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Magnetic Connectivity of the Ecliptic Plane within 0.5 au: Potential Field Source Surface Modeling of the First *Parker Solar Probe* Encounter

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Abstract

We compare magnetic field measurements taken by the FIELDS instrument on board *Parker Solar Probe* (*PSP*) during its first solar encounter to predictions obtained by potential field source surface (PFSS) modeling. Ballistic propagation is used to connect the spacecraft to the source surface. Despite the simplicity of the model, our results show striking agreement with *PSP*'s first observations of the heliospheric magnetic field from ∼0.5 au (107.5 $R_\odot$) down to 0.16 au (35.7 $R_\odot$). Further, we show the robustness of the agreement is improved both by allowing the photospheric input to the model to vary in time, and by advecting the field from *PSP* down to the PFSS model domain using in situ *PSP*/Solar Wind Electrons Alphas and Protons measurements of the solar wind speed instead of assuming it to be constant with longitude and latitude. We also explore the source surface height parameter (R_{SS}) to the PFSS model, finding that an extraordinarily low source surface height (1.3–1.5 $R_\odot$) predicts observed small-scale polarity inversions, which are otherwise washed out with regular modeling parameters. Finally, we extract field line traces from these models. By overlaying these on extreme ultraviolet images we observe magnetic connectivity to various equatorial and mid-latitude coronal holes, indicating plausible magnetic footpoints and offering context for future discussions of sources of the solar wind measured by *PSP*.

Unified Astronomy Thesaurus concepts: Heliosphere (711); Solar corona (1483); Solar wind (1534); Solar magnetic fields (1503); Magnetic fields (994); Open source software (1866); Interplanetary magnetic fields (824)

1. Introduction

Parker Solar Probe (*PSP*; Fox et al. 2016) is a NASA mission intended to revolutionize our understanding of the solar corona by becoming the first spacecraft to measure its outer layers in situ. The fundamental science objectives are to (1) trace the flow of energy that heats and accelerates the solar corona and solar wind, (2) determine the structure and dynamics of the plasma and magnetic fields at the sources of the solar wind, and (3) explore mechanisms that accelerate and transport energetic particles.

Central to its science goals is *PSP*'s record-breaking orbit. *PSP* launched on 2018 August 12 and, after its first Venus gravity assist, entered into the closest-grazing heliocentric orbit ever reached by an artificial satellite. On 2018 November 6 *PSP* completed its first perihelion pass at 35.7 $R_\odot$ from the Sun traveling at almost 100 km s^{-1}. Future Venus gravity assists will eventually asymptote these numbers to 9.86 $R_\odot$ closest approach at over 200 km s^{-1} in 2024 December. A unique outcome of achieving this rapid orbital velocity is that *PSP* briefly reached a greater angular velocity than the Sun's equator. This means it moves very slowly relative to the local corotating magnetic structure and samples the same solar meridians multiple times in the same orbit (Figure 1(C)).

PSP carries a suite of four scientific instruments. The electromagnetic fields investigation (FIELDS; Bale et al. 2016) probes in situ electric and magnetic fields and plasma waves, the spacecraft potential, quasithermal noise, and low-frequency radio waves. The Solar Wind Electrons Alphas and Protons (SWEAP; Kasper et al. 2016) investigation provides distribution functions and density, velocity, and temperature moments of the most abundant species in the solar wind. The Integrated Science Investigation of the Sun (McComas et al. 2016) observes energetic electrons, protons and heavy ions from 10 s of keV to 100 MeV. The Wide-Field Imager for Solar Probe Plus (Vourlidas et al. 2016) is a white light imager that

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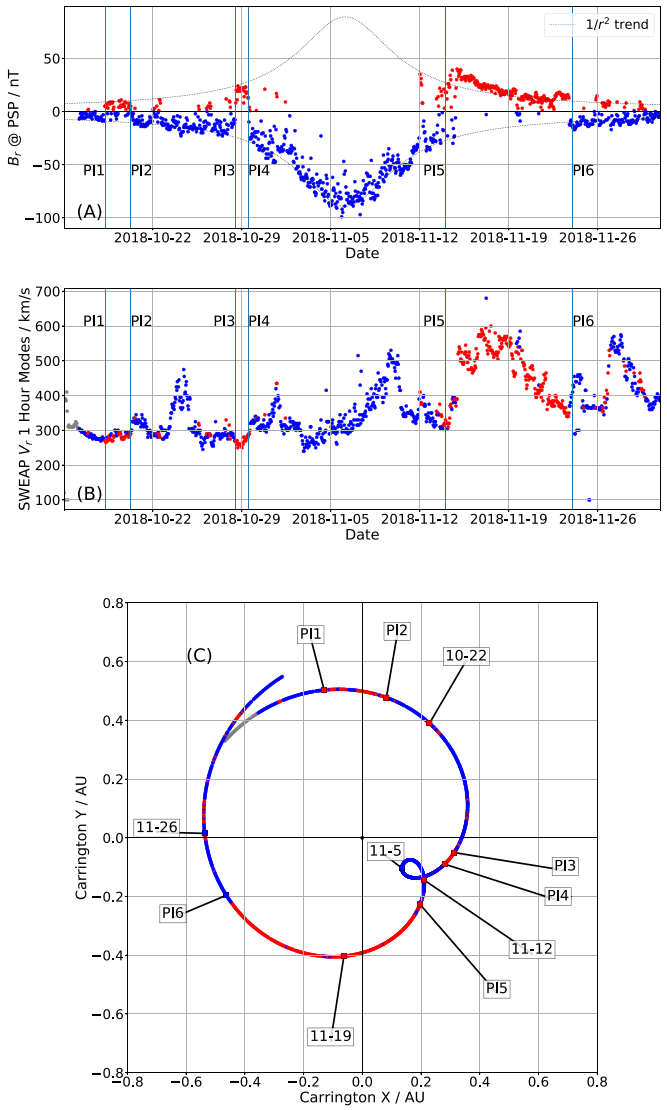

Figure 1. Exposition of 1 hr modes (see the main text) of *PSP* data taken during Encounter 1 (E1) from 2018 October 15 to 2018 November 30. Perihelion occurred on 2018 November 6. Panel (A): *PSP*/FIELDS radial magnetic field (B_r) versus time colored by magnetic polarity: positive (red) means radially outward, while negative (blue) means radially inward. A $1/r^2$ dotted trend line shows the zeroth-order behavior. Panel (B): *PSP*/Solar Wind Electrons Alphas and Protons radial proton bulk velocity (V_{SW}) vs. time colored by measured magnetic polarity. Panel (C): the E1 orbital trajectory of *PSP* plotted in Carrington (solar-corotating) coordinates. The trajectory is colored by the measured magnetic polarity, demonstrating the apparent spatial structure of the magnetic sectors sampled.

observes structures in the solar wind, such as shocks, approaching and passing the spacecraft. In this work we utilize FIELDS DC magnetic field data and SWEAP proton velocity moments.

A major source of contextual information for spacecraft in situ measurements of the solar wind comes from global coronal and heliospheric modeling. Modeling techniques of varying complexity (see, e.g., the review by Wiegelmann et al. 2017) have been developed using historical measurements (both remote and in situ) as boundary conditions. *PSP* provides unique constraints on such models given it is sampling entirely new regions of the heliosphere. It is therefore of great interest to compare *PSP* observations to these models both to contextualize the measurements and to improve the models themselves.

In this work we compare *PSP* magnetic field observations with predictions made using the widely used potential field source surface (PFSS) model (Altschuler & Newkirk 1969; Schatten et al. 1969; Hoeksema 1984; Wang & Sheeley 1992). PFSS employs two key assumptions: (1) a current-free corona, which is a special case of force-free models that require an assumption of very low plasma beta (meaning magnetic pressure dominates over thermal pressure); (2) a spherical source surface of heliocentric radius R_{SS} at which field lines are enforced to be radial, simulating the role of the solar wind in opening these field lines out to interplanetary space. Despite these assumptions PFSS compares well to more modern magnetohydrodynamic (MHD) models (e.g., Riley et al. 2006) and is widely used due to its computational simplicity and associated high resolution.

With zero currents the magnetic field, $\boldsymbol{B}$, in PFSS may be expressed as a scalar potential, Φ_B, such that $B = -\nabla\Phi_B$. By the "no-monopole" Maxwell equation, this potential must obey the Laplace equation, $\nabla^2\Phi_B = 0$, for which solutions are very well understood. PFSS solves for the field in an annular volume of radial extent $1R_\odot \leqslant r \leqslant R_{SS}$. Boundary conditions are the measured radial magnetic field at the photosphere ($1\,R_\odot$) and the requirement that the tangential components of $\boldsymbol{B}$ vanish at R_{SS}. This allows the problem to be uniquely solved as a spherical harmonic decomposition, resulting in a full 3D magnetic field model between the photosphere and source surface. The solution is steady state and represents a low-energy bound on more general force-free models (Régnier 2013).

PFSS models have historically been used to predict magnetic polarity at 1 au (e.g Hoeksema 1984), observe coronal structure during solar eclipses (e.g., Altschuler & Newkirk 1969), and identify open field line regions associated with coronal holes (e.g., Wang & Ko 2019). In addition, even though PFSS only models the magnetic field directly, Wang & Sheeley (1990) showed an inverse correlation between the divergence rate of PFSS field lines with observed solar wind speed at 1 au, indicating that coronal magnetic field topology plays an important role in the acceleration of the solar wind. This observation has since been refined by Arge & Pizzo (2000) and Arge et al. (2003, 2004) into the modern Wang–Sheeley–Arge (WSA) model which assimilates PFSS and a Schatten current sheet model (Schatten 1972), and is currently used operationally in space weather predictions and hosted by NASA's Coordinated Community Modeling Center (CCMC, http://ccmc.gsfc.nasa.gov).

