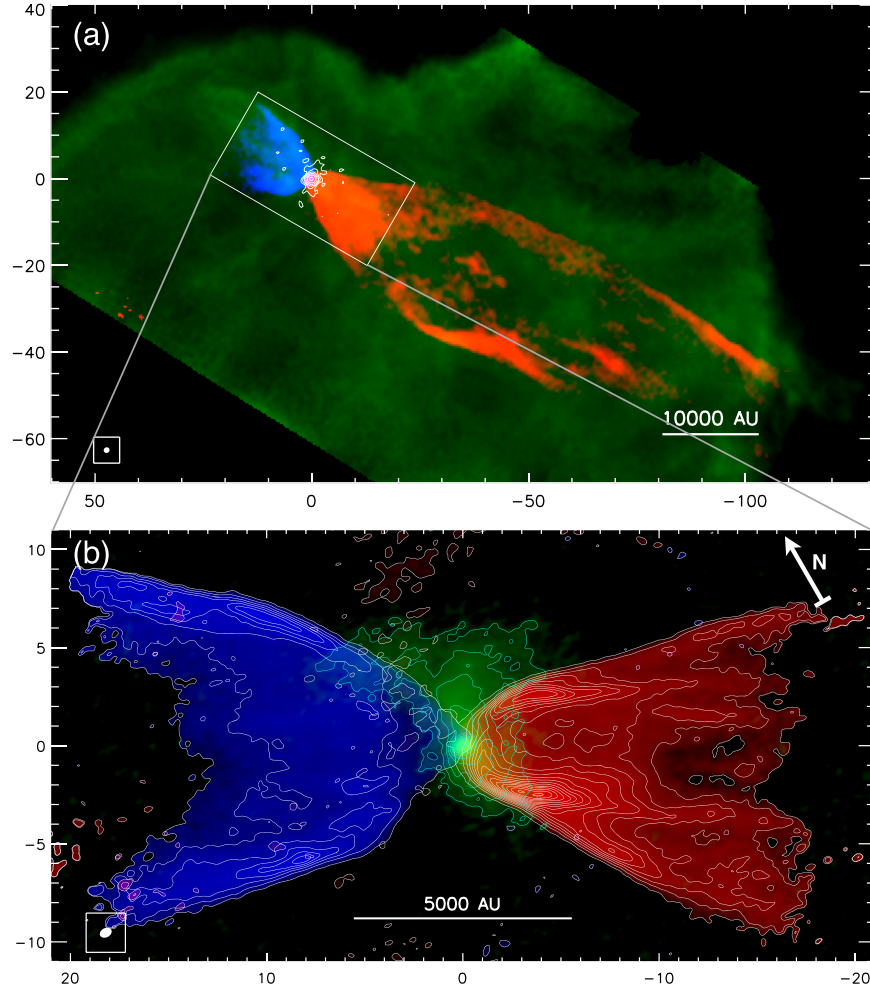


Figure 1. (a): ^{12}CO (1–0) integrated intensity map of the HH 46/47 molecular outflow at large scales from Figure 1 of Paper II. The red, blue, and green color scales show emission integrated over the velocity ranges from 1 to 10 km s^{-1} , from -10 to -1 km s^{-1} , and from -0.6 to 0.6 km s^{-1} (relative to the cloud velocity), respectively. The synthesized beam is $1''.37 \times 1''.31$. The white contours show the 100 GHz continuum emission. (b): ^{12}CO (2–1) integrated intensity map of the HH 46/47 molecular outflow overlaid on the 1.3 mm dust continuum emission. The blueshifted lobe is integrated from $v_{\text{out}} = -35$ to -10 km s^{-1} , and the redshifted lobe is integrated from $v_{\text{out}} = +10$ to $+50$ km s^{-1} . The contours start at 5σ and have intervals of 30σ ($1\sigma = 4.3$ $\text{mJy beam}^{-1} \text{ km s}^{-1}$ for the blueshifted lobe and $1\sigma = 4.9$ $\text{mJy beam}^{-1} \text{ km s}^{-1}$ for the redshifted lobe). The green color shows the 1.3 mm continuum emission with contour levels of $(2'' \times 5\sigma) \times 5\sigma$ ($n = 0, 1, 2, 3, \dots, 8$) with $1\sigma = 0.021$ mJy beam^{-1} . The image are rotated by 30° counterclockwise. The synthesized beam of the ^{12}CO map is $0''.67 \times 0''.48$, and is shown in the lower left corner of the panel.

high-velocity sides of Sb1 and Sb2 are clearly separate, the low-velocity sides of these two shells seem to have merged. We expect the high-velocity side of a shell seen in PPV space to correspond to the front side of the blueshifted shell or the back side of the redshifted shell as the expanding motion of the outflow shell, in addition to the outflowing motion, is contributing to the observed line-of-sight velocities. On the other hand, at a particular line-of-sight velocity these shells have shapes similar to ellipses, partial ellipses, or parabolas (see Figures 3 and 4). Therefore, structures seen in different positions and velocities can come from a single coherent structure.

In addition to the velocity field within one shell, the overall velocities of the shells are different from each other, which is shown by their different opening directions in the PPV space. For example, in the redshifted lobe, shell Sr1 is generally faster than Sr2 (i.e., the velocity of the Sr1 shell at any distance from the protostar is higher than that of the Sr2 shell at the same distance), while in the blueshifted lobe, shell Sb1 is generally faster than Sb2 (see also Figure 5). The shape of the shells is similar, but some have different widths. In the redshifted lobe,

shell Sr1 is much narrower than shell Sr2 (see also Figure 4). Because shell Sr1 is faster and narrower than Sr2, the two shells intersect in PPV space (Figure 2(a)). In the blueshifted lobe, however, the shells appear to have similar widths. At low velocities (in the lower part of the two PPV diagrams in Figure 2), the emission becomes complex and has many substructures, therefore it cannot be clearly identified as being part of one of the shells identified at higher velocities.

In Figures 3 and 4 we plot the channel maps of the ^{12}CO (2–1) emission. Significant emission in the blueshifted lobe is detected up to about $v_{\text{out}} = -35$ km s^{-1} , even though the spectral window covers velocities up to $v_{\text{out}} = -99$ km s^{-1} . In this outflow lobe, the emission moves away from the central source as the velocity increases. At blueshifted outflow velocities of about $v_{\text{out}} = -35$ km s^{-1} the emission is found at the edge of our map. Thus, it is probable that there exists higher velocity outflow emission beyond the border of our map. In the redshifted lobe, the emission is still quite strong at the edge of the spectral window, which only covers up to outflow velocities of about $v_{\text{out}} = +51$ km s^{-1} . Hence, we suspect the redshifted lobe extends to even higher velocities.

