



The DEIMOS 10K Spectroscopic Survey Catalog of the COSMOS Field*

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Abstract

We present a catalog of 10,718 objects in the COSMOS field, observed through multi-slit spectroscopy with the Deep Imaging Multi-Object Spectrograph (DEIMOS) on the Keck II telescope in the wavelength range $\sim 5500\text{--}9800\text{ \AA}$. The catalog contains 6617 objects with high-quality spectra (two or more spectral features), and 1798 objects with a single spectroscopic feature confirmed by the photometric redshift. For 2024 typically faint objects, we could not obtain reliable redshifts. The objects have been selected from a variety of input catalogs based on multi-wavelength observations in the field, and thus have a diverse selection function, which enables the study of the diversity in the galaxy population. The magnitude distribution of our objects is peaked at $I_{AB} \sim 23$ and $K_{AB} \sim 21$, with a secondary peak at $K_{AB} \sim 24$. We sample a broad redshift distribution in the range $0 < z < 6$, with one peak at $z \sim 1$, and another one around $z \sim 4$. We have identified 13 redshift spikes at $z > 0.65$ with chance probabilities $< 4 \times 10^{-4}$, some of which are clearly related to protocluster structures of sizes > 10 Mpc. An object-to-object comparison with a multitude of other spectroscopic samples in the same field shows that our DEIMOS sample is among the best in terms of fraction of spectroscopic failures and relative redshift accuracy. We have determined the fraction of spectroscopic blends to about 0.8% in our sample. This is likely a lower limit and at any rate well below the most pessimistic expectations. Interestingly, we find evidence for strong lensing of $\text{Ly}\alpha$ background emitters within the slits of 12 of our target galaxies, increasing their apparent density by about a factor of 4.

Key words: catalogs – galaxies: distances and redshifts – surveys

Supporting material: machine-readable tables

1. Introduction

The Cosmic Evolution Survey (COSMOS; Scoville et al. 2007b) is a galaxy survey designed to probe the formation and evolution of galaxies, star formation, and active galactic nuclei (AGNs), both over cosmic time (redshift $z = 0.5\text{--}6$) and as a function of the local galaxy environment defined by the dark matter and its large-scale structure (LSS). It is designed to be representative of the large-scale structure of the universe defined by the dark matter scaffolding. The survey covers a 2 deg^2 equatorial field with multi-wavelength imaging and spectroscopy from X-ray to radio wavelengths by most of the major space-based telescopes (*Hubble*, *Spitzer*, *GALEX*, *XMM*, *Chandra*, *Herschel*, *NuStar*) and large ground-based observatories (Keck, Subaru, VLA, ESO-VLT, UKIRT,

NOAO, CFHT, JCMT, ALMA and others). Over 2 million galaxies are detected in deep optical images (Ilbert et al. 2009), and 1.2 million in the NIR (Laigle et al. 2016), spanning 75% of the age of the Universe. A subset of the field has also been selected as part of the “Cosmic Assembly Near-infrared Deep Extragalactic Legacy Survey” (CANDELS), and surveyed deeper in the NIR with *Hubble* (Nayyeri et al. 2017). Given the depth and resolution of these data, COSMOS also provides unprecedented samples of rare objects and structures at high redshifts with greatly reduced cosmic variance, compared to other, smaller area deep surveys. Comprehensive descriptions of the large-scale structure in the COSMOS Field have been presented by Scoville et al. (2007a, 2013), and Massey et al. (2007).

Spectroscopic identifications are an essential part of every large-scale cosmological survey, since redshifts and distances are required for practically all physical applications of the data. Massive spectroscopy campaigns have been undertaken in the COSMOS field over the last decade (e.g., Lilly et al. 2007, 2009; Trump et al. 2009; Coil et al. 2011; Le Fèvre et al. 2015;

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Table 1
Subsamples of Objects Included in Slit Masks

Sample	Total	Unique	$Q = 2$	$Q = 1.5$	$Q = 1$	$Q = 0$	Success	$\langle I_{AB} \rangle$	$\langle K_{AB} \rangle$
X-ray	1237	589	697	305	50	185	56.3%	23.0	21.0
high- z	2007	1878	564	468	141	834	28.1%	25.2	23.7
MIPS	3001	1588	2201	420	45	335	73.3%	22.0	20.4
VLA	1238	136	827	231	27	171	66.8%	22.5	20.5
<i>Herschel</i>	787	171	626	118	3	40	79.5%	22.2	20.8
OVV	363	314	255	63	8	37	70.2%	23.5	22.2
O II	262	238	184	45	1	32	70.2%	24.2	22.9
PL AGN	69	29	38	18	2	11	55.1%	23.0	21.6
Filler	3188	3188	2158	384	69	577	67.7%	22.0	20.9
Serendipitous	966	810	763	134	12	57	78.9%	22.3	21.3
Total	10718		6617	1798	279	2924	61.7%	22.9	21.4

Note. Shown here is the total number of objects and the number of unique objects in each subsample and each quality class Q (see below). Many objects fall in more than one subsample, as discussed in the text. Success is the percentage of high-quality spectra ($Q = 2$) compared to the total in each subsample. The last two columns are the mean I- and K-magnitudes of the subsample.

Silverman et al. 2015). Here, we describe the more than 100 multi-object spectroscopy observations of a total of more than 10,000 objects taken with the medium-resolution Deep Imaging Multi-Object Spectrograph (DEIMOS; Faber et al. 2003) on the Keck II telescope during the decade 2006–2016. Although the DEIMOS redshifts have already been used for a wide variety of fundamental astrophysical studies, this paper is the first comprehensive presentation of the entire spectroscopic data.

DEIMOS spectra have, e.g., been used to calibrate the excellent photometric redshifts for galaxies (Ilbert et al. 2009, 2013; Laigle et al. 2016) and AGNs (Salvato et al. 2009, 2011; Marchesi et al. 2016a) in the COSMOS field. They have also been used to characterize optical counterparts of *Spitzer* MIPS (Kartaltepe et al. 2010a, 2010b) and *Herschel* PACS sources (Lee et al. 2013), as well as X-ray sources discovered by *XMM-Newton* (e.g., Brusa et al. 2007, 2009, 2010; Elvis et al. 2012) and *Chandra* (e.g., Civano et al. 2012; Marchesi et al. 2016a, 2016b) and to determine the most accurate cosmic evolution of AGN luminosity functions so far (e.g., Masters et al. 2012; Miyaji et al. 2015). The 3D spatial correlation function of galaxies and AGNs, as well as their Halo occupation functions, can provide strong constraints on the type of dark matter halos, in which the objects typically reside (Gilli et al. 2009; Allevato et al. 2011, 2012, 2014, 2016). Optical spectroscopy has been used to determine the black hole masses in broad emission line objects and thus to measure the Eddington ratios of their accretion flows (Suh et al. 2015). The spectra have been used to identify high-redshift Ly α emitters (Mallery et al. 2012) and high-redshift proto-clusters in the COSMOS field (Capak et al. 2011), allowing a study of the metal properties of high-redshift galaxies (e.g., Faisst et al. 2016), as well as designing high-redshift ALMA surveys (e.g., Capak et al. 2015). Finally, the DEIMOS spectra will be used to calibrate future surveys with ESA’s Euclid and the NASA Wide-Field Infrared Survey-Telescope (WFIRST) missions (Masters et al. 2017), with far-reaching cosmological implications.

In this paper we present the overall observations in Section 2 and the spectral analysis in Section 3. The description of the released catalog is presented in Section 4, while a comparison of various metrics between different selection functions is discussed in Section 5. Examples of scientific applications of

the catalog are presented in Sections 6–7 and the summary in Section 8 concludes the paper. Throughout this work we adopt a Λ -cosmology with $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (Spergel et al. 2003), and all magnitudes are given in the AB system.

2. Observations

The spectroscopic observations in the COSMOS field were conducted with DEIMOS on the Keck II telescope. The field of view (FOV) of DEIMOS is approximately $16 \times 4 \text{ arcmin}^2$, which allows placement of slit masks in the field for multi-object spectroscopy of ~ 60 – 100 objects. The slit masks for our observations were prepared with the IRAF “dsimulator” software provided by the Keck observatory.¹⁵ For most masks we used a minimum slit length of $10''$ with a gap between slits of $0''.35$. The slit width was typically $1''$. The samples of objects placed on the slit masks were selected according to different criteria, depending on scientific objectives of the corresponding program. Table 1 gives a summary of the different subsamples we have grouped our targets in.

1. Spectroscopy of *Spitzer* mid-IR ($24 \mu\text{m}$ and $70 \mu\text{m}$) MIPS sources (PI: Kartaltepe). This contains a sample of ULIRGs selected through their mid-IR *Spitzer*/MIPS detections (see Kartaltepe et al. 2010a).
2. Deep spectroscopy of high-redshift candidates selected through a variety of broadband and narrowband photometry (see also Mallery et al. 2012); this subsample is denoted “high- z ” (PI: Capak). These objects are potential Lyman break galaxies (LBGs) and Ly α emitters (LAEs) at $z \sim 4.2$, $z \sim 4.8$, and $z \sim 5.6$ selected from intermediate and narrowband Subaru SuprimCam observations.
3. A subset of narrow- and intermediate-band excess sources selected possible [O II] emission line objects in the redshift range $0.3 < z < 1.6$; this is called the “O II” subsample (Takahashi et al. 2007). The authors discriminated against other possible strong emission lines using the broadband colors.
4. Spectroscopy of the optical counterparts of X-ray sources selected from the *XMM-Newton* (Hasinger et al. 2007)

¹⁵ <http://www2.keck.hawaii.edu/inst/deimos/dsim.html>

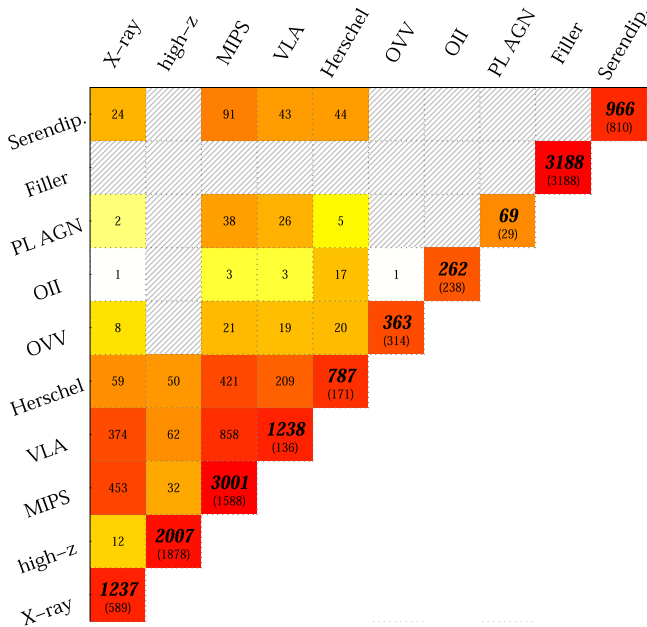


Figure 1. Sample cross-correlation matrix.

and *Chandra* (e.g., Civano et al. 2012; Marchesi et al. 2016a, 2016b) surveys (PI: Hasinger); this is the “X-ray” subsample.

- Additional slits were assigned to other sources with lower space density. These include radio sources (“VLA”), *Herschel* PACS sources (“*Herschel*”), and power-law AGN candidates (PL AGN, Donley et al. 2012; Chang et al. 2017).
- One subsample of spectroscopic targets consists of optically variable sources (“OVV”), as defined in Salvato et al. (2009). In the course of the determination of photometric redshifts for AGNs, the object photometry is corrected for variability. Using the same procedure, potentially variable objects have been selected here for follow-up spectroscopy.
- Slits in empty mask areas were assigned to “Filler” targets; these were predominantly drawn from a z-band magnitude-limited sample weighted to photometric redshifts at $z > 0.8$, where the zCOSMOS program (Lilly et al. 2009) had difficulties due to fringing. The zCOSMOS BRIGHT sample contains spectroscopy for a sample of 10,644 galaxies with $I_{AB} < 22.5$ mag. Another sample of filler targets was subthreshold *Chandra* X-ray detections.
- Serendipitous sources were picked up on the target slits and form the “Serendipitous” subsample.

This sample selection is not unique in each subsample, because a particular object can appear in several subsamples. There was also an evolution of the various subsamples over the years. When, e.g., deeper X-ray or VLA observations became available across the whole COSMOS field, some objects were added to the respective subsamples. Figure 1 shows the overlap matrix between the various subsamples. The diagonal shows the total number and number of unique objects in each subsample (also noted in Table 1), while the off-diagonal elements show the respective overlap with other subsamples. The entries are color-coded, with increasing sample sizes from white, to yellow, to orange, to red.

