

Deep late-time observations of the supernova impostors SN 1954J and SN 1961V

Rachel A. Patton,¹★ C. S. Kochanek^{1,2} and S. M. Adams³

¹Department of Astronomy, The Ohio State University, 140 West 18th Ave, Columbus OH 43210, USA

²Centre for Cosmology and AstroParticle Physics, The Ohio State University, 191 West Woodruff Avenue, Columbus OH 43210, USA

³Cahill Centre for Astrophysics, California Institute of Technology, Pasadena CA 91125, USA

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ABSTRACT

SN 1954J in NGC 2403 and SN 1961V in NGC 1058 were two luminous transients whose definitive classification as either non-terminal eruptions or supernovae remains elusive. A critical question is whether a surviving star can be significantly obscured by dust formed from material ejected during the transient. We use three lines of argument to show that the candidate surviving stars are not significantly optically extinct ($\tau \lesssim 1$) by dust formed in the transients. First, we use SED fits to new *HST* optical and near-IR photometry. Secondly, neither source is becoming brighter as required by absorption from an expanding shell of ejected material. Thirdly, the ejecta masses implied by the H α luminosities are too low to produce significant dust absorption. The latter two arguments hold independent of the dust properties. The H α fluxes should also be declining with time as t^{-3} , and this seems not to be observed. As a result, it seems unlikely that recently formed dust can be responsible for the present faintness of the sources compared to their progenitors, although this can be verified with the *James Webb Space Telescope*. This leaves three possibilities: (1) the survivors were misidentified; (2) they are intrinsically less luminous; (3) SN 1954J and SN 1961V were true supernovae.

Key words: supernovae: general – supernovae: individual: SN 1954J, SN 1961V.

1 INTRODUCTION

There are a subset of massive stars that appear to experience violent eruptions, some of which are spectroscopically similar to Type IIn supernovae (SNe; Schlegel 1990; Filippenko 1997) due to their moderate line widths ($v_{\text{ej}} \lesssim 2000 \text{ km s}^{-1}$). The brightest of these eruptions are not easily distinguishable from the faintest SNe, leading to the term SN ‘impostor’ (Van Dyk, Filippenko & Li 2002) and the potential for misclassification. The only certain difference between the most luminous eruptions and the least luminous SNe is that these eruptions are non-terminal. These eruptions remain poorly understood and so accurate classification of these events is critical to understanding outburst mechanisms and rates.

Two of the most famous examples are SN 1954J in NGC 2403 and SN 1961V in NGC 1058. SN 1954J was originally classified as an SN but was later identified as the luminous blue variable (LBV) V12 and reclassified as an eruption (Tammann & Sandage 1968). The progenitor remained fairly quiescent at $M_B \approx -6.6$ mag until it became highly variable a few years prior to its eruption, sometimes fluctuating by as much as two magnitudes over a few days. The peak of the transient was not observed due to its proximity to the

Sun. When next observed, its magnitude had increased from $M_B \approx -8.5$ mag to $M_B \approx -11.3$ mag before settling at $M_B \approx -5.6$ mag, a full magnitude fainter than the progenitor (Tammann & Sandage 1968). The survivor never regained its pre-eruption luminosity and has not significantly varied since the eruption (Smith, Humphreys & Gehrz 2001; Van Dyk et al. 2005; Kochanek, Szczygiel & Stanek 2012; Humphreys et al. 2017). Van Dyk et al. (2005) resolved the region around SN 1954J into four stars (see Fig. 1) and identified star 4 as the most likely survivor candidate due to its strong H α emission. Follow-up spectroscopy by Humphreys et al. (2017) shows no significant change in the H α emission between 2014 and 2017.

SN 1961V was originally classified as a ‘Type V’ supernovae (Zwicky 1964), and its true nature remains disputed. The progenitor of SN 1961V was one of the brightest stars in NGC 1058. Pre-transient, it had $M_B \approx -12$ mag until brightening to $M_B \approx -14$ mag around a year prior to the peak. In the following months it brightened to $M_B \approx -16$ mag before reaching a peak at $M_B \approx -18$ mag in 1961 December. The star’s brightness decreased over a few months before briefly plateauing at $M_B \approx -13$ mag and then fading to $M_B \approx -11.5$ mag where it remained for 4 yr until dropping below the point of visibility in 1968 (see the various summaries in Goodrich et al. 1989; Humphreys & Davidson 1994; Humphreys et al. 1999; Kochanek et al. 2011; Smith et al. 2011).

* E-mail: patton.502@osu.edu

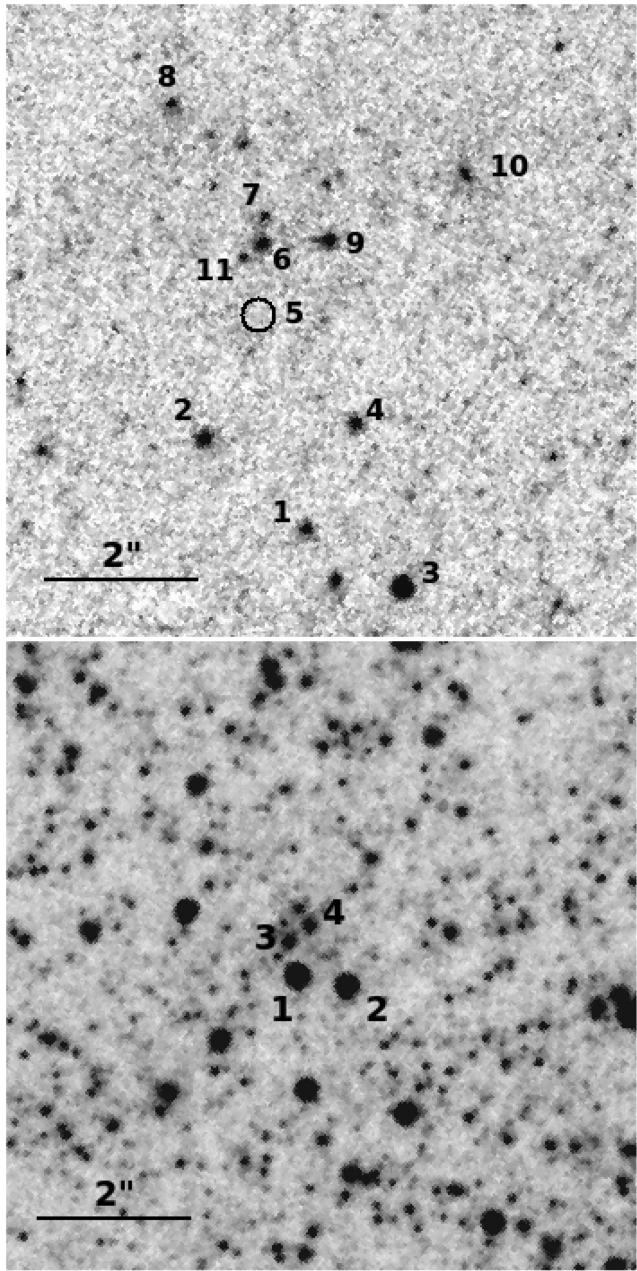


Figure 1. *HST* WFC3 *F814W* images of the regions around SN 1961V (top) and SN 1954J (bottom). Stars 7 and 4 are believed to be associated with SN 1961V and SN 1954J, respectively. The stars are labelled following Filippenko et al. (1995) for SN 1961V and Van Dyk et al. (2005) for SN 1954J. Star 5 from Filippenko et al. (1995), whose position is indicated by the labelled circle, is not visible in any of our data.