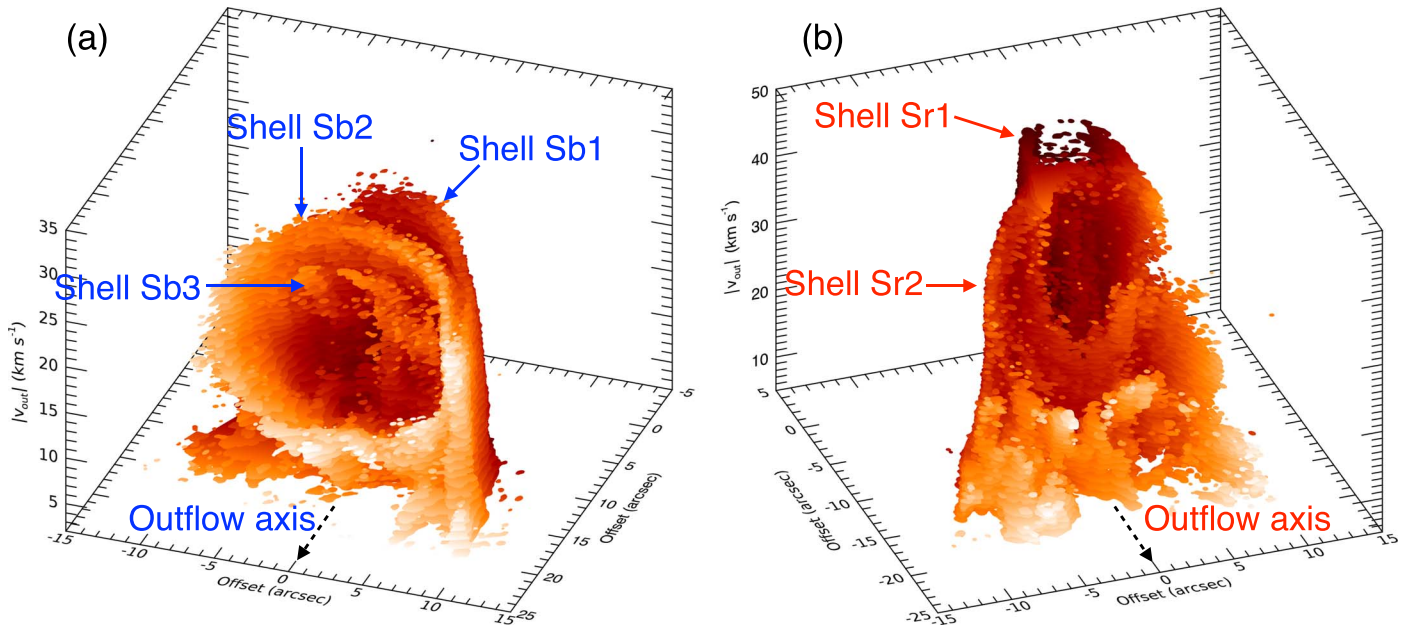


Figure 2. (a): Position–position–velocity diagram for the ^{12}CO (2–1) blueshifted lobe. Channels at $|v_{\text{out}}| > 2 \text{ km s}^{-1}$ with emission higher than 5σ ($1\sigma = 1.6 \text{ mJy beam}^{-1}$ for a channel width of 0.3 km s^{-1}) are included. The emission outside of the outflow cavity is not included. The position of the outflow source, at offset (0, 0), is at the back, and the outflow direction is toward the reader. The color is selected to emphasize the layered structure. (b): Same as panel (a) but for the redshifted lobe. In this panel, data with velocities $|v_{\text{out}}| > 5 \text{ km s}^{-1}$ are included.

The shell structures identified in the PPV diagrams are clearly seen in these channel maps, and are labeled in the figures. In the blueshifted outflow, shells Sb1 and Sb2 can be easily distinguished at velocities $v_{\text{out}} = -33$ to -20 km s^{-1} . Shell Sb3 is seen at velocities $v_{\text{out}} = -24$ to -21 km s^{-1} . At $v_{\text{out}} = -18$ to -9 km s^{-1} , the emission inside the cavity is actually from a structure different from Sb3 (best seen in Figure 5), which we label Sb4. At these relatively lower outflow velocities, the shells Sb1, Sb2, and Sb3 appear to merge together and show a full elliptical shape. It is not clear whether the far side of the ellipse corresponds to the low-velocity side of one of the Sb1, Sb2, or Sb3 shells, or a structure produced from the combination of these three shells (also see Figure 5). At velocities of $|v_{\text{out}}| < 20 \text{ km s}^{-1}$, additional emission appears close the central source (which we label Sb5), showing a cone shape rather than an elliptical shape. The major structures Sb1, Sb2, and Sb3, all shift to the northeast (i.e., left in Figure 3) and become wider as the outflow velocity increases.

In the redshifted outflow, the two main shell structures Sr1 and Sr2 are best seen in the outflow velocity range from $+30$ to $+50 \text{ km s}^{-1}$. As discussed above, the widths of the two main shells are quite different. Although in general, the bulk of the emission shifts away from the source as the outflow velocity increases, the position of the narrower shell (Sr1), does not change much. As discussed above, the two shells intersect in the PPV space and this is most clearly seen in the $v_{\text{out}} = 39$ to 45 km s^{-1} channels in Figure 4. At low velocities (e.g., $|v_{\text{out}}| < 15 \text{ km s}^{-1}$), the Sr1 shell can still be discerned even though a significant amount of material fills the outflow cavity.

Figure 5 shows the position–velocity (PV) diagrams along the outflow axis and perpendicular to the outflow axis. They correspond to the intersections of the PPV diagram with different position–velocity planes. As discussed above, a shell in the PPV space has a high-velocity side and a low-velocity side, which becomes evident in the PV diagrams. We label

different structures in Figure 5 with “H” or “L” to indicate the high- and low-velocity sides of the same shell. In the redshifted lobe, pairs of high- and low-velocity structures of the same shell are easily identified (Sr1H/Sr1L and Sr2H/Sr2L). There is also emission between the Sr2H and Sr1L structures, and emission at larger distances with low velocities, which cannot be identified as part of a shell. In the blueshifted lobe, while the high-velocity sides of the Sb1, Sb2, and Sb3 shells are easily distinguished, their corresponding low-velocity walls are not so clear. It is unclear whether the structure labeled “Sb1L, Sb2L, Sb3L” corresponds to the low-velocity side of one of the three shells (Sb1, Sb2 or Sb3), or if this structure is produced by the merger of the low-velocity side of all three shells. It is also unclear whether Sb4 and Sb5 are separate structures or the low- and high-velocity sides of a single shell.