Compared to other, more homogeneous sample selections for multi-object spectroscopy, e.g., the magnitude-selected samples observed with VIMOS on the VLT (Lilly et al. 2009; Le Fèvre et al. 2015), our overall spectroscopic sample is much more heterogeneous. This, however, has the advantage that some of the magnitude-selection biases are washed out, and that our sample spans an unprecedented range in redshifts and magnitudes, as well as other characteristics (see below).

The observing log, covering the years 2007–2017, is shown in Table 2. The full list can be accessed online on the official COSMOS webpage.¹⁶ The typical seeing for these observations was $\sim 0''.7$ – $1''.2$. For some of the observations the 600ZD or 830G gratings were used, with blue blocking filters GG455 and OG550, respectively. The 600ZD grating yields a wavelength coverage of ~ 4800 – 10000 Å with a dispersion of 0.65 Å/pixel and a spectral resolution of $R \sim 2000$. The 830G grating yields a wavelength coverage of ~ 6700 – 10500 Å with a dispersion of 0.47 Å/pixel and a spectral resolution of $R \sim 2700$. The wavelength coverage varies with the position of the respective slit on the mask. The spectral resolution is sufficient to, e.g., distinguish the [O II] 3727 Å doublet emission line from a skewed Ly α profile and thus provides secure redshifts even in single emission line cases. Each mask was observed with a total integration time of 1–2 hr, typically split into four exposures with an ABBA dither pattern of $\pm 2''$. Depending on exposure times and spectral content we reach limiting magnitudes of $I_{AB} = 23.5$ – 25 . The last column in Table 2 shows the number of slits assigned for each mask, as well as the number of successful high-quality spectroscopic IDs, and serendipitous sources (see below).

3. Spectral Analysis

Most of the raw data were reduced using our specifically modified version of the DEEP2 data reduction pipeline. The original DEEP2 pipeline (*spec2d*; Cooper et al. 2012b; Newman et al. 2013) consists of the bias removal, flat-fielding, slit-tilt correction, cosmic ray rejection, sky subtraction, and wavelength calibration. The modified version accounts for dithering, removes the ghosting on the grating data, and corrects for variable slit losses and errors in the alignment introduced by the dithering. The flux calibration was then applied using the existing multi-wavelength photometry available on the COSMOS field.

For most of the spectroscopic analysis and redshift identification we used the interactive IDL program “*SpecPro*” developed for viewing and analyzing astronomical spectra in the context of the COSMOS survey (Masters & Capak 2011). Its interactive design lets the user simultaneously view spectroscopic, photometric, and imaging data, allowing for rapid object classification and redshift determination. The spectroscopic redshift can be determined by automated cross-correlation of continuum and spectral features against a variety of spectral templates or by manually overlaying common emission and absorption features on the 1D and 2D spectra. Stamp images and the spectral energy distribution (SED) of a source can be displayed with the interface, with the positions of prominent photometric features indicated on the SED plot. *SpecPro* also displays the slit position on the 2D spectra and stamp images, and enables the re-extraction of 1D spectra from the 2D data. This is in particular important for serendipitous

¹⁶ <http://cosmos.astro.caltech.edu>

Table 2
List of Observed Slit Masks

Name	R.A. (J2000)	Decl. (J2000)	PA ($^{\circ}$)	Date	UTC	Exp. (hr)	Airmass	Grating	Filter	Spectra
clae1	9:59:49	+02:07:43.7	6.5	2007 Jan 14	9:09:53	2	1.79	830G	OG550	75/56/5
clae2	10:01:36	+02:31:52.9	66.1	2007 Jan 14	15:11:22	2	1.28	830G	OG550	77/66/7
clae3	10:00:53	+02:35:14.7	82.9	2007 Jan 15	10:31:03	2	1.25	830G	OG550	80/77/16
clae4	10:02:24	+01:55:20.5	74.7	2007 Jan 15	14:19:21	2	1.14	830G	OG550	71/62/9
clae5	10:00:41	+01:33:42.1	94.6	2007 Jan 16	10:55:26	2	1.17	830G	OG550	74/64/11
clae6	9:59:48	+01:38:16.2	-15.9	2007 Jan 16	15:49:07	1	1.54	830G	OG550	72/78/20
clae7	10:01:35	+01:45:44.9	10.7	2007 Jan 17	11:19:07	2	1.12	830G	OG550	75/62/11
clae8	10:00:58	+01:52:32.1	60.1	2007 Jan 17	12:27:01	2	1.05	830G	OG550	77/43/1
clae9	10:00:05	+01:54:13.0	20.7	2007 Jan 18	8:02:42	2	2.69	830G	OG550	69/64/14
clae10	10:03:05	+01:55:25.7	0.1	2007 Jan 18	15:02:28	2	1.31	830G	OG550	71/70/15
B7	9:59:41	+02:27:47.7	26.0	2007 Jan 21	7:40:47	1	2.96	600ZD	GG455	117/58/0
F7	9:59:41	+02:27:48.0	26.0	2007 Jan 21	9:51:39	1.2	1.31	600ZD	GG455	122/61/0
...										
m3be	10:02:04	+02:39:52.5	113.5	2015 Feb 22	6:41:50	1	1.75	600ZD	GG455	62/48/4
m40be	10:00:37	+01:37:11.4	113.5	2015 Feb 22	7:59:35	1	1.27	600ZD	GG455	67/77/13
m28be	9:58:20	+02:09:11.3	113.5	2015 Feb 22	9:10:57	1	1.09	600ZD	GG455	60/50/5
m4bd	9:59:27	+02:51:25.1	113.5	2015 Feb 22	10:23:16	1	1.05	600ZD	GG455	75/76/12
m37be	10:01:24	+01:36:44.6	113.5	2015 Feb 22	10:49:58	1	1.06	600ZD	GG455	67/72/10
m7be	10:02:30	+02:31:02.2	113.5	2015 Feb 22	12:37:02	1	1.27	600ZD	GG455	68/71/10
mn46	9:59:40	+01:35:24.5	113.0	2015 Nov 10	13:36:49	1	1.70	600ZD	GG455	69/74/12
mn47	9:58:47	+02:38:22.1	32.0	2015 Nov 10	15:23:12	1	1.15	600ZD	GG455	68/63/6
mn44	9:57:42	+02:18:09.5	148.0	2015 Nov 11	13:29:48	1	1.69	600ZD	GG455	68/66/6
mn45	9:58:50	+02:22:06.3	48.0	2015 Nov 11	14:42:46	1	1.27	600ZD	GG455	67/72/9
mn48	10:01:51	+02:46:41.4	165.0	2015 Nov 11	15:08:48	1	1.19	600ZD	GG455	71/60/6
mn50	10:02:40	+02:07:27.5	46.0	2016 May 30	5:55:25	1	1.22	600ZD	GG455	72/62/8
mn51	9:58:11	+02:07:06.9	52.0	2017 Oct 27	13:52:14	1.3	2.14	600ZD	GG455	74/57/6

Note. The last column shows the number of spectra in each mask. The first entry is the number of slits assigned, the second entry is the number of high-quality redshifts ($Q \geq 1.5$) obtained, and the third entry is the number of serendipitous sources in each mask. The full list can be retrieved at <http://cosmos.astro.caltech.edu>.

(This table is available in its entirety in machine-readable form.)

sources on the slits, which can be analyzed and positionally identified using *SpecPro*. Such serendipitous detections were later identified in the photometric master catalogs (Ilbert et al. 2009; Laigle et al. 2016). For the quality assessment of the spectroscopic redshifts, we employed a scheme based on the zCOSMOS spectroscopic survey (Lilly et al. 2009), where $Q_f = 4$ is the highest-quality spectrum with several identified spectral features, and $Q_f = 3$ corresponds to a lower quality but still reliable spectroscopic identification with at least two spectral features or a single characteristically shaped emission line like a skewed Ly α or a double-humped [O II] line. $Q_f = 2$ and $Q_f = 1$ are lower-quality spectra with decreasing reliability. $Q_f = 9$ is based on a single high-significance emission line. If there are broad lines in the spectrum, the quality flag is increased by 10 (i.e., 14, 13, 12, 11, 19). If a source were detected serendipitously, the quality flag was increased by 20 (i.e., 24, 23, 22, 21, 29). Every spectrum was reviewed by at least two co-authors to find a consolidated solution in case of ambiguities.

For further discussion we define a more comprehensive quality flag “ Q ”, which combines spectroscopic and photometric redshift information, following Zheng et al. (2004). The Q_f flags 3, 4, 13, 14, 23, 24 are given the value $Q = 2$, signaling reliable spectroscopic identification. The Q_f flags 1, 2, 9, 11, 12, 19, 21, 22, 29 are given the value $Q = 1$ for an uncertain spectroscopic identification. However, if the photometric redshift value for a $Q = 1$ source is matching with the uncertain spectroscopic redshift within an interval

$\delta_z/(1+z) < 0.1$, where $\delta_z = |z_{\text{spec}} - z_{\text{phot}}|$, we raise the quality flag to $Q = 1.5$. An unsuccessful redshift measurement yields $Q = 0$. Figure 2 shows the magnitude distribution for samples with different spectroscopic qualities $Q = 2, 1.5, 1, 0$ as a function of I_{AB} and K_{AB} magnitudes.

As Table 1 shows, there is some correlation between the success rate for a particular subsample and its median i- and k-band magnitudes. The success rate, however, also depends on the redshift distribution in each subsample. Figure 3 shows both the optical (I_{AB}) and near-infrared (K_{AB}) magnitude distributions of the different spectroscopically identified ($Q = 2$) subsamples listed in Table 1 as stacked histograms (removing duplications). The high- z subsample (orange) contains the faintest objects, followed by the X-ray subsample (red). Figure 4 shows the corresponding redshift distributions. The X-ray-detected AGNs fill in the “redshift desert” ($1.5 < z < 3$) known for normal galaxies, thanks to their typically strong emission lines bluer than 3000 Å. The high-redshift sample completes the range up to $z \leq 6$.

Figure 5 shows the distribution of redshift versus K_{AB} magnitude. The total DEIMOS sample is shown in gray and each panel shows the sub-type of the selected targets as listed in Table 1. The unique sources are presented in black, while the shared subsample sources are overplotted in yellow. In red we show the sources with broad emission lines, indicating the presence of an AGN. These sources (mainly QSOs) are clearly dominant at higher redshifts in the X-ray subsample, compared to the lower-luminosity AGNs at lower redshifts. In

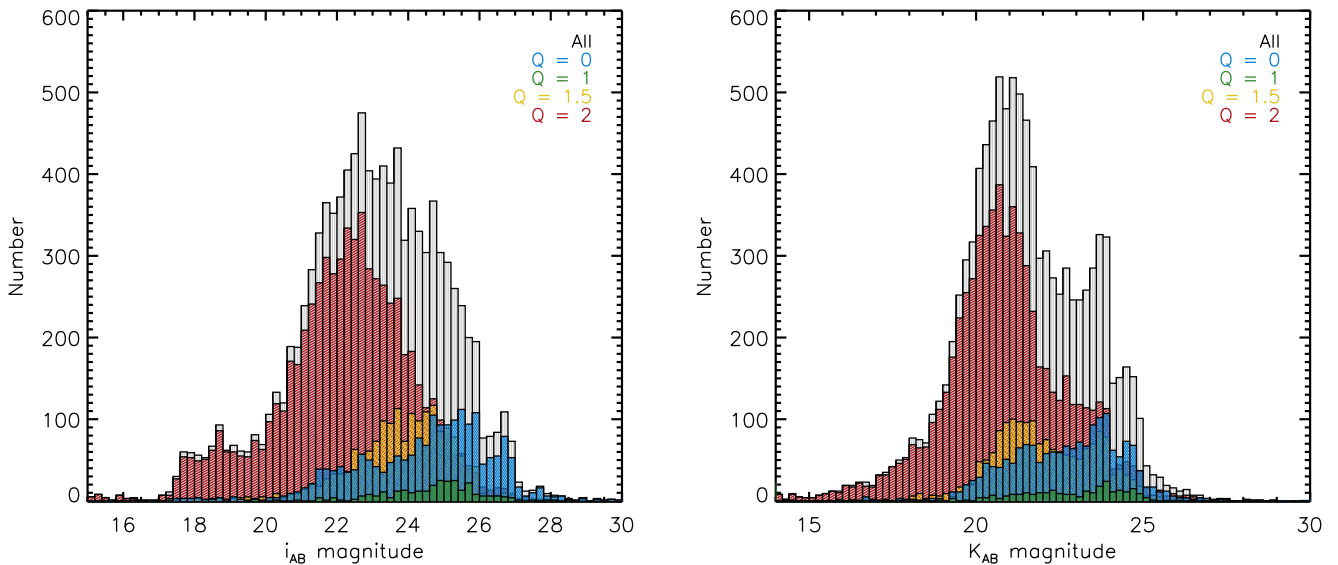


Figure 2. Histogram of the magnitude distribution of objects with different spectral quality classes (see the text). The highest-quality spectra (red) cover brighter optical/NIR magnitudes. Sources for which we failed to assign a redshift (in blue) are faint in the observed magnitudes. The gray histogram shows all sources.