The light curve of SN 1961V is peculiar and not typical of either an LBV outburst or a supernova. Goodrich et al. (1989) argued that it was best explained by a non-terminal eruption: first an S Doradus phase, where the luminosity is roughly constant but a drop in the stellar temperature makes the source optically brighter in the years prior to the peak, followed by a major LBV outburst in 1961. Many have tried to identify a survivor at the location of SN 1961V calculated by Klemola (1986). In Fig. 1, star 6 (Filippenko et al. 1995), star 11 (Van Dyk et al. 2002), and star 7 (Chu et al. 2004) have all been identified as potential survivors or companions to the

progenitor based on their red colours and $V > 24$ mag. Star 7 is presently the preferred candidate due to its $\text{H}\alpha$ emission.

However, there is strong evidence that SN 1961V was an actual, albeit odd, supernova. Branch & Cowan (1985), Cowan, Henry & Branch (1988), Stockdale et al. (2001), and Chu et al. (2004) all detected a fading non-thermal radio source, more like a supernova than an eruption, near the location of star 7. A reanalysis of the radio data by Van Dyk et al. (2005) found a 2σ discrepancy between the centre of the radio emission and the location of SN 1961V, which they use to argue that the radio emission was not associated with SN 1961V. Smith et al. (2011) argue that it was an SN based on the transient's energetics. The post-transient spectra showed strong hydrogen emission with velocity widths of about 2000 km s^{-1} (Zwicky 1964; Branch & Greenstein 1971) consistent with a Type IIn SN or an eruption. Assuming the ejecta of a non-terminal eruption coalesced into a dusty shell, we would expect to see IR emission from the dust re-radiating the star's light. Kochanek, Szczygiel & Stanek (2011) analysed archival *Spitzer* images of SN 1961V and found no evidence of IR emission. However, Van Dyk & Matheson (2012) argue for obscuration by foreground dust which contributes no IR emission. Kochanek et al. (2012) pointed out that this solution makes the energetics argument for an SN by Smith et al. (2011) even more compelling.

Here we examine both sources with new optical and near-IR *HST* observations along with continued monitoring of SN 1954J with the Large Binocular Telescope (LBT). Our goal is to determine if these two sources are less optically luminous than their progenitors due to obscuration by dust formed in the transients. We present our data and models in Section 2, a discussion of each event in Sections 3 and 4, and a summary in Section 5.

2 DATA AND MODELS

Images of SN1954J and SN1961V were taken in 2013 October and December using WFC3 on *HST* under program 13477 (PI: Kochanek). Both SN 1954J and SN 1961V were imaged a total of 12 times in four filters. For SN 1954J, the exposures were 3×430 s in *F475W*, 3×430 s in *F814W*, 2×699 s + 2×49 s in *F110W*, and 2×799 s in *F160W*. For SN 1961V, the exposures were 3×381 s in *F475W*, 3×381 s in *F814W*, 2×499 s + 2×99 s in *F110W*, and 2×799 s in *F160W*. All of the individual exposures were dithered to control for hot pixels and cosmic rays. The pixel scales of the WFC3/UVIS and IR detectors are $0''.04/\text{pix}$ and $0''.13/\text{pix}$, respectively. Fig. 1 shows the *F814W* images of 8 arcsec regions around SN 1954J and SN 1961V. The stars are labelled according to Filippenko et al. (1995) for SN 1961V and Van Dyk et al. (2005) for SN 1954J. Star 5 from Filippenko et al. (1995) near SN 1961V is not visible in any of our filters.

We processed the data using DOLPHOT (Dolphin 2016), separately optimizing the parameters for the UVIS and IR data following the procedures laid out in the manual.¹ DOLPHOT simultaneously fits point sources to each image using the appropriate PSF for each filter, providing magnitudes and parameters, such as sharpness and crowding. All magnitudes are on the Vega scale and presented in Table 1. Our magnitudes for SN 1954J are consistent with those in Humphreys et al. (2017), who analysed the same data.

We adopt a distance of 3.3 Mpc to NGC 2403 (Freedman & Madore 1988) and Galactic reddening of $E(B - V) = 0.04$ mag (Schlegel, Finkbeiner & Davis 1998), and correct for both prior

¹<http://americano.dolphinsim.com/dolphot/>

Table 1. *HST* photometry of stars near SN 1954J and SN 1961V.

Event	Object	<i>F475W</i>	err	<i>F814W</i>	err	<i>F110W</i>	err	<i>F160W</i>	err
SN 1954J	Star 1	24.157	0.021	20.714	0.006	19.573	0.002	18.618	0.002
	Star 2	24.101	0.020	20.820	0.006	19.710	0.002	18.729	0.002
	Star 3	23.190	0.012	23.073	0.021	23.109	0.018	22.940	0.034
	Star 4 ^a	24.084	0.022	22.630	0.017	22.023	0.008	21.591	0.012
SN 1961V	Star 6	26.399	0.089	23.752	0.033	22.695	0.012	21.823	0.012
	Star 7 ^a	25.808	0.058	25.198	0.103	24.418	0.045	23.831	0.056
	Star 11	27.946	0.302	25.218	0.089	24.174	0.035	23.310	0.036

^aThe most likely eruption survivor.

to any modelling. We adopt a distance of 10 Mpc to NGC 1058 (Boroson 1981) and Galactic reddening of $E(B - V) = 0.06$ mag (Schlegel et al. 1998) and again make the requisite reddening corrections before modelling. We assume a minimum error of 10 per cent on the photometry. We model the stars' spectral energy distributions (SEDs) in two ways: a simple model with additional foreground extinction by an $R_V = 3.1$; Cardelli, Clayton & Mathis (1989) extinction law and with DUSTY (Ivezic & Elitzur 1997; Ivezic, Nenkova & Elitzur 1999; Elitzur & Ivezic 2001). In both cases we model the stars using Castelli & Kurucz (2003) model atmospheres convolved with the appropriate filter functions. The parameters and their uncertainties are determined using Markov Chain Monte Carlo (MCMC) methods. We do not include the *Spitzer* mid-IR data analysed in Kochanek et al. (2011, 2012) in our SEDs because no IR emission was observed at the site of SN 1961V and only upper limits were found at the site of SN 1954J.

The simple model minimizes the fit statistic:

$$\chi^2 = \sum_i \left(\log L_i - \log L_i^{\text{mod}}(L_*, T_*) + 0.4R_\lambda E(B - V) \right)^2 / \sigma_i^2, \quad (1)$$

where L_i and L_i^{mod} are the observed and model band luminosities (λL_λ), to estimate the total luminosity L_* , temperature T_* , and additional foreground reddening $E(B - V)$ beyond the Galactic contribution. The χ^2 is calculated and minimized in log space. DUSTY solves for radiation transfer from the star through a dusty shell. We embed DUSTY in an MCMC driver (Adams & Kochanek 2015) that allows for a dusty circumstellar shell plus a variable amount of foreground extinction. This model has the V-band optical depth τ_V and dust temperature T_d as additional parameters. We again optimize the same fit statistic (equation 1) but in this model, the band luminosities, $L_i^{\text{mod}}(L_*, T_*, \tau_V, T_d)$ have more parameters. We fix the dust temperature T_d at 50 K for all our models since the optical and near-IR bands we use have no contribution from the dust emission. For all DUSTY models we treat the dust as purely silicate as this is the type of dust expected from massive stars (Speck et al. 2000). Changes in dust composition have only modest effects on our results (see Kochanek et al. 2012, for further discussion on the effects of dust composition).