Here, we report our results obtained from the use of a simple PFSS model and a ballistic propagation model (Nolte & Roelof 1973) to connect the spacecraft to the PFSS model domain, and use this to explain features of large-scale magnetic structure observed in *PSP*'s first solar encounter. In Section 2 we introduce the data taken by *PSP* used in this work as well as using them to extrapolate magnetic polarity structure out to 1 au and compare to measurements by the Magnetic Field Investigation (MFI; Lepping 1995) on board the *Wind* spacecraft (Harten & Clark 1995). In Section 3 the implementation of PFSS modeling and the procedure to connect those results to the in situ measurements of *PSP* are described. Section 4 lists the major results from this work: (1) general successful prediction of in situ timeseries measurements; (2) improvements to modeling through time-evolving magnetospheric inputs and use of *PSP*/SWEAP solar wind velocity measurements; (3) recovery of smaller-scale

structure consistent with measurements when the source surface height parameter (R_{SS}) of the PFSS model is dramatically lowered; (4) identification of mid-latitude and equatorial coronal holes as potential sources of the solar wind *PSP* measured in its first encounter. We conclude in Section 5 with a discussion of the results and interpretation with particular attention to address the limitations of this modeling method in light of its simplifications, and reference concurrent and future modeling work.

2. *PSP* Data

2.1. Time Series and Spatial Distribution

We begin by presenting the data used in this work from the FIELDS and SWEAP instruments on board *PSP* measured during the first solar encounter (E1) from 2018 October 15 to November 30. From FIELDS we use measurements of the radial component of the magnetic field (B_r) and from SWEAP we use the radial component of the proton velocity moment (V_{SW}). Since the focus of this work is the large-scale solar wind structure, we first pre-process these data to remove transients and rapid fluctuations such as the newly observed $\delta B/B \sim 1$ magnetic "switchbacks" and velocity spikes (see, e.g., Bale et al. 2019; Kasper et al. 2019; Dudok de Wit et al. 2020; Horbury et al. 2020). To do this we bin the full-cadence data into hourly segments, generate a histogram of the data in each bin, and take the modal value (the value corresponding to the peak of the histogram).

The resulting time series are shown in Figure 1. B_r is shown in panel (A), in which we identify a $1/r^2$ overall trend (dotted line) and colorize by polarity with the convention of red for radially outward ($B_r > 0$) and blue for radially inward ($B_r < 0$). This convention will be followed in all plots in this paper. The time of perihelion at $35.7\,R_\odot$ is easily identified by the occurrence of maximum field strength on 2018 November 6. The time series shows generally negative polarity indicating the orbital trajectory was mostly on the southward side of the heliospheric current sheet (HCS), and a number of clear multi-hour excursions into positive polarity on 2018 October 20 (bounded by PI1 and PI2), 2018 October 29 (bounded by PI3 and PI4), and from 2018 November 14 (PI5) through to 2018 November 23 (PI6). A detailed look at the nature of these inversions as *PSP* crossed the HCS is given in Szabo et al. (2020). The positive spike on 2018 November 13 is attributed to small coronal mass ejections as discussed in McComas et al. (2019), Korreck et al. (2020), Nieves-Chinchilla et al. (2020), Giacalone et al. (2020), and Mitchell et al. (2020).

V_{SW} is shown in panel (B) and colorized by the magnetic polarity for ease of comparison with the B_r time series. We observe a mostly slow wind ($<500\,\mathrm{km\,s^{-1}}$) which is generally uncorrelated with the magnetic polarity inversions, except for PI5 which is coincident with a sudden transition to a moderately fast wind stream peaking at about $600\,\mathrm{km\,s^{-1}}$.

Panel (C) of Figure 1 demonstrates the spatial distribution of magnetic polarities observed during *PSP*'s E1. We plot the spacecraft trajectory (projected on to the solar equatorial plane) in Carrington coordinates (i.e., corotating with the Sun's equator) and color it according the polarity of B_r. In this corotating frame, *PSP* starts in the upper left quadrant and tracks clockwise (i.e., retrograde). As its radial distance from the Sun decreases, its angular velocity catches up to that of the Sun, eventually reaching corotation prior to perihelion; it briefly rotates faster than the Sun before transitioning back to sub-corotational speeds as it climbs to higher altitudes. This results in the small loop centered near the 2018 November 5 date label. It should also be noted that this changing angular speed means the rate of change of longitude of *PSP* varied dramatically; two weeks out of the six week interval shown were spent between the longitudes of the corotation loop. The extent and location of the positive polarity inversions in the time series are identified as the red regions of the trajectory. Clearly the PI1–PI2 and PI3–PI4 intervals occurred over a small range of solar longitudes ($<10^\circ$), while the interval between PI5 and PI6 was protracted over a much larger region, spanning about 90° of solar longitude.

2.2. Parker Spiral Stream Structure and Comparison to 1 au Measurements

In the previous section, we showed the in situ data of B_r and V_{SW} measurements and their orbital context. The orbital context allows us to directly assign solar wind parameters along a narrow path through a complex 3D medium. In this section we seek to extrapolate these measurements to infer a solar wind stream structure to connect these data out to 1 au and compare it to *Wind*/MFI observations of the magnetic field polarity.

To accomplish this, we use the Parker spiral (Parker 1958) approximation in which we assume each point along the *PSP* trajectory is threaded by a Parker spiral field line with a curvature determined by the in situ solar wind speed, giving the spiral the following 2D parameterization of longitude and radius (ϕ, r):

$$\phi(r) = \phi_0 - \frac{\Omega_\odot}{V_{SW}}(r - r_0) \quad (1)$$

where ϕ_0, r_0 are the longitude and radial distance of *PSP* at the time of the V_{SW} measurement. $\Omega_\odot$ is the solar rotation rate which we calculate from the equatorial rotation period of 24.47 days, assuming the latitudinal offset of *PSP* ($<5^\circ$) is small enough to not consider differential rotation. This equation shows the qualitative dependence of the spiral field lines of solar wind speed: faster V_{SW} gives straighter field lines (smaller $d\phi/dr$), slower wind speed produces more curved field lines (larger $d\phi/dr$).

Each spiral field line is colored by *PSP* measured polarity. The result is plotted in Figure 2, again in Carrington coordinates. The date-labeled *PSP* trajectory is shown in the context of the stream structure out to 1 au where we plot the trajectory of the *Wind* spacecraft, located at the Earth L1 point, for a similar date range. At a 12 hr cadence, we plot the polarity of the radial magnetic field measured by the *Wind*/MFI instrument. Guided by a similar convention from Hoeksema et al. (1983), to assign this 12 hr interval a single polarity, we take all measurements at a minute cadence from ± 6 hr of the measurement time and assign a positive polarity (red) if more than 7/12 of all data values are positive, and negative polarity (blue) if more than 7/12 of all data values are negative. If neither of these criteria is satisfied, the field is designated "mixed" for that interval and colored gray.

The choice of generating field lines at a constant time interval means that, as the relative angular velocity of the spacecraft gets smaller, the field lines appear to bunch together more. We emphasize this is purely a sampling effect and does

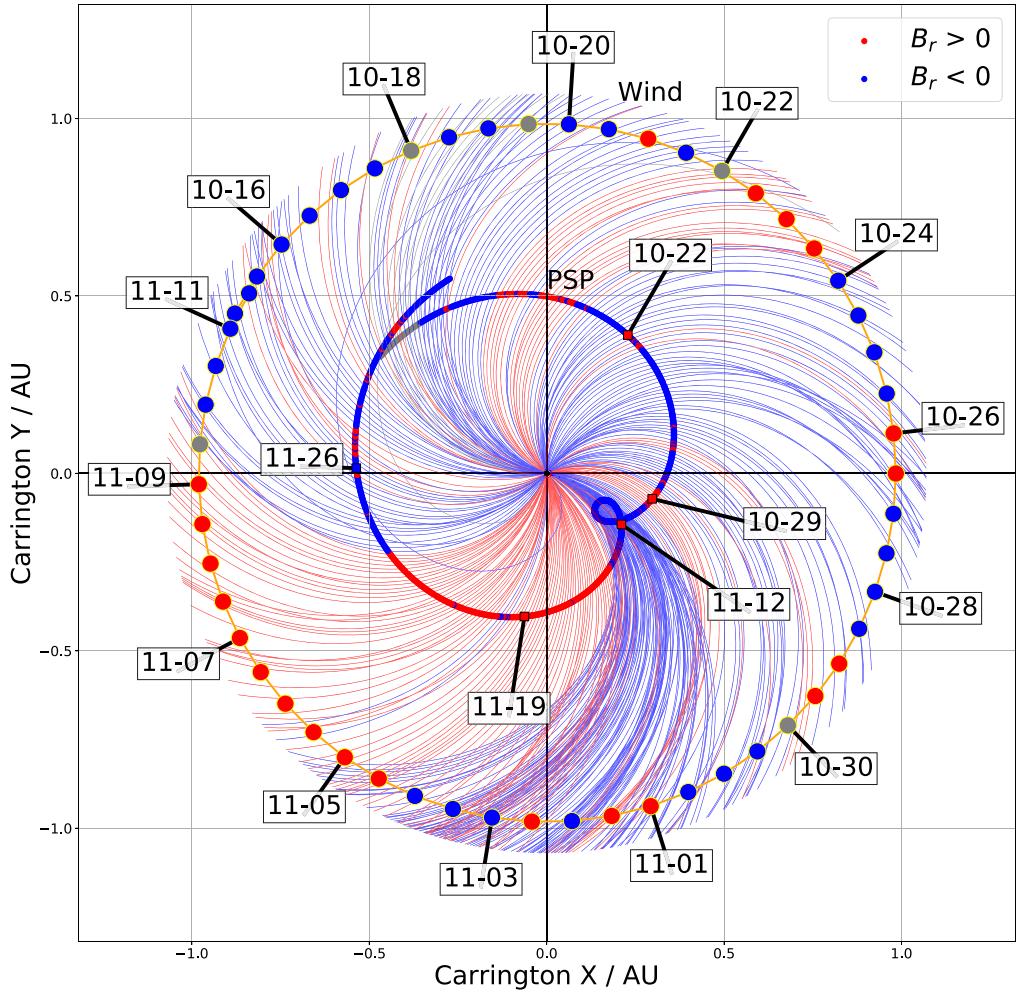

Figure 2. Magnetic polarity sector structure implied by the *PSP* measurements extrapolated to 1 au. Parker spiral field lines are initialized by the SWEAP V_{SW} measurements and colored by the FIELDS B_r polarity. These are propagated out to 1 au. Measurements of the magnetic polarity by *Wind*/MFI are shown at a 12 hourly cadence. For each time, positive polarity is designated if the field is >7/12 positive for ±6 hr of this time and similar for negative, otherwise the polarity is designated mixed and colored gray. As in Figure 1(C), the coordinates corotate with the Sun.

not indicate anything physical about the field. It is further compounded by the corotation loop which means several field lines are generated at the same Carrington longitude.