X-ray-selected AGN samples there is a well known effect of a significant increase in the fraction of broad-line objects with increasing X-ray luminosity. This can be partially understood in terms of the difficulty of detecting broad lines of weak AGNs against the host galaxy light, but also due to a strong dependence of the obscured AGN fraction on X-ray luminosity (see, e.g., Hasinger 2008). The other subsamples contain only very few broad-line objects (the X-ray objects have been removed from all other subsamples). It is interesting to see how each of the multi-wavelength photometric selections corresponds to a characteristic distribution in this graph. The [O II] emission line selection worked well, yielding a narrow redshift range. The high-redshift photometric selection was rather successful, with the majority of the objects at $z > 3$, but with a significant fraction of lower-redshift interlopers, possibly associated with mis-identified color breaks and/or photometric noise. The mid- and far-infrared selections (MIPS and *Herschel*), as well as the VLA radio sources, have a very similar redshift versus magnitude pattern.

Figure 6 shows the stellar mass versus SFR relation for the 10 subsamples color-coded by their redshifts. The relation for star-forming galaxies at different redshifts (same color code) derived from *Herschel* photometry by Schreiber et al. (2015) is shown as a dashed line in each panel. The stellar mass and SFR for our galaxies are obtained by matching the spectroscopic sample to the COSMOS2015 catalog (Laigle et al. 2016) with a radius of $1''.5$. We only consider galaxies with quality flags $Q = 2$ and further restrict our sample to galaxies for which the photometric and spectroscopic redshifts agree within the uncertainty of the photometric redshifts, i.e., $|z_{\text{spec}} - z_{\text{phot}}|/(1 + z_{\text{spec}}) < 0.03$. The latter selection is important to provide the appropriate stellar masses and SFRs for our galaxies, since the COSMOS2015 catalog is based on photometric redshifts. In all panels we only use objects without broad emission lines. A full SED fitting using the spectroscopic redshift for all our galaxies is beyond the scope of this paper but will be published in a future work. We first note that the bulk of galaxies in our sample are consistent with being on the star-forming main-sequence. However, the sample also shows large diversity. In particular, there is a non-negligible fraction

of galaxies that are 2–3 mag below the main sequence at a fixed stellar mass, hence they are considered quiescent in star formation. Such galaxies are not seen in the power-law AGN (“PL-AGN”) and “O II” subsamples. In the former, the sample size might be too small to detect them. The latter is by definition targeting galaxies with [O II] emission via narrow-band and color-excess selections, and is thus biased to star-forming systems. The subsample of color and narrowband selected high-redshift galaxies (“High- z ”) contains some contamination ($< 20\%$) from low redshifts (see also Figure 3). Note that the accurate multi-band photometric redshifts in the COSMOS2015 catalog would have picked out these interlopers as opposed to the color and narrowband selections applied here. About two-thirds of the contaminants are star-forming and about one-third are quiescent galaxies. These could be mistaken as high-redshift galaxies because of their red color caused by dust or old stellar populations or also pure photometric scatter. We checked the spectra of the quiescent galaxies. Almost all of them seem to show strong H + K features, indicating that they are quiescent or in a post-starburst phase.

Figures 7 and 8 show some examples of high-quality spectra to illustrate different elements of the spectral classification applied here.

1. The object irsgv51416 (Figure 7(A)) is an emission line galaxy at $z = 0.753$. It has several interesting features. First, the [O III] 4960 Å and 5008 Å doublet emission lines show a broad, double-humped structure, indicating a rapidly rotating gas disk in this galaxy. The difference between the blue and the red wing of the double-humped line profile corresponds to a velocity difference of 770 km s^{-1} , or a rotation velocity of $\sim 390 \text{ km s}^{-1}$. This is the only object in the DEIMOS sample showing such a rapidly rotating disk. Simultaneously, this object shows strong Ca-H&K 3940 Å and 3960 Å absorption lines representative of an old stellar population, together with strong Balmer absorption lines of $H\delta$ 4103 Å, $H\epsilon$ 3971 Å, Hf 3888 Å, etc., corresponding to an A-star population,

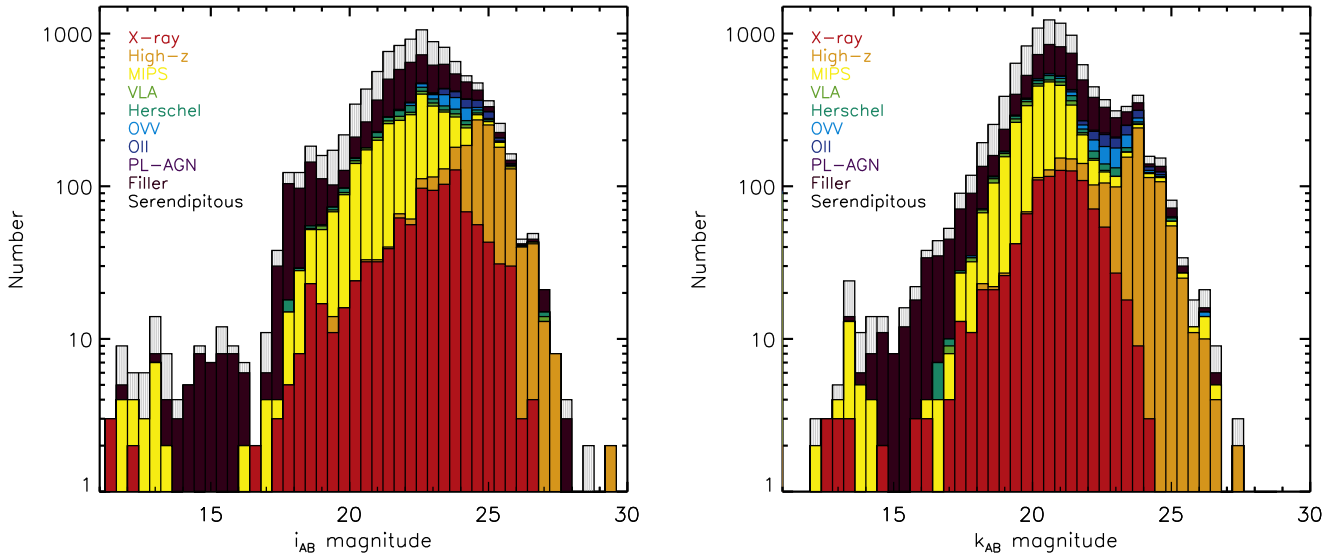


Figure 3. Stacked histograms for I_{AB} and K_{AB} magnitude distributions for the different subsamples of sources with reliable redshift ($Q = 2$). The high-redshift candidates (orange) show the faintest optical/NIR magnitudes.

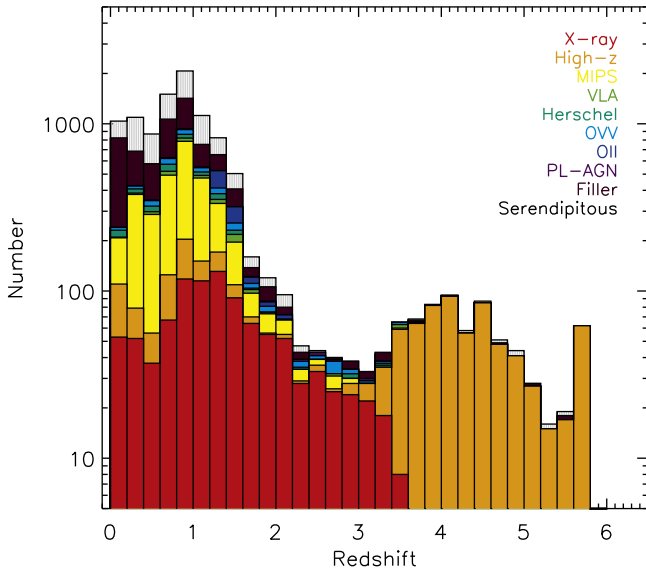


Figure 4. Redshift distribution of the different subsamples of objects on the DEIMOS masks.

indicative of more recent, quenched star formation. Together with the relatively weak [O II] 3728 Å emission line, this object therefore qualifies as a textbook example of a so-called “E+A” galaxy. Interestingly, this object is not detected at X-ray, infrared, or radio wavelengths. In the whole DEIMOS sample we detected about 30 objects with E+A features, but often together with stronger [O II] 3728 Å emission.

2. The object in Figure 7(B), hzoi-747644, has been selected as a strong [O II] 3728 Å emission line candidate through its intermediate-band-filter excess, paired with significant emission in the bluer bands. Indeed, a strong [O II] 3728 Å line has been detected in its spectrum, which shows the characteristic double-humped line profile (see the inset). In the bluer part of the spectrum we also see significant absorption lines of Mg II 2799 Å, Mg I 2852 Å, and

Si II 2881 Å. The total DEIMOS catalog contains more than 1000 detections of the [O II] 3728 Å doublet, which is therefore by far the most abundant emission line in our sample.

3. The object civ_2862 in Figure 7(C), an X-ray-selected broad-line QSO at $z = 2.179$, is interesting, because in addition to broad C IV 1549 Å, C III] 1909 Å, and Mg II 2799 Å lines, plus significant absorption features, it contains a second spectrum of an emission line object at $z = 0.209$. The *HST* ACS image of this object shows the QSO as a point-like object (actually close to a bright star), and a separate galaxy, which apparently is in the foreground of the QSO. It is possible that the QSO is slightly gravitationally lensed by the galaxy, but there are no other lensed images of the QSO. This object is one of the best examples of a superposition of spectra of two objects in the same slit. We have found quite a number of double spectra in our sample, which will be discussed in Section 7.
4. Object chandra_917 in Figure 7(D) is an example of a standard X-ray-selected broad-line QSO at $z = 3.097$, with some absorption features superposed on the broad lines. Figure 8(A) shows another broad-line object at $z = 3.317$, in this case selected through its optical variability.
5. Figures 8(B)–(D) shows Ly α emitters selected from our high- z samples. Although often only a single strong Ly α 1216 Å line is visible, the line can be identified through the characteristic skewed emission line shape (see, e.g., Laursen 2010) caused by the neutral intergalactic medium at high redshifts (see insert in Figure 8(B)). In the case of the object Rd-816509 we also clearly see the continuum drop blueward of the Ly α line.

