



# An ALMA/*HST* Study of Millimeter Dust Emission and Star Clusters

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## Abstract

We present results from a joint ALMA/*HST* study of the nearby spiral galaxy NGC 628. We combine the *Hubble Space Telescope* (*HST*) Legacy ExtraGalactic UV Survey (LEGUS) database of over 1000 stellar clusters in NGC 628 with ALMA Cycle 4 mm/submillimeter observations of the cold dust continuum that span  $\sim 15$  kpc<sup>2</sup> including the nuclear region and western portions of the galaxy’s disk. The resolution— $1''.1$  or approximately 50 pc at the distance of NGC 628—allows us to constrain the spatial variations in the slope of the millimeter dust continuum as a function of the ages and masses of the nearby stellar clusters. Our results indicate an excess of dust emission in the millimeter, assuming a typical cold dust model for a normal star-forming galaxy, but little correlation of the dust continuum slope with stellar cluster age or mass. For the depth and spatial coverage of these observations, we cannot substantiate the millimeter/submillimeter excess arising from the processing of dust grains by the local interstellar radiation field. We detect a bright unknown source in NGC 628 in ALMA bands 4 and 7 with no counterparts at other wavelengths from ancillary data. We speculate this is possibly a dust-obscured supernova.

*Unified Astronomy Thesaurus concepts:* [Spiral galaxies \(1560\)](#); [Star clusters \(1567\)](#); [Interstellar medium \(847\)](#)

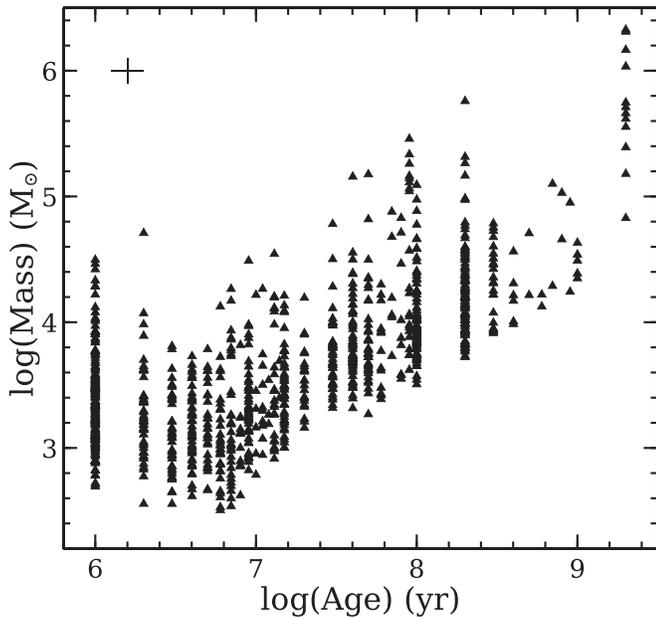
## 1. Introduction

Many galaxies are factories of current star formation that deplete the gas of the interstellar medium (ISM). Dense regions of gas condense into stars and star clusters which exhaust the galaxy of gas; in some instances, the gas is replenished through mergers or by the infall of external gas. Star formation, in turn, affects the evolution of galaxies by returning metals, energy, and momentum into the ISM and intergalactic medium. The cycling of gas into and out of galaxies regulates galaxy growth and its turbulence regulates star formation. To understand galaxy evolution, it is crucial to understand the physical processes that determine the evolution. Emission from stars and star clusters, along with the dust and gas of the ISM, all contribute to the shape of a galaxy’s spectral energy distribution (SED). By studying spatially resolved SEDs, we can begin to understand the small and large-scale properties of galaxies and the astrophysical processes underlying their evolution.

One unresolved question of using SEDs to study the star-forming environments of galaxies is an observed excess at millimeter/submillimeter wavelengths. This excess has been primarily seen in low-metallicity star-forming galaxies like the Magellanic Clouds (i.e., Galliano et al. 2005; Galametz et al. 2011; Planck Collaboration et al. 2011; Gordon et al. 2014; Izotov et al. 2014; Hermelo et al. 2016; Dale et al. 2017). Dust emission in the far-infrared regime of the SED is often described as a modified blackbody,  $S_\nu \propto B_\nu(T_{\text{dust}})\nu^\beta$ , where  $\beta$  is a measured effective grain emissivity parameter that empirically ranges between  $\sim 0.8$  and 2.5. This emissivity has been seen to vary significantly not only between galaxies

(Galametz et al. 2012), but also within galaxies (Kirkpatrick et al. 2014). In some cases, a modified blackbody extrapolated from far-infrared photometry cannot fully account for the emission beyond  $\sim 500 \mu\text{m}$ , hence the so-called “millimeter/submillimeter excess” (Galliano et al. 2003; Bot et al. 2010; Planck Collaboration et al. 2011). Proposed causes of this excess include very cold dust at  $T_{\text{dust}} < 7$  K (Galliano et al. 2005); fluctuations in the cosmic microwave background radiation as found in the LMC (Planck Collaboration et al. 2011); thermal free-free emission of ionized gas (Izotov et al. 2014); nonthermal emission from polycyclic aromatic hydrocarbons or nanoparticles (Lisenfeld et al. 2002; Meny et al. 2007; Coupeaud et al. 2011; Hensley & Draine 2017); and dust emissivity variations can be caused by processing of the dust that results in a flatter  $\beta$ , which we explore here (Gordon et al. 2014; Hermelo et al. 2016).