It is important to note that *PSP* and *Wind* sampled the same solar longitudes at different times due to the differing orbits. In displaying the picture in Figure 2, we are assuming the structure shown is essentially fixed in time as the spacecraft travel through it making measurements. Even so, we see the dominant features in *PSP* data bear out at 1 au using this simple picture. The PI1–PI2 interval connects via the spiral to positive polarity at 1 au. The PI3–PI4 interval merges into slowing wind which also mixes the negative polarity at perihelion with the onset of the positive polarity and fast wind during PI5. Although, as noted above, this region of the plot is made more complicated by the perihelion loop, the boundaries where field lines overlap further out than *PSP* represent likely locations where the Parker spiral assumption breaks down and field lines would bend into compressions or rarefactions. The more complex picture of interacting streams corresponds to a greater mix of positive and negative polarity measured as *Wind* sampled this region. The continuous period of positive polarity between PI5 and PI6 at *PSP* is also seen by *Wind*, with the Parker spiral very accurately predicting the transition from positive back to negative (PI6) by *Wind* on 2018 November 10, well in advance of *PSP* traversing the same structure on 2018 November 23.

This good correspondence between *PSP* and 1 au measurements suggests the polarity inversions observed at *PSP* are relatively stable large-scale features which must have an origin in open flux emerging from the solar corona. In addition, the implementation of the Parker spiral model lays the groundwork for the discussion in the next section on connecting *PSP* measurements inward to the corona where they can be compared to modeling results. As mentioned above, we also note that the stream structure implies magnetic pile-ups (and therefore divergence from the Parker spiral picture) further out than *PSP*, suggesting the locations of stream interaction regions that are studied in Allen et al. (2020) and Cohen et al. (2020). However, tracing the field lines inwards, the streams should interact and distort each other much less.

3. Modeling Method

Having introduced *PSP* data from E1, their spatial context, and salient features, we now introduce our PFSS modeling procedure and method of producing resulting time series

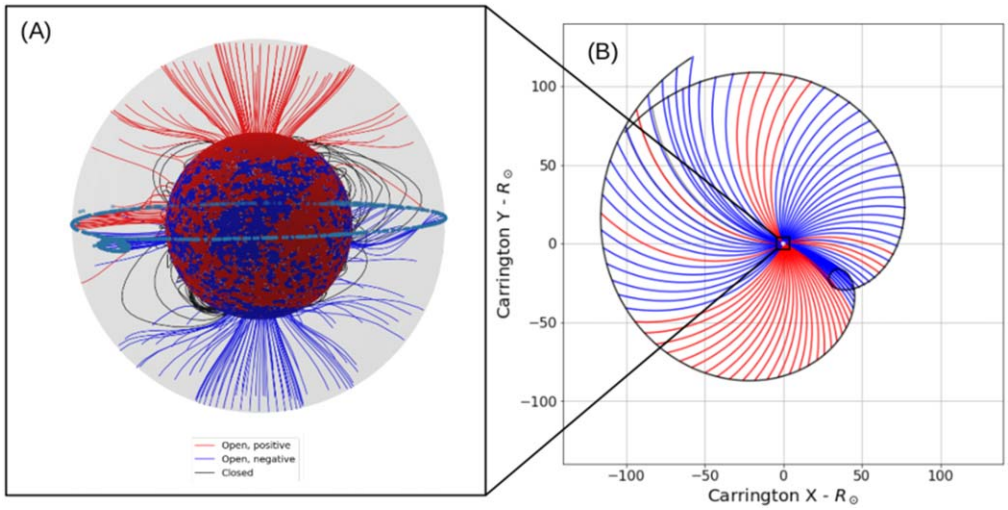

Figure 3. Schematic of process to connect *PSP* measurements to PFSS modeling. Panel (A): PFSS model output from *pfsspy*. The synoptic magnetogram input is shown as the photospheric (inner) boundary. The model domain is bounded at at the exterior by the source surface (gray surface). Field lines initialized by a uniform grid at the photosphere are shown. Panel (B): the outer boundary of the model is connected to the orbital position of *PSP* via an ideal Parker spiral magnetic field line. With some choice of solar wind speed, this maps the *PSP* trajectory to a locus of latitudes and longitudes at the source surface. This is illustrated by the near equatorial blue data points on the source surface in panel (A).

predictions at *PSP*. This procedure is schematically illustrated in Figure 3, and derives from Stansby et al. (2019a).

3.1. PFSS Implementation

The PFSS model is run in Python using the open source *pfsspy* package (Yeates 2018; Stansby 2019). This code is available online and is very flexible, fast, and easy to use. As input, it takes synoptic maps of the radial magnetic field at the photosphere and a value for the source surface height parameter R_{SS}. From these data it produces a full 3D magnetic field within the modeled volume, as well as a utility to trace individual magnetic field lines through the model solution. This input/output mapping is illustrated in Figure 3(A) which shows selected 3D field lines produced by *pfsspy*, the extent of the model, and the photospheric map that seeds the model. The gray surface illustrates the spherical source surface at which the field is constrained to be radial. As shown by the coloring of the field lines, they may either be open (red or blue depending on polarity) or closed (black). Open field lines are those which by definition intersect the source surface. Where they connect to the inner boundary of the model they indicate the probable locations of coronal holes (see Section 4.2.1). During solar minimum, most open field lines emerge from large polar coronal holes.

For the input magnetogram, there are a number of possible sources of data. In this work we initially considered the Global Oscillation Network Group (GONG; Harvey et al. 1996) zero-corrected data product (Clark et al. 2003), and the Heliospheric Magnetic Imager (HMI; Scherrer et al. 2012) vector magnetogram data product (Hoeksema et al. 2014). GONG is measured from a network of ground-based observatories and is operationally certified as input to a number of space weather prediction models. HMI, an instrument on board the *Solar Dynamics Observatory* (*SDO*; Pesnell et al. 2012), is higher resolution and does not suffer atmospheric effects. Both of these have the limitation that they rely purely on observations and so cannot account for evolution on the far side of the Sun until that part of the Sun rotates into view. With this in mind, we have also used the Air Force Data Assimilative Photospheric Flux Transport (ADAPT) modeled magnetogram (Arge et al. 2010) evaluated with GONG input, and the DeRosa/LMSAL modeled magnetogram (Schrijver & De Rosa 2003) (based on HMI data) to compare results. ADAPT and LMSAL make use of surface flux transport models into which new observations are assimilated. This procedure therefore models the far side evolution, implying a more accurate global picture of the photosphere. In practice, on discriminating between PFSS outputs from different magnetogram inputs using *PSP* data, we find little impact on our conclusions. We find GONG maps produce smooth predictions combining maps from one day to the next (Section 4.1), and require no pre-processing. While this smoothness may be a product of low resolution and atmospheric effects, it results in good clarity in displaying the features discussed in Section 4 without changing the conclusions. ADAPT maps resulted in very similar predictions but with some small fluctuation in the flux strength prediction from one day to the next which can be interpreted as model uncertainty. The HMI data include some missing days in the magnetogram record and does not include reconstruction of unobserved polar regions. The DeRosa/LMSAL model reconstructs the polar region and produces very similar predictions to the other models at lower source surface heights. For source surfaces much higher than $2.0\,R_\odot$ some deviation from the observations and other models takes place such as predicting constant positive field prior to October 29 (see Appendix A). This adds to the evidence we build in the results below that taking lower source surface heights in general is necessary for the best agreement between PFSS modeling and the observations.

In addition, in certain parts of this work, we use a model from a single date to represent the entire encounter (Figures 6(C), 7, 8(A)). The extra model evolution of ADAPT or the DeRosa model actually makes this presentation difficult since times earlier than the model evaluation have changed significantly and no longer agree with what *PSP* measured at that time. For

GONG, longitudes earlier in time are frozen after they go out of view and due to fortuitous orbital alignment the "older" parts of the model agree better with corresponding *PSP* measurements. For this reason, and in the absence of any strong effects on conclusions, PFSS model results shown in this paper use GONG magnetograms unless otherwise stated. Further discussion of this choice and comparisons of time series predictions using different magnetograms are included in Appendix A.

3.2. Ballistic Propagation Model

PFSS only models the coronal magnetic field out to the source surface at a couple of solar radii. *PSP* on the other hand made in situ measurements at radii down to a minimum of $35.7\,R_\odot$ during its first encounter. To connect the model domain outwards to *PSP*'s orbit we use ballistic propagation. Proposed by Nolte & Roelof (1973), this technique assumes an arbitrary point in the heliosphere can be connected inward to the corona by an ideal Parker spiral field line. The curvature of the spiral is driven by the solar wind speed measured out in the heliosphere. This model implicitly assumes that this measured wind speed (and therefore spiral curvature) remains constant all the way down into the corona. While this is not an accurate picture of the real solar wind, the dominant correction to this (acceleration of the solar wind) is counterbalanced by the effect of corotation, meaning that the real footpoint of the field line is in fact close to where the ballistic propagation assumption puts it. Nolte & Roelof (1973) conclude the error in longitude of this method is within 10°. The implications of this error are discussed in Section 4.2.2.

Much like in Section 2.2, we assign a series of Parker spiral field lines to the orbit of *PSP* but this time propagate them *inward* to derive a longitude and latitude on the source surface to which each position along the *PSP* trajectory is connected to. The spiral field lines initialized from the *PSP* orbit are shown in Figure 2(B) and the resulting locus of coordinates on the source surface is indicated by the near equatorial blue scatter points in panel (A). At each of these points on the source surface B_r is obtained from the model and then multiplied by the value r_{SS}^2/r_{PSP}^2 at that time to project this model field out to *PSP*. Finally, since the total unsigned magnetic flux calculated from different magnetograms generally do not agree with each other, rather than assume one particular magnetogram is correctly normalized, we scale our results by a constant factor so that their peak magnitude is equal to that measured by *PSP* at perihelion (Figure 1(A)). As shown in Figure 6 and discussed in Section 5, this factor is a function of source surface height and ranges from O(10) to order unity over the range of source surface heights considered in this work.