Circumstellar absorption due to dust formed in the transient requires that the dust optical depth is time variable (see Kochanek et al. 2012). The dust forms as the ejecta cools to form a dusty circumstellar shell. As time passes, the optical depth drops as $\tau = \tau_0(t_0/t)^2$ due to its geometric expansion. If the shell fragments or becomes clumpy, the drop in optical depth will accelerate. A clump directly in our line of sight does not prevent a steady drop in the effective optical depth because most of the escaping emission is scattered light from a broad region across the shell. As the optical

depth drops, we should see the surviving star brighten with time, as is the case with η Carinae (see Humphreys & Davidson 1994). That η Carinae had an extended period of constant optical flux means that it must have also been forming a dusty, optically thick wind in that period (Kochanek et al. 2012). A constant optical depth requires a steady dust forming wind and emission by hot dust, which is not seen for either SN 1954J or SN 1961V (Kochanek et al. 2011, 2012).

Consider a source with intrinsic band luminosity L inside a dusty shell of optical depth τ_0 at time t_0 . The observed band luminosity is then

$$L_{\text{obs}}(t) = L e^{-\tau_0(t_0/t)^2}. \quad (2)$$

If we detect the source at two epochs, t_0 and t_1 , we can determine the optical depth without reference to the object's SED or the type of dust since

$$\frac{L_{\text{obs}}(t_1) - L_{\text{obs}}(t_0)}{L_{\text{obs}}(t_0)} = e^{\tau_0(1 - (t_0/t_1)^2)} - 1. \quad (3)$$

Given the optical depth τ_0 , we then know the true luminosity. Detection of a change in luminosity determines τ_0 , and limits on changes in luminosity set upper bounds on τ_0 .

Crowding, particularly in ground-based observations, means it is only possible to measure changes in luminosity. In particular, difference imaging eliminates crowding, but all you measure is $\Delta L(t_1, t_0) = L_{\text{obs}}(t_1) - L_{\text{obs}}(t_0)$. Given a long enough light curve, one can still determine both L and τ_0 , but our observations have too short a time baseline to do so. If a change in luminosity is observed, then there is a joint constraint on L and τ_0 . An upper limit on the change in luminosity ΔL_{max} , provides an upper limit on the luminosity of the source,

$$L < \frac{\Delta L_{\text{max}}}{e^{-\tau_0(t_0/t_1)^2} - e^{-\tau_0}}, \quad (4)$$

as a function of the optical depth.

For SN 1954J, we have *HST* observations separated by roughly 10 yr, with $t_0 \approx 50$ and $t_1 \approx 60$, corresponding to a fractional change in optical depth of $t_0^2/t_1^2 \approx 0.7$. We also have *R*-band LBT data covering a similar baseline with $t_0 \approx 54$ and $t_1 \approx 64$ so again $t_0^2/t_1^2 \approx 0.7$. For SN 1961V we have *HST* observations separated by roughly 20 yr, with $t_0 \approx 40$ and $t_1 \approx 60$, so the fractional change in the optical depth, $t_0^2/t_1^2 \approx 0.4$, is large. This will make it difficult to invoke a significant optical depth without also requiring an intrinsically low-luminosity source.

Finally, both star 4 near SN 1954J and star 7 near SN 1961V have strong, broad H α emission, which Van Dyk & Matheson (2012) and Humphreys et al. (2017) use to argue that both transients had survivors. For a fully ionized, constant density, thin hydrogen shell

of mass M and radius $R = vt$, the H α recombination luminosity is

$$L_{H\alpha} = \frac{M^2 \alpha_{H\alpha} E_{H\alpha}}{4\pi \Delta R^3 m_p^2}, \quad (5)$$

where $\alpha_{H\alpha} \simeq 10^{-13} \text{ cm}^3 \text{ s}^{-1}$, $E_{H\alpha} = 1.89 \text{ eV}$, m_p is the proton mass, Δ is the fractional shell thickness, v is the expansion velocity, and t is the expansion time. This assumes that the outbursts were single events with no continuous outflows. For SN 1961V, this assumption follows a previous analysis by Van Dyk & Matheson (2012). For SN 1954J, we consider the possibility of dense stellar wind below. Note that the H α luminosity is not constant, but decays as $L_{H\alpha} \propto R^{-3} \propto t^{-3}$ with time. If the H α emission comes from photoionizing the ejecta, then the luminosity provides an estimate of the mass with

$$M \simeq 0.1 v_3^{3/2} t_{50}^{3/2} L_{36}^{1/2} \Delta_{0.1}^{1/2} \text{ M}_\odot, \quad (6)$$

where $t = 50 t_{50}$ yr, $v = 1000 v_3 \text{ km s}^{-1}$, $L_{H\alpha} = 10^{36} L_{36} \text{ erg s}^{-1}$, and $\Delta = 0.1 \Delta_{0.1}$.

Any dust formed in the ejecta is mixed with the hydrogen, so the dust optical depth of the ejecta is

$$\tau = \frac{M\kappa}{4\pi R^2}, \quad (7)$$

where κ is the dust opacity. We can solve for the total H α luminosity in terms of the optical depth τ and ejection velocity v to find that

$$L_{H\alpha} = \frac{4\pi v t \tau^2 \alpha_{H\alpha} E_{H\alpha}}{\Delta \kappa^2 m_p^2}. \quad (8)$$

Note that this scaling only depends on dust properties like size distribution and composition only through the opacity. This does not account for dust absorption in the shell. While there is no analytic expression for a shell, the observed luminosity from a cube of side $2R$ is smaller by $(1-e^{-2\tau})/2\tau$, which goes to $1/2\tau$ in the limit of large optical depth. The τ^{-1} scaling is generic because you only see radiation from the regions where the optical depth to the observer is $\tau \lesssim 1$, which is a fraction τ^{-1} of the overall volume. Thus, in the high optical depth limit

$$L_{H\alpha} \simeq \frac{4\pi v t \tau \alpha_{H\alpha} E_{H\alpha}}{\Delta \kappa^2 m_p^2}. \quad (9)$$

Solving for τ we find,

$$\tau = 0.07 \kappa_2 t_{50}^{-1/2} v_3^{-1/2} L_{36}^{1/2} \Delta_{0.1}^{1/2}, \quad (10)$$

for $\tau \leq 1$, and

$$\tau \simeq 0.00 \kappa_2^2 t_{50}^{-1} v_3^{-1} L_{36} \Delta_{0.1}, \quad (11)$$

for $\tau \geq 1$, where $\kappa = 100 \kappa_2 \text{ cm}^2 \text{ g}^{-1}$. If the material is in a uniform sphere, then the optical depth is larger by $(3/\Delta)^{1/2}$ or $(3/\Delta)$ for the low and high τ cases, respectively. We also note that the dust destruction time in such a photoionized nebula is $\gg 10^2$ yr (see e.g. Draine 1995). Dust formed in the ejecta will continue to be present if the ejecta is photoionized at a later time.

Finally, we consider a scenario put forth by Humphreys et al. (2017), where the H α emission comes from a dense stellar wind of a binary companion. The H α luminosity of the wind is

$$L_{H\alpha} = \frac{\dot{M}^2 \alpha E_{H\alpha}}{4\pi v_w^2 \mu^2 m_p^2 R_*}. \quad (12)$$

where \dot{M} is the mass loss rate, v_w is the wind speed, and R_* is the stellar radius. The Thomson optical depth of the wind is

$$\tau_T = \frac{\dot{M} \kappa_T}{4\pi R_* v}, \quad (13)$$

where the Thomson opacity is $\kappa_T \simeq 0.5 \text{ cm}^2 \text{ g}^{-1}$. Humphreys et al. (2017) want $\tau_T \simeq 1$ to explain the shape of the line profile of SN 1954J. The H α luminosity is then

$$L_{H\alpha} = \frac{4\pi R_* \alpha E_{H\alpha} \tau_T^2}{\mu^2 m_p^2 \kappa_T^2}, \quad (14)$$

in terms of its Thomson optical depth. Finally, we can express the stellar radius $R_*^2 = L_*/4\pi\sigma T_*^4$ in terms of the stellar luminosity L_* and effective temperature T_* , to find that

$$L_{H\alpha} = 84 L_{*2}^{1/2} T_{*2}^{-2} \tau_T^2 \text{ L}_\odot \quad (15)$$

for $L_* = 100 L_{*2} \text{ L}_\odot$, $T_* = 20000 T_{*2} \text{ K}$, $\kappa = 0.5 \text{ cm}^2 \text{ g}^{-1}$, $\alpha = 10^{-13} \text{ cm}^3 \text{ s}^{-1}$, and $\mu = 1$.