Kirkpatrick et al. (2014) study the spatial variations of  $\beta$  in the Key Insights on Nearby Galaxies: A Far-Infrared Survey with *Herschel* (KINGFISH) sample of nearby star-forming galaxies, and find a radial dependence for  $\beta$  which flattens toward the outskirts of the galaxies. Those authors test the possibility that  $\beta$  is modified by different heating sources, specifically young versus old stellar populations. The two stellar populations both contribute to the interstellar radiation field (ISRF) with the younger stars providing comparatively more ionizing photons than the less energetic photons from the old population. They find a correlation between  $\beta$  and the heating from the old stellar population as quantified by the ratio of luminosities at 3.6 and  $500 \mu\text{m}$ . This correlation is interpreted to mean the submillimeter slope is flattened due



**Figure 1.** Distribution of the ages and masses of the LEGUS classified star clusters in the central field of NGC 628. A typical error bar is given in the top left corner. Data from Grasha et al. (2015), Adamo et al. (2017).

to inefficient heating of the dust grains by the old stellar population. The inefficient heating could allow a cold dust emission component to exist at wavelengths beyond that of the dust emission peak, a scenario that will manifest as a shallower  $\beta$ .

The *Hubble Space Telescope* (*HST*) treasury program Legacy ExtraGalactic UV Survey (LEGUS) was designed to provide a novel catalog of star clusters in 50 nearby galaxies and their properties (see Calzetti et al. 2015). Given the *HST* resolution and proximity of the LEGUS targets, the galaxies have been resolved into their main stellar components: stars, star clusters, and associations. The LEGUS data set provides star clusters’ measured properties including age, mass, and color excess for a broad range of cluster populations across the LEGUS targets (see Adamo et al. 2017). In the LEGUS cluster catalog, the average uncertainty for both a cluster’s age and mass is 0.1 dex.

Grasha et al. (2019) make use of the rich LEGUS data set by exploring the spatial relation between star clusters and giant molecular clouds (GMCs) in the spiral galaxy NGC 5194 (M51). The authors find a spatial correlation between young ( $\leq 10$  Myr) star clusters and the GMCs which gives a timescale for disassociation of the star clusters from the GMCs of  $\sim 4$ –6 Myr. A similar study from Matthews et al. (2018) compares *HST* and ALMA observations and determines that by  $10^{6.7}$  yr, some star clusters will have lost all of their molecular material from which they were formed. In this study, we look to link the star clusters, not with their associated gas, but with the dust leftover from the cluster formation.

In this work, we combine the LEGUS star cluster catalog with dust continuum observations from the Atacama Large Millimeter/Submillimeter Array (ALMA) of the nearby spiral galaxy NGC 628 (M74). NGC 628 is nearly face-on ( $i \sim 25^\circ$ ), located at a distance of 9.9 Mpc (Calzetti et al. 2015), and provides an excellent testbed for these unprecedented ALMA observations with the wealth of LEGUS star clusters and ancillary data available. Figure 1 shows the distribution of the

star cluster ages and masses provided by LEGUS in NGC 628. We make use of the LEGUS cluster catalog with the averaged aperture corrections, Milky Way extinction, and Padova-AGB stellar evolution track.<sup>10</sup> By combining the high-resolution ALMA dust continuum maps with the star cluster data from *HST*, we can begin to study how the local ISRF generated by the star clusters may be affecting the dust emissivity.

In Section 2, we describe the new observations from ALMA. In Section 3, we review the analysis carried out, and in Section 4, we present the essential results and discuss their implications. Finally, in Section 5, we summarize our findings.

## 2. ALMA Observations

The ALMA Cycle 4 observations in band 7 (343 GHz, 0.87 mm) and band 4 (145 GHz, 2.1 mm) were carried out in 2016–2017 (ID = 2016.1.01435.S3, PI = D. Dale). Our band 7 observations consisted of 137 pointings with 43 12 m antennas with a baseline ranging from 15.1 m to 2.6 km. The phase and amplitude were calibrated using sources J0006-0623 and J0121+1149. Our band 4 observations consisted of 23 pointings with 44 12 m antennas with a baseline ranging from 18.6 m to 1.1 km. The phase and amplitude were calibrated using sources J0006-0623, J0139+1753, and J0238+1636.

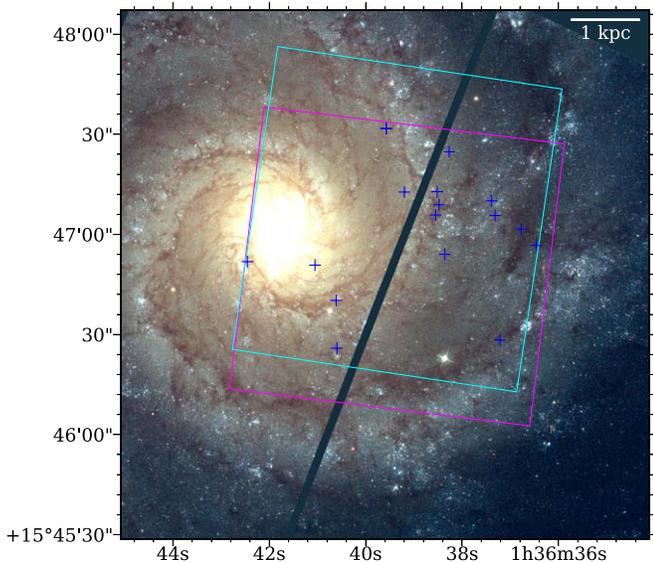
The data were reduced using the Common Astronomy Software Applications (CASA) package, version 5.4.0-68 (McMullin et al. 2007). Images were constructed and cleaned using the CASA task `TCLEAN` in the mosaic imaging mode. A Briggs weighting is used with the “robust” parameter set to 2.0 to increase the sensitivity to extended emission at the cost of angular resolution. A  $u$ - $v$  taper of  $1''.2$  and  $1''.0$  is applied to the band 7 and band 4 images, respectively. This gives a restoring beam size of  $1''.12 \times 1''.04$  with position angle  $39^\circ.73$  for band 7 with a maximum recoverable scale of  $7''.2$ . For band 4, the restoring beam is  $1''.12 \times 1''.08$  with position angle  $29^\circ.63$  and a maximum recoverable scale  $13''.8$ . The angular resolution of  $1''.1$  corresponds to a spatial resolution of about 50 pc assuming a distance of 9.9 Mpc to NGC 628 (Calzetti et al. 2015). At these resolutions, the  $1\sigma$  sensitivity is  $225 \mu\text{Jy } \text{bm}^{-1}$  for band 7 and  $31.7 \mu\text{Jy } \text{bm}^{-1}$  for band 4. Continuum maps were generated using all four spectral windows (SPWs 17, 19, 21, 23) in each band which were then corrected for the primary beam attenuation and used for the analysis in this paper. The sensitivity was measured on the images before the primary beam correction.