4. Results

We now present the results of comparison of the above modeling to *PSP* data.

4.1. Time Series Prediction and Comparison

In Figure 4, we overlay time series predictions using GONG magnetograms and a source surface height $R_{SS} = 2.0\,R_\odot$ (the choice of this value is discussed in Section 4.2). The model results are plotted as lines while the B_r data are plotted unchanged from Figure 1(A) as a scatterplot colored by polarity. We compare a number of modeling variations across the four panels.

1. In each panel a number of modeling realizations are plotted as faint colored lines. Each realization is generated using a GONG magnetogram from a different date spaced apart by 3 days over the encounter. We overplot a "time-integrated" model as a thick solid black line where these individual realizations are combined: for each faint line only the data from ±1.5 days from the date of the relevant magnetogram are used and stitched together sequentially. In terms of scaling, the models are stitched together first and then multiplied by a constant factor to match the peak measured B_r. For panels A and B this value is 6.90 while for panels C and D it is 6.73.
2. From left to right (A,C versus B,D) we multiply the prediction and data by r_{PSP}^2 to compare the models and data without the $1/r^2$ scaling, which dominates the overall shape of the time series in the raw data.
3. We demonstrate the impact of using the measured solar wind velocity (bottom panels, C and D) when generating Parker spirals to connect *PSP* to the source surface versus just assuming a constant slow solar wind speed of $360\,\mathrm{km\,s^{-1}}$ (top panels, A and B).

We observe an overall good agreement with a generally negative polarity field, predominantly varying with $1/r^2$ as predicted by the Parker spiral model, and with a significant excursion into positive polarity bounded by PI5 and PI6. However, we note here that PI1, PI2, PI3, and PI4 are lacking in these predictions. This matter is discussed further in Section 4.2.2 where we recover these polarity inversions through tuning of the source surface height parameter.

Examining the ensemble of models over different observations we see some scatter about the data, and in particular an unobserved "bump" in the field around 2018 November 4. This small disagreement with the data is mitigated by the time-integrated model which in all cases traces closer to the real data and ignores this unobserved bump. From this, we infer that the time evolution of the input magnetograms is important on the timescale over which *PSP* observations took place but treating it as a time series of steady state models mitigates the lack of dynamics in PFSS modeling.

Comparing the models with and without the $1/r^2$ scaling, we observe that the negative interval around perihelion infers a near constant (flat) magnetic field at the source surface. This is consistent with the fact that this time interval corresponds to times when *PSP* was corotating or moving very slowly compared to the solar surface and so was likely crossing flux tubes very slowly and observing a very slow change in the source region (see Section 4.3).

We also observe that using solar wind speed measurements to propagate field lines causes a subtle improvement to the model. For example, the timing of maximum field strength agrees better with the data, the relative amplitude of the peak compared to the trough improves, and the near perihelion field profile flattens in panel D versus B. The timing and steepness of the PI5 polarity inversion also improves.

Meanwhile, PI6 demonstrates a limitation of PFSS: the prediction is for a smooth, protracted transition through $\mathrm{Br} = 0$ whereas the data show a sharp transition. MHD modeling

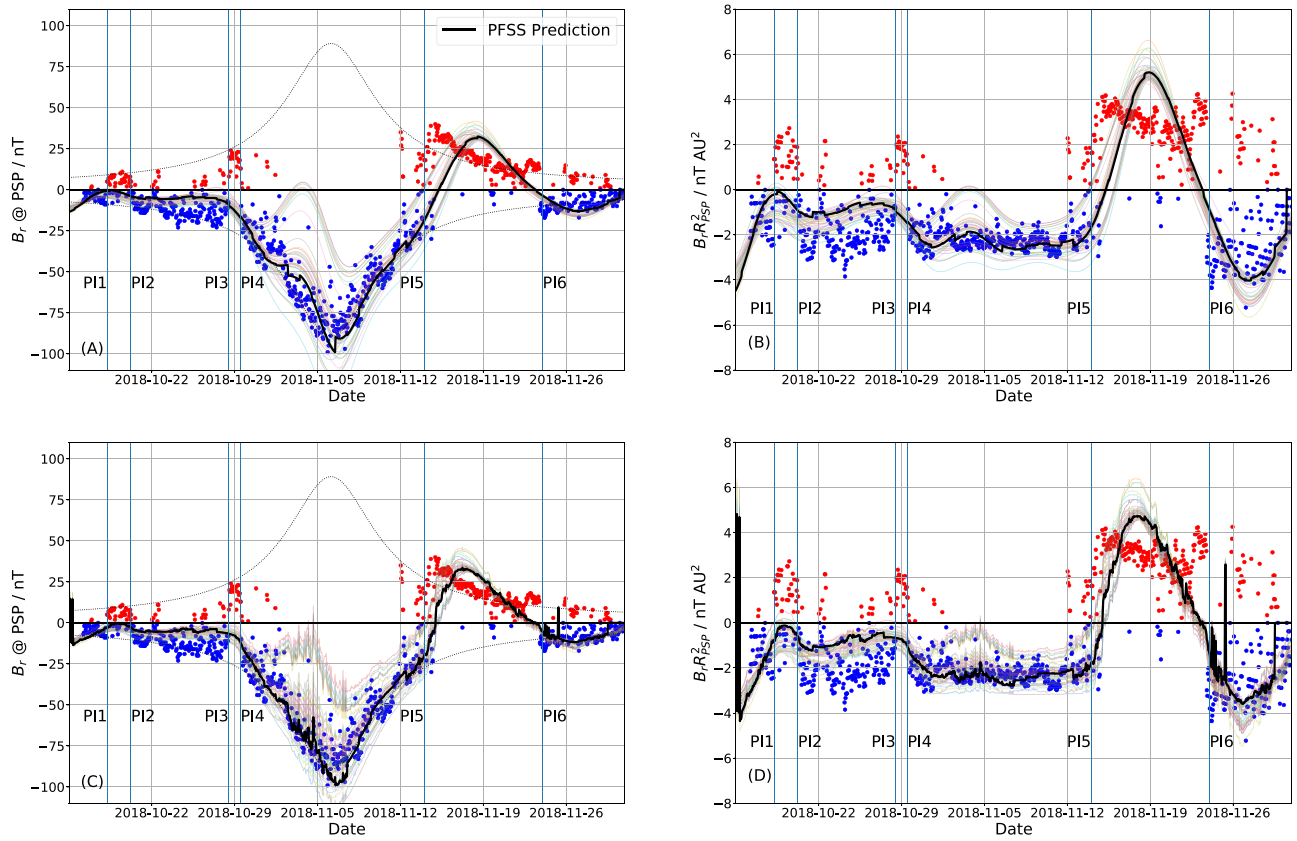

Figure 4. Comparison of PFSS predictions of B_r with observations. The left-hand column (A,C) shows predictions and measurements in situ at *PSP*, while the right-hand column (B,D) scales out the $1/r^2$ dependence to compare predictions near the source surface. The top row (A,B) shows predictions made assuming a constant wind speed of 360 km s^{-1}, while the bottom row (C,D) shows the results of using the SWEAP V_{SW} measurements to connect the source surface to *PSP*. The faint colored lines in each plot indicate an ensemble of predictions made using updating magnetograms at a cadence of 3 days. The solid black line indicates a synthesis of these predictions by combining segments from each prediction ±1.5 days from the date of that photospheric map. Models all use GONG data and a source surface height of $R_{ss} = 2.0\,R_\odot$.

(Réville & Brun 2017) has shown that further latitudinal evolution beyond the source surface sharpens the location of polarity inversions by homogenizing the radial field in latitude, which better matches *Ulysses* observations (Smith & Balogh 1995; Smith 2011). This could explain this discrepancy.

Additionally, even after time integration and correction for varying solar wind speed, the general negative field predicted prior to PI3 is approximately a factor of two weaker than measurements suggest at the source surface (panel D). This indicates the same A/r^2 scaling (with a constant A) of the PFSS-derived field magnitude is not globally applicable to the whole time series, particularly as *PSP* gets closer to the polarity inversion line (PIL; see Section 4.2.2) where PFSS predicts a drop-off in field strength. The sharpening effect of non-radial expansion outside the source surface is also a likely explanation for this discrepancy as it predicts the field strength remains constant closer to the PIl than PFSS suggests.

4.2. Impact of Variation of Source Surface Height Parameter

In this section, the impact on the results of varying the source surface height are investigated. While this is a numerical modeling parameter, it does have physical consequences on the predictions of the model affecting the total open flux, the apparent size of coronal holes, and the complexity of the PIL which seeds the HCS in other models such as WSA (Arge et al. 2003).

4.2.1. Coronal Hole Distribution

In Figure 5 we begin our discussion of the source surface height parameter by comparing the footpoints of open field lines to synoptic maps of extreme ultraviolet (EUV) emission from the solar corona. These data are assembled from data observed during Carrington rotation 2210 by *STEREO*/EUVI (Wuelser et al. 2004) observing at 193 Å and *SDO*/AIA (Lemen et al. 2012) observing at 193 Å, these emissions are produced by Fe XII, a highly ionized state of iron that is excited around 1000,000 K. The brightness of these maps indicates the density of plasma at the 1000,000 K isotherm which is approximately a surface of constant height in the lower corona. Dark regions identify coronal holes that are the probable locations of open field lines which allow plasma to escape outward into the heliosphere, resulting in an underdense region.

By comparing the locations of footpoints of open field lines implied by the PFSS model, we can evaluate how accurately the model reproduces the observed coronal hole distribution. We use model results using the GONG 2018 November 6 map (centered on perihelion) and vary the source surface from $2.5\,R_\odot$ down to $1.5\,R_\odot$. The open footpoints are generated by initializing field lines from a uniform grid at the source surface where the field lines are all open by construction. We then use the *pfsspy* field line tracing utility to propagate each line down into the model and find its point of intersection with the lower boundary of the model.