Adding dusts obscuration has an interesting consequence due to the $L_{H\alpha} \propto L_{*2}^{1/2}$ scaling of equation (15). Let $L_{H\alpha 0}$ and L_{*0} be the H α luminosity inferred from the line flux and the stellar luminosity inferred from equation (15) assuming there is no dust. Suppose we add a dust optical depth τ_α at the wavelength of H α . Then the true line luminosity is $L_{H\alpha} = L_{H\alpha 0} \exp(-\tau_\alpha)$ and we must have $L_* = L_{*0} \exp(-2\tau_\alpha)$. This means that the expected flux of the star, $F_* \propto L_* \exp(-\tau_\alpha) \propto L_{*0} \exp(-\tau_\alpha)$ actually increases rather than decreases as you try to obscure the star behind dust because of how the H α luminosity depends on L_* when you hold the Thomson optical depth of the wind constant. Note that if we consider the flux of the star near H α this is completely independent of dust properties.

3 SN 1954J IN NGC 2403

All four stars in the environment of SN 1954J are well fit as single stars. Table 2 lists the best-fitting parameters for each star and Fig. 2 shows each star's SED after being corrected for the model foreground extinction. We assumed the same foreground extinction for the archival and current band luminosities. Humphreys et al. (2017) argue that an additional A_V of 0.8–0.9 beyond Galactic is necessary to properly fit each star, which is consistent with our findings. Stars 1 and 2 appear to be red giants and star 3 is an A or B star.

The candidate counterpart to SN 1954J, star 4, has an intermediate temperature. Humphreys et al. (2017) note that their blackbody fit gives $T \approx 6600 \text{ K}$, while their spectroscopic estimate is 5000 K. Our fits based on model atmospheres have uncertainties large enough to be consistent with either estimate ($T_* \approx 6600_{-1500}^{+3300} \text{ K}$). If we constrain its temperature to lie within 10 per cent of their spectroscopic temperature by including a penalty of $(T_* - 5000)^2/\sigma_T^2$ with $\sigma_T = 500 \text{ K}$ in the fit statistic, the difference in the fit statistic between the two models ($\Delta\chi^2 = 0.42$) shows that the SED is consistent with the spectroscopic temperature (see Table 2).

We again modelled star 4 with DUSTY, this time including the temperature prior and allowing the circumstellar optical depth (τ_v) to vary. If we allow both the foreground and circumstellar extinction to vary, we, not surprisingly, find degenerate solutions. If we fix the foreground extinction to $E(B-V) = 0.25$ to match the neighbouring stars, we find $\tau_v = 0.03_{-0.03}^{+0.73}$, consistent with the results from Kochanek et al. (2012), while if we fix $E(B-V) = 0$, we find $\tau_v = 1.26_{-1.08}^{+0.75}$.

Humphreys et al. (2017) ultimately propose that the system consists of two stars, a ‘cold’ star with $T_C \simeq 5000 \text{ K}$ and $\log L_* \simeq 4.6$ and a ‘hot’ star with $T_H \simeq 20000 \text{ K}$ and $\log L_* \simeq 5.3$. To explore this model further, we constrained the temperature of the cooler star by our temperature prior of $T_C = 5000 \pm 500 \text{ K}$ and consider fixed temperatures for the hot star of $T_H = 10000, 15000, 20000$, and

Table 2. Best-fitting DUSTY models.

Event	Object	T_{eff} (K)	$\log L_*$	$E(B-V)$ (mag)	χ^2/dof	Priors
SN 1954J	Star 1	3730^{+3230}_{-190}	$4.63^{+0.89}_{-0.10}$	$0.25^{+1.20}_{-0.19}$	0.077	No T_* prior
	Star 2	3810^{+12760}_{-240}	$4.59^{+2.03}_{-0.11}$	$0.26^{+1.60}_{-0.20}$	0.049	
	Star 3	12470^{+1480}_{-3280}	$4.18^{+0.98}_{-0.39}$	$0.14^{+0.19}_{-0.11}$	0.392	T_* prior
	Star 4 ^a	6590^{+3290}_{-1540}	$4.01^{+0.62}_{-0.34}$	$0.39^{+0.41}_{-0.35}$	0.047	
	Star 4 ^a	5230^{+290}_{-350}	$3.71^{+0.06}_{-0.07}$	$0.09^{+0.07}_{-0.08}$	0.468	
SN 1961V	Star 6	4740^{+760}_{-880}	$4.58^{+0.19}_{-0.31}$	$0.56^{+0.23}_{-0.52}$	0.000	All bands
	Star 7 ^a	5860^{+10570}_{-180}	$3.78^{+1.40}_{-0.00}$	$0.04^{+0.70}_{-0.01}$	8.228	
	Star 7 ^a	7510^{+830}_{-770}	$3.78^{+0.06}_{-0.06}$	$0.06^{+0.03}_{-0.05}$	0.00	
	Star 11	4830^{+1010}_{-950}	$4.01^{+0.25}_{-0.30}$	$0.59^{+0.32}_{-0.50}$	0.046	Optical bands only

^aThe most likely eruption survivor.

Errors reflect the 90 per cent confidence interval. The extinction represents additional absorption beyond the Galactic contribution.