In both bands, a “footprint” is defined as  $\geq 80\%$  of the primary beam coverage. Figure 2 shows rectangular approximations for the two footprints overlaid on a 3-color image using *HST* LEGUS observations. For the analysis carried out here, we focus on the region where the footprints overlap which corresponds to about  $1.8 \text{ arcmin}^2$  or  $15.1 \text{ kpc}^2$ . This area encompasses the center of the galaxy out to  $0.27 R_{25}$  ( $\sim 4 \text{ kpc}$ ), where  $R_{25}$  is the traditional  $25 \text{ mag arcsec}^{-2}$  isophotal radius.

## 3. Analysis

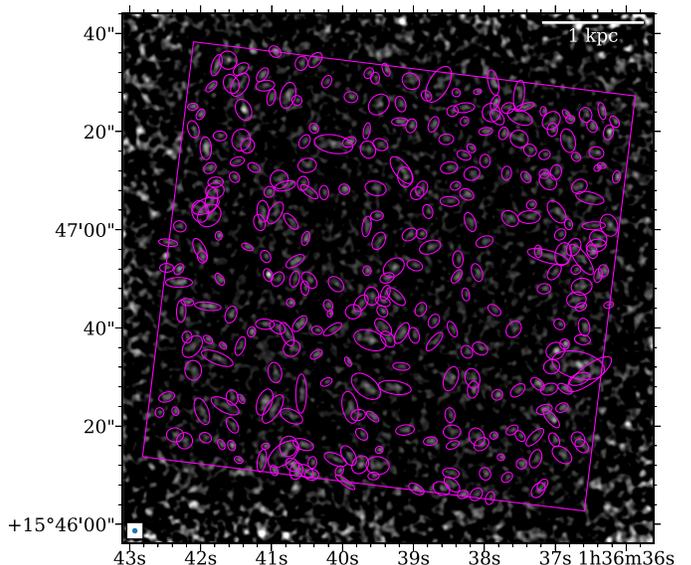
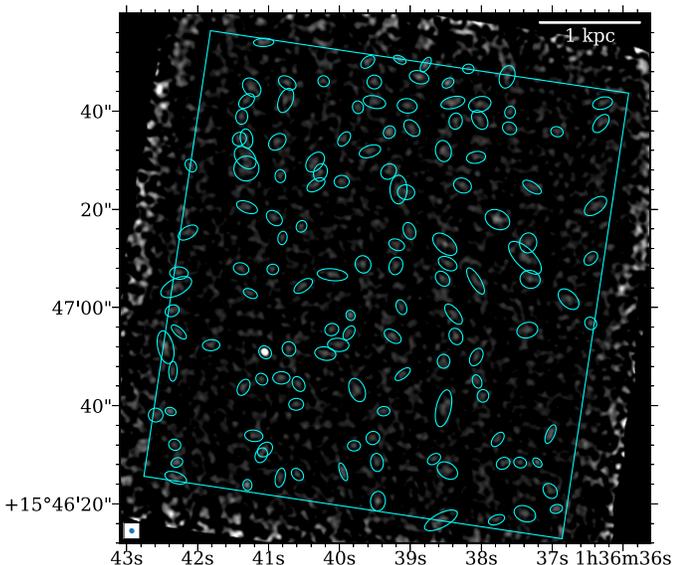
Millimeter sources for each band are identified using SExtractor (Bertin & Arnouts 1996) with the requirement that each source must have 50 contiguous pixels with flux greater than  $2\sigma$  above the background. With the pixel scale of  $0''.06$  per pixel, 50 pixels is about  $420 \text{ pc}^2$ . Only the sources, highlighted

<sup>10</sup> The star cluster catalog is available at [legus.stsci.edu](http://legus.stsci.edu).



**Figure 2.** Three-color image of NGC 628 using *HST* F435W, F555W, and F814W observations from LEGUS. Overlaid are the 80% primary beam coverage footprints from the ALMA band 7 (cyan) and band 4 (magenta) maps. The center locations of the 16 mm sources as shown in Table 1 are marked by the blue crosses.

by the SExtractor ellipses, that lie within each image’s 80% footprint are kept as shown in Figure 3. The photometry on each millimeter source yields flux-to-sigma-flux ratios inside the SExtractor ellipses of greater than 7; Figure 4 shows the distribution of this ratio for the sources. A noise for each ellipse is estimated by randomly placing ellipses with the same size and shape as the source found by SExtractor and calculating the standard deviation of the ensemble fluxes for that particular ellipse. A new signal-to-noise ratio (S/N) is then calculated by dividing a source’s flux by this noise estimate; the distribution of which is shown in Figure 4. In combining the source identification results from each band, we restrict the final sample to have source centroids in one band fall within the beam size of sources detected in the other band. This yields a final tally of 16 dual-band sources, as shown in Figure 5, with



**Figure 3.** ALMA band 7 (0.87 mm) continuum map is given on the left. The 80% footprint is outlined with the box and sources identified using SExtractor are shown. The beam size is given in the lower left corner. The ALMA band 4 (2.1 mm) continuum map is given on the right with the footprint and sources shown.