4. Discussion

4.1. Shell Model Fitting

To be more quantitative, we fit the morphology and kinematics of the outflow shells with expanding parabolas. Following the method by Lee et al. (2000), the morphology and velocity of a single expanding parabolic shell can be described (in cylindrical coordinates) as

$$\left(\frac{z}{R_0}\right) = \left(\frac{R}{R_0}\right)^2, \quad v_z = \frac{z}{t_0}, \quad v_R = \frac{R}{t_0}, \quad (1)$$

where the z -axis is along the outflow axis, the R -axis is perpendicular to the z -axis, and v_z and v_R are the velocities in the directions of z and R (i.e., the forward velocity and the expansion velocity), respectively. The free parameters in this model are the inclination i between the outflow axis and the plane of the sky, parameter R_0 , which determines the width of the outflow shell, and t_0 , which determines the velocity distribution of the outflow shell. Note that the characteristic

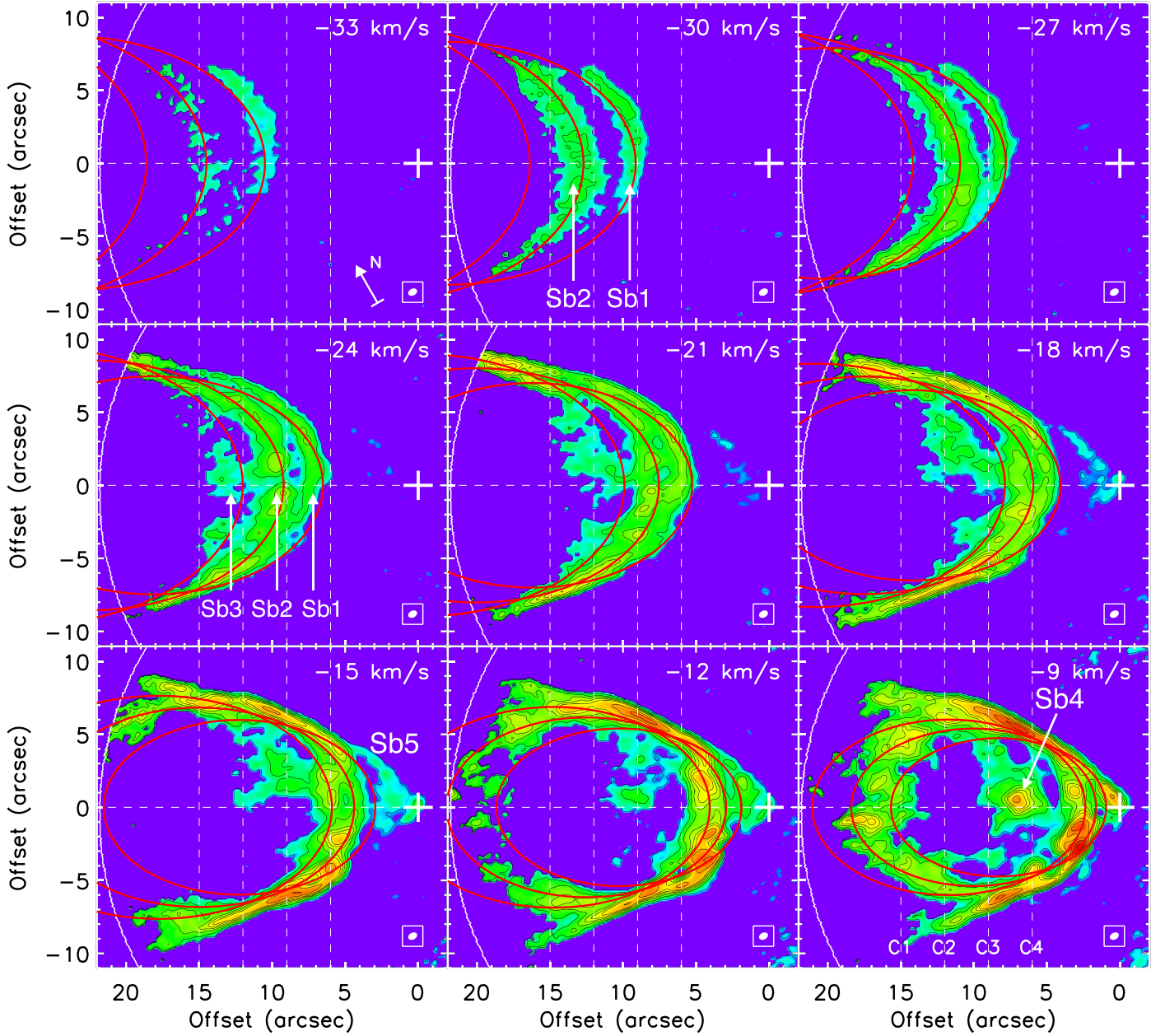


Figure 3. Channel maps of the ^{12}CO (2–1) emission of the blueshifted lobe. The central velocity of each 3 km s^{-1} -wide channel is indicated in the upper right corner of each panel. The maps are rotated by 30° so that the outflow axis is along the x -axis. The contours are at levels of 5σ , 20σ , 40σ , 60σ , 80σ , 100σ , 150σ , 200σ , 250σ , 300σ , and 350σ with $1\sigma = 0.52 \text{ mJy beam}^{-1}$. The white crosses mark the position of the 1.3 mm continuum peak. The white curve to the left of each panel indicates the edge of the field of view with a primary beam gain greater than 0.1. The horizontal and vertical dashed lines show the cuts for the position–velocity diagrams in Figure 5. The major shell structures are labeled. Red ellipses are the “best-fit” models for shells Sb1, Sb2, and Sb3.

radius R_0 is just the radius of the shell at $z = R$, i.e., $\theta = 45^\circ$. Also, t_0 can be considered the dynamical age of the shell. We further define a characteristic velocity $v_0 \equiv R_0/t_0$, which is the outflowing velocity v_z or expanding velocity v_R at $z = R$. In such a model, the half-opening angle of the shell at the height of z is $\tan \theta_{\text{open}}(z) = \sqrt{R_0/z}$.

Such a model predicts an elliptical shape of emissions in the channel maps, and as the channel velocity increases the elliptical structure becomes wider and shifts further away from the central source. The model also predicts a parabolic shape in the PV diagram along the outflow axis and elliptical shapes in the PV diagrams perpendicular to the outflow axis. Such behaviors are indeed consistent with our observations. The

same model was used to explain the blueshifted lobe in Paper II, in which we did not have enough spatial resolution and sensitivity to detect the multiple shell structures.

We fit the shells Sb1 and Sb2 in the blueshifted lobe and shells Sr1 and Sr2 in the redshifted lobe by comparing the model described above with the observed emission distributions in both channel maps and PV diagrams. Here, we only focus on the location of the emission in space and velocity, and do not attempt to reproduce the intensity distribution. To reduce the number of free parameters, we adopt a constant inclination angle for shells in the same lobe. To perform the fitting, the inclination i is searched within a range $30^\circ \leq i \leq 45^\circ$ with an interval of 5° , the parameter t_0 is searched within a range from 0.05 to $1 \text{ arcsec km}^{-1} \text{ s}$ with an