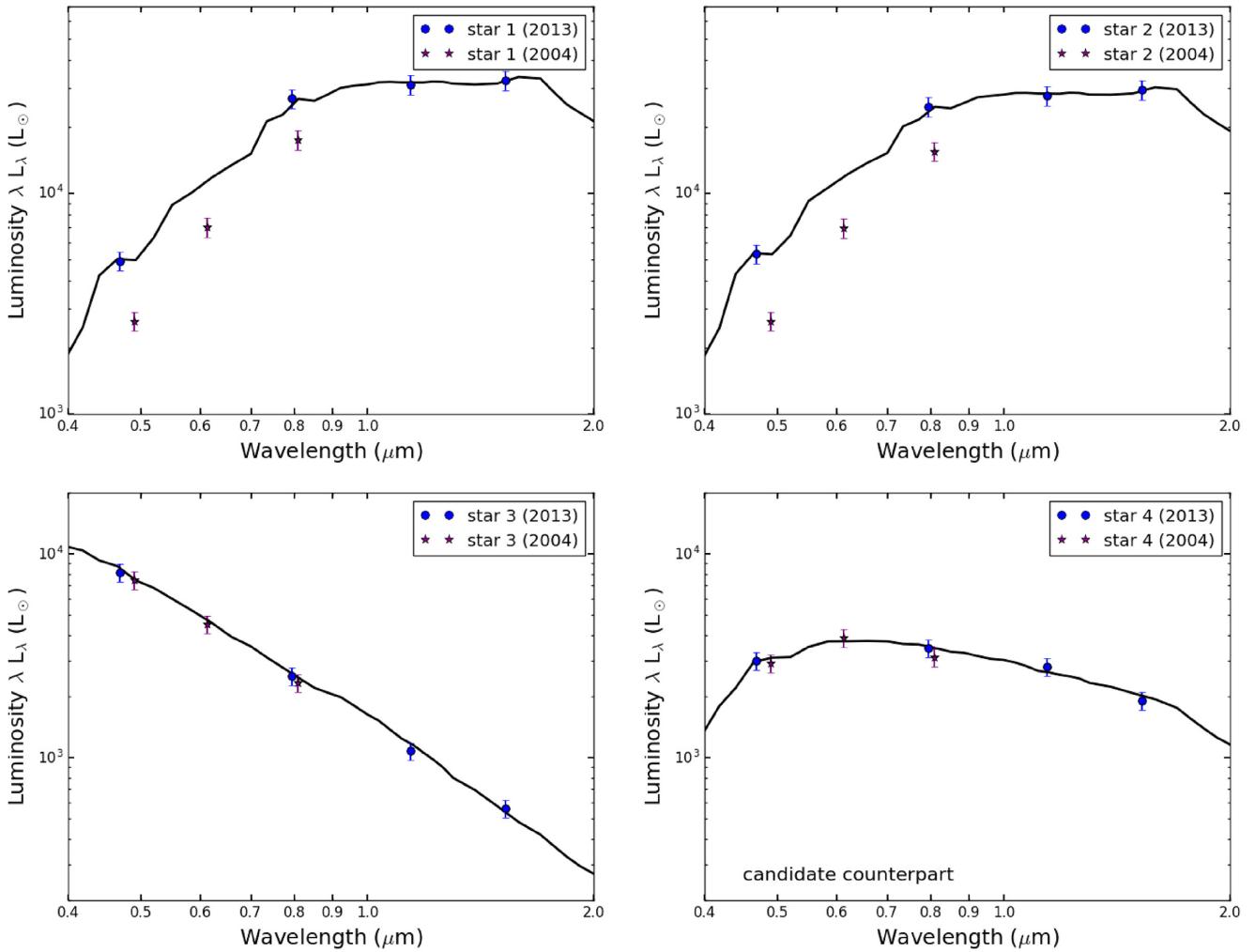


Figure 2. The SEDs of stars 1, 2, 3, and 4 near SN 1954J for our current data (circles) and the archival ACS/WFC $F475W$, $F606W$, and $F814W$ photometry (stars) from Humphreys et al. (2017). The model SED for star 4 includes the temperature prior $T_* = 5000 \pm 500$ K. The $F475W$ and $F814W$ points are offset by ± 0.01 μm from the mean filter wavelength so that they do not overlap. These models assume the foreground extinction for the best-fitting single star models (i.e. Table 2).

Table 3. 2-star models of star 4 near SN 1954J.

Model type	T_H (K)	T_C (K)	$\log L_H$	$\log L_C$	$E(B-V)$ (mag)	χ^2/dof
2 stars fixed T_H and T_C	10 000	5180^{+470}_{-450}	$0.30^{+2.76}_{-0.16}$	$3.70^{+0.11}_{-0.10}$	$0.08^{+0.15}_{-0.08}$	0.233
2 stars fixed T_H and T_C	15 000	5190^{+480}_{-460}	$0.40^{+3.22}_{-0.21}$	$3.70^{+0.12}_{-0.09}$	$0.08^{+0.18}_{-0.08}$	0.233
2 stars fixed T_H and T_C	20 000	5200^{+440}_{-460}	$1.00^{+2.57}_{-0.82}$	$3.70^{+0.11}_{-0.10}$	$0.08^{+0.13}_{-0.08}$	0.233
2 stars fixed T_H and T_C	25 000	5180^{+470}_{-430}	$1.06^{+2.76}_{-0.90}$	$3.70^{+0.11}_{-0.09}$	$0.08^{+0.15}_{-0.07}$	0.233

Note. Here H refers to the hot star and C to the cooler star. Errors reflect the 90 per cent confidence interval.

25 000 K. The luminosities of the two stars were free to vary, but with a conservative penalty in the fit statistic of $(L_H/L_C)^2$ for the $F475W$ and $F814W$ band luminosities, if $L_H > L_C$. We know from the Humphreys et al. (2017) spectrum that the observed optical flux has to be dominated by the cooler star, so we (conservatively) should not allow the hot star to dominate the optical emission. Both stars are subject to the same extinction.

The results for these models are shown in Table 3. One can never rule out the presence of a hot star, but all of our models require the hot star to have a negligible luminosity which cannot produce sufficient ionizing photons to explain the line emission. If we try to put in a hot star with a luminosity of $L_H = 10^{5.3} L_\odot$ and $T_H = 20 000$ K, as proposed by Humphreys et al. (2017), we find unacceptable fits with $\chi^2 = 4.4$.

While we agree with Humphreys et al. (2017) on the amount of dust, either foreground or circumstellar, needed to reproduce the stellar colours, Humphreys et al. (2017) then add $A_V = 2.5$ of grey dust to allow star 4 to be intrinsically more luminous. They argue that the grains formed in the eruption might be large enough to be effectively grey. Such dust cannot be identified in the SED fits, but it is constrained by the time variability of the source and the $H\alpha$ luminosity.

As discussed in Section 2, we can use variability to constrain the optical depth of any dust formed in the transient. For SN 1954J, we can do this in two ways. First, we have archival R -band images from the Large Binocular Telescope from 2008 March through 2018 May (Gerke, Kochanek & Stanek 2015; Adams et al. 2017). We see no variability at its location. To set a limit on the variability, we selected a nine-point grid with 2 arcsec offsets from the location of SN 1954J and measured the brightness in each epoch at each point over the 9 yr time span. The average slope at the site of SN 1954J was 30 ± 90 counts yr^{-1} . Taking our 1σ upper limit as the estimate, the average slope was 120 counts yr^{-1} , which at $0.36 L_\odot$ per count, corresponds to approximately $43 L_\odot \text{ yr}^{-1}$, or an upper limit of $\Delta L = 430 L_\odot$ over the 10 yr baseline.

We know from the HST data and the SED fits to star 4 that the observed R band luminosity in 2013.9 was $L_R = 7527 L_\odot$. The LBT data span from 2008.5 ($t_a \simeq 54$ yr) to 2018.4 ($t_b \simeq 64$ yr), so the change in luminosity should be

$$\Delta L = L_R e^{\tau_{HST}} [e^{-\tau_{HST}(t_{HST}/t_b)^2} - e^{-\tau_{HST}(t_{HST}/t_a)^2}], \quad (16)$$

where $t_{HST} \simeq 60$ yr. The observed lack of variability implies an upper limit on the R -band optical depth of $\tau_{HST} < 0.16$ at R -band at the time of the HST observations. The second approach is to compare our new HST photometry to that from Humphreys et al. (2017), who present photometry updated from the initial results reported in Van Dyk et al. (2005). They report $F475W$ and $F814W$ magnitudes of 24.11 and 22.73 mag, respectively, as shown in Fig. 2. These correspond to changes of 0.03 and 0.10 mag, both in the sense of the star brightening. If we treat these as upper limits on the star

becoming brighter, equation (3) implies $\tau_{F475W} < 0.07$ and $\tau_{F814W} < 0.25$ at the time of the first HST epoch near the end of 2013. The lack of variability strongly implies that star 4 cannot be surrounded by a dusty expanding medium with any significant optical depth.

Finally, Van Dyk et al. (2005) argued that star 4 was the most likely survivor of SN 1954J due to its strong $H\alpha$ emission. Like the lack of optical variability, the lack of $H\alpha$ variability is difficult to reconcile with material ejected in 1954. Recall that $H\alpha$ emission coming from the ionized ejecta should fall off like t^{-3} (equation 5). Humphreys et al. (2017) report no significant change in the $H\alpha$ flux between 2014.0 and 2017.1, a period over which we would expect the $H\alpha$ flux to drop by ~ 14 per cent. More significantly, the $H\alpha$ flux found by Humphreys et al. (2017) appears to be significantly greater than that reported by Van Dyk et al. (2005) at 2002.9, a period over which the $H\alpha$ luminosity should have dropped by a factor of ~ 2 . Van Dyk et al. (2005) did report poor seeing conditions, which may account for the difference in luminosity.