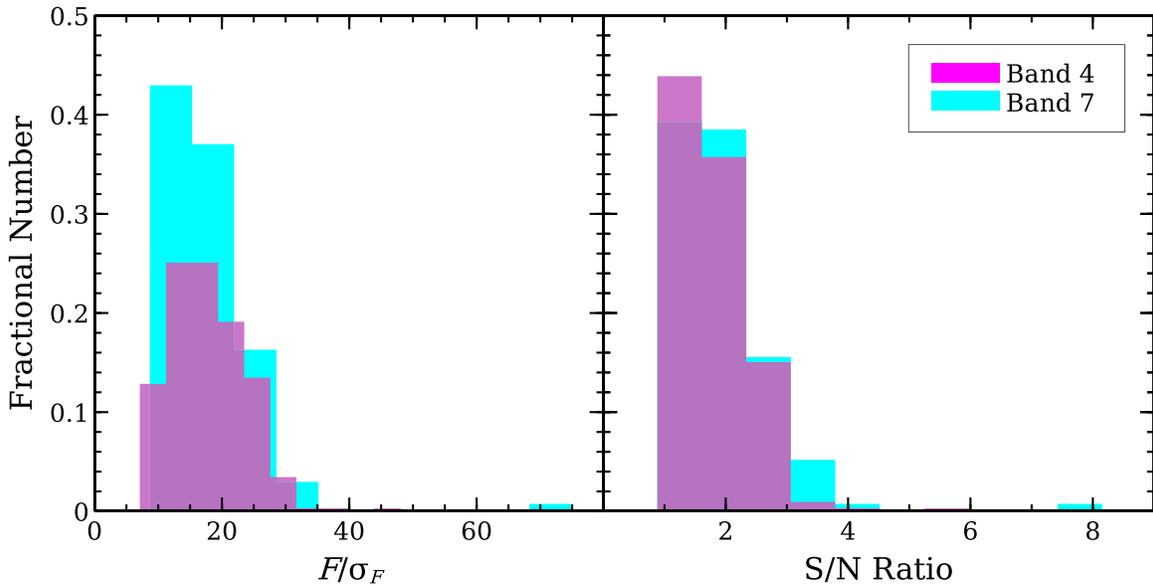
the location and fluxes given in Table 1. Given the density of sources found by SExtractor, there is a possibility of chance alignment between each image. To quantify this, we randomly place our same SExtractor identified source apertures across each image, count how many overlap, and repeat this process 10,000 times. We find an average of 15 uncorrelated sources randomly overlap, with a standard deviation of four sources.

Figure 5 also shows the position of the LEGUS star clusters in NGC 628. We identify the three LEGUS star clusters that are projected to lie closest to each millimeter source. Because the *Gaia* mission (Gaia Collaboration et al. 2016) provides superior astrometry, we recomputed the projected separations using the *Gaia* astrometry for the star cluster centroids (i.e., the *HST* LEGUS images were redrizzled onto the *Gaia* reference frame). The improved astrometry resulted in changes of only  $\sim 0''.1$  in decl. and  $\sim 0''.01$  in R.A., which do not significantly affect the results presented in Section 4. Additionally, we use the fluxes obtained from SExtractor to compute the millimeter/submillimeter slope via

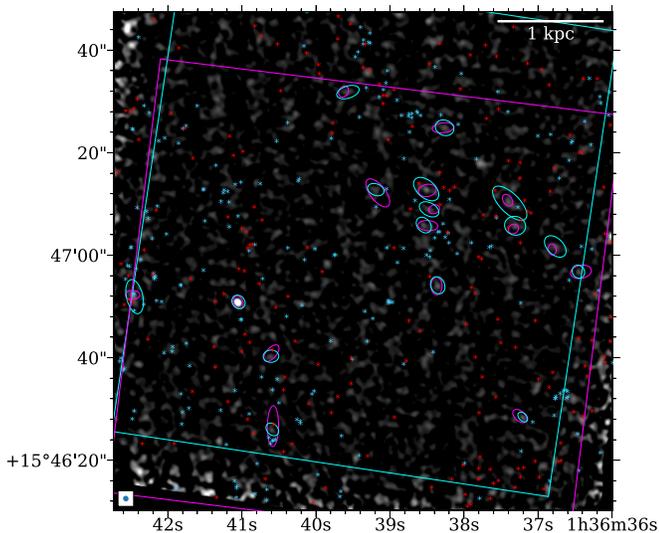
$$S(0.87 \text{ mm}/2.1 \text{ mm}) = \frac{\Delta \log_{10} F (\text{Wm}^{-2})}{\Delta \log_{10} \lambda (\text{m})}. \quad (1)$$