4. The DEIMOS Redshift Catalog

Table 3 shows the DEIMOS spectral identifications and redshifts in the COSMOS field. The full catalog can be

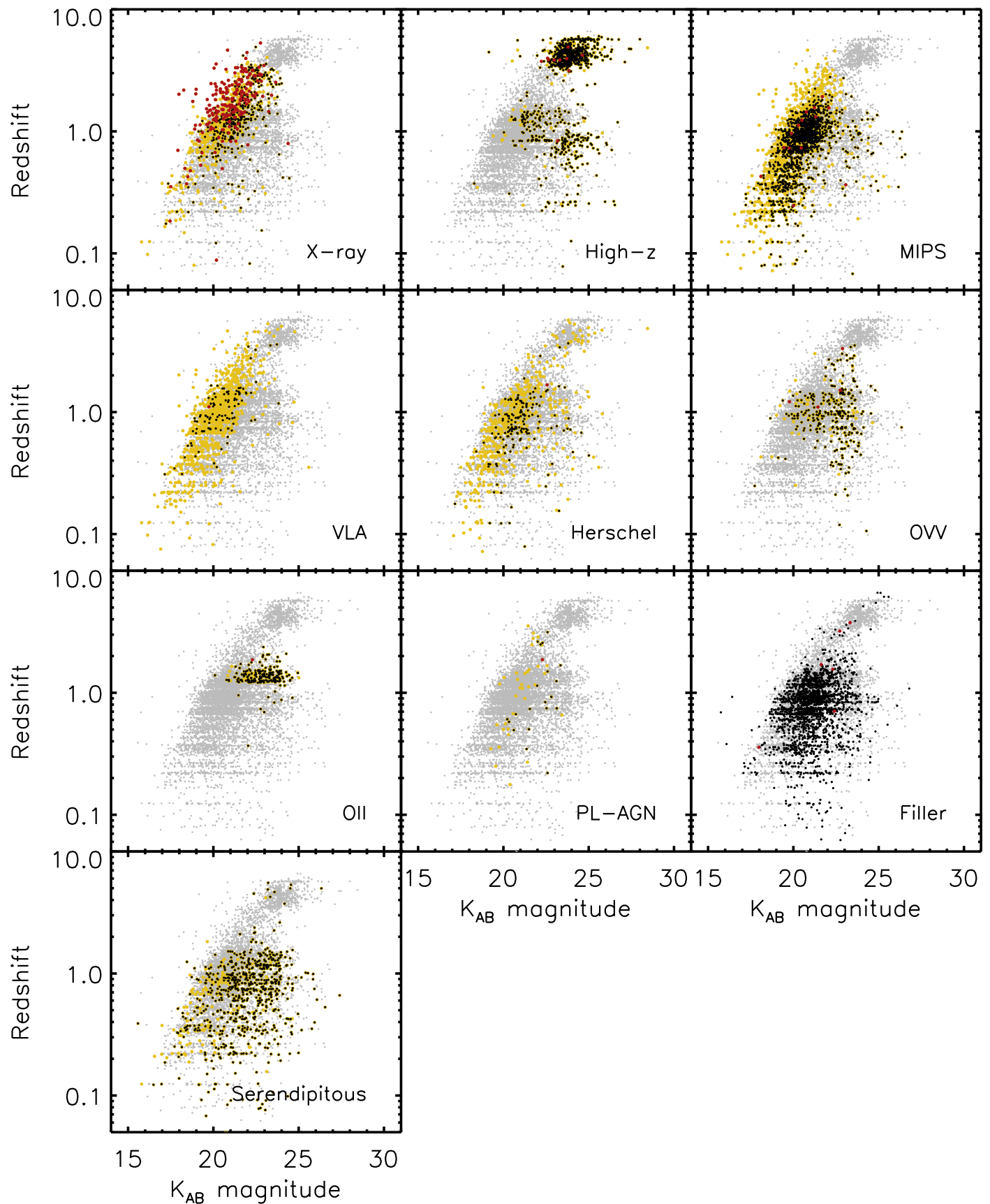


Figure 5. Redshift vs. K_{AB} magnitude for the different subsamples of objects with $Q \geq 1.5$ on the DEIMOS masks. The red asterisks are objects of the particular subsample with broad emission lines. The black points give the unique objects of this subsample, while the yellow data points show the objects overlapping with other subsamples. The underlying gray points in each panel show the total DEIMOS catalog.

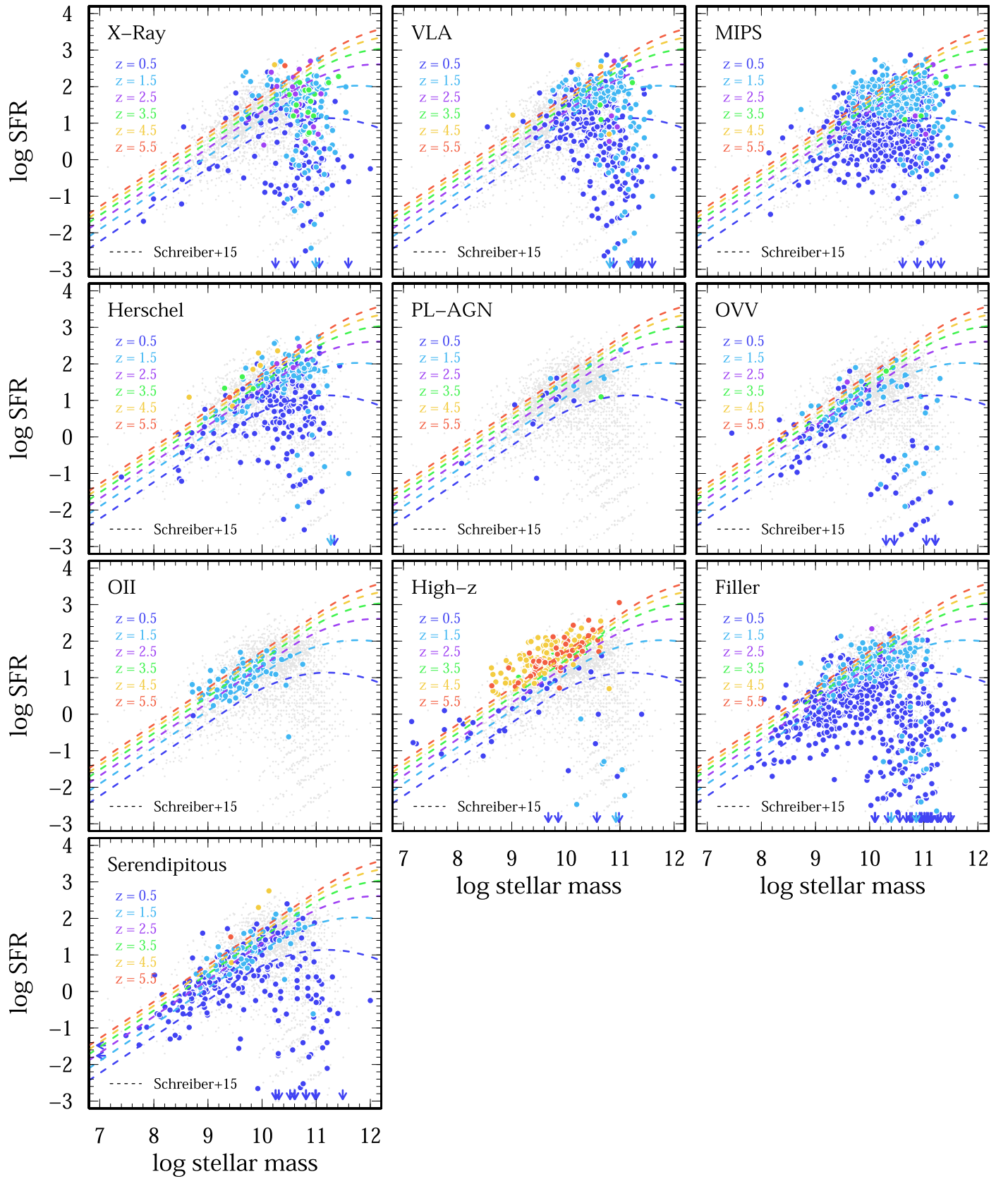


Figure 6. Star formation rate vs. stellar mass for high-quality spectroscopic and photometric subsamples in the DEIMOS catalog (see the text), compared to the star-forming main sequence at different redshifts. The dashed lines are derived from *Herschel* photometry by Schreiber et al. (2015).

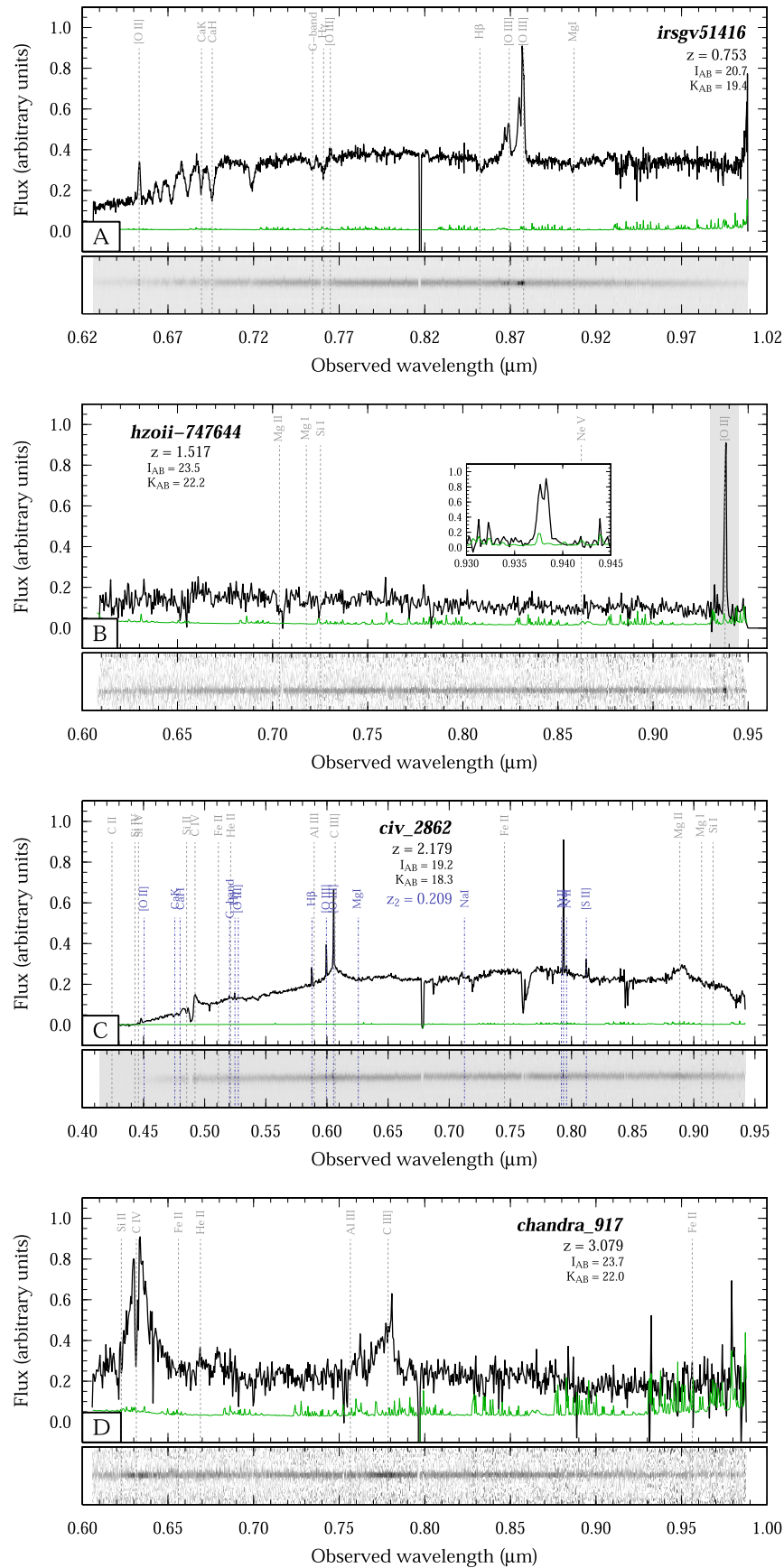


Figure 7. Sample of different spectra in the range $0.75 < z < 3.1$. Fluxes are in arbitrary units. The green line shows the noise level due to the sky background subtraction.