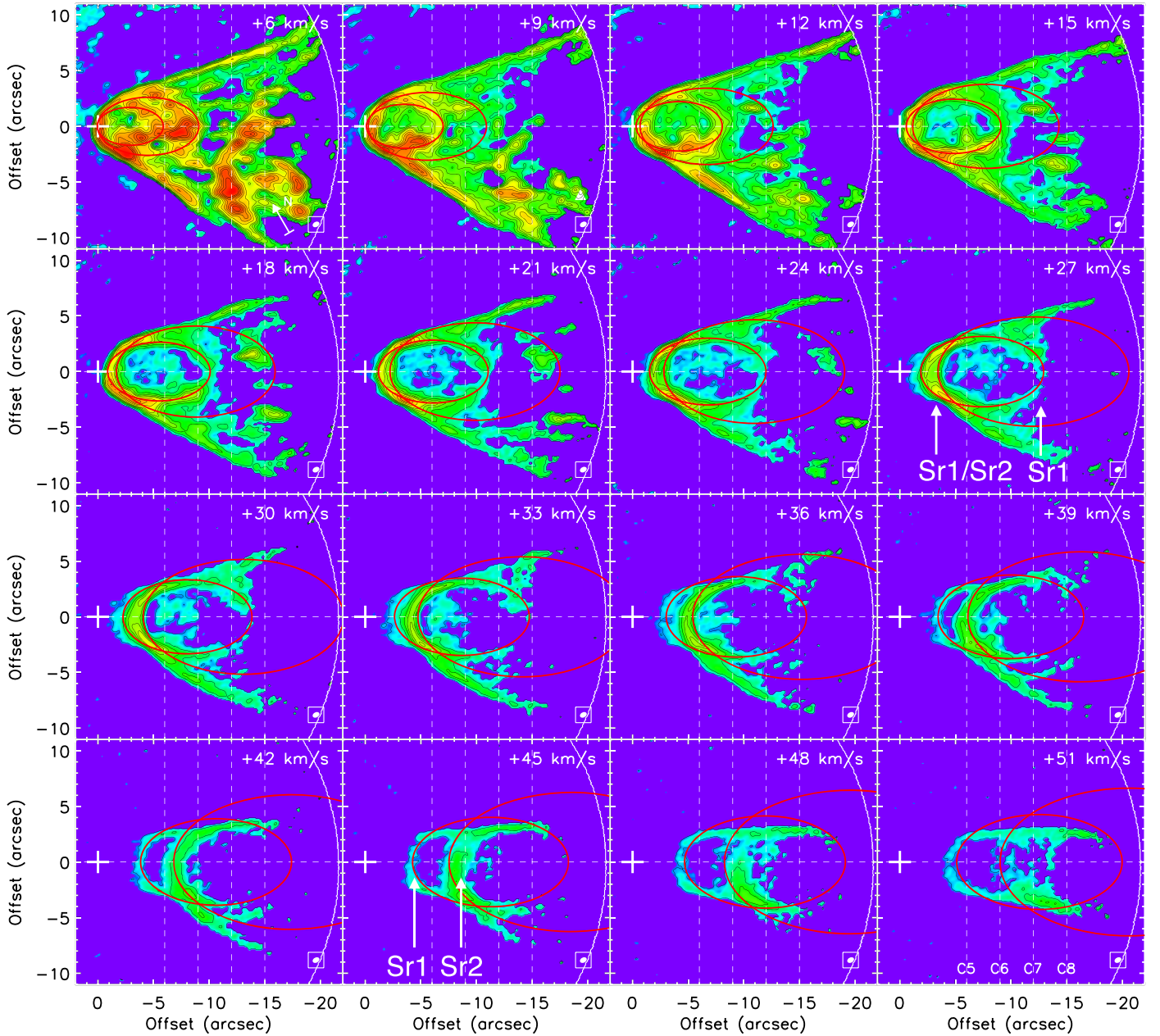


Figure 4. Same as Figure 3, but for the redshifted lobe.

interval of $0.05 \text{ arcsec km}^{-1} \text{ s}$, and the parameter R_0 is searched with an interval of $0''.01$ within a range from $1''$ to $3''$. Furthermore, in fitting the shells Sb1 and Sb2, we assume that the low-velocity walls of these two shells have merged into one structure (labeled as “Sb1L, Sb2L, Sb3L” in Figure 5).

The best-fit models are selected by visually comparing the model curves (which are shown by the red curves in Figures 3 and 4, and the blue and red curves in Figure 5) with the observed distribution of the outflow emission. The parameter values of what we consider the “best-fit” models are listed in Table 1, including the characteristic velocities $v_0 = R_0/t_0$ and the shell half-opening angles θ_{open} at $z = 15''$. The fitted inclinations are $i = 40^\circ$ and $i = 35^\circ$ between the outflow axis and the plane of sky for the blueshifted and redshifted outflows, respectively. These are consistent with the values derived from observations of the optical (blueshifted) jet by Eisloffel & Mundt (1994) and

Hartigan et al. (2005), which are $34^\circ \pm 3^\circ$ and $37.5^\circ \pm 2.5^\circ$, respectively.

In the redshifted outflow, the Sr1 and Sr2 shells are fit with $t_0 = 0.15$ and $0.25 \text{ arcsec km}^{-1} \text{ s}$ and $R_0 = 1''.3$ and $1''.9$. These models describe the two shells relatively well, especially the Sr1 shell. The model fit to Sr2 is not as good as that of Sr1, especially at higher velocities. This is partly due to the fact that Sr2 is slightly asymmetric with respect to the outflow axis (seen more clearly in the channel maps at $v_{\text{out}} = +33$ to $+45 \text{ km s}^{-1}$ Figure 4). Sr2 also appears tilted (or skewed) in the PV diagrams perpendicular to the outflow axis (panels h and i of Figure 5). These features may be caused by rotation or a slight change in the outflow direction, which the models do not take into account. The fitted values of t_0 of 0.15 and $0.25 \text{ arcsec km}^{-1} \text{ s}$ correspond to timescales of 3.2×10^2 and $5.3 \times 10^2 \text{ yr}$ assuming a source distance of 450 pc, which can be considered the dynamical ages of these two shells, result in