We can also use the observed $H\alpha$ luminosity of $L_{H\alpha} = 1.3 \times 10^{36}$ erg s^{-1} (Humphreys et al. 2017) to constrain the dust optical depth. Following the arguments in Section 2, we find that the amount of dust associated with the $H\alpha$ emission can be at most

$$\tau \approx 0.073 \kappa_2 t_{50}^{-1/2} v_3^{-1/2} \Delta_{0.1}^{1/2}, \quad (17)$$

in the low-optical depth limit and

$$\tau \approx 0.004 \kappa_2^2 t_{50}^{-1} v_3^{-1} \Delta_{0.1}, \quad (18)$$

in the high-optical depth limit. Here we set $t_{50} = 1.26$, corresponding to the optical depth in 2017. Both of these results are consistent with our upper limits on circumstellar absorption from the variability, and neither implies a high optical depth. If additional absorption is placed outside the $H\alpha$ emitting region, the limits on τ increase as $L_{H\alpha}^{1/2} (L_{H\alpha})$ in the low- (high-) optical depth limit. However, such dust would have to be pre-existing and not formed in the transient.

Finally, in the hot binary companion with a wind scenario used by Humphreys et al. (2017) to explain the broad asymmetric wings of star 4's $H\alpha$ profile with Thompson scattering, equation (15) means that $L_{H\alpha} = 340 L_\odot$ corresponds to a stellar luminosity of $L_* = 1600 L_\odot$ for a Thompson optical depth of $\tau_T = 1$ and no dust extinction. This already appears to rule out this hypothesis, as it is very difficult to see how a star could emit a fifth of its total luminosity in $H\alpha$ emission. The luminosity is also above our limits for the luminosity of a hot companion (Table 3). If we now add extinction, not only are the luminosity limits on a companion violated even further, but the constraint of a constant $H\alpha$ luminosity would actually force the optical continuum flux from the companion upwards rather than downwards, in further conflict with the data. In summary, our SED fits, the lack of variability, and the $H\alpha$ luminosity and its evolution in either a shell or a wind scenario all suggest that star 4 has negligible circumstellar extinction associated with the transient or any present

day wind. Aside from the SED fits, these constraints apply even to grey dust.

4 SN 1961V IN NGC 1058

Table 1 also includes the photometry of stars 6, 7, and 11 (see Fig. 1) near SN 1961V, where star 7 is believed to be the counterpart of SN 1961V. The SEDs of the stars are shown in Fig. 3. Both the archival and current band luminosities have been corrected for foreground Galactic extinction. We modelled the stars with DUSTY, with the results in Table 2 and Fig. 3. Stars 6 and 11 are well fit, but star 7 is not, as might be expected given the SED. The issue appears to be crowding in the WFC3/IR images, where stars 7 and 6 are not well resolved. This is reflected in the DOLPHOT crowding parameter for star 7. Crowding measures how much brighter a star would be in magnitudes for a given filter if all of the stars in the image were not fit simultaneously. Star 7 has crowding corrections of 0.501 and 0.712 mag in $F110W$ and $F160W$, respectively, compared to 0.050 and 0.099 mag in $F475W$ and $F814W$. This likely makes the near-IR photometry of star 7 unreliable.

Since we were unable to ‘fix’ the photometry, we first worked to obtain upper limits on the luminosity of star 7. We calculate the total stellar luminosity normalized to the star’s luminosity at $F475W$ over a grid of temperature and $E(B - V)$. We let temperature run from 3500 to 25 000 K in increments of 500 K. We let the reddening run from $E(B - V) = 0$ to 1 in increments of 0.05 mag and from 1 to 2 in increments of 0.1 mag. Given the total luminosity we calculate the band luminosities for the other three filters and set $\chi^2 = \sum_i (L_i/L_{\text{mod}})^2$, keeping only the points with $\chi^2 \leq 12$, corresponding to the model exceeding (on average) the flux in each of the other three filters by a factor of two. We then repeat the process, this time normalizing to the luminosity in $F160W$.

Fig. 4 shows the allowed luminosities as a function of $E(B - V)$ and fixed temperature for both the blue- and red-normalized models. The progenitor had $L_* \approx 10^{6.4} L_\odot$ (Goodrich et al. 1989), so we expect the survivor to be comparably luminous. Only the hottest stars with the highest reddening come close to reaching the progenitor luminosity. The measured stellar luminosity and $E(B - V)$ for stars 6 and 11 are shown for comparison. In these models, hot stars can be luminous because they need only pass through the $F475W$ band point, while lying below all the other bands. In the red-normalized case, with the right amount of extinction, the observed SED can be matched to the extinct Rayleigh–Jeans tail of a hot star, allowing the very high total luminosities for the red-normalized case.

However, we trust both the $F475W$ and $F814W$ band luminosities and so should have our models fit both. Unlike the IR data, there are no flags indicating any problems, and star 7 is nicely visible and isolated in both images. If we ignore star 7’s near-IR photometry and just model the optical photometry with DUSTY (see Fig. 3), we find much tighter constraints on the star’s temperature, luminosity, and foreground extinction. Table 2 shows the results. The luminosity is required to be nearly three orders of magnitude less than that of the progenitor of SN 1961V.

Fig. 3 also shows the $F606W$ data from 1994 (PI: Illingworth under program 5446) and the $F450W$ and $F814W$ data from 2001 (PI: Smartt under program 9042) reported in Van Dyk et al. (2002). Star 7 appears to have systematically faded by 1.7 mag in $F475W$. It would require an increasing optical depth of $\Delta\tau_{F475W} = 1.74$ to explain this change with dust. In $F814W$, there is a systematic fading of 1.4 mag, which corresponds to a change in optical depth of $\Delta\tau_{F814W} = 1.17$. For comparison, stars 4, 6, and 9 show no

significant changes, and any changes for stars 2 and 11 (Fig. 3) are far less significant. However, the quality of the old data is poor and the region around SN 1961V is only barely on the detector (see figs 2 and 3 in Van Dyk et al. 2002). Examining the images, we had difficulty convincing ourselves that the fading is actually real, but it is at least clear that there is no significant brightening of the source. If we pretend that the change in magnitude is reversed and that the source could have at most brightened by 1.7 and 1.4 mag in $F475W$ and $F814W$, we find that $\tau_{F475W} < 1.02$ and $\tau_{F814W} < 0.89$ based on equation (4).

Van Dyk & Matheson (2012) argue for the existence of a survivor of SN 1961V based partially on the $H\alpha$ emission. If the $H\alpha$ emission comes from the ionized ejecta, then $L_{H\alpha}$ should fall off with time, in this case, by a factor of 4.4 between the last two spectra taken of star 7. We examined the reported $H\alpha$ fluxes from Goodrich et al. (1989) and Chu et al. (2004) looking for signs of variability, but we could not identify any directly comparable estimates of the $H\alpha$ flux.

Using the most recent measurement of $L_{H\alpha}$, we can still place limits on the dust optical depth. Van Dyk & Matheson (2012) report that $L_{H\alpha}$ for SN 1961V is $6.5 \times 10^{36} \text{ erg s}^{-1}$, implying that

$$\tau \approx 0.19\kappa_2 v_3^{-1/2} \Delta_{0.1}^{1/2}, \quad (19)$$

in the low-optical depth limit, and

$$\tau \approx 0.04\kappa_2 t_{50}^{-1} v_3^{-1} \Delta_{0.1}, \quad (20)$$

in the high-optical depth limit using equations (9) and (10). We have set $t_{50} = 0.82$ corresponding to the optical depth in 2002. As before, adding additional extinction external to the $H\alpha$ emission can weaken these limits as $L_{H\alpha}^{1/2}$ ($L_{H\alpha}$) for the low (high) optical depths. However, as emphasized by Kochanek et al. (2012), adding any additional distant extinction essentially forces the event to be a supernova on energetic grounds.