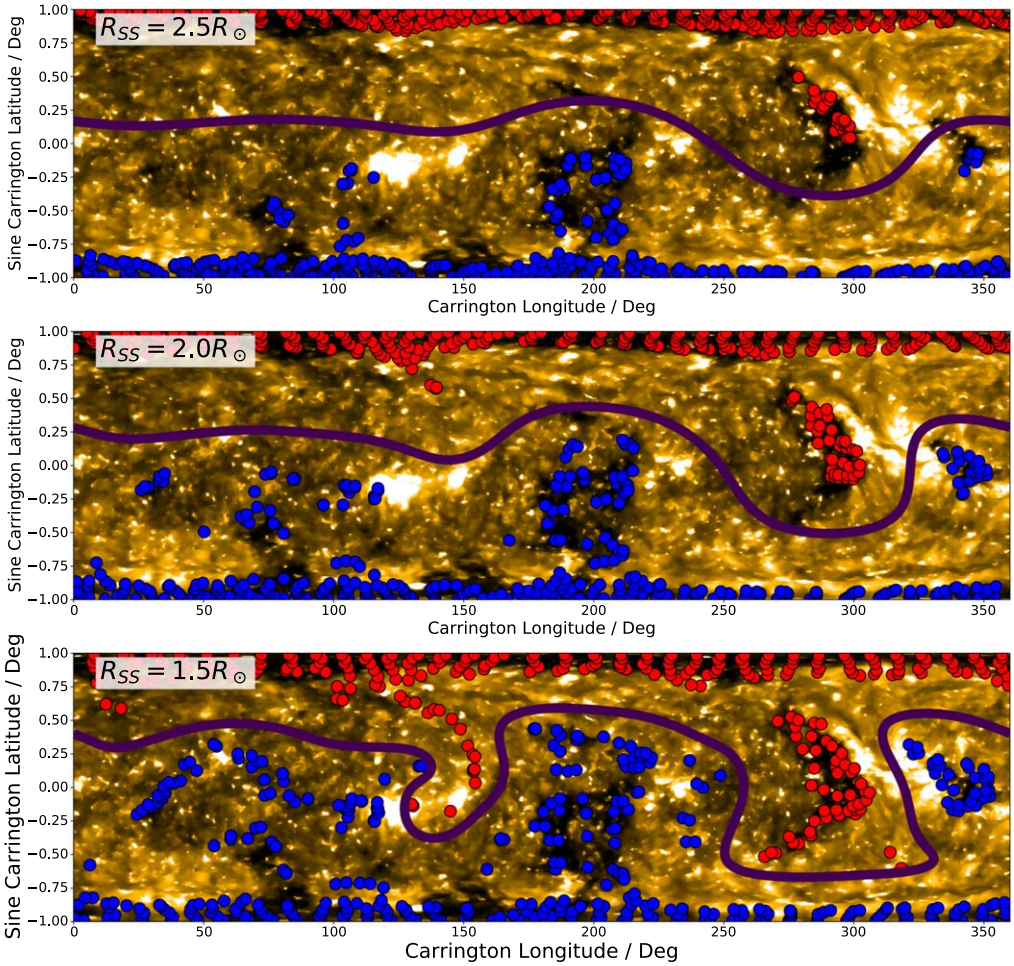

Figure 5. Distribution of PFSS predicted open field line footpoints as a function of source surface height. For the values of source surface height 2.5 $R_\odot$, 2.0 $R_\odot$, 1.5 $R_\odot$, field lines are initialized on a uniform grid at the source surface and mapped down to the photosphere. By definition, field lines initialized at the source surface are open and so this mapping shows the PFSS prediction of the source regions of open magnetic flux. For context, these mappings are overlaid on a synoptic map of the 193 Å emission synthesized from *STEREO*/EUVI and *SDO*/AIA. At this wavelength, dark regions imply low-density plasma regions in the lower solar corona, indicating the presence of coronal holes where open magnetic field lines allow plasma to evacuate into interplanetary space. The dark line overplotted is the PFSS neutral line at the source surface and can be seen to warp more at lower source surface heights.

From comparing dark regions with the modeled footpoint locations we see a generally sensible solar minimum model result: the majority of open field lines map to the north and south polar coronal holes, with some equatorward extensions picked out along with individual isolated mid-latitude and equatorial coronal holes.

However, we observe for the canonical 2.5 $R_\odot$ that these isolated coronal holes are very underexpanded in the PFSS model, with particular emphasis on the small negative equatorial coronal hole around 340° longitude and the more extended positive polarity one just below 300° longitude. Lowering the source surface height to 2.0 $R_\odot$ we find a more reasonable filling of these two features. Going further and examining a source surface height of 1.5 $R_\odot$, we start to overexpand the mid-latitude and equatorial coronal holes and produce some footpoints without obvious coronal hole correspondence.

It should be noted that a precise match between dark EUV images and the footpoints of open field lines is not expected since the EUV coronal hole boundary is somewhat wavelength dependent. Nevertheless, it appears 2.0 $R_\odot$ is a reasonable height for a globally consistent PFSS model with regard to observed coronal hole locations. This inference is consistent with the parameter chosen for the predictive modeling work for *PSP* E1 by Riley et al. (2019) to make PFSS results agree with MHD modeling as closely as possible without overexpanding coronal holes. Lee et al. (2011) also found that the canonical 2.5 $R_\odot$ source surface height resulted in underexpanded coronal holes for a similarly quiet solar minimum Carrington rotation. More recently, Nikolić (2019) has made the same observation for PFSS extractions of GONG maps from 2006 to 2018. Further evidence for this lower source surface height is presented in Appendix B where we compute a cost function as a function of source surface height and show the "optimum" source surface height is significantly lower than 2.5 $R_\odot$ for both GONG and ADAPT evaluations. We use these findings to justify our use of this parameter value in our comparison in Figure 4, and to inform our further investigation of variation of this parameter.

4.2.2. Impact on Timeseries Predictions

We next examine the results of varying the source surface height parameter on the time-integrated best fits from Section 4.1. The results are summarized in Figure 6. Panel (A) shows time series predictions for source surface heights

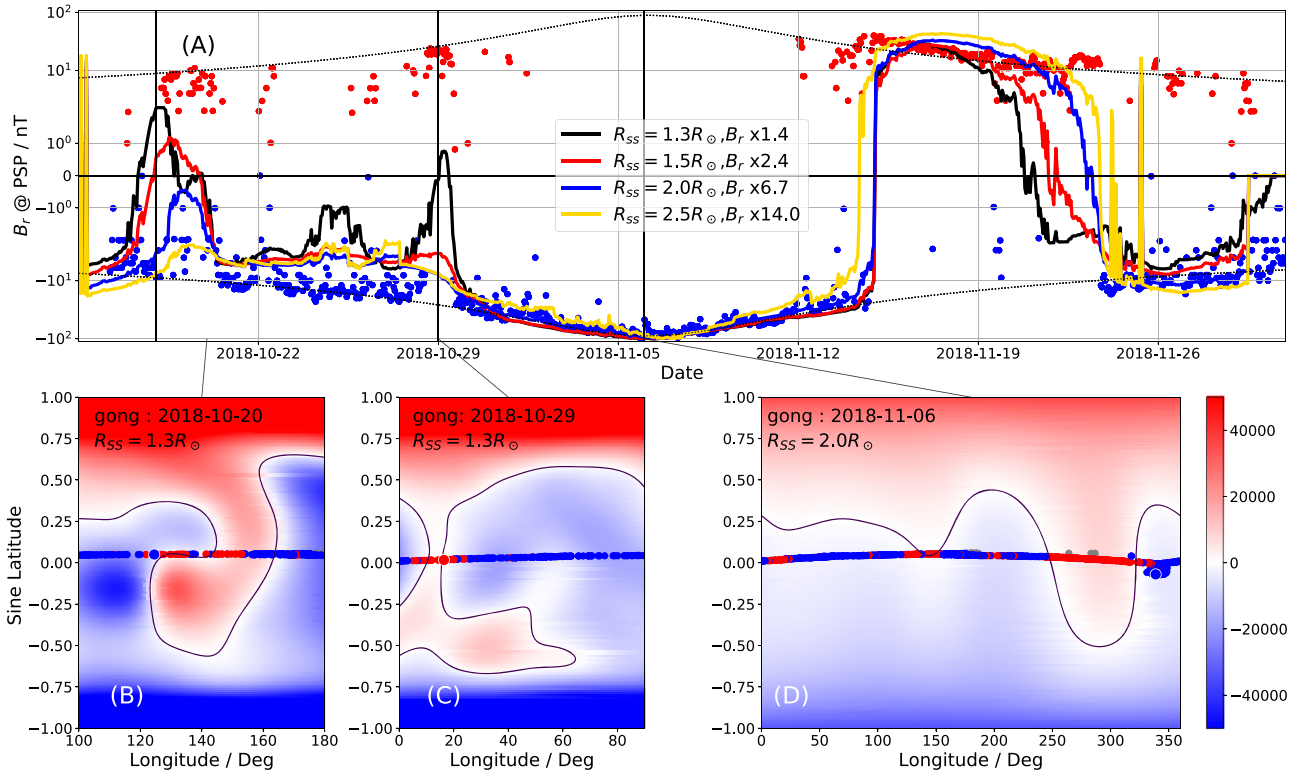

Figure 6. PFSS results as a function of source surface height. Panel (A): comparisons of model predictions for different source surface heights. Predictions use time-evolving magnetograms and SWEAP measurements. The scaling factors required to get the peak field strengths of the predictions and data to match for each source surface height are indicated in the legend. 1.3 $R_\odot$ and 2.0 $R_\odot$ models are as shown in Bale et al. (2019) Figure 1(b). Panels (B)–(D) show the magnetic field strength at the source surface predicted by PFSS projected onto longitude vs. sine(latitude). A black contour indicates the polarity inversion line (PIL). Overlaid on this are the source surface footpoints of *PSP* colored by the measured polarity. Panels (B) and (C) show zoom-ins on PIL structure consistent with small-scale positive polarity inversions measured on 2018 October 18 and 29 respectively for a source surface height of 1.3 $R_\odot$. Panel (D) shows the map for a $R_{SS} = 2.0\,R_\odot$ model over the whole Sun. *PSP* starts at approximately 180° longitude and tracks from right to left as time passes. The lines between panels indicate the time on the time series when the model shown was evaluated.

ranging from 2.5 $R_\odot$ down to 1.3 $R_\odot$. The B_r values are shown on a symmetric log scale to emphasize the polarity inversion features. The scaling factors applied to match the model peak field with the measured peak field are indicated in the legend. These range from 14.0 at the canonical $R_{SS} = 2.5\,R_\odot$ down to order unity for $R_{SS} = 1.3\,R_\odot$. Panels (B–D) are shown to contextualize the time series. In each of these, a colormap of magnetic field strength at the source surface is plotted along with the polarity inversion line (contour of $B_r = 0$) in black. On top of this, the *PSP* trajectory ballistically propagated down to the source surface is shown and colored by *measured* polarity. Where the measured color matches that of the source surface below it, the model and data are in agreement. Panels B and C show model results evaluated from GONG magnetograms on 2018 October 20 and 29 respectively, both for the extremely low source surface height 1.3 $R_\odot$ and zoomed in to a specific part of the Sun to highlight specific PIL topology associated with the PI1–PI2 and PI3–PI4 intervals. Panel D meanwhile offers a more global view with the model at 2.0 $R_\odot$ showing the entire encounter. *PSP* began the encounter at approximately 180° longitude and tracks in the direction of decreasing longitude in time, approximately reaching the position it started from at the end of the time interval considered.