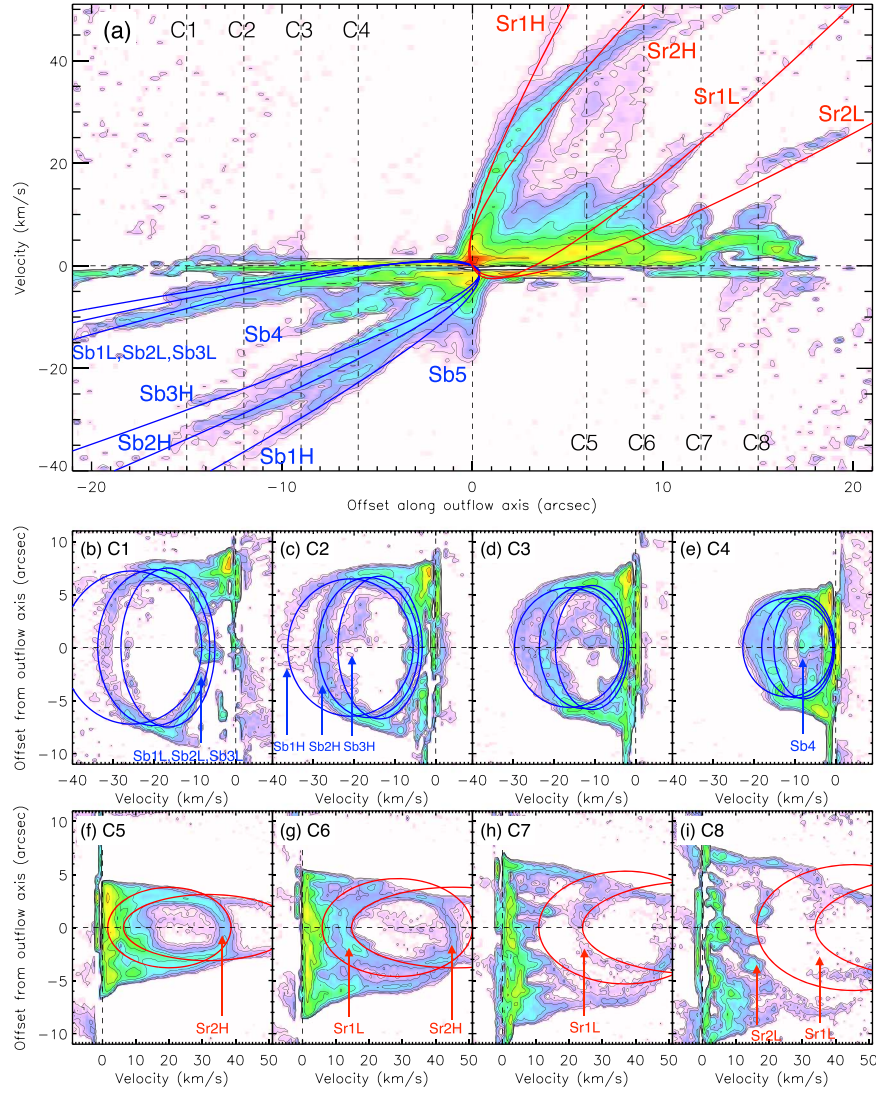


Figure 5. (a): Position–velocity diagram of the ^{12}CO (2–1) emission along the outflow axis with a cut width of $1''$. (b)–(i): Position–velocity diagrams of the ^{12}CO (2–1) emission along $1''$ -wide cuts perpendicular to the outflow axis. The cuts are shown in Figures 3 and 4. The offset positions of the cuts perpendicular to the outflow axis are also shown in panel (a). The contours are at levels of $(2'') \times 5\sigma$ ($n = 0, 1, 2, 3, \dots$) with $1\sigma = 0.56 \text{ mJy beam}^{-1}$. The major structures are labeled. The red and blue curves are the “best-fit” models for shells Sb1, Sb2, and Sb3 (blue lines) and shells Sr1 and Sr2 (red lines).

an age difference between the Sr1 and Sr2 shells of $2.1 \times 10^2 \text{ yr}$.

In the blueshifted side, the parameters for the best-fit models for shells Sb1 and Sb2 are $t_0 = 0.55$ and $0.70 \text{ arcsec km}^{-1} \text{ s}$ and $R_0 = 2''.6$ and $2''.7$. The fitted values of t_0 correspond to timescales of 1.2×10^3 and $1.5 \times 10^3 \text{ yr}$, which result in an age difference between the Sb1 and Sb2 shells of $3.2 \times 10^2 \text{ yr}$. If we assume that the time interval between shells Sb1 and Sb2 is the same as the interval between Sb2 and Sb3, then we can estimate a value for t_0 for shell Sb3 of $0.85 \text{ arcsec km}^{-1} \text{ s}$ (by adding $\Delta t_{0,B} \equiv t_{0,\text{Sb2}} - t_{0,\text{Sb1}}$ to the estimated t_0 for Sb2). With this assumption, we find that using a value of $R_0 = 2''.8$ results in a model that agrees fairly well to shell Sb3. The widths of the three shells, parameterized by R_0 , are slightly different, the fastest shell (Sb1) being the narrowest and the slowest shell (Sb3) being the widest. Varying shell widths are needed for the model to fit the velocity gradient seen at the edge of the blue lobe cavity (e.g., at position offsets of about $6''$ on both sides of outflow axis) in the PV diagrams perpendicular to the

outflow axis (panels (b)–(e) of Figure 5), where the emission becomes wider at lower outflow velocities.

From the fitted models, it can be seen that the blueshifted shells in general are wider, slower, and much older than the redshifted shells. On each side, the faster, younger shells are also narrower than the slower, older shells. However, in the blueshifted side, the shells have very similar widths (which can also be seen from the half-opening angles listed in Table 1), while on the redshifted side, the two shells have clearly different widths. Furthermore, the three blueshifted shells can be explained by outflow shells of different dynamical ages, with similar age differences among consecutive shells.

4.2. Origin of the Multiple Shell Structure

The parabolic outflowing shells can be produced by entrainment by a wide-angle wind (Li & Shu 1996; Lee et al. 2000). In such models, the molecular outflow is swept up by a radial wide-angle wind with force distribution $\propto 1/\sin^2(\theta)$, where θ is the polar angle relative to the outflow axis. Such a wind interacts with a flattened ambient core with density

Table 1
Parameters of the Fitted Parabolic Shells

Shell	i^a	R_0 (arcsec)	t_0 (arcsec km $^{-1}$ s)	Age (10^3 yr) b	v_0 (km s $^{-1}$) c	θ_{open} ($z = 15''$) d
Sb1	40°	2.6	0.55	1.2	4.7	22°6
Sb2		2.7	0.70	1.5	3.9	23°0
Sb3		2.8	0.85	1.8	3.3	23°3
Sr1	35°	1.3	0.15	0.32	8.7	16°4
Sr2		1.9	0.25	0.53	7.6	19°6

Notes.

^a Inclination angle between the outflow axis and the plane of sky. The same value is used for shells in the same lobe in order to reduced the numbers of free parameters.

^b Dynamical age calculated from t_0 assuming a distance of $d = 450$ pc.

^c Characteristic velocity of the shell defined as $v_0 \equiv R_0/t_0$.

^d Half-opening angle of the fitted shell at a height of $z = 15''$, $\tan \theta_{\text{open}} = \sqrt{R_0/15''}$.

distribution $\propto \sin^2(\theta)/r^2$, and instantaneously mixes with shocked ambient gas. The resultant swept-up outflowing shell is then a radially expanding parabola with a Hubble-law velocity structure.

Because each shell can be well fit with the wide-angle wind entrainment model, it is natural to explain the multiple shell structure as being formed by the entrainment of ambient circumstellar material by multiple outbursts of a wide-angle wind. One outburst of the wide-angle wind may not be able to entrain all the material to clear up the cavity, therefore the later outbursts will continue to entrain material to form subsequent shells. In such a scenario, the time intervals between successive shells, which are 2.1×10^2 yr for the redshifted outflow and 3.2×10^2 yr for the blueshifted outflow, can be considered the time interval between wind outbursts. These estimated outburst intervals are consistent with those seen in the episodic knots of HH 46/47 and in other sources. In the HH 46/47 outflow, an outburst interval of about 300 yr was estimated from the knots observed along the jet (see Paper I). Plunkett et al. (2015) estimated outburst intervals to range from 80 to 540 yr, with a mean value of 310 yr for a young embedded source in the Serpens South protostellar cluster. We thus suggest that in HH 46/47 the multiple shell structures may arise from the same high accretion rate episodes, which is reflected in both the jet and wide-angle wind components of the outflow. In fact, in Paper I, the identified jet knots R1 and R2 are found to have dynamical ages of 360 and 650 yr, close to the ages of shells Sr1 (320 yr) and Sr2 (530 yr), suggesting that the episodicity seen in the jet and the wide-angle outflows may be caused by the same outburst events.