#### 4. Results and Discussion

Figure 6 shows the distribution of the projected physical separations between the 16 overlapping millimeter sources and their associated nearby star clusters. The star clusters are split into two age bins: younger than 10 Myr and older than 10 Myr. The average separation between the 16 mm sources and star clusters is  $\sim 161 \pm 21$  pc and  $\sim 172 \pm 13$  pc, respectively, for the younger and older clusters. Because all the sources found in band 7 have high S/N, we include the separations between these sources and their nearby star clusters in Figure 6 as the shaded regions. It is possible we detect these sources in band 7, but not in band 4, because their flux falls below the detection limit at the longer wavelength. To quantify this, we adopt the average dust continuum slope of  $-3.92$  and find 57 of the 135 band 7 sources would be above the band 4 detection limit of  $31.7 \mu\text{Jy} \text{bm}^{-1}$ . Using the slope of the brightest source (Section 4.1), we find only nine of the 135 sources would be



**Figure 4.** Distribution of the flux to the error in flux ratio for each millimeter source found in bands 4 (319 sources, magenta) and 7 (135 sources, cyan). Ratios range from  $\sim 7$  to 35. The brightest source, discussed in Section 4.1, has a  $F/\sigma_F$  of  $\sim 31$  in band 4 and 75 in band 7. The distribution of the total S/N for each millimeter source is given on the left. The S/N for band 4 ranges from  $\sim 0.6$  to 5.4 with a mean of 1.8. For band 7, the S/N ranges from  $\sim 0.9$  to 8.2 with a mean of 1.9.



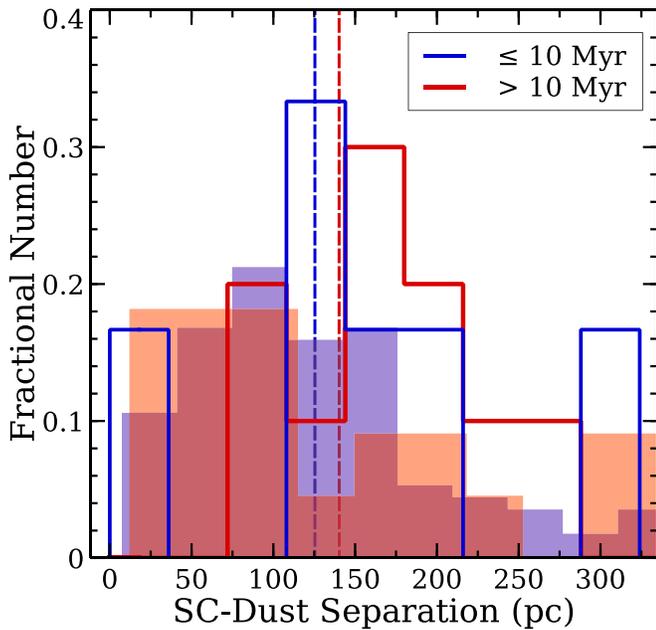
**Figure 5.** ALMA band 7 (0.87 mm) map is shown in grayscale with the 16 overlapping sources identified (band 7 ellipses shown in cyan and band 4 ellipses shown in magenta). The 80% primary beam footprints (band 7 in cyan and band 4 in magenta) are also overlaid to show the  $1.8 \text{ arcmin}^2$  ( $15.1 \text{ kpc}^2$ ) region where the footprints overlap. The position of the LEGUS star clusters are shown as star symbols differentiated according to age. Blue stars represent star clusters 10 Myr and younger. Red stars represent star clusters older than 10 Myr.

bright enough to detect in band 4. The new average separation between the millimeter sources and the star clusters is  $\sim 125 \pm 7 \text{ pc}$  and  $\sim 140 \pm 22 \text{ pc}$ , respectively, for the younger and older clusters. This slight difference in projected separations between young and old star clusters, albeit statistically insignificant, still echoes a similar result by Grasha et al. (2019) for the spiral galaxy NGC 5194: the younger LEGUS star clusters have, on average, smaller separations from the nearest GMCs than do the older star clusters. Again, given the uncertainties in these average separations, there is no statistical difference between the typical projected distance to old and young star clusters.

**Table 1**  
Location and Fluxes of the 16 Overlapping Millimeter Sources

R.A. (J2000)	Decl. (J2000)	$\log_{10}F(0.87 \text{ mm})$ ( $\text{W m}^{-2}$ )	$\log_{10}F(2.1 \text{ mm})$ ( $\text{W m}^{-2}$ )
1:36:36.45	+15:46:56.77	$-17.52 \pm 0.06$	$-18.54 \pm 0.03$
1:36:36.76	+15:47:01.69	$-17.09 \pm 0.05$	$-18.98 \pm 0.04$
1:36:37.20	+15:46:28.41	$-17.50 \pm 0.05$	$-18.45 \pm 0.03$
1:36:37.31	+15:47:05.73	$-17.05 \pm 0.05$	$-18.89 \pm 0.04$
1:36:37.38	+15:47:10.12	$-16.84 \pm 0.05$	$-18.59 \pm 0.03$
1:36:38.26	+15:47:24.85	$-17.30 \pm 0.05$	$-18.68 \pm 0.04$
1:36:38.35	+15:46:54.10	$-17.29 \pm 0.05$	$-18.79 \pm 0.04$
1:36:38.47	+15:47:08.93	$-17.33 \pm 0.05$	$-19.06 \pm 0.05$
1:36:38.51	+15:47:12.88	$-17.02 \pm 0.05$	$-18.72 \pm 0.04$
1:36:38.54	+15:47:05.86	$-17.38 \pm 0.05$	$-18.68 \pm 0.04$
1:36:39.19	+15:47:12.78	$-17.32 \pm 0.05$	$-18.36 \pm 0.03$
1:36:39.57	+15:47:31.80	$-17.29 \pm 0.05$	$-19.05 \pm 0.05$
1:36:40.59	+15:46:25.99	$-17.58 \pm 0.06$	$-18.47 \pm 0.03$
1:36:40.60	+15:46:40.27	$-17.49 \pm 0.06$	$-18.45 \pm 0.03$
1:36:41.04	+15:46:50.86	$-16.69 \pm 0.04$	$-18.47 \pm 0.03$
1:36:42.45	+15:46:51.90	$-16.88 \pm 0.05$	$-19.05 \pm 0.06$
	Sum	$-15.95 \pm 0.01$	$-17.45 \pm 0.01$