We also considered the wind scenario from Humphreys et al. (2017) for SN 1961V. With $L_{H\alpha} = 1700 L_\odot$, equation (15) implies $L_* = 41 000 L_\odot$ for $\tau_T = 1$, $T_* = 20 000 \text{ K}$ and no dust. Reducing T_* to 7500 K to reflect the best-fitting SED model (optical bands only) gives $L_* = 810 L_\odot$, although such a star is also too cool to ionize the wind. As discussed in Section 2 and for SN 1954J, while adding dust increases the stellar luminosity, doing so at fixed $H\alpha$ flux also means that the observed continuum flux of the star increases rather than decreases as the dust optical depth is increased, violating the observed flux limits on the star.

5 SUMMARY

SN 1954J and SN 1961V were peculiar events, and the wealth of literature on the two over the past 60 yr is a testament to the difficulty of classifying them. The fundamental challenge is that the candidate counterparts have present day optical magnitudes significantly fainter than the progenitors. We know from *Spitzer* observations that neither source can be obscured by hot dust forming in a present day wind (Kochanek et al. 2011, 2012). This leaves three options for reconciling the fluxes if these are the surviving stars. First, the stars can become obscured by cold dust formed in the transient. Secondly, the star can be comparably luminous but have a much higher temperature so that bolometric corrections allow the optical fluxes to be larger. Thirdly, the star could have become intrinsically fainter. We are currently unaware of any mechanism which would cause the surviving star to become intrinsically fainter, so the latter explanation is the least likely of the three.

Here we have used four different probes of the optical depth of any dust formed in the transient: SED fits, photometric variability,

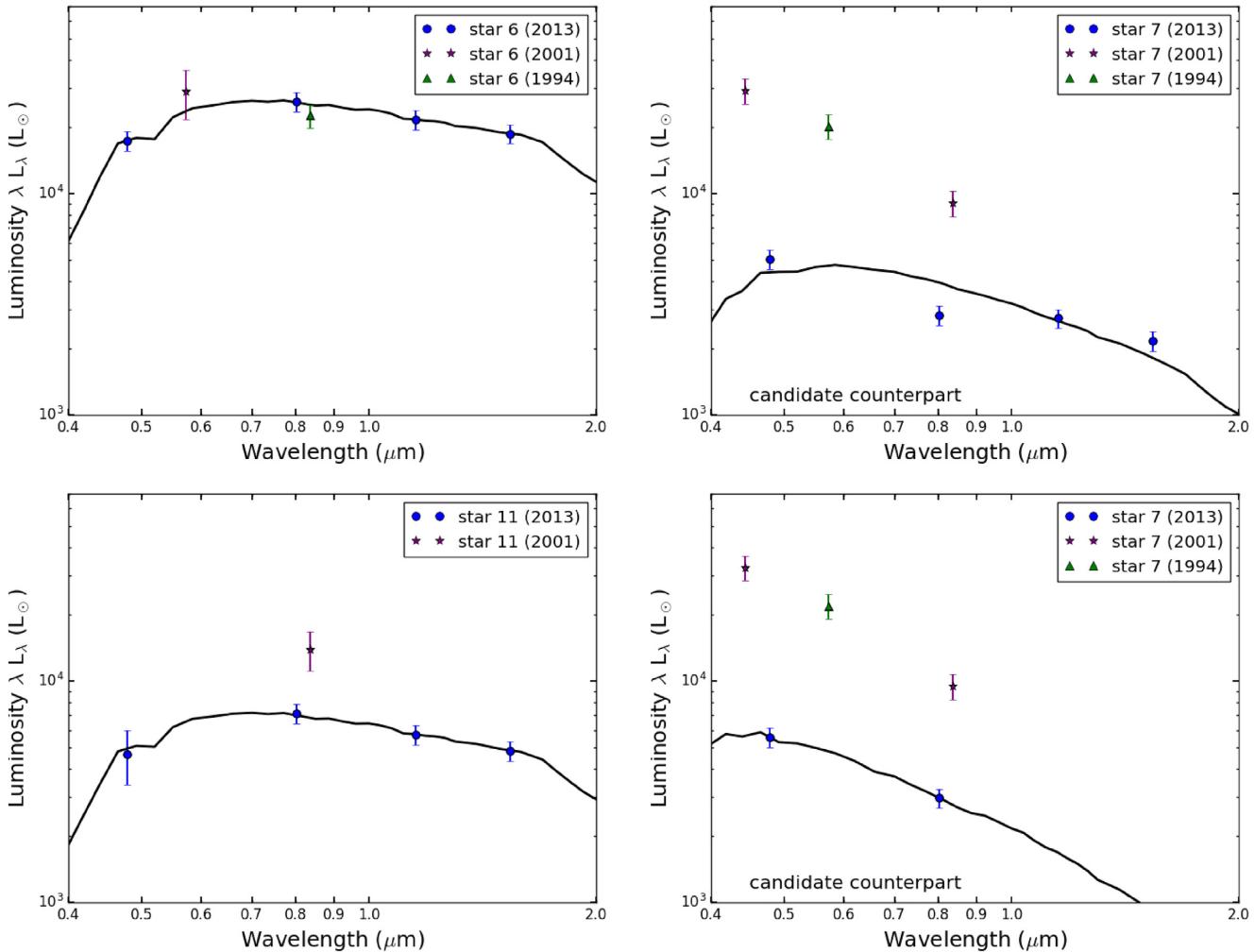


Figure 3. The measured SED of stars 6 (top left), star 11 (bottom left), and candidate counterpart star 7 (right) near SN 1961V with our current data (circles) and the archival photometry (stars and triangles) from Van Dyk et al. (2002). In the top right panel we show the best fit to the full SED, while the bottom panel shows the fit to just the optical bands. Star 7 appears to have significantly faded, but this may be an artefact (see the text).

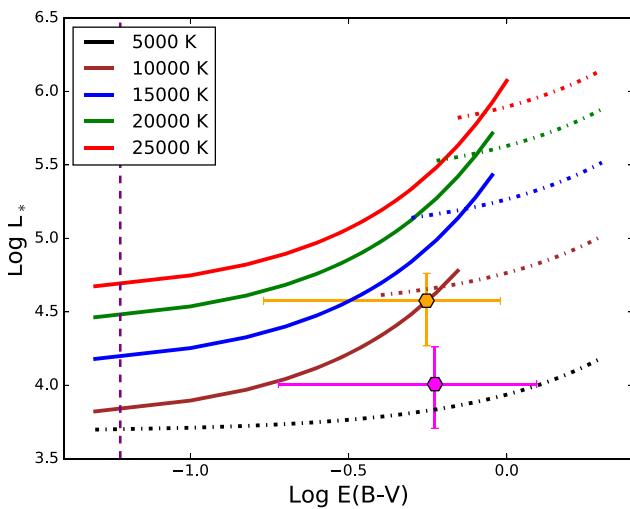


Figure 4. Constraints on the luminosity and $E(B-V)$ of star 7 from our photometry normalized at either $F475W$ (solid) or $F110W$ (dot-dashed). Stars 11 (magenta) and 6 (orange) are shown for comparison. $E(B-V)$ is the extinction in excess of the Galactic contribution.

and the $H\alpha$ luminosity and its evolution. All four probes imply little dust optical depth associated with the transient for both sources, far too little to significantly modify the luminosity of either star. This includes any significant grey opacity as invoked by Humphreys et al. (2017) in order to bring the luminosity of star 4 near SN 1954J back to its pre-transient level. While our SED fits have no sensitivity to grey dust, the time variability, and the $H\alpha$ luminosity limits hold regardless of the extinction curve.