We note that the dynamical ages of these shells estimated here may not accurately reflect their true ages. If the outburst happens in a short time compared to the dynamical timescale of the outflow shell, the shell entrained during such an outburst event will decelerate due to the interaction with the surrounding material. Therefore, it is likely that the estimated dynamical ages are upper limits of their true ages. The time intervals are also likely to be upper limits. Yet, the similarity between the time intervals estimated here and those estimated from the jet supports the scenario that the different observed shell structures are produced in multiple outburst events.

If the wide-angle shells on both sides are caused by the same accretion bursts in the disk, then a shell in the blueshifted side should correspond to a shell in the redshifted side. However, it is difficult to identify such pairs in our data because the entrainment is affected by significantly different environments

with which each lobe is interacting. The dynamical ages of the identified shells in the blueshifted lobe are significantly larger than those of the shells in the redshifted lobe and also significantly higher than the time interval between shells ($\Delta t_{0,B}$). It is therefore likely that the Sb1, Sb2, or Sb3 shells on the blueshifted lobe are not caused by the most recent outburst events. There may not be much molecular gas left in the blue lobe cavity in order for the youngest outburst to entrain any material to form a shell detectable in CO (unlike the redshifted lobe, see below), as previous outflow bursts may have cleared the cavity. Observations of the optical jet on the blueshifted side found that the furthestmost jet knot (HH 47D) has a dynamical age of 1.3×10^3 yr (Hartigan et al. 2005), which is similar to the age of shell Sb1, and other knots closer to the protostar have much younger ages. This supports that Sb1, Sb2, and Sb3 shells are not caused by recent, but by relatively old outburst events. It is possible that the Sb4 and Sb5 structures in the blueshifted side are caused by the most recent outburst, but due to the lack of ambient cloud material, the CO emission associated with these shell is only concentrated in the region close to the protostar.

Unlike the shells in the blue lobe, the youngest shell in the red lobe (Sr1) has an age similar to the outburst interval. Hence, the Sr1 shell may be the product of the most recent outburst. The red lobe is immersed in the dense part of the parent core and therefore there is still abundant material inside this cavity. Also, the fact that Sr1 is significantly narrower than shell Sr2, is consistent with the scenario where Sr1 has only formed recently. Because the narrower and newer shells are faster than the older and wider shells (see Section 4.1), they are expected to collide with older shells in the future. Based on the sizes R_0 and velocities (v_0 ; assumed to be constant), Sb1 will catch up with Sb2 in $(R_{0,\text{Sb2}} - R_{0,\text{Sb1}})/(v_{0,\text{Sb1}} - v_{0,\text{Sb2}}) = 2.7 \times 10^2$ yr, Sb2 will merge with Sb3 in 3.6×10^2 yr, and Sr1 will reach Sr2 in 1.2×10^3 yr. The real catch-up timescales should be shorter than these, as the outer shells are likely to slow down due to the interaction with the dense ambient material. This may explain why only two or three shells can be detected on both sides, as the shells may only survive for a few outburst periods before they collide with the old shells and form the outflow cavity walls seen in the low-velocity channels.

In order to further explore whether the observed outflow shells are caused by entrainment/interaction with the envelope or are being directly launched from the disk, we estimate the mass and momentum rates of these shells from the ^{12}CO (2–1) emission. To obtain the gas mass, we assume optically thin

emission and adopt an abundance of ^{12}CO of 10^{-4} relative to H_2 and a gas mass of 2.34×10^{-24} g per H_2 molecule. Following Paper II, we adopt an excitation temperature of $T_{\text{ex}} = 15$ K. An excitation temperature of 50 K would increase the mass estimate by a factor of 1.5. In each velocity channel, we only include the primary-beam-corrected emission above 3σ and within a primary beam response greater than 0.2 relative to the phase center. We include all the emission associated with the outflow, except the emission at outflow velocities less than 2 km s^{-1} in order to avoid possibly adding emission from core material to our outflow mass estimate.

We estimate a total mass of 5.6×10^{-3} and $1.0 \times 10^{-2} M_{\odot}$ and momentum of 4.4×10^{-2} and $8.7 \times 10^{-2} M_{\odot} \text{ km s}^{-1}$ in the shells of the blue and redshifted outflow lobes, respectively. Here, in calculating the total momenta, we use the velocity of each channel and multiply by the outflow mass of that channel. These are very likely lower limits due to the optically thin assumption, unaccounted low-velocity outflowing material, and possible higher excitation temperatures (e.g., Dunham et al. 2014). Using the estimated ages of the oldest shells on both sides (1.8×10^3 yr for Sb3 and 5.3×10^2 yr for Sr2), the time-averaged mass outflow rates are 3.1×10^{-6} and $1.8 \times 10^{-5} M_{\odot} \text{ yr}^{-1}$ for the blueshifted and redshifted outflow lobes, respectively. And the time-averaged momentum injection rates are 2.4×10^{-5} and $1.6 \times 10^{-4} M_{\odot} \text{ yr}^{-1} \text{ km s}^{-1}$ for the blue-shifted and redshifted outflow lobes, respectively. Note that these rates are averaged over the outflow age, and the mass loss and momentum injection rates during each outburst are expected to be significantly higher than these values.

The above estimates for the mass outflow rates are one to two orders of magnitude larger (or even larger given that the values quoted above are lower limits) than most estimates of the mass-loss rate for the HH 46/47 protostellar jet using optical and IR atomic line emission, which range between 0.3 and $5 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ (e.g., Hartigan et al. 1994; Antoniucci et al. 2008; Garcia Lopez et al. 2010; Mottram et al. 2017).¹² Moreover, if we assume a mass-loss rate of $\sim 10^{-7} M_{\odot} \text{ yr}^{-1}$ and a velocity of about 100 km s^{-1} (e.g., Morse et al. 1994; Hartigan et al. 2005) for the wind launched by the disk, this then leads to a momentum injection rate of approximately $10^{-5} M_{\odot} \text{ yr}^{-1}$. These results show that the observed ^{12}CO (2–1) shells have mass-loading rates that are one to two orders of magnitudes higher than the mass-loss rates of the jet (or wind) launched from the disk, but have momentum injection rates similar to the jet/winds directly launched from the disk. This is consistent with the scenario in which the observed ^{12}CO (2–1) shells are mostly made of ambient material that was entrained by the wind launched from the disk, in a momentum-conserving interaction, and not of material that was directly launched from the disk. This is also consistent with theoretical simulations that show that only 25%–30% of the mass in a molecular outflow is directly launched from the disk and the rest of the mass is entrained material (e.g., Offner & Chaban 2017).