Finally, we track the millimeter continuum emission with the nearby star cluster properties. Figure 7 gives the 16 sources' millimeter/submillimeter continuum slopes as a function of the nearest three star clusters' ages and masses averaged together. Overplotted in Figure 7 is the reference slope  $S(0.87 \text{ mm}/2.1 \text{ mm})_{\text{ref}} = -4.81$  provided by the dust model SED described in Section 4.1; slopes larger than this value imply a flatter dust emissivity than the reference dust template. Thirteen of the 16 sources exhibit millimeter/submillimeter slopes larger than this reference value. Also plotted is the slope of the sum of all 16 sources at  $S(0.87 \text{ mm}/2.1 \text{ mm})_{\text{sum}} = -3.92 \pm 0.3$ , which implies a small but real excess for the sources. The uncertainty on this slope is approximated by the difference between  $S(0.87 \text{ mm}/2.1 \text{ mm})_{\text{sum}}$  and the slope  $S(0.87 \text{ mm}/2.1 \text{ mm})_{\text{stack}}$  obtained after stacking all 16 sources with  $30''$  cutouts and



**Figure 6.** Histogram of the physical separations between the 16 overlapping millimeter sources and their nearby star clusters shown by the solid lines. The histogram of separations between all millimeter sources found in band 7 and the nearby star clusters is shown as the shaded regions. Star clusters younger than 10 Myr are given in blue and clusters older than 10 Myr are shown in red. The average star cluster-millimeter source separations, for all band 7 sources, are given as colored dashed lines. The older star clusters, on average, lie slightly further from the dust clouds which is in agreement with the results for star clusters and giant molecular clouds in NGC 5194 from Grasha et al. (2019).

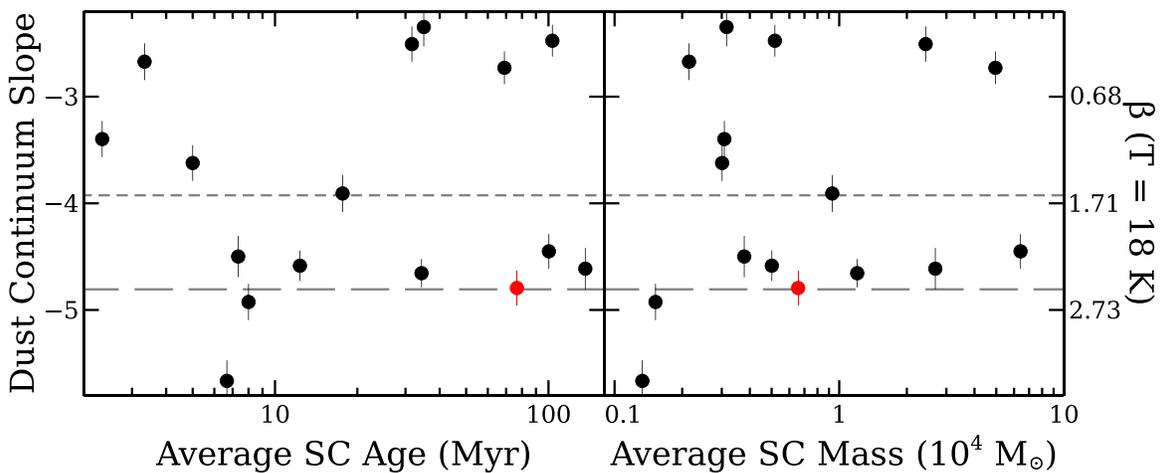
extracting the photometry from the 0.87 mm and 2.1 mm stacks using the same approach outlined in Section 3. There is one caveat to this excess measurement: recent laboratory measurements of dust grain emissivity show certain dust compositions can lead to a flatter millimeter/submillimeter slope than is expected by modified blackbody models (see e.g., Demyk et al. 2017a, 2017b).

We find no correlation with either star cluster mass or star cluster age, the latter being a proxy for the hardness and/or intensity of the local ISRF. This null result implies the radiation

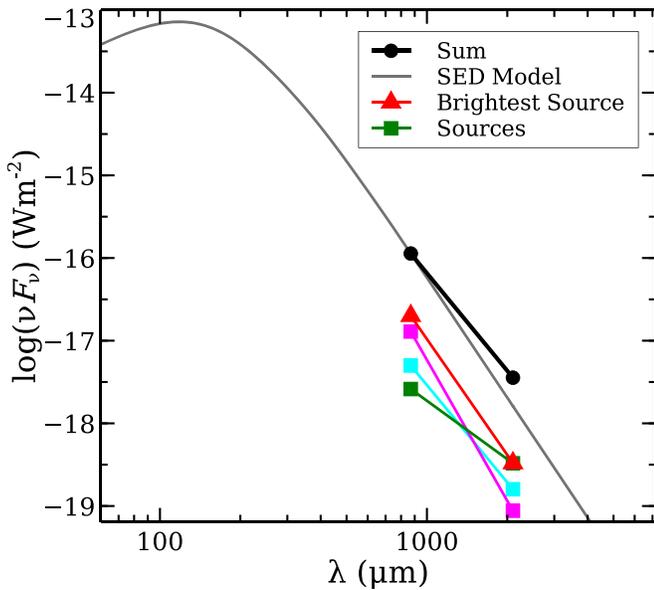
fields produced by the nearest star clusters do not directly affect the shape of the millimeter/submillimeter continuum through the processing of dust grains and their emissivity properties. However, two provisos should be clarified. First, it is possible that more sensitive millimeter/submillimeter continuum observations that detect many more sources could yield different results. This is possibly due to the dust associated with the star clusters having already been dispersed by the time we can observe the clusters at optical wavelengths. Second, as noted in the 1, millimeter/submillimeter excesses are primarily observed in low-metallicity environments. Our observational footprint only extends to a galactocentric distance of  $\sim 4$  kpc, where the H II region metal abundance for NGC 628 is still approximately solar (Moustakas et al. 2010).