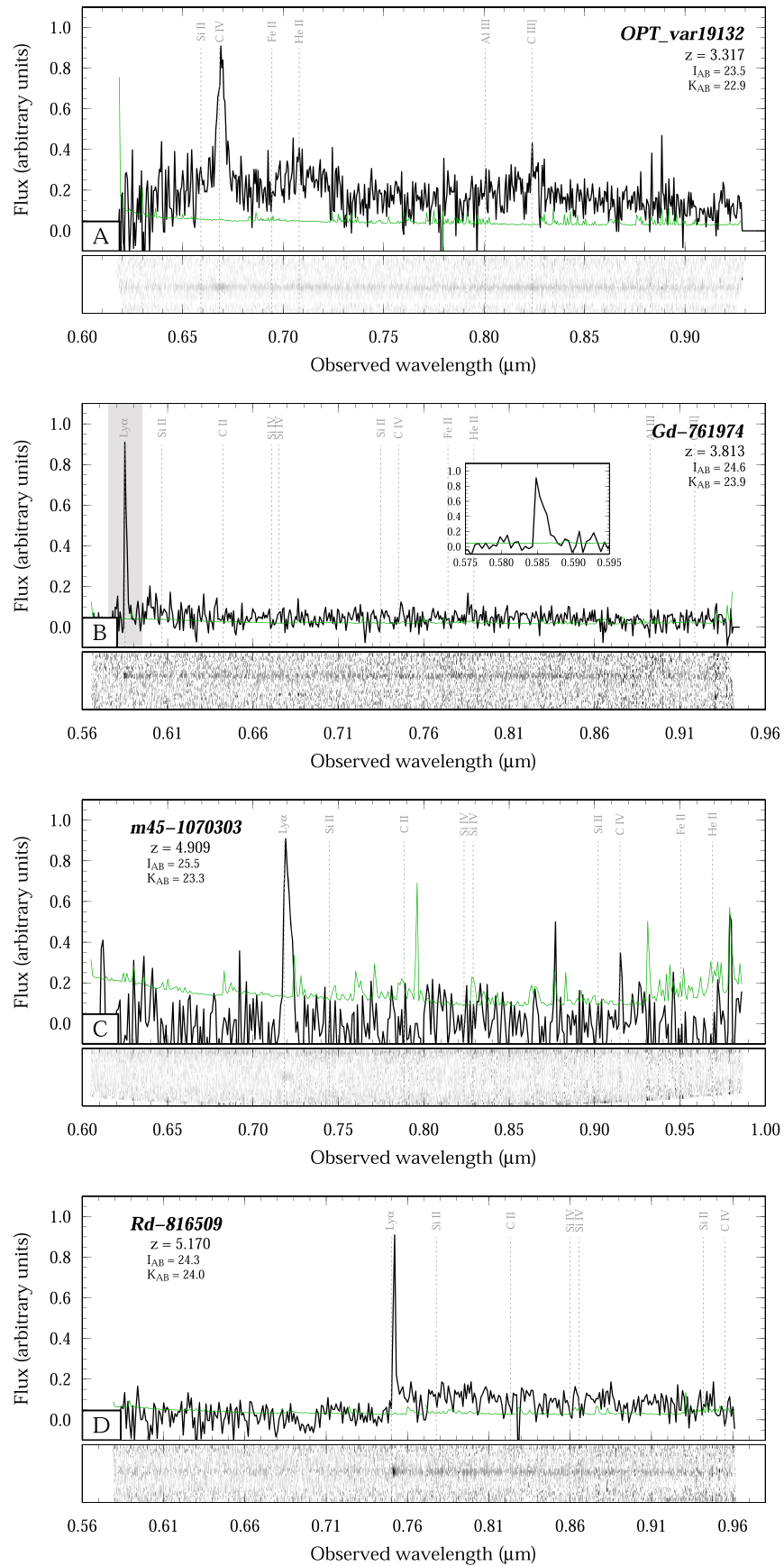


Figure 8. Sample of different spectra in the range $3.3 < z < 5.2$, as in Figure 7.

Table 3
List of Spectral Identifications

ID	R.A. (J2000)	Decl. (J2000)	Sel	i_{AB} (Mag)	k_{AB} (Mag)	Redshift	Q_f	Q	Remarks
C1785365	149.358553	2.750301	512	20.11	20.21	0.5655	2	1.5	H+K?
C1343700	149.366394	2.37669396	128	22.89	21.61	0.708	3	2	[O II], [O III]
C1347270	149.3690949	2.34408808	128	21.73	20.14	0.708	3	2	[O III]d
C1784226	149.37375	2.7765083	2	18.17	17.16	0	4	2	M-star
C1785765	149.3740997	2.7369051	2	21.37	19.63	0.549	2	1.5	[O II], H + K?
C1344930	149.3744965	2.41525006	2	18.36	15.95	0	4	2	star
C1344754	149.3755035	2.42347503	128	24.28	21.75	1.079	1	1.5	H γ , [O III]?
C1346553	149.377594	2.36321092	2	18.49	18.19	0	4	2	star
C1346616	149.3811035	2.36196303	512	21.84	20.52	0.772	4	2	[O II], [O III]d
C1345095	149.3820953	2.40950799	128	22.5	21.04	0.93	3	2	[O II], H β
C1345098	149.38211	2.4109753	1	17.39	16.55	0	22	1.5	star
L709816	149.3825989	2.37418294	512	22.82	21.73	3.367	14	2	Ly α , C IV(br), C III]
L941303	149.386634	2.743066	512	21.62	19.26	0.737	4	2	[O II], H + K, H β , MgI
L738187	149.3869019	2.41814995	128	23.37	21.97	1.099	3	2	[O II]d, K
...									
L365372	150.79555	1.8459245	1	22.53	24.45	0.842	24	2	[O II], H + K, H β , [O III]d
C1796700	150.7964783	2.9122829	128	21.1	23.56	0.298	4	2	H β , [O III]d, H α , N II, [S II]
C1797283	150.8010864	2.9061069	128	20.86	23.55	0.925	4	2	[O II], H + K, H β , [O III]?
C1793829	150.8035583	2.9437439	256	25.67			0	0	
C1795460	150.8054047	2.925	256	25.59		1.19	2	1.5	[O II], H, G?
C1797605	150.8085327	2.9029269	128	19.85	23.55	0.365	14	2	H β (br), [O III]d, H α (br)N II, [S II]
L1041939	150.8104248	2.8838029	128	23.16	23.55	1.203	2	1.5	[O II]
C1798722	150.8167114	2.8923731	128	20.97	23.55	0.924	4	2	[O II], H γ , H β ?
C1798723	150.817749	2.89088893	129	19.82	23.55	0.249	24	2	H β , [O III]d, H α , N II, [S II]
L1046705	150.8190918	2.8908889	8	23.7	23.55	1.352	1	1.5	[O II]?
C1797314	150.8242493	2.9058919	128	21.09	23.55	0.918	4	2	[O II], K, H γ , H β , [O III]
L1038944	150.8267365	2.879705	128	23.48	23.55		0	0	
	150.8373566	2.930975	256				0	0	
C1794871	150.8376007	2.9305582	1	23.92	23.55	1.365	3	2	[O II]d

Note. The full catalog can be retrieved from <http://cosmos.astro.caltech.edu> and in the online journal.

(This table is available in its entirety in machine-readable form.)

accessed online on the official COSMOS webpage¹⁷ and through the VizieR service. The first column gives an object identifier from the major two photometric catalogs. An “L” in front of the number refers to the red multi-band, band-selected catalog of Laigle et al. (2016). A “C” in front of the number refers to the i-band-selected catalog of Capak et al. (2007) and Ilbert et al. (2009).¹⁸ If an object is not present in either of these catalogs, it does not have an identifier. Columns two and three give the J2000 right ascension and declination for each object. The fourth column gives an identifier, in which the subsample of a particular object is present. It is a decimal representation of a binary flag¹⁹ containing the X-ray, high-z, MIPS, VLA, *Herschel*, OVV, O II, PL AGN, Filler, and Serendipitous flags following the order in Table 4 and Figure 1. The fifth and sixth columns give the I_{AB} and K_{AB} AB magnitudes, based on the ultra-deep Subaru Hyper Suprime-Cam (Tanaka et al. 2017) and UltraVista (Laigle et al. 2016), the Subaru Suprime-Cam (Ilbert et al. 2009), and the *Hubble* ACS (Koekemoer et al. 2007) photometric catalogs. Because of field-coverage, bright star cutouts, blending, or other confusion issues, not all objects in the spectroscopic catalog are covered by a single photometric catalog, thus we have to refer to various imaging

data sets. The seventh column gives the spectroscopic redshift z . The eighth column gives the spectroscopic quality flag ($Q_f = 0, 1, 2, 3, 4, 9, 11, 12, 13, 14, 19$), following the original zCOSMOS scheme (Lilly et al. 2009), where values 11–19 indicate broad emission lines (see above). The ninth column gives the reduced comprehensive spectral quality flag (see above). Finally, the tenth column gives remarks for most objects, in particular indicating the spectral features detected, e.g., the Ly α and Balmer lines (H α , H β , H γ , ...) of hydrogen, or the Mg II line, as well as the [C IV], C III], C II], [O II], [O III], N II, and [S II] emission lines. A “d” behind an emission line designation indicates a detected line doublet. A “br” behind an emission line refers to a broad emission line profile. An “abs” behind a line indicates its appearance in absorption rather than emission. H&K and G correspond to the Ca–H 3940 Å and Ca–K 3960 Å absorption lines and the G 4304 Å absorption band, respectively. Other prominent absorption lines are MgI 5175 Å and NaI 5892 Å. Finally, “E+A” features indicate the forest of spectral emission and absorption features (“ringing”) between the [O II] line and Ca–H&K, characteristic of post-starburst (E+A) galaxies.

5. Comparison between Redshift Surveys

There is a large variety of spectroscopic surveys in the COSMOS field shown in Table 4, which used different instruments and spectral resolution, different wavebands, and

¹⁷ <http://cosmos.astro.caltech.edu>

¹⁸ <http://cosmos.astro.caltech.edu/page/photom>

¹⁹ sel=512*X + 256* hiz+128*M + 64*VLA+32*H + 16*OVV+8*OII +4*PL+2*Fil+1*ser.

Table 4
Comparison with Other Spectroscopic Surveys

Sample	Total	$Q_f = 4$	$Q_f = 3$	$Q_f = 2$	$Q_f = 1$	$Q_f = 9$
3D- <i>HST</i>	69/0.0/3.4	69/0.0/3.4
DEIMOS-C3R2	38/2.6/1.1	31/0.0/1.1	5/0.0/0.2	1/0.0/0.0	1/100/–	...
FMOS-14	12/0.0/0.2	11/0.0/0.2	...	1/0.0/0.4
FMOS-15	155/12.9/7.4	62/1.6/2.8	50/6.0/6.7	11/63.6/1.3	32/28.1/13.3	...
FMOS-16	70/4.3/7.5	70/4.3/7.5
FOCAS	12/0.0/1.0	12/0.0/1.0
FORS2-11	41/2.4/0.6	41/2.4/0.6
FORS2-15	140/20.0/11.0	54/3.7/5.5	...	21/52.4/14.0	9/100/–	52/11.5/11.7
Gemini-S	29/3.4/0.9	19/0.0/0.2	9/11.1/1.6	1/100/–
IMACS	194/33.0/9.8	114/21.9/8.3	17/41.2/1.4	39/48.7/12.2	24/54.2/15.5	...
IRS	22/0.0/7.5	15/0.0/3.8	5/0/0.5	...	2/0.0/22.4	...
LRIS	45/11.1/1.5	13/0.0/1.5	8/0.0/2.3	8/12.5/1.5	16/25.0/1.1	...
MMT	35/0.0/4.3	35/0.0/4.3
MOSFIRE MOSDEF	12/16.7/1.6	12/16.7/1.6
PRIMUS	1657/18.4/10.9	824/4.6/5.8	309/27.8/11.5	437/41.2/14.0
SDSS-DR14	82/2.4/1.0	82/2.4/1.0
VIMOS-LEGA-C	198/1.0/3.2	198/1.0/3.2
WFC3 grism	20/0.0/9.0	10/0.0/2.4	9/0.0/4.7	1/0.0/39.6
zBRIGHT	1146/5.9/4.6	268/2.2/0.7	413/1.7/2.6	229/5.7/3.8	134/29.1/12.1	79/3.8/4.6
zDEEP	112/40.2/2.4	28/0.0/1.2	21/9.5/1.0	19/52.6/2.3	40/77.5/5.4	4/50.0/1.6
hCOSMOS	452/0.9/0.3	452/0.9/0.3
DEIMOS	568/3.4/2.7	493/1.8/0.9	29/17.2/6.4	19/15.8/2.7	13/7.7/12.8	8/12.5/3.0

Note. Q_f is the quality class in each of the spectral surveys (see the text). The three entries in each cell (1/2/3) are (1) total number of objects, (2) percentage of outliers ($\delta z/(1+z) > 0.05$), and (3) redshift accuracy $<\delta z/(1+z)>$ in units of 10^{-3} . References for the spectroscopic surveys are as follows: 3D-*HST* (Momcheva et al. 2016), DEIMOS-C3R2 (Masters et al. 2017), FMOS-15 (Kartaltepe et al. 2015), FMOS-16 (T. Nagao 2018, private communication), FORS2-11 (George et al. 2011), FORS2-15 (Comparat et al. 2015), Gemini-S (Balogh et al. 2014), IMACS (Trump et al. 2009), LRIS (C. Casey 2018, private communication), MMT (Prescott et al. 2006), PRIMUS (Coil et al. 2011), SDSS-DR14 http://www.sdss.org/dr14/data_access/, VIMOS-LEGA-C (van der Wel et al. 2016), zBRIGHT (Lilly et al. 2009), zDEEP (S. Lilly et al. 2018, private communication), hCOSMOS (Damjanov et al. 2018). The row in boldface is for emphasis.