As discussed above, the observed shells are consistent with wide-angle wind entrainment with multiple outburst events. Although the same episodicity is also seen in the jet, we think

these shells are unlikely to be formed by jet entrainment. In fact, multiple layer structures were identified in the extended redshifted lobe of the HH 46/47 molecular outflow observed in ^{12}CO (1–0), which were identified to be associated with several jet–bow-shock events (see Paper II). Those structures are at much lower velocities ($v_{\text{out}} < 10 \text{ km s}^{-1}$), have a different morphology, and are found at much larger distances from the source ($> 50''$) compared to the shells reported here. Thus, they are unlikely associated with the high-velocity shells discussed in this paper. In some cases, shells at the base of the outflow cavities indeed are found to be connected to the jet–bow shocks far away from the central sources (e.g., Lee et al. 2015). However, it is unclear whether the morphology and kinematics of the shells observed here (which are consistent with those expected for radially expanding parabolic shells), can be also explained by jet–bow-shock entrainment. More theoretical simulations and models are needed to test whether such shells can be formed solely by jet–bow-shock entrainment.

4.3. Implications for Evolution of Protostellar Outflow

The opening angle of protostellar outflows appears to increase with the source’s age; the outflow cavity gradually widens as the source evolves (e.g., Arce & Sargent 2006; Seale & Looney 2008; Velusamy et al. 2014; Hsieh et al. 2017). Outflows are therefore thought to be able to disperse the parent core, terminate the accretion phase, and regulate the core-to-star formation efficiency (e.g., Machida & Hosokawa 2013; Offner & Arce 2014; Offner & Chaban 2017). There are several ways that an outflow can widen as it evolves. In one scenario the outflow cavity widens as the envelope material is continuously entrained by the protostellar jet and/or wide-angle disk wind. In this model one would expect the recently accelerated material to be inside the older, previously entrained material. The newest and faster shell will soon reach the outer, slower shells and transfer momentum to them and the outflow cavity walls. This way, in general, the outflow cavity opening angle will be increasing with time. The observed multiple shell structure in HH 46/47 appears to be consistent with this picture.

In the second scenario the observed outflow is mostly composed of material that is directly launched from the disk (e.g., Machida & Hosokawa 2013). In this model, the disk slowly grows in size, and with it the launching region, at the base of the outflow, slowly widens. This in turn produces the outflow cavity that gradually becomes less collimated. In this scenario at least a part of the recently launched material is expected to be outside of the previously launched material, which is launched from the new outer regions of the disk. Such a model, however, is not consistent with the observations presented here, in which the molecular outflow is made of entrained material and the material entrained by the most recent outflow episode is inside of the older shells. However, we note that these two scenarios are not mutually exclusive.

It is also possible that the outflow cavity widens as the outflow changes direction over time (e.g., Offner et al. 2016; Lee et al. 2017b), which can be caused by a change in the angular momentum direction of the accretion flow, binary interaction, and/or jet precession. However, in the case of HH 46/47, despite the existence of a binary system at the center, the main outflow appears to be symmetric and not affected by a secondary outflow (Papers I and II). Also the precession of the jet appears to be much smaller than the opening angle of the

¹² These values are consistent with the “typical” mass-loss rate value for winds in Class I sources of $\sim 10^{-7} M_{\odot} \text{ yr}^{-1}$ (e.g., Hartigan et al. 1994; Podio et al. 2006, 2012; Mottram et al. 2017). The mass-loss rate estimate for the HH 46/47 jet of $(2\text{--}9) \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ quoted by Nisini et al. (2015) is an outlier compared to other measurements in the literature.

outflow cavity (Paper I), which also indicates that a changing outflow direction may not be the dominant cause of the widening of the outflow in HH 46/47.

4.4. Implications for Wind Launching

The spatial resolution of our current observation is too low to resolve the launching region of the wide-angle wind that entrained the observed molecular outflow. However, the highly coherent properties of the observed outflow shells can provide some constraints on the wind-launching mechanism. In the outflow entrainment scenario, in order to form such a coherent shell structure in each outburst, the wind launched from the disk toward different polar angles needs to be well coordinated. Such coordination can be naturally understood if the launching area in each outburst is a narrow region on the disk. It is very likely that the duration of each outburst is significantly shorter than the interval between outburst events, given that the observed outflow shell from each outburst is very well-defined. We therefore assume the duration of each outburst $\Delta t_{\text{outburst}}$ is $\sim 20\%$ – 30% of the outburst intervals or about 60 yr (see Section 4.1). If we use the sound speed c_s as the characteristic speed of accreting material in the disk moving inward, we obtain a length scale of $\Delta R = c_s \Delta t_{\text{outburst}} = (6 \text{ au}) \sqrt{T/40 \text{ K}}$, which we can use as a proxy for the width of the launching region on the disk. If the width of the outflow-launching region is much larger than this, it would then be hard for the wind launching at different stellocentric radii to be coordinated enough to form such coherent shells. This estimate of the size for the outflow-launching region is consistent with recent observational studies that deduce a relatively narrow range of radii for the outflow-launching regions (e.g., Bjerkeli et al. 2016; Hirota et al. 2017; Zhang et al. 2018). Note that this is different from the classical picture of a disk wind that is launched over a wide range of stellocentric radii (e.g., Blandford & Payne 1982).

Because accretion variability is believed to be caused by various instabilities in the accretion disk, it is possible that such instabilities can affect particular regions on the disk to enhance the mass or momentum of the launched wind during the outburst. We can further use the outburst interval of 200–300 yr to obtain a characteristic radius by comparing the outburst interval with the Keplerian orbital timescales. The resultant radius is 10–13 au, assuming a mass of $0.3 M_{\odot}$ for the central object (see Paper II). This can be used as a characteristic radius for the disk instability and outflow-launching zone. Note that this source contains a protobinary system with a separation of $0''.26$ (120 au) at the center, therefore the estimated characteristic radius of 10–13 au indicates that the outflow-launching region is likely on a circumstellar disk around the primary.

5. Conclusions

We present ALMA ^{12}CO (2–1) observation of the HH 46/47 molecular outflow, in which we have detected multiple wide-angle outflowing shell structures in both the blueshifted and redshifted lobes. These shells are found to be highly coherent in position–position–velocity space, extending to $\gtrsim 40\text{--}50 \text{ km s}^{-1}$ in velocity and 10^4 au in space, with well-defined morphology and kinematics. We argue that these structures are formed by the entrainment of circumstellar gas by a wide-angle wind with multiple outburst events. The intervals between consecutive outbursts are found to be $(2\text{--}3) \times 10^2 \text{ yr}$,

consistent with the timescale between outburst events in the jet powered by the same protostar. Our results provide strong evidence that wide-angle disk winds can be episodic, just like protostellar jets.

The authors thank Nami Sakai and Satoshi Yamamoto for valuable discussions. This paper makes use of the following ALMA data: ADS/JAO.ALMA#2015.1.01068.S. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada) and NSC and ASIAA (Taiwan) and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO, and NAOJ. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. Y.Z. acknowledges support from RIKEN Special Postdoctoral Researcher Program. H.G.A. acknowledges support from his NSF grant AST-1714710. D.M. and G.G. acknowledge support from CONICYT project Basal AFB-170002.

Software: CASA (McMullin et al. 2007).