The last spectra taken of stars 4 and 7 were in 2017 and 2002, respectively. New spectra of these targets could verify any $H\alpha$ variability. By 2020, we would expect to see any $H\alpha$ luminosity from the ionized ejecta fade by ~ 13 per cent for SN 1954J and ~ 66 per cent for SN 1961V. While we argue that the data already rule out a significant dust optical depth associated with the ejecta, there is also a problem with the required ejecta mass. For a thin shell, expanding at velocity v , the shell mass needed to produce an optical depth of τ is

$$M = \frac{4\pi v^2 t^2}{\kappa} \tau = 1.6\kappa_2^{-1} v_3^2 t_{50}^2 \tau M_\odot, \quad (21)$$

which means that both SN 1954J and SN 1961V require ejecta masses of $\sim 10 M_\odot$ to have a significant amount of absorption

($\tau \simeq 3$). This is not impossible, as it roughly matches estimates for η Carinae (e.g. Morris et al. 1999). Note, however, that the mass implied by the $H\alpha$ luminosity (equation 6) is just $0.2v_3^{3/2} \Delta_{0.1}^{1/2} M_\odot$ for both SN 1954J and SN 1961V. As time passes, all these variability arguments become stronger since they have a minimum scaling that is quadratic in time. As noted earlier, fragmentation of the ejecta only accelerates the evolution. *JWST* observations at $\sim 20 \mu\text{m}$ would also end any further speculation about dust.

Our DUSTY models favour relatively low temperatures of 5000–7000 K for both stars, so at least the observed stars cannot use a hot temperature to conceal a luminosity comparable to the progenitor. There is some room to have a hot companion, although not one as luminous as the progenitor. Particularly for very hot stars (Wolf–Rayet) stars, it is also important to remember that stellar atmosphere models greatly underestimate the optical luminosities because they do not include reprocessing of the UV radiation by the stellar wind into optical emission and emission lines (see e.g. Groh et al. 2013). It is marginally possible to have the hot companion suggested by Humphreys et al. (2017) for SN 1954J. *HST* observations at shorter wavelengths would easily constrain these possibilities.

The $H\alpha$ emission lines still seem a relatively compelling reason for the association of these stars with the transients. However, in the case of SN 1954J, Humphreys et al. (2017) argue against the $H\alpha$ emission arising from ejecta produced in the transient. They instead require a hot companion star with a dense stellar wind (see section 4.4 of their paper). The apparent lack of evolution, at least in the case of SN 1954J, and the lack of any obvious source of ionizing photons suggests that the line emitting material may not be associated with the observed transients. Instead, it could be material ejected in some earlier event that was then photoionized in 1954 or 1961. By placing the material further away, the evolution is slowed and it is easier to have long recombination times. For example, a solar mass of material at $R \sim 1/3$ pc in a thin shell ($\Delta = 0.1$) can produce roughly $L_{H\alpha} \simeq 10^{36} \text{ erg s}^{-1}$ with a recombination time of about 200 yr. At 1000 km s $^{-1}$, the shell would have been ejected nearly 300 yr ago. Shells, on smaller scales, quickly require a source of ionizing photons because the recombination times become shorter. We also found that the wind scenario considered by Humphreys et al. (2017) also requires the present day stars to be much less luminous than the progenitors, independent of the dust properties.

This appears to leave only the options of survivors that are intrinsically less luminous than their progenitors, or that these stars are not the survivors. While we have not explicitly carried out the full calculations, it seems clear that we would reach the same conclusions for any of the historical progenitor candidates. Continued monitoring of both the broad-band fluxes and the emission lines should steadily strengthen these conclusions.

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REFERENCES

- Adams S. M., Kochanek C. S., 2015, *MNRAS*, 452, 2195
 Adams S. M., Kochanek C. S., Gerke J. R., Stanek K. Z., 2017, *MNRAS*, 469, 1445
 Boroson T., 1981, *ApJS*, 46, 177
 Branch D., Cowan J. J., 1985, *ApJ*, 297, L33
 Branch D., Greenstein J. L., 1971, *ApJ*, 167, 89
 Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245
 Castelli F., Kurucz R. L., 2003, preprint ([arXiv:0405087](https://arxiv.org/abs/0405087))
 Chu Y.-H., Gruendl R. A., Stockdale C. J., Rupen M. P., Cowan J. J., Teare S. W., 2004, *AJ*, 127, 2850
 Cowan J. J., Henry R. B. C., Branch D., 1988, *ApJ*, 329, 116
 Dolphin A., 2016, Astrophysics Source Code Library. record ascl:1608.013
 Draine B. T., 1995, *Ap&SS*, 233, 111
 Elitzur M., Ivezić Ž., 2001, *MNRAS*, 327, 403
 Filippenko A. V., 1997, *ARA&A*, 35, 309
 Filippenko A. V., Barth A. J., Bower G. C., Ho L. C., Stringfellow G. S., Goodrich R. W., Porter A. C., 1995, *AJ*, 110, 2261
 Freedman W. L., Madore B. F., 1988, *ApJ*, 332, L63
 Gerke J. R., Kochanek C. S., Stanek K. Z., 2015, *MNRAS*, 450, 3289
 Goodrich R. W., Stringfellow G. S., Penrod G. D., Filippenko A. V., 1989, *ApJ*, 342, 908
 Groh J. H., Meynet G., Georgy C., Ekström S., 2013, *A&A*, 558, 24
 Humphreys R. M., Davidson K., 1994, *PASP*, 106, 1025
 Humphreys R. M., Davidson K., Smith N., 1999, *PASP*, 111, 1124
 Humphreys R. M., Davidson K., Van Dyk S. D., Gordon M. S., 2017, *ApJ*, 848, 86
 Ivezić Ž., Elitzur M., 1997, *MNRAS*, 287, 799
 Ivezić Ž., Nenkova M., Elitzur M., 1999, Astrophysics Source Code Library, record ascl:9911.001
 Klemola A. R., 1986, *PASP*, 98, 464
 Kochanek C. S., Szczygiel D. M., Stanek K. Z., 2011, *ApJ*, 737, 76
 Kochanek C. S., Szczygiel D. M., Stanek K. Z., 2012, *ApJ*, 758, 131
 Morris P. W. et al., 1999, *Nature*, 402, 502
 Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
 Schlegel E. M., 1990, *MNRAS*, 244, 269
 Smith N., Humphreys R. M., Gehrz R. D., 2001, *PASP*, 113, 692
 Smith N., Li W., Silverman J. M., Ganeshalingam M., Filippenko A. V., 2011, *MNRAS*, 415, 773
 Speck A. K., Barlow M. J., Sylvester R. J., Hofmeister A. M., 2000, *A&AS*, 146, 437
 Stockdale C. J., Rupen M. P., Cowan J. J., Chu Y.-H., Jones S. S., 2001, *AJ*, 122, 283
 Tammann G. A., Sandage A., 1968, *ApJ*, 151, 825
 Van Dyk S. D., Filippenko A. V., Chornock R., Li W., Challis P. M., 2005, *PASP*, 117, 553
 Van Dyk S. D., Matheson T., 2012, *ApJ*, 746, 179
 Van Dyk S. D., Filippenko A. V., Li W., 2002, *PASP*, 114, 700
 Zwicky F., 1964, *ApJ*, 139, 514

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