target selections, as well as target redshift ranges. The multitude of COSMOS surveys therefore provides a unique opportunity to compare independent redshift measurements, and to cross-calibrate their respective individual quality assessments. For this purpose we have used our internal master catalog where we have collected all data available in the literature (M. Salvato 2018, in preparation). In Table 4, the quality entries in the individual surveys have been reduced to four basic quality classes, i.e., $Q_f = 1-4$, as well as the 9 for single line spectra. For those surveys with a more elaborate quality flag (e.g. zBRIGHT, zDEEP, VUDS, DEIMOS) the integer value of the Q_f quality flag has been taken modulo 10 for the purpose of this comparison (e.g., $Q_f = 14.5, 24.5$ etc. were reduced to 4). Objects in each individual survey and quality class was then compared to the overlapping objects in the DEIMOS $Q_f = 4$ class. In general, the overlap between different surveys should be small, because there is some degree of coordination between the different redshift surveys. However, because of the long time interval covered by our observations, because not all surveys were coordinated, and because there are serendipitous or filler targets, which can overlap with other surveys, there is a sizable sample of objects with spectra in more than one survey. Here, we restrict the comparison only to surveys, where there is a meaningful statistical sample (>15) of overlapping objects. Because two imperfect samples are compared to each other in each case, this procedure only provides information on the combined accuracy of the two samples in question. The true accuracy of the individual sample should be somewhat better. Table 4 shows the results of this analysis. The best spectroscopic

surveys in comparison to the highest-quality objects in our DEIMOS sample, with less than 2.5% spectroscopic failures and redshift difference less than 1.5×10^{-3} after removing the outliers, are the highest-quality class ($Q_f = 4$) for DEIMOS-C3R2, FMOS-14, FOCAS, FORS2-11, Gemini-S, LRIS, SDSS, zBRIGHT, zDEEP, and hCOSMOS. The quality comparison for the DEIMOS sample itself was done only against these highest-quality surveys ($Q_f = 4$). There is an overlapping sample of 568 DEIMOS objects with the combination of $Q_f = 4$ objects from DEIMOS-C3R2, FMOS-14, FOCAS, FORS2-11, Gemini-S, LRIS, SDSS, zBRIGHT, zDEEP, and hCOSMOS. Differences in quality are partially due to the distribution of redshifts and magnitudes and telescope/instruments involved. However, the time-consuming visual inspection and vetting of every redshift by several independent collaborators also plays an important role in the highest-quality surveys. The DEIMOS survey accuracy is among the best spectroscopic samples in the COSMOS field, with less than 2% spectroscopic failures and a redshift accuracy better than 10^{-3} .

6. Protoclusters and Large-scale Structures

Galaxies in the centers of nearby rich clusters and groups are passive, with little or no ongoing star formation (e.g., Baldry et al. 2006). Models for the evolution of cluster galaxies at low redshift based on the studies of cluster galaxy colors and luminosity functions point to galaxy populations that had a violent phase of star formation at high redshifts ($z > 2$) and have been passively evolving ever since (Mancone et al. 2010).

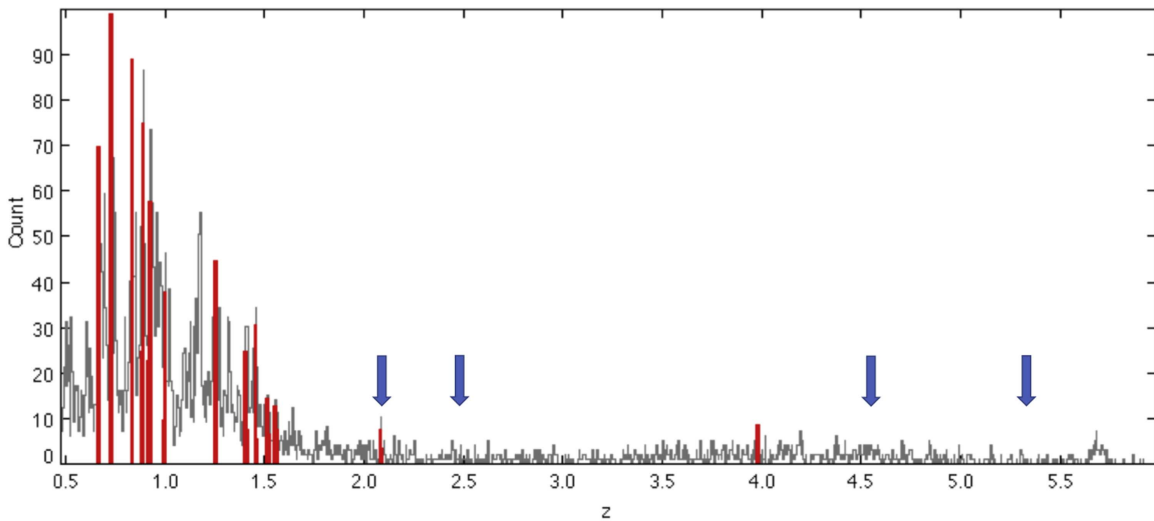


Figure 9. Redshift distribution of the catalog objects in bins of $\Delta z = 0.005$. Thirteen redshift spikes in the range $0.6 < z < 4.0$ have been colored in red. The blue arrows indicate previously reported protoclusters discussed in the text.

On the other hand, actively star-forming galaxies, luminous and ultraluminous infrared galaxies, as well as AGNs, are typically found in the outskirts of nearby clusters. A statistical study of clusters in the redshift range $1 < z < 1.5$ shows that the fraction of star-forming galaxies is systematically suppressed in the cluster centers and increases to the value of the field galaxies toward the cluster outskirts (Brodwin et al. 2013). This is true for clusters at redshifts $z < 1.4$, while at higher redshifts the trend reverses and the fraction of star-forming galaxies increases significantly toward the cluster center to values above the field galaxy ratio (Peng et al. 2010; Scoville et al. 2013; Darvish et al. 2016).

The highest-redshift bona fide cluster of galaxies has been discovered at $z \sim 2.5$ (Wang et al. 2016), with a violently starbursting core. At $z > 1.4$ there are about 5 known clusters, of which 2 have masses above $10^{14} M_{\odot}$ (Papovich et al. 2010; Mantz et al. 2016). Massive clusters at higher redshifts are extremely rare, because very few regions in the universe had time to virialize by that cosmic time. Cosmological simulations including cold dark matter and baryons (e.g., Springel et al. 2005) suggest that most massive clusters of galaxies have started their life at high redshift as overdensities of massive galaxies—“protoclusters”—which have not had time to virialize and therefore occupy regions in space about two orders of magnitudes larger than local clusters. We therefore have to look at protoclusters (e.g., Capak et al. 2015) to study the early violent formation phase of the progenitors of today’s cluster galaxies.

Significant spikes in the redshift distribution of X-ray-selected samples were first identified at $z \leq 1.1$ in the *Chandra* Deep Field North (Barger et al. 2002) and South (Hasinger 2002; Gilli et al. 2003; Szokoly et al. 2004), but are also present in the photometric redshift distribution of the X-ray sources in the Lockman Hole (e.g., Fotopoulou et al. 2012), XMM-XXL (Georgakakis et al. 2017), and STRIPE82X (Ananna et al. 2017). They point to AGNs in superclusters or sheet-like structures of the cosmic web. The question of whether AGN activity or star formation is enhanced in these structures compared to the field, remains open. With the advent of large photometric and spectroscopic surveys in the COSMOS field, a number of rich protocluster

structures have been identified at higher redshifts, which are believed to be the progenitors of some of the most massive clusters of galaxies in the local Universe. One of the more prominent of those is a massive protocluster of galaxies at a redshift of $z \sim 5.3$, with a size of more than 13 Mpc and containing a luminous quasar as well as a galaxy with a large amount of molecular gas (Capak et al. 2011). Another massive protocluster in the COSMOS field was recently discovered at $z = 4.57$ with $\log(M_h/M_{\odot})_{z=0} \sim 14.5\text{--}15$ by Lemaux et al. (2017). A third interesting object is a massive, distant protocluster at $z = 2.47$, found serendipitously during a spectroscopic redshift survey of dusty star-forming galaxies detected by *Scuba-2* in the COSMOS field. This structure may be seen in a phase of violent star formation (Casey et al. 2015), and may be connected to a large overdensity of Ly α emitters found in the HETDEX pilot survey (Adams et al. 2011) of the COSMOS field. Finally, a large-scale structure around $z \sim 2.2$ discovered in the Z-FOURGE photometric redshift survey in COSMOS (Spitler et al. 2012), could be confirmed spectroscopically as a Virgo-like cluster ancestor at $z = 2.095$ (Yuan et al. 2014). Both the $z = 2.47$ and the $z = 2.09$ feature are also present in the list of 36 candidate 15 Mpc-scale protocluster structures identified at redshifts $z = 1.6\text{--}3.1$ in the COSMOS field through photometric redshifts (Chiang et al. 2014). Diener et al. (2013) used the zDEEP sample to search for groups of galaxies in the COSMOS field in the redshift range $1.8 < z < 3$ within a physical distance of 500 Mpc and a velocity difference of 700 km s^{-1} . They identified 42 candidate groups with 3–5 members, and a comparison to mock catalogs indicates that most of them should be in large-scale structures, which later may merge into single groups, but almost none of them should already be virialized.

We analyzed our sample of spectroscopic redshifts to look for potential protocluster targets for future follow-up. Since the spectroscopic sample is very heterogeneous, it is not possible to perform a rigorous treatment. Figure 9 shows the redshift distribution of the DEIMOS sample in bins of $\Delta z = 0.005$. These redshift bins on one hand are about 5 times larger than the intrinsic redshift uncertainty in the catalog, and on the other hand correspond to velocity differences between 630 km s^{-1} (at $z = 0.8$) and 300 km s^{-1} (at $z = 4$), appropriate for the selection

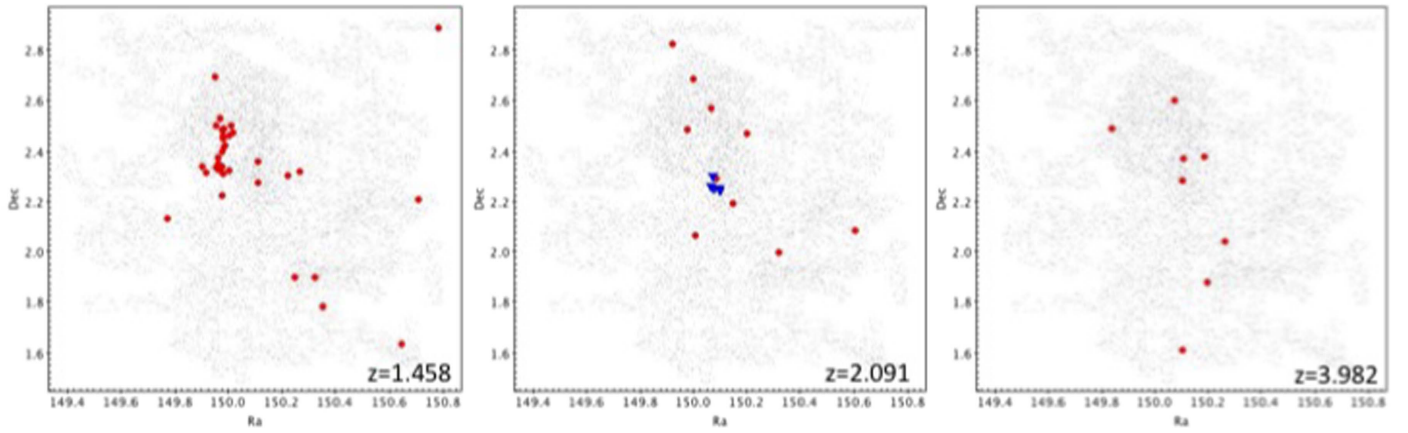


Figure 10. Spatial distribution of the objects in the redshift spikes at $z = 1.458$ (left), $z = 2.091$ (middle), and $z = 3.982$ (right). The centers of large galaxy overdensities at $z = 2.095$ identified by Yuan et al. (2014) are shown by blue triangles in the middle panel. The gray points show all objects in the DEIMOS redshift catalog.

Table 5
Spectroscopically Identified Redshift Spikes of Large-scale Structure

Spike	#Members	$\langle z \rangle$	σ_z	Expected	P_{Poisson}	Comment
z0667	69	0.6670	0.0013	25.2	7.1×10^{-11}	several filaments
z0732	98	0.7321	0.0014	28.2	6.4×10^{-11}	protocluster, several filaments
z0837	88	0.8373	0.0015	37.5	4.7×10^{-11}	two clumps, filaments
z0891	98	0.8910	0.0015	37.0	5.6×10^{-11}	two clumps
z0925	79	0.9253	0.0013	35.2	1.5×10^{-10}	...
z1001	46	1.0014	0.0015	16.1	9.1×10^{-10}	...
z1257	53	1.1257	0.0014	19.6	3.5×10^{-10}	...
z1408	31	1.4080	0.0015	14.7	1.4×10^{-4}	...
z1458	35	1.4579	0.0014	14.5	3.6×10^{-6}	one filament
z1518	20	1.5185	0.0013	5.9	4.0×10^{-6}	...
z1559	19	1.5586	0.0016	6.0	1.8×10^{-5}	...
z2091	10	2.0906	0.0013	2.3	1.4×10^{-4}	Virgo ancestor
z3982	8	3.9816	0.0013	1.7	3.9×10^{-4}	...