ORCID iDs

Yichen Zhang  <https://orcid.org/0000-0001-7511-0034>
 Héctor G. Arce  <https://orcid.org/0000-0001-5653-7817>
 Diego Mardones  <https://orcid.org/0000-0002-5065-9175>
 Sylvie Cabrit  <https://orcid.org/0000-0002-1593-3693>
 Michael M. Dunham  <https://orcid.org/0000-0003-0749-9505>
 Stella S. R. Offner  <https://orcid.org/0000-0003-1252-9916>
 Alejandro C. Raga  <https://orcid.org/0000-0002-0835-1126>

References

- Antoniucci, S., Nisini, B., Giannini, T., & Lorenzetti, D. 2008, *A&A*, **479**, 503
- Arce, H. G., Borkin, M. A., Goodman, A. A., Pineda, J. E., & Halle, M. W. 2010, *ApJ*, **715**, 1170
- Arce, H. G., Mardones, D., Corder, S. A., et al. 2013, *ApJ*, **774**, 39
- Arce, H. G., & Sargent, A. I. 2006, *ApJ*, **646**, 1070
- Bacciotti, F., Mundt, R., Ray, T. P., et al. 2000, *ApJ*, **537**, 49
- Bai, X.-N., Ye, J., Goodman, J., et al. 2016, *ApJ*, **818**, 152
- Bjerkeli, P., van der Wiel, M. H. D., Harsono, D., et al. 2016, *Natur*, **540**, 406
- Blandford, R. D., & Payne, D. G. 1982, *MNRAS*, **199**, 883
- Dunham, M. M., Arce, H. G., Mardones, D., et al. 2014, *ApJ*, **783**, 29
- Dunham, M. M., & Vorobyov, E. I. 2012, *ApJ*, **747**, 52
- Eisloffel, J., & Mundt, R. 1994, *A&A*, **284**, 530
- García López, R., Nisini, B., Eisloffel, J., et al. 2010, *A&A*, **511**, A5
- Hartigan, P., Heathcote, S., Morse, J. A., Reipurth, B., & Bally, J. 2005, *AJ*, **130**, 2197
- Hartigan, P., Morse, J. A., & Raymond, J. 1994, *ApJ*, **436**, 125
- Hirota, T., Machida, M. N., Matsushita, Y., et al. 2017, *NatAs*, **1**, 0146
- Hsieh, T.-H., Lai, S.-P., & Belloche, A. 2017, *ApJ*, **153**, 173
- Königl, A., & Pudritz, R. E. 2000, in *Protostars and Planets IV*, ed. V. Mannings et al. (Tucson, AZ: Univ. Arizona Press), 759
- Kwan, J., & Takemaru, E. 1995, *ApJ*, **454**, 382
- Lee, C.-F., Hirano, N., Zhang, Q., et al. 2015, *ApJL*, **805**, 186L
- Lee, C.-F., Ho, P. T. P., Li, Z.-Y., et al. 2017a, *NatAs*, **1**, 0152
- Lee, C.-F., Ho, P. T. P., Palau, A., et al. 2007, *ApJ*, **670**, 1188
- Lee, C.-F., Mundy, L. G., Reipurth, B., Ostriker, E. C., & Stone, J. M. 2000, *ApJ*, **542**, 925
- Lee, J. W. Y., Hull, C. L. H., & Offner, S. S. R. 2017b, *ApJ*, **834**, 201
- Li, Z.-Y., & Shu, F. H. 1996, *ApJ*, **472**, 211
- Louvet, F., Dougados, C., Cabrit, S., et al. 2018, *A&A*, **618**, 120
- Machida, M. N., & Hosokawa, T. 2013, *MNRAS*, **431**, 1719
- Matt, S., Winglee, R., & Böhm, K.-H. 2003, *MNRAS*, **345**, 660
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, in *ASP Conf. Ser. 376, Astronomical Data Analysis Software and Systems XVI*, ed. R. A. Shaw, F. Hill, & D. J. Bell (San Francisco, CA: ASP), 127

- Micono, M., Davis, C. J., Ray, T. P., Eisloffel, J., & Shetrone, M. D. 1998, [ApJL](#), **494**, L227
- Morse, J. A., Hartigan, P., Heathcote, S., et al. 1994, [ApJ](#), **425**, 738
- Mottram, J. C., van Dishoeck, E. F., Kristensen, L. E., et al. 2017, [A&A](#), **600**, 99
- Nisini, B., Santangelo, G., Giannini, T., et al. 2015, [ApJ](#), **801**, 121
- Noriega-Crespo, A., Morris, P., Marleau, F. R., et al. 2004, [ApJS](#), **154**, 352
- Offner, S. S. R., & Arce, H. G. 2014, [ApJ](#), **784**, 61
- Offner, S. S. R., & Chaban, J. 2017, [ApJ](#), **847**, 104
- Offner, S. S. R., Dunham, M. M., Lee, K. I., et al. 2016, [ApJL](#), **827**, L11
- Plunkett, A. L., Arce, H. G., Corder, S. A., et al. 2013, [ApJ](#), **774**, 22
- Plunkett, A. L., Arce, H. G., Mardones, D., et al. 2015, [Natur](#), **527**, 70
- Podio, L., Bacciotti, F., Nisini, B., et al. 2006, [A&A](#), **456**, 189
- Podio, L., Kamp, I., Flower, D., et al. 2012, [A&A](#), **545**, 44
- Raga, A., & Cabrit, S. 1993, [A&A](#), **278**, 267
- Reipurth, B. 2000, [AJ](#), **120**, 1449
- Reipurth, B., & Bally, J. 2001, [ARAA](#), **39**, 403
- Reipurth, B., & Heathcote, S. 1991, [A&A](#), **246**, 511
- Schwartz, R. D. 1977, [ApJL](#), **212**, L25
- Seale, J. P., & Looney, L. W. 2008, [ApJ](#), **675**, 427
- Shang, H., Shu, F. H., & Glassgold, A. E. 1998, [ApJL](#), **493**, L91
- Shu, F. H., Najita, J. R., Shang, H., & Li, Z.-Y. 2000, in *Protostars and Planets IV*, ed. V. Mannings et al. (Tucson, AZ: Univ. Arizona Press), 789
- Tabone, B., Cabrit, S., Bianchi, E., et al. 2017, [A&A](#), **607**, L6
- Tafalla, M., Santiago-García, J., Hacar, A., & Bachiller, R. 2010, [A&A](#), **522**, 91
- Tafalla, M., Su, Y.-N., Shang, H., et al. 2016, [A&A](#), **597**, 119
- van Kempen, T. A., van Dishoeck, E. F., Güsten, R., et al. 2009, [A&A](#), **501**, 633
- Velusamy, T., Langer, W. D., & Marsh, K. A. 2007, [ApJL](#), **668**, L159
- Velusamy, T., Langer, W. D., & Thompson, T. 2014, [ApJ](#), **783**, 6
- Vorobyov, E. I., & Basu, S. 2015, [ApJ](#), **805**, 115
- Zhang, Y., Arce, H. G., Mardones, D., et al. 2016, [ApJ](#), **832**, 158
- Zhang, Y., Higuchi, A., Sakai, N., et al. 2018, [ApJ](#), **864**, 76
- Zhu, Z., Hartmann, L., Gammie, C. F., et al. 2010, [ApJ](#), **713**, 1134