Examining each measured polarity inversion in turn we observe that the timing of PI5 is accurately predicted by all models regardless of the source surface height. PI6 is well predicted for $R_{SS} \geqslant 2.0\,R_\odot$, but for lower R_{SS} it is predicted to occur too early. This is discussed further in Figure 7.

Conversely, PI1, PI2, PI3, and PI4 are entirely missed in predictions using $R_{SS} \geqslant 2.0\,R_\odot$. Only at the very lowest source surface heights do these features convincingly appear in the prediction.

Panel (D) shows the overall distribution of magnetic field at the source surface for 2.0 $R_\odot$. An overall flat PIL generally skewed north of the equator explains why the near equatorial *PSP* mainly connected to negative polarity. A major southward warp of the PIL between 250° and 330° longitude explains *PSP*'s major excursion into positive polarity between PI5 and PI6 which is generally consistent with the shape of the HCS inferred by Szabo et al. (2020). However the amplitude of the warps we infer and hence distance from the current sheet are larger than expected at the radius of *PSP*. This is expected since modeling beyond the source surface (either WSA's Schatten current sheet or flow dynamics in MHD) has the effect of flattening the PFSS-derived PIL as it evolves into the HCS (see Figure 2 of Szabo et al. 2020). However, in spite of the discrepancy in distance to the current sheet, the timing of crossings and overall shape of the predicted magnetic time series are largely unaffected and remain consistent with other models.

Panels B and C show that when the source surface is lowered to 1.3 $R_\odot$ thin tenuous southward extrusions of the PIL develop at the correct longitudes at which *PSP* observed the PI1–PI2 and PI3–PI4 intervals. However, to achieve this we find that the model must be generated using magnetograms from very close

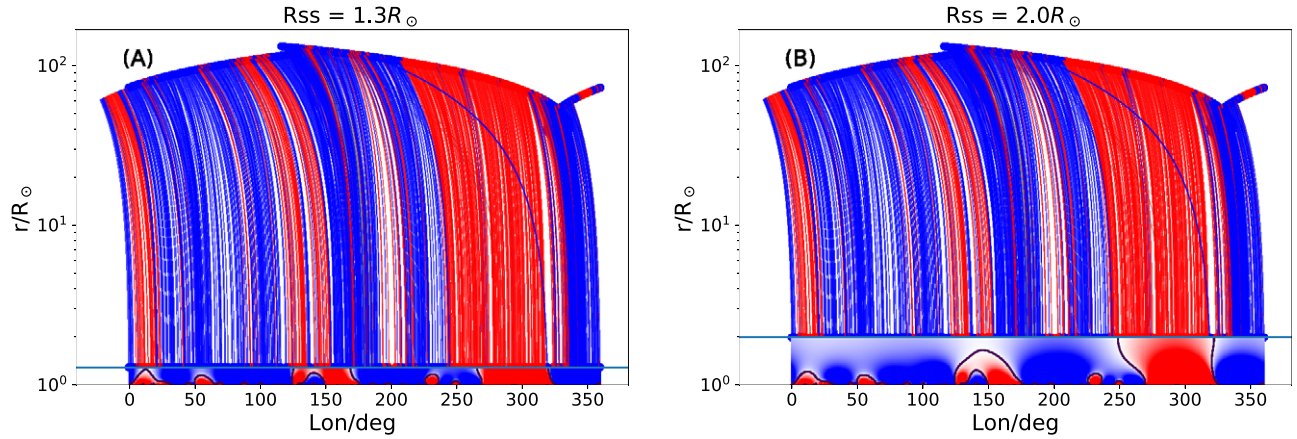

Figure 7. Illustration of impact of lowering source surface height on the model. A comparison of the PFSS model and connection to *PSP* is shown in a plot of log(r) vs. radius. Projecting everything into the ecliptic plane, we plot the *PSP* trajectory colored by measured polarity, then plot the Parker spiral field lines down to the source surface, also colored by measurements. Below the source surface we plot a colormap of an equatorial cut through the model colored by B_r, with black lines indicating contours of $B_r = 0$. Panel (A) shows the results at $R_{SS} = 1.3\,R_\odot$ while panel (B) shows the results at $R_{SS} = 2.0\,R_\odot$. Comparisons between model and observations are made at the source surface.

in time to when *PSP* was at that location. This suggests these features evolved quickly.

In the case of the PI1–PI2 interval, and the PI2 transition in particular, the predictions at the source surface can be traced to a distinct photospheric feature: a simple dipolar active region lies directly below the longitude where *PSP* observed positive polarity. The implied connection is discussed further in Figure 8(B).

These model results and correspondence to data are compared in an alternative format in Figure 7. Here, confining all data to the solar equatorial plane, we plot the *PSP* trajectory colored by measured polarity as log(radius) versus longitude. For each point along the trajectory we trace the Parker spiral used to connect it to the source surface and also color this by data. Below the source surface we plot a colormap of an equatorial cut through a PFSS model evaluated using GONG data from 2018 November 6 (perihelion). The comparison between model and observations is made at the source surface by comparing the model color just below and the observation color just above. By plotting log(radius) we are able to display both the interplanetary scale of the Parker spiral field lines and the coronal scale PFSS model. In panel (A) we plot the low source surface height model ($R_{SS} = 1.3\,R_\odot$), and panel (B) shows the high source surface model ($R_{SS} = 2.0\,R_\odot$).

Examining the evolution of the equatorial cut with radius, it can clearly be seen how higher-order structures initialized at the photosphere smooth out and become simpler with increasing altitude. A general picture in the equatorial cut shown here is an overall negative polarity with small perturbations of positive polarity, most notably at 10°, 55°, 150°, and 300° longitude. Examining the high source surface case, panel (B), we see how only the 300° longitude feature intersects the source surface and the others close over and revert back to negative. The 300° longitude feature also exhibits expansion in its longitudinal extent with altitude. On comparison with the observations, we notice this evolution is a vital element to explaining polarity inversion timings observed by *PSP*.

In comparison, panel (A) shows how the low source surface height is required for the "low-lying" structures at 10° and 150° longitude to be opened to interplanetary space. We also see how this source surface is likely the lowest possible since taking it further down would open up structures at 55° and 240° longitude, which have no observational support. However, while the lower-height model recovers these small structures better, it intersects the 300° feature before it has expanded sufficiently, and so produces a prediction of positive to negative polarity inversion much earlier than was observed (Figure 6(A)). These competing model features make the limitations of a single height source surface apparent: with this constraint our ability to globally match all observed features is hindered.

Figure 7 is also useful in making a qualitative assessment of the uncertainty associated with ballistic propagation. According to Nolte & Roelof (1973), this method has an estimated uncertainty in projected longitude of ±10°. Figure 7 allows visualization of the extent to which a shift of 10° could shift projected footpoints relative to the observed structure in the PFSS model. For example, the westward expansion of the 300° feature from 1.3 to $2.0\,R_\odot$ is approximately 20° and so a 10° shift of the observed transition from positive to negative polarity could not explain the discrepancy in the two source surface height models. The blue sector from 330° to 360° longitude indicates the extent of the source surface region which maps down to the coronal hole we infer *PSP* was connected to at perihelion (Section 4.3). This is similarly larger in extent than the uncertainty and so we can state this connection is robust within uncertainty. With the smaller-scale positive polarity features, which are of the order of 10° in extent, a 10° shift could just drive the footpoints of measured positive polarity out of consistency; however, it would not be sufficient to drive them to connect to any of the other positive polarity regions so the connectivity we infer remains the most likely source of those measurements. Overall, we argue that 10° can affect the details of the agreement of model and data but not our overall conclusions. In addition, with *PSP* orbiting much closer to the corona than the 1 au, this may result in a lower longitude error as the magnitude of the longitude correction is smaller. The excellent matching of observations and modeling to much less than 10° may be evidence of this.