#### 4.1. Bright Unknown Source

The brightest source in our ALMA maps is found at  $1^{\text{h}}36^{\text{m}}41^{\text{s}}.04$ ,  $+15^{\circ}46'50.86''$  (J2000). It is only found in our ALMA bands 7 and 4 observations; there is no emission found in archival *GALEX* ultraviolet (observed 2003), *HST* optical (observed 2013), *HST* near-infrared (observed 2005), ground-based H $\alpha$  (observed 2001), *Spitzer* mid-infrared (observed 2004), archival ALMA bands 6 and 3 continuum maps (observed 2013 and 2015), or CO (2–1) emission line maps (observed 2013). In order to determine if this source is a part of NGC 628 and not a background source like a quasar, we determine the most likely photometric redshift using the observed 0.87 mm and 2.1 mm photometry and the star-forming SED templates of Dale et al. (2014). The best match is a standard cool dust template SED at redshift zero. The particular template we adopt is the “ $\alpha = 2.5$ ” model, where  $\alpha$  is the exponent in a power-law distribution of different localized infrared SEDs. Fitting a modified blackbody to the “ $\alpha = 2.5$ ” model yields a blackbody temperature  $T = 23$  K and  $\beta = 1.8$ . Hence, given the information at hand, this millimeter source is likely within NGC 628. The Rayleigh–Jeans tail of this model dust SED is shown in Figure 8 with the the brightest source’s flux shown as red triangles. Also shown are three additional example sources to illustrate the range of dust continuum slopes we measure. The photometry for all 16 sources is summed and shown as black circles; the model SED



**Figure 7.** Dust continuum slope for each of the 16 sources as calculated in Equation (1) plotted as a function of the average star cluster age (left) and mass (right). Shown in red is the brightest source as discussed in Section 4.1. There is no correlation between the continuum slope and either star cluster age or mass. The long-dashed line designates the slope above which is an excess in dust millimeter emission as determined by the SED given in Figure 8. The short-dashed line marks the slope of the sum of all 16 sources at  $-3.92 \pm 0.3$ , which is also shown in Figure 8. See Section 4 for a description of how the uncertainty on the slope was estimated. Also given are the  $\beta$  values assuming a temperature of 18 K for a dust continuum slope of  $-3$ ,  $-4$ , and  $-5$ .



**Figure 8.** Rayleigh-Jeans tail of a galaxy SED model with cold dust ( $\alpha = 2.5$  from Dale et al. 2014) at a redshift of zero is shown in gray. The colored squares are a selection of three sources that illustrate the full range of dust continuum slopes measured in our sample. The brightest source, as discussed in Section 4.1, is shown as the red triangles. The black circles are the sum of the fluxes from all 16 sources. The model’s flux has been shifted to match the sum at 0.87 mm. This particular SED model was chosen because it best fit the slope of the brightest source in our sample. Warmer dust templates from Dale et al. (2014) did not fit as well. The sum shows indicates a slight excess at 2.1 mm.

has been scaled to match the flux of the sum at 0.87 mm. For the combined 16 sources, we find a slightly shallower slope compared to the model SED implying a perceptible millimeter/submillimeter excess over what is expected from the cool SED dust template.

We can only speculate on the nature of this peculiar source; it may be a compact and cold infrared dark cloud that is unresolved with our ALMA observations. If this was the case, we would expect to observe the source in the archival ALMA band 6 ( $\sim 1.2$  mm, 243 GHz) continuum map taken in 2013 with a comparable beam size of about  $1''$  and a sensitivity of  $850 \mu\text{Jy } \text{bm}^{-1}$ . Interpolating the source’s flux in band 7 down to the band 6 wavelength gives an expected flux of  $\sim 1800 \mu\text{Jy } \text{bm}^{-1}$ . The source should have been detected in the band 6 data if it existed at the time of the observation, but it is not detected. Therefore, we can rule out the infrared dark cloud explanation. It could alternatively have resulted from a supernova event triggered before the ALMA observations but after the collection of ancillary data mentioned above were taken (follow-up ground-based optical observations are planned). It is possible for this source to be a dust-obscured supernova, as it would only take an  $A_V$  on the order of a few magnitudes to dim the event enough to not be detectable in optical surveys (see e.g., Jencson et al. 2017). Ultimately, deeper imaging at many wavelengths coupled with sensitive spectroscopic information (e.g., the *James Webb Space Telescope* (*JWST*)) is needed to more fully understand this enigmatic source. However, if it is indeed a transient event, it may not still be there if observed later with *JWST*.