Note. Redshift spike members (column 2) are measured in a redshift interval $\Delta z = 0.005$ and the expected number of objects (column 5) is derived from a redshift interval of objects $\Delta z = 0.1$ (excluding the redshift spike). No correction for the relative area subtended by the structure is made. P_{Poisson} is the Poisson probability of obtaining the observed number of members given the expectation value. Only features with $P_{\text{Poisson}} \leq 4 \times 10^{-4}$ ($\sim 3\sigma$) have been retained. The last column gives some comments about the geometry of the structures.

of groups in large-scale structures (see above). Several significant redshift spikes are seen in this figure at redshift $z < 1.3$, corresponding to large-scale structure identified previously in deep fields (see above), and also in the COSMOS field (see, e.g., Lilly et al. 2009). Here, we concentrate on so far unexplored redshift spikes at $z > 0.8$. Table 5 lists 13 significant redshift spikes in the range $0.6 < z < 4$, which are also colored in red in Figure 9. In order to guarantee sufficiently low false positive likelihoods, we required a minimum of eight members per $\Delta z = 0.005$ redshift bin, and compared the detected number of members with the number expected from a $\Delta z = 0.1$ bin, reduced by the objects detected in the redshift spike. Since the redshift ranges covered by the narrowband filters are typically smaller than or comparable to the reference redshift range of $\Delta z = 0.1$, the narrowband filters create an artificial bias in this analysis, which could be misinterpreted as redshift spikes. We therefore excluded the redshift ranges selected by prominent emission lines in the narrowband filters. For the two Subaru Suprime-Cam narrowband filters NB711 and NB816, respectively, these are the redshift ranges $0.895 < z < 0.914$ and $1.164 < z < 1.208$ for the [O II] line, as well as the ranges $4.809 < z < 4.901$ and $5.637 < z < 5.771$ for the Ly α line,

respectively. Indeed, one redshift spike was detected at $\langle z \rangle = 1.176$, which we excluded here. Table 5 gives the number of objects detected in each spike, the average redshift, and its standard deviation, as well as the expected number of objects and the Poisson likelihood for a statistical chance occurrence of the number of members. No correction for the relative areas subtended by the redshift spike features on the sky were made at this point. The last entry in Table 5 is a comment on the spatial distribution of the feature.

Figure 10 shows the sky distribution of the objects in three of the 13 redshifts spikes discussed above. In the case of the $z = 1.458$ spike there is a strong concentration of ~ 18 objects in an elongated structure of $\sim 12 \times 4$ arcmin² ($\sim 6 \times 2$ Mpc²). In our sample this is the best example of a protocluster, which may turn into a rich cluster in the future. The DM halo mass of these structures can be estimated through mock catalogs from simulations (e.g., Diener et al. 2013). This, however, is beyond the scope of this paper and will be done in a future publication. The middle panel of Figure 10 shows a redshift spike of ten objects at $z = 2.091$. Six of these are X-ray sources, and four have been detected by *Spitzer* and *Herschel* in the mid-IR. Because of the difficult redshift range, most of

their redshift (8 out of 10) identification qualities are relatively low ($Q = 1.5$). Nevertheless, this structure is spatially consistent (both in redshift and sky distribution) with belonging to the Virgo-like cluster ancestor at $z = 2.095$ discussed above. The centers of the subclumps identified by (Yuan et al. 2014) are indicated by blue triangles in the figure. The right panel shows the redshift spike of eight objects at $z = 3.982$, the least significant and most dispersed of our candidates. Follow-up observations and comparisons with other spectroscopic catalogs are necessary to confirm the nature of these concentrations.

7. Spectroscopic Confusion and Ly α Lensing

Spectroscopic blending of galaxies at different redshifts in the same slit is a key source of uncertainty in future weak lensing cosmology experiments such as Euclid, LSST, and WFIRST (e.g., Newman et al. 2015; Dawson et al. 2016; Rhodes et al. 2017). Specifically, blended objects skew the photometric redshift distribution of weak lensing tomographic bins, leading to biases in the mean redshift distribution that must be corrected at the 0.2% level for stage IV dark energy experiments. Even if objects are at the same redshift, the blending leads to biases in shape measurement that need to be corrected. Therefore, constraining the fraction of blends with different and similar redshifts is key for these experiments. Estimates of possible blends, based on photometric counts, range from 1% in DEEP2 (Newman et al. 2015) to as high as 50% (Schneider 2014) at the depths of WFIRST, so distinguishing the number of blended objects that have measurably different redshifts is important.

In our spectroscopic sample we find 43 objects with evidence of 2 different redshifts in one spectrum. This is based on an interactive visual inspection of the spectra, and thus affected by biases. Typically, only strong emission line spectra can be identified as these kind of interlopers. In some of these cases we can clearly see the presence of two galaxies in the same slit from the *HST* ACS images. Compared to our sample of 5515 high-quality spectra of galaxies with magnitudes $19 < I < 25.3$ (the LSST lensing depth), this yields a fraction of 0.8% observed spectroscopic interlopers. This should be viewed as a lower limit since there is no guarantee we would obtain a redshift for the second source if it does not have strong emission features in the spectral range covered.

Another constraint on the potential fraction of interlopers with different redshift can be obtained from the percentage of catastrophic outliers in the comparison between spectroscopic and photometric redshifts. For this comparison we use the 3700 galaxies with the highest spectral quality ($Q = 2$) and magnitudes $19 < I < 25.3$, for which galaxy templates yield the best photometric redshift solution determined by Laigle et al. (2016) or Ilbert et al. (2009). Because of the additional photometric redshift model selection, this is a subsample of the 5515 high-quality spectra in the same magnitude range discussed above. Of these, 198 or 5.4% have catastrophic photometric redshift outliers $|\delta z / (1 + z)| > 0.1$. On one hand, this is likely a lower limit, because the two blended sources need to have roughly similar brightness and a large enough redshift difference, in order to yield a significant error on the photometric redshift. On the other hand, the fraction of photometric blends could be larger than that of spectroscopic blends, because the spectroscopic slit reduces the cross section for a blending impact. Also, we compare here the fraction of

real overlapping sources with the fraction of catastrophic errors in photo- z . The latter are due to a mix of two different aspects—one is the projected overlap, which may disturb the photometry, and the other is the inevitable degeneracy in the photo- z solution, which may lead to catastrophic outliers. This degeneracy is dependent on the depth and number of filters available, so it is difficult to draw any general conclusion from these numbers. Both of the effects discussed above, however, give an upper limit constraint on the true fraction of photometric confusion. Interestingly, the fraction of catastrophic photometric redshift outliers among the 50 identified spectroscopic blends is about 11%, about twice that in the general high-quality sample. These fractions of spectroscopic blends can be compared to the expected rate of 14% for all blends, including those at the same redshift (Dawson et al. 2016).

The combined redshift distribution of the spectroscopic blends contains a surprising number of high-redshift sources. About 28% of the objects (12 out of 43) have redshifts in the range $4 < z < 6$, all identified with Ly α emitters. We can compare this with the sample of all serendipitous sources detected in our target slits, which represents the best approximation of a fair blind spectroscopic survey of field galaxies. We have 682 serendipitous galaxies with high-quality spectra in our sample. Of these, only 8 objects (1.2%) are at redshifts $z > 3.7$, again, all of them Ly α emitters. In order to understand this large excess of blended Ly α emitters, we can look at the relative surface density of these objects in the different samples.

The surface density of serendipitous Ly α emitters can be estimated from the total blind spectroscopic area in our survey, which is the number of slits (about 7000 slits have been analyzed for serendipitous sources), multiplied by the average slit length ($\sim 12''$) times the average slit width ($\sim 1''$). This yields a blind survey area of ~ 23 arcmin², and thus a serendipitous Ly α source density of ~ 1250 deg⁻² over the redshift range $3.5 < z < 6$. In comparison, the LALA survey (Dawson et al. 2007) has spectroscopically identified 73 Ly α emitters in the redshift range $4.37 < z < 4.57$ in a field of 0.7 deg². In this particular redshift shell the surface density is thus ~ 100 deg⁻². Extrapolating this to the redshift range $3.7 < z < 6$ using the Ly α emitter luminosity function compiled by Faisst et al. (2014), yields a surface density of Ly α emitters of ~ 1000 deg⁻², comparable to our blind spectroscopic survey. This is in stark contrast to the 12 Ly α emitters blended with other galaxies in our survey. Again, the survey area can be estimated by the number of slits searched (in this case corresponding to the total number of reasonable quality spectra in our sample, i.e., ~ 7800), multiplied by the average size of objects on the slits (estimated to be $\sim 4''$) times the slit width $1''$. This yields a survey area of 8.7 arcmin², and thus a surface density of ~ 5000 deg⁻², about a factor of 4 higher than that in the blind field survey.

One possible interpretation of this excess is lensing and thus magnification bias of background Ly α emitters by the foreground galaxies targeted on the spectroscopic slits. One particular example, where we could localize the background Ly α emitter about $1''$ offset from the center of the putatively lensing foreground galaxy, is shown in Figure 11. The median Ly α flux of our blended source sample is $\sim 10^{-17}$ erg cm⁻² s⁻¹. Around this flux, the slope of the cumulative number counts of $4 < z < 6$ Ly α emitters is

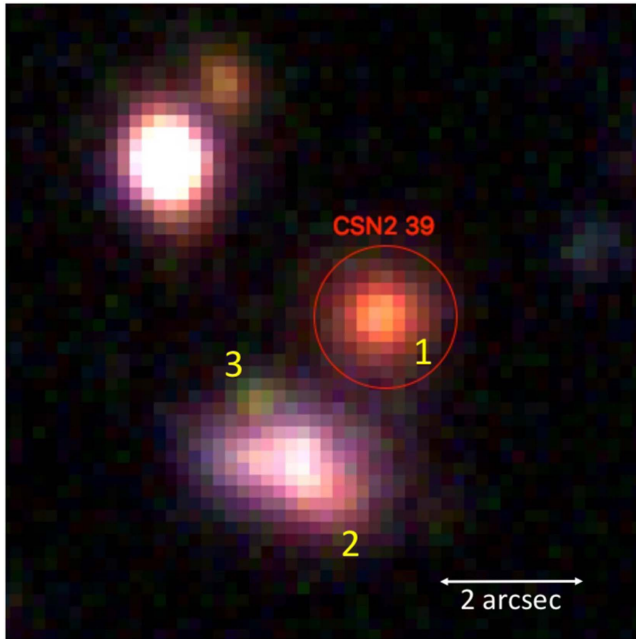


Figure 11. Example of a $z = 4.46$ $\text{Ly}\alpha$ emitter (object 3), putatively lensed by a $z = 0.75$ galaxy (object 2), serendipitously detected in a slit targeted at object 1, a $z = 1.22$ emission line galaxy. The angular separation of the lensed object from the lens is $\sim 1''$. The RGB image is composed of the ultra-deep Subaru HSC g- (blue), r- (green), and i-band (red) data in the COSMOS field (Tanaka et al. 2017). The $\text{Ly}\alpha$ line is seen only in the r-band.

estimated to roughly $\Gamma \sim -1.4$, again, extrapolating the luminosity function in Faisst et al. (2014). A factor of 4 increase of the surface density of $\text{Ly}\alpha$ emitters thus roughly corresponds to an effective flux limit a factor of ~ 2.7 lower. A significant part of this could indeed be due to strong lensing, if the lensed objects are situated well within the Einstein radii of their lenses. The average strong lensing magnification over the Einstein radius is about a factor of 2. At about 40% of the Einstein radius the magnification is 2.7. Most of the foreground galaxies of our blended $\text{Ly}\alpha$ emitters have stellar masses $8 < \log(M/M_\odot) < 11$ and redshifts $0.7 < z < 1.5$. Together with the redshifts of the lensed objects in the range $4 < z < 6$, and assuming a mass-to-light ratio ~ 200 and a fraction of 10% for the dark matter mass within the projected Einstein radius, this yields Einstein radii in the range $0.1 < R_E < 2''.3$, with a median around $0''.9$. Since our slit width is typically $1''$, it is conceivable that we detect strongly lensed objects well within the Einstein radii of their lenses. There are, however, aspects of caution: some fraction of the putative $\text{Ly}\alpha$ blends could be artifacts, like e.g. cosmic rays or sky subtraction issues. Also, assigning the interloper designation to the higher of the two blended redshifts could introduce some redshift bias.