4.3. Implied Field Line Mappings

Having built confidence in the modeling approach by observing good time series predictions, as well as identifying

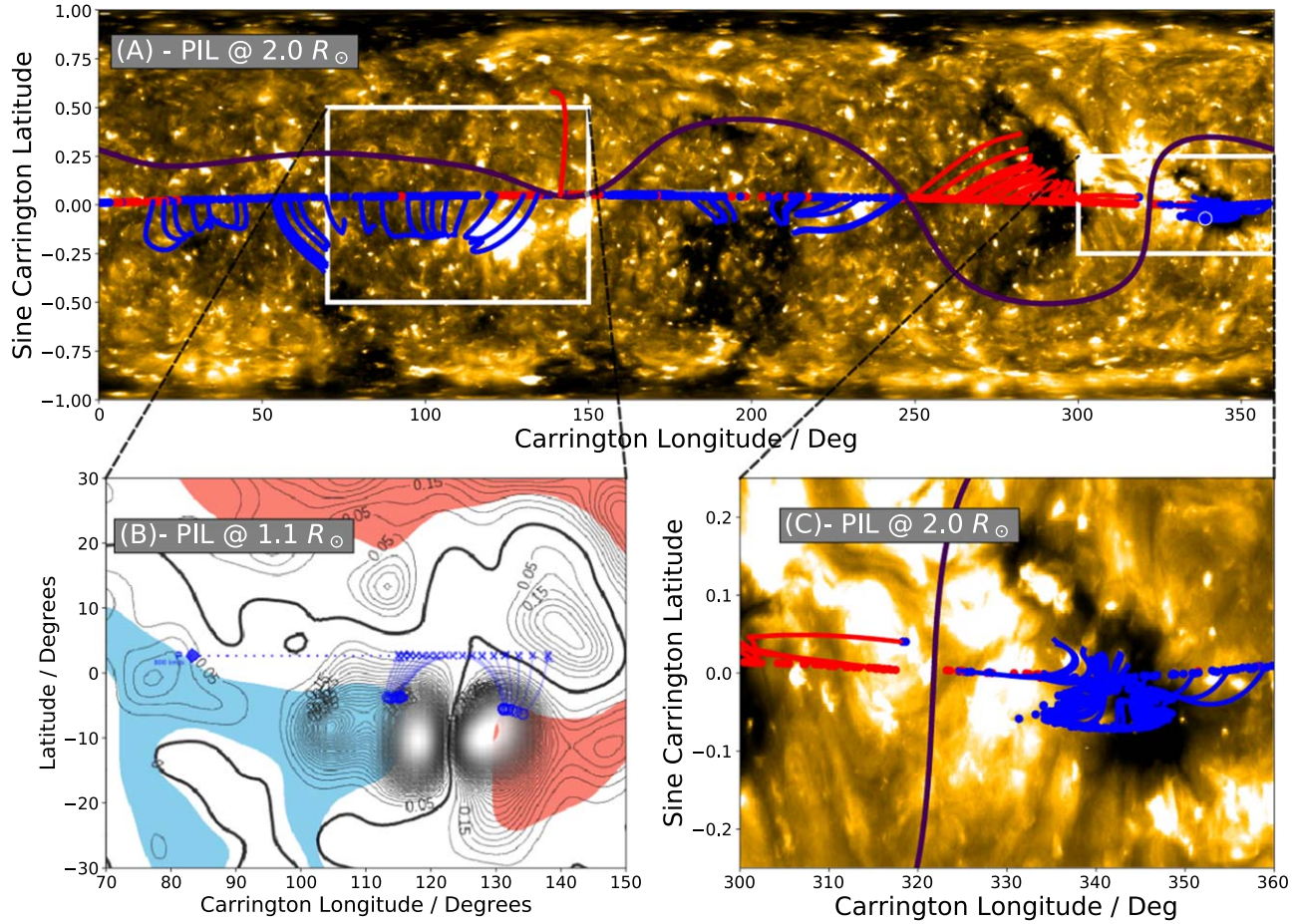

Figure 8. PFSS predictions of photospheric connectivity of *PSP*. PFSS field line traces are initialized at the source surface footpoints of *PSP* and propagated down through the corona to the solar surface. These are contextualized with the same 193 Å map from Figure 5, and the polarity inversion line at the source surface from the model. Panel (A) shows a synoptic view of the whole encounter with the model evaluated at $R_{SS} = 2.0\,R_\odot$. Panel (B) shows a zoom-in to the 2018 October 20 polarity inversion with the model evaluated at $R_{SS} = 1.4\,R_\odot$ using the DeRosa/LMSAL surface flux transport model. Contours of B^2 are shown and the PIL evaluated at $1.1\,R_\odot$ is plotted in bold. Coronal hole regions are shown as red and blue shading. The blue diamond at 85° longitude indicates the Carrington longitude of *PSP*; the crosses indicate the footpoints of *PSP* at the source surface. The circles indicate magnetic footpoints at $1.1\,R_\odot$ and the crosses and circles are connected by field line extrapolations. Panel (C) shows a zoom-in of the connectivity around perihelion with $R_{SS} = 2.0\,R_\odot$. The field line mappings indicate connectivity to an equatorial negative-polarity coronal hole preceding a polarity inversion. Field lines shown in panels (A) and (C) are from same model evaluation depicted in Bale et al. (2019, Figures 1(C) and (D)).

its response to variation in input parameters, we now explore the implications of the model results for the sources of the solar wind measured by *PSP* during its first encounter.

Similar to how open field footpoints were generated for Figure 5, we use the *PSP* trajectory projected down to the source surface to initialize a series of field lines which we propagate down into the model corona using the *pfsspy* field tracing routine. The results of this tracing with $R_{SS} = 2.0\,R_\odot$ are summarized in Figure 8(A) where we overplot the field lines colored by their polarity in the model on top of the same EUV data from Figure 5. We also show the *PSP* trajectory projected on to the source surface and colored by *measured* polarity as in Figures 6 (B), (C), and (D).

In Figures 8(B) and (C), we examine the field line traces of two subregions of the Sun of particular interest. Panel (B) shows results around the active region which appears associated with the PI1–PI2 interval. The modeling results shown here are using the DeRosa/LMSAL model evaluated on 2018 October 20 with a source surface height of $1.4\,R_\odot$, the highest value of R_{SS} for which the model and measurements agree (see Panasenco et al. 2020). The faint contours of $|B|^2$ illustrate the topology of the active region, while the bold contour depicts the PIL at $1.1\,R_\odot$, indicating the polarity structure at the base of the corona. Crosses mark the positions on the source surface *PSP* connects to and track from right to left in time. The circles indicate the field footpoints these map to. It can clearly be seen that in this time interval the footpoints jump from the positive to negative area of the active region and that the PFSS-inferred neutral line is guided along the neutral line of this active region. It is difficult to argue that PFSS accurately captures the magnetic topology around such non-potential structures, but it is compelling that the longitude of this region matches with the projected longitude of *PSP* at the same time it measured PI2, implying at the very least an association of the observed feature with the measurements. This fine tuning of the crossing timing and more modeling with the DeRosa/LMSAL model are expanded on in more detail in Panasenco et al. (2020) in which the authors explore optimizing source surface height for various polarity inversion case studies, suggest a resulting fitted non-spherical source surface height profile, and study the propagation and loss of Alfvènicity of slow wind streams en route to 1 au.

Panel (C) zooms in on the region of the Sun where *PSP* was located during the two periods of corotation either side of perihelion. The results here are shown for $R_{SS} = 2.0\,R_\odot$ but the qualitative conclusions here are largely independent of this choice. The major implication here is that for the entire "loop" part of the trajectory *PSP* was connected very stably to the small equatorial coronal hole at around 340° longitude. This suggests *PSP* was connected to a region of the Sun of less than 10° angular extent for over two weeks, over which time its radial distance varied significantly. Therefore data from this interval can be interpreted as a time series of evolving solar wind from a single source convolved with changing sampling radius. Panel (C) also shows striking agreement between the location of the PIL at approximately 320° longitude and the point at which *PSP* measurements indicated a polarity inversion from negative to positive.

Other general observations from Figure 8(A) indicate that *PSP* connected to another (but this time positive) equatorial coronal shortly after perihelion, but with much more rapid footpoint motion tracking northward over time. Comparison to SWEAP measurements suggests this configuration is responsible for the longest interval of fast wind observed in E1 (Figure 1(B)). After crossing back from positive to negative polarity on November 23, it continued to be connected to mid-latitudes via a very large negative coronal hole.

5. Discussion

Overall, despite its simplicity, our PFSS + ballistic propagation scheme produces compelling predictions of the B_r time series *PSP* measured during its first encounter. The large-scale features were well reproduced, including the flat source surface field convolved with $1/r^2$ variation, the predominantly negative polarity, and the timing of significant excursions into positive polarity. The model results are likely distorted by unmodeled chromospheric currents just above its lower boundary, as well as field re-organization beyond the source surface; however, empirically it appears these distortions were not great enough to destroy this overall good agreement. The applicability of PFSS is likely enhanced by the extremely low activity on the Sun at the time of observations, which would result in fewer strong current systems. The impact of latitudinal reorganization between the source surface and $10\,R_\odot$ (Réville & Brun 2017) is likely mitigated by the near equatorial sampling of the spacecraft. *PSP*'s orbital inclination is ∼4° relative to the solar rotation axis, which is even less than the Earth–ecliptic inclination of ∼7°.25. Even so, we do see PI6 is predicted to be a smoother transition than observations suggest and this sharpening can be explained by modeling beyond the source surface.

We also found the time evolution of the input photospheric maps was important to take into account, resulting in more robust and accurate time series predictions compared to those generated with a single magnetogram. The slow angular velocity of *PSP* compared to the surface of the Sun around perihelion likely plays a role in this finding, meaning since *PSP* tracked very slowly around the Sun in the corotating frame, the photospheric magnetic field changed significantly in this time interval. Since PFSS is a time-independent model it is not possible to include dynamics; however, we have shown that this limitation can be mitigated by treating the model as a snapshot of assumed static coronal structure sampled by the spacecraft for a small timestep, and chaining together these snapshots to approximate dynamical evolution.

Next, we discussed the impact of varying the source surface height parameter and found that, on lowering it, new polarity inversion features emerged in the prediction which were consistent with observations. However, while this improved the predictions at some heliographic locations it worsened it at others, suggesting there is no clear "best" source surface height (see also Appendix B). This highlights the limitation of having a spherically symmetric source surface boundary condition. This qualitative finding is unsurprising given the lack of apparent spherical symmetry in observations of the outer corona, and development of PFSS since its inception has sought to generalize beyond a spherical boundary (e.g., Levine et al. 1982). Comparisons to MHD modeling have suggested a boundary of near radial field which is far from spherical (e.g., Riley et al. 2006). This latter result, however, generally showed the dominant perturbation to sphericity was a latitudinal effect at solar minimum with the source surface elongated at high latitudes into a prolate spheroid shape. Given *PSP*'s near-equatorial orbit and limited connectivity to high latitudes, producing model results that agree with observations is unlikely to contain information about the solar magnetic field at high latitudes. Instead, our results here are suggestive of variation in the height with *longitude* and perhaps with localized perturbations to the source surface below the canonical radius over specific magnetic structures as explored in Panasenco et al. (2020).