## 5. Summary

We present results of a joint ALMA/*HST* study of the millimeter/submillimeter emission covering  $\sim 15 \text{ kpc}^2$  of the

nucleus and disk of the nearby spiral galaxy NGC 628 in an attempt to understand how variations in dust emissivity may be affected by local star clusters. We detect 16 sources that spatial overlap in bands 4 and 7, and measure the slope of their dust continua over this 0.87–2.1 mm wavelength baseline. We note, however, that a significant fraction of these spatial overlaps may arise by chance. We find the closest star clusters, as given by LEGUS, to each millimeter source and our analysis suggests that younger star clusters lie closer to the dust than older clusters. This result is in agreement with Grasha et al. (2018) with ALMA-LEGUS observations of NGC 7793 and Grasha et al. (2019) in M51 who find younger star clusters lie closer on average to GMCs than the older clusters. Tracking the millimeter/submillimeter continuum slopes with the nearby star cluster ages and masses gives no correlation, suggesting that the hardness and/or intensity of local ISRF generated by the star clusters is not responsible for the flatter dust emissivity we observe. However, we have detected a relatively small number of millimeter/submillimeter sources, and mostly probed the nuclear and central disk regions of NGC 628 where the metallicity is approximately solar and there is a relative dearth of young massive star clusters (Shabani et al. 2018). It would be interesting to carry out more sensitive continuum observations further out in the disk where the metallicity is lower and any excess dust emission is more likely to appear.

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*Facilities:* ALMA, *HST* (ACS, UVIS).

*Software:* APLpy (Robitaille & Bressert 2012), Astropy (Astropy Collaboration et al. 2013; Price-Whelan et al. 2018), CASA (McMullin et al. 2007), NumPy (van der Walt et al. 2011), SExtractor (Bertin & Arnouts 1996).

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## References

- Adamo, A., Ryon, J. E., Messa, M., et al. 2017, *ApJ*, **841**, 131  
 Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, **558**, A33  
 Bertin, E., & Arnouts, S. 1996, *A&AS*, **117**, 393  
 Bot, C., Ysard, N., Paradis, D., et al. 2010, *A&A*, **523**, A20  
 Calzetti, D., Lee, J. C., Sabbi, E., et al. 2015, *AJ*, **149**, 51  
 Coupeaud, A., Demyk, K., Meny, C., et al. 2011, *A&A*, **535**, A124  
 Dale, D. A., Cook, D. O., Roussel, H., et al. 2017, *ApJ*, **837**, 90  
 Dale, D. A., Helou, G., Magdis, G. E., et al. 2014, *ApJ*, **784**, 83  
 Demyk, K., Meny, C., Leroux, H., et al. 2017a, *A&A*, **606**, A50  
 Demyk, K., Meny, C., Lu, X. H., et al. 2017b, *A&A*, **600**, A123  
 Gaia Collaboration, Prusti, T., de Bruijne, J. H. J., et al. 2016, *A&A*, **595**, A1  
 Galametz, M., Kennicutt, R. C., Albrecht, M., et al. 2012, *MNRAS*, **425**, 763  
 Galametz, M., Madden, S. C., Galliano, F., et al. 2011, *A&A*, **532**, A56  
 Galliano, F., Madden, S. C., Jones, A. P., et al. 2003, *A&A*, **407**, 159  
 Galliano, F., Madden, S. C., Jones, A. P., Wilson, C. D., & Bernard, J. P. 2005, *A&A*, **434**, 867  
 Gordon, K. D., Roman-Duval, J., Bot, C., et al. 2014, *ApJ*, **797**, 85  
 Grasha, K., Calzetti, D., Adamo, A., et al. 2015, *ApJ*, **815**, 93  
 Grasha, K., Calzetti, D., Adamo, A., et al. 2019, *MNRAS*, **483**, 4707  
 Grasha, K., Calzetti, D., Bittle, L., et al. 2018, *MNRAS*, **481**, 1016  
 Hensley, B. S., & Draine, B. T. 2017, *ApJ*, **836**, 179  
 Hermelo, I., Relaño, M., Lisenfeld, U., et al. 2016, *A&A*, **590**, A56  
 Izotov, Y. I., Guseva, N. G., Fricke, K. J., Krügel, E., & Henkel, C. 2014, *A&A*, **570**, A97  
 Jencson, J. E., Kasliwal, M. M., Johansson, J., et al. 2017, *ApJ*, **837**, 167  
 Kirkpatrick, A., Calzetti, D., Kennicutt, R., et al. 2014, *ApJ*, **789**, 130  
 Lisenfeld, U., Israel, F. P., Stil, J. M., & Sievers, A. 2002, *A&A*, **382**, 860  
 Mathews, A. M., Johnson, K. E., Whitmore, B. C., et al. 2018, *ApJ*, **862**, 147  
 McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, in *ASP Conf. Ser. 376, Astronomical Data Analysis Software and Systems XVI*, ed. R. A. Shaw, F. Hill, & D. J. Bell (San Francisco, CA: ASP), 127  
 Meny, C., Gromov, V., Boudet, N., et al. 2007, *A&A*, **468**, 171  
 Moustakas, J., Kennicutt, R. C. J., Tremonti, C. A., et al. 2010, *ApJS*, **190**, 233  
 Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2011, *A&A*, **536**, A17  
 Price-Whelan, A. M., Sipőcz, B. M., Günther, H. M., et al. 2018, *AJ*, **156**, 123  
 Robitaille, T., & Bressert, E. 2012, *APLpy: Astronomical Plotting Library in Python*, Astrophysics Source Code Library, version. 1.1.1, ascl:1208.017  
 Shabani, F., Grebel, E. K., Pasquali, A., et al. 2018, *MNRAS*, **478**, 3590  
 van der Walt, S., Colbert, S. C., & Varoquaux, G. 2011, *CSE*, **13**, 22