8. Summary

We present a catalog of 10,718 objects in the COSMOS field observed through multi-slit spectroscopy with DEIMOS on the Keck II telescope. The objects have been selected from a variety of input catalogs based on multi-wavelength observations in the field, and thus have a diverse selection function. We have a success rate of 62% for high-quality spectra in the overall field and obtain a broad redshift distribution up to $z < 6$ with peaks at $z \sim 1$ and $z \sim 4$.

A direct object-to-object comparison with a multitude of other spectroscopic samples in the same field shows that our DEIMOS sample is among the best samples in terms of the fraction of discrepant spectroscopic redshifts ($\leq 1.8\%$), and relative redshift accuracy of $\langle \delta z / (1+z) \rangle \leq 9 \times 10^{-4}$.

We have identified 13 redshift spikes at $z > 0.65$ with chance probabilities $< 4 \times 10^{-4}$, some of which are clearly related to protocluster structures of sizes > 10 Mpc.

We have determined the fraction of spectroscopic blends to be about 0.8% of our sample. This is likely a lower limit and at any rate well below the most pessimistic expectations. Interestingly, we find evidence for strong lensing of $\text{Ly}\alpha$ background emitters within the slits of 12 of our target galaxies, increasing their apparent surface density by about a factor of 4.

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Facilities: Keck:2.

Software: Dsimulator (<http://www2.keck.hawaii.edu/inst/deimos/dsim.html>), spec2d (Cooper et al. 2012a; Newman et al. 2013), SpecPro (Masters & Capak 2011).

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References

- Adams, J. J., Blanc, G. A., Hill, G. J., et al. 2011, *ApJS*, **192**, 5
- Allevato, V., Civano, F., Finoguenov, A., et al. 2016, *ApJ*, **832**, 70
- Allevato, V., Finoguenov, A., Cappelluti, N., et al. 2011, *ApJ*, **736**, 99
- Allevato, V., Finoguenov, A., Civano, F., et al. 2014, *ApJ*, **796**, 4
- Allevato, V., Finoguenov, A., Hasinger, G., et al. 2012, *ApJ*, **758**, 47
- Ananna, T. T., Salvato, M., LaMassa, S., et al. 2017, *ApJ*, **850**, 66
- Baldry, I. K., Balogh, M. L., Bower, R. G., et al. 2006, *MNRAS*, **373**, 469
- Balogh, M. L., McGee, S. L., Mok, A., et al. 2014, *MNRAS*, **443**, 2679
- Barger, A. J., Cowie, L. L., Brandt, W. N., et al. 2002, *AJ*, **124**, 1839
- Brodwin, M., Stanford, S. A., Gonzalez, A. H., et al. 2013, *ApJ*, **779**, 138
- Brusa, M., Civano, F., Comastri, A., et al. 2010, *ApJ*, **716**, 348
- Brusa, M., Comastri, A., Gilli, R., et al. 2009, *ApJ*, **693**, 8
- Brusa, M., Zamorani, G., Comastri, A., et al. 2007, *ApJS*, **172**, 353
- Capak, P., Aussel, H., Ajiki, M., et al. 2007, *ApJS*, **172**, 99
- Capak, P. L., Carilli, C., Jones, G., et al. 2015, *Natur*, **522**, 455
- Capak, P. L., Riechers, D., Scoville, N. Z., et al. 2011, *Natur*, **470**, 233
- Casey, C. M., Cooray, A., Capak, P., et al. 2015, *ApJL*, **808**, L33
- Chang, Y.-Y., Le Floc'h, E., Juneau, S., et al. 2017, *MNRAS*, **466**, L103
- Chiang, Y.-K., Overzier, R., & Gebhardt, K. 2014, *ApJL*, **782**, L3
- Civano, F., Elvis, M., Brusa, M., et al. 2012, *ApJS*, **201**, 30
- Coil, A. L., Blanton, M. R., Burles, S. M., et al. 2011, *ApJ*, **741**, 8
- Comparat, J., Richard, J., Kneib, J.-P., et al. 2015, *A&A*, **575**, A40
- Cooper, M. C., Griffith, R. L., Newman, J. A., et al. 2012b, *MNRAS*, **419**, 3018
- Cooper, M. C., Newman, J. A., Davis, M., Finkbeiner, D. P., & Gerke, B. F. 2012a, spec2d: DEEP2 DEIMOS Spectral Pipelinem, Astrophysics Source Code Library, ascl:[1203.003](#)
- Damjanov, I., Zahid, H. J., Geller, M. J., Fabricant, D. G., & Hwang, H. S. 2018, *ApJS*, **234**, 21
- Darvish, B., Mobasher, B., Sobral, D., et al. 2016, *ApJ*, **825**, 113
- Dawson, S., Rhoads, J. E., Malhotra, S., et al. 2007, *ApJ*, **671**, 1227
- Dawson, W. A., Schneider, M. D., Tyson, J. A., & Jee, M. J. 2016, *ApJ*, **816**, 11
- Diener, C., Lilly, S. J., Knobel, C., et al. 2013, *ApJ*, **765**, 109
- Donley, J. L., Koekemoer, A. M., Brusa, M., et al. 2012, *ApJ*, **748**, 142
- Elvis, M., Hao, H., Civano, F., et al. 2012, *ApJ*, **759**, 6
- Faber, S. M., Phillips, A. C., Kibrick, R. I., et al. 2003, *Proc. SPIE*, **4841**, 1657
- Faisst, A. L., Capak, P., Carollo, C. M., Scarlata, C., & Scoville, N. 2014, *ApJ*, **788**, 87
- Faisst, A. L., Capak, P. L., Davidzon, I., et al. 2016, *ApJ*, **822**, 29
- Fotopoulou, S., Salvato, M., Hasinger, G., et al. 2012, *ApJS*, **198**, 1
- Georgakakis, A., Salvato, M., Liu, Z., et al. 2017, *MNRAS*, **469**, 3232
- George, M. R., Leauthaud, A., Bundy, K., et al. 2011, *ApJ*, **742**, 125
- Gilli, R., Cimatti, A., Daddi, E., et al. 2003, *ApJ*, **592**, 721
- Gilli, R., Zamorani, G., Miyaji, T., et al. 2009, *A&A*, **494**, 33
- Hasinger, G. 2002, arXiv:[astro-ph/0202430](#)
- Hasinger, G. 2008, *A&A*, **490**, 905
- Hasinger, G., Cappelluti, N., Brunner, H., et al. 2007, *ApJS*, **172**, 29
- Ilbert, O., Capak, P., Salvato, M., et al. 2009, *ApJ*, **690**, 1236
- Ilbert, O., McCracken, H. J., Le Fèvre, O., et al. 2013, *A&A*, **556**, A55
- Kartaltepe, J. S., Sanders, D. B., Le Floc'h, E., et al. 2010a, *ApJ*, **709**, 572
- Kartaltepe, J. S., Sanders, D. B., Le Floc'h, E., et al. 2010b, *ApJ*, **721**, 98
- Kartaltepe, J. S., Sanders, D. B., Silverman, J. D., et al. 2015, *ApJL*, **806**, L35
- Koekemoer, A. M., Aussel, H., Calzetti, D., et al. 2007, *ApJS*, **172**, 196
- Laigle, C., McCracken, H. J., Ilbert, O., et al. 2016, *ApJS*, **224**, 24
- Laursen, P. 2010, PhD thesis, Univ. Copenhagen
- Le Fèvre, O., Tasca, L. A. M., Cassata, P., et al. 2015, *A&A*, **576**, A79
- Lee, N., Sanders, D. B., Casey, C. M., et al. 2013, *ApJ*, **778**, 131
- Lemaux, B. C., Tomczak, A. R., Brodwin, M., et al. 2017, *MNRAS*, **472**, 419
- Lilly, S. J., Le Brun, V., Maier, C., et al. 2009, *ApJS*, **184**, 218
- Lilly, S. J., Le Fèvre, O., Renzini, A., et al. 2007, *ApJS*, **172**, 70
- Mallery, R. P., Mobasher, B., Capak, P., et al. 2012, *ApJ*, **760**, 128
- Mancone, C. L., Gonzalez, A. H., Brodwin, M., et al. 2010, *ApJ*, **720**, 284
- Mantz, A. B., Allen, S. W., Morris, R. G., et al. 2016, *MNRAS*, **463**, 3582
- Marchesi, S., Civano, F., Elvis, M., et al. 2016a, *ApJ*, **817**, 34
- Marchesi, S., Civano, F., Salvato, M., et al. 2016b, *ApJ*, **827**, 150
- Massey, R., Rhodes, J., Ellis, R., et al. 2007, *Natur*, **445**, 286
- Masters, D., & Capak, P. 2011, *PASP*, **123**, 638
- Masters, D., Capak, P., Salvato, M., et al. 2012, *ApJ*, **755**, 169
- Masters, D. C., Stern, D. K., Cohen, J. G., et al. 2017, *ApJ*, **841**, 111
- Miyaji, T., Hasinger, G., Salvato, M., et al. 2015, *ApJ*, **804**, 104
- Momcheva, I. G., Brammer, G. B., van Dokkum, P. G., et al. 2016, *ApJS*, **225**, 27
- Nayyeri, H., Hemmati, S., Mobasher, B., et al. 2017, *ApJS*, **228**, 7
- Newman, J. A., Abate, A., Abdalla, F. B., et al. 2015, *ApJ*, **63**, 81
- Newman, J. A., Cooper, M. C., Davis, M., et al. 2013, *ApJS*, **208**, 5
- Papovich, C., Momcheva, I., Willmer, C. N. A., et al. 2010, *ApJ*, **716**, 1503
- Peng, Y.-j., Lilly, S. J., Kovač, K., et al. 2010, *ApJ*, **721**, 193
- Prescott, M. K. M., Impey, C. D., Cool, R. J., & Scoville, N. Z. 2006, *ApJ*, **644**, 100
- Rhodes, J., Nichol, R. C., Aubourg, É., et al. 2017, *ApJS*, **233**, 21
- Salvato, M., Hasinger, G., Ilbert, O., et al. 2009, *ApJ*, **690**, 1250
- Salvato, M., Ilbert, O., Hasinger, G., et al. 2011, *ApJ*, **742**, 61
- Schneider, M. 2014, NASA WPS Proposal, [14-WPS14-4](#)
- Schreiber, C., Pannella, M., Elbaz, D., et al. 2015, *A&A*, **575**, A74
- Scoville, N., Arnouts, S., Aussel, H., et al. 2013, *ApJS*, **206**, 3
- Scoville, N., Aussel, H., Benson, A., et al. 2007a, *ApJS*, **172**, 150
- Scoville, N., Aussel, H., Brusa, M., et al. 2007b, *ApJS*, **172**, 1
- Silverman, J. D., Kashino, D., Sanders, D., et al. 2015, *ApJS*, **220**, 12
- Spergel, D. N., Verde, L., Peiris, H. V., et al. 2003, *ApJS*, **148**, 175
- Spitler, L. R., Labbé, I., Glazebrook, K., et al. 2012, *ApJL*, **748**, L21
- Springel, V., White, S. D. M., Jenkins, A., et al. 2005, *Natur*, **435**, 629
- Suh, H., Hasinger, G., Steinhardt, C., Silverman, J. D., & Schramm, M. 2015, *ApJ*, **815**, 129
- Szokoly, G. P., Bergeron, J., Hasinger, G., et al. 2004, *ApJS*, **155**, 271
- Takahashi, M. I., Shioya, Y., Taniguchi, Y., et al. 2007, *ApJS*, **172**, 456
- Tanaka, M., Hasinger, G., Silverman, J. D., et al. 2017, arXiv:[1706.00566](#)
- Trump, J. R., Impey, C. D., Elvis, M., et al. 2009, *ApJ*, **696**, 1195
- van der Wel, A., Noeske, K., Bezanson, R., et al. 2016, *ApJS*, **223**, 29
- Wang, T., Elbaz, D., Daddi, E., et al. 2016, *ApJ*, **828**, 56
- Yuan, T., Nanayakkara, T., Kacprzak, G. G., et al. 2014, *ApJL*, **795**, L20
- Zheng, W., Mikles, V. J., Mainieri, V., et al. 2004, *ApJS*, **155**, 73