The successful prediction of new observed features at low R_{SS} implies small, short-lived, magnetically open structures that persist out to interplanetary space are normally smoothed out by PFSS modeling but can be recaptured by investigating lower source surface heights. Figure 2 suggests such small-scale features may even be measurable at 1 au. One of these features appeared associated with a dipolar active region with a pronounced neutral line. Active regions are typically highly non-potential and dynamic, requiring MHD or nonlinear force-free modeling approaches. Riley et al. (2006) found PFSS modeling (with the canonical $R_{SS} = 2.5\,R_\odot$) lacking in modeling fields in the vicinity of such a feature. Nevertheless *PSP* observed a polarity inversion from positive to negative at exactly the time low source surface height PFSS modeling suggests connectivity would have switched from the positive to negative lobes of the active region (Figure 8(B)). This implies that, in spite of the non-potential details, PFSS can be useful in associating such structures with observations in the inner heliosphere. The possibility of open field lines connecting to active regions has been discussed before, for example in Švestka et al. (1977).

A brief discussion of the scaling factor as a function of R_{SS} shown in Figure 6 is also warranted. We note that for $R_{SS} \geqslant 2.0$ a $1/r^2$ scaling from the source surface out into the heliosphere produces underpredicts the in situ field strength by an order of magnitude. This was first observed early in the history of PFSS by Levine et al. (1977). A simple interpretation is that non-radial evolution beyond the source surface is important in connecting it far out to the heliosphere. At the source surface, PFSS predicts stronger field at the poles than at the equator and it is known from *Ulysses* observations (Smith & Balogh 1995; Smith 2011) that this relaxes to a latitudinally independent state further out. The simplest configuration for this to happen is for the polar field to decrease faster than $1/r^2$

and for the equatorial field to drop more slowly than $1/r^2$, which is consistent with an underprediction of the field at low solar latitudes. It is possible that this underprediction could be mitigated with a Schatten current sheet model (Schatten 1972) as used in the WSA model, and this will be the subject of future work. However, from the results plotted in Figure 4, the overall shape of the B_r time series is consistent to within this scaling factor from which we conclude this "extra" field strength is only weakly dependent on solar longitude, with some worsening agreement at higher *PSP* radii (early and late in the encounter). That the scaling factor drops to order unity with decreasing source surface height demonstrates how the radial field inside the PFSS model domain drops faster than $1/r^2$ since the dominant component is a dipolar $1/r^3$ field. Flux which opens up to the solar wind very low down in the corona likely escapes purely radially, which may explain why $1/r^2$ predicts the correct magnitude at *PSP* for these cases.

Lastly, in addition to the active region connectivity discussed above, we presented a synoptic view of the connectivity implied by PFSS during the whole encounter. The relatively low $2.0\,R_\odot$ source surface height suggests predominantly equatorial and mid-latitude connectivity as opposed to deep within polar coronal holes. This is consistent with *PSP*'s predominantly slow wind observations (Figure 1(B)). However we note here that raising the source surface height to its canonical value can change the connectivity very early in the encounter to the southern (negative) polar coronal hole. Given the time interval examined is dominated by the perihelion loop, our inference of a "best" source surface height is likely skewed toward the corotation interval and so statements of connectivity earlier or later in the encounter at $2.0\,R_\odot$ may be weakened. For example, Réville et al. (2020) see polar connectivity at intervals prior to October 29 with MHD modeling results, while Szabo et al. (2020) examine an ensemble of different models and establish that *PSP* was very close to the HCS at this time and thus may be observing streamer belt plasma.

Of particular interest is the connectivity in the two week interval surround perihelion (Figure 8(C)) which shows throughout the corotational period that *PSP* was magnetically connected to an isolated, negative polarity equatorial coronal hole. This was also predicted prior to the encounter by Riley et al. (2019) and now appears well corroborated by Kim et al. (2020), Réville et al. (2020), and Szabo et al. (2020). The coronal hole is approximately 10° in longitudinal extent at the 193 Å isosurface which implies a linear distance scale of the order of 100 Mm. Therefore without assuming detailed knowledge of the PFSS-derived footpoints we can state the source region for the sampled solar wind from this two week period (2018 October 30 to November 14) was confined to this 100 Mm size region, and thus variation in the in situ data from this time can be interpreted as measurements of dynamics of the solar wind emitted by an approximately fixed source on the Sun (see, e.g., Bale et al. 2019). This connectivity will also lead to interesting discussion on sources of the slow solar wind, which is very much an active area of research (e.g., Wang & Ko 2019).

In summary, we have presented a first attempt at global coronal and inner heliospheric modeling to contextualize observations made by *PSP* in its historic first solar encounter. Our potential field-based modeling scheme is extremely simplistic and it will be vital to make detailed further comparison with concurrent or future global modeling work, e.g., MHD (Kim et al. 2020; Réville et al. 2020) or other PFSS models (Panasenco et al. 2020; Szabo et al. 2020). Nevertheless we report various pieces of evidence that suggest the limitations of our modeling have been mitigated sufficiently to claim real and useful contextual information of *PSP*'s magnetic connectivity. In particular, we find *PSP* spent the 14 days surrounding perihelion connected to a small negative equatorial coronal hole and may have also sampled open flux tubes associated with an active region around 2018 October 20, prior to perihelion.

The data discussed in this work come from the first of many planned encounters for *PSP*. At the time of writing, encounters 2 and 3, which both sample down to the same perihelion distance of $35.7\,R_\odot$, have taken place. Following encounter 3, *PSP* will perform a gravity assist with Venus and begin to probe deeper into the solar atmosphere, eventually reaching below $10\,R_\odot$. Repeating the analysis of this work on these future encounters will be interesting and will, for example, allow us to control for the impact of solar activity on the modeling results.

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Appendix A
Choice of Magnetogram Source

As mentioned in Section 3, although we show throughout this paper results using GONG zero-point-corrected magnetograms, we did consider a range of other possible sources. Here

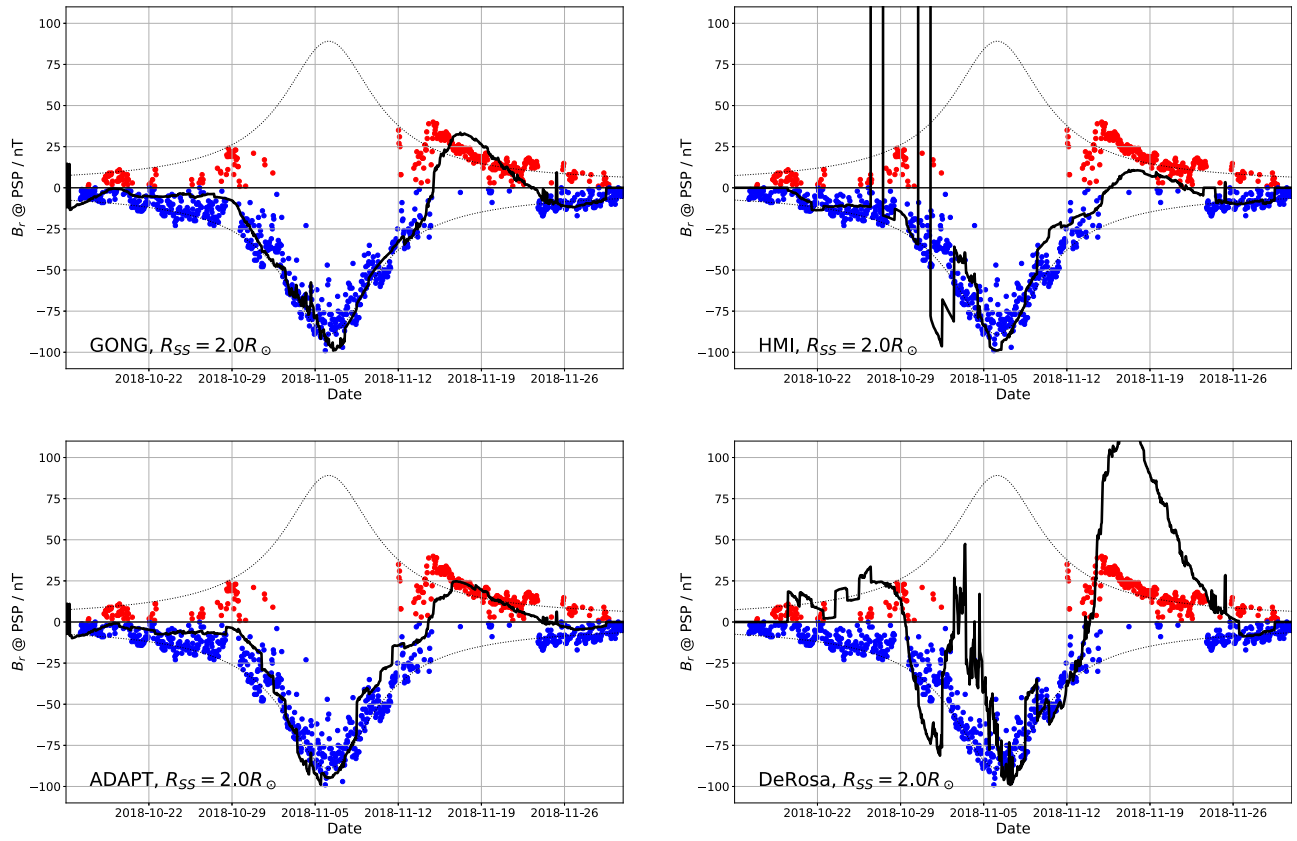

Figure 9. Comparison of time series predictions using different magnetograms. The "best" GONG time series prediction from Figure 4 is shown here in comparison to the same procedure applied to magnetograms from HMI, ADAPT, and the DeRosa/LMSAL model. GONG produces the smoothest prediction on time integration but the general picture of negative polarity, $1/r^2$ variation, and the times of polarity inversions are not strongly perturbed by choice of magnetogram.

we show that our results are largely independent of the choice but that GONG produces the most compelling predictions.

In Figure 9, we compare the GONG results from Figure 4 with those obtained by an identical procedure on different magnetogram outputs. On the top row are results from GONG and HMI that are both purely observational data products. Below are results from ADAPT and the DeRosa/LMSAL models that are surface flux transport models which assimilate the GONG and HMI data respectively. On visual comparison, we see GONG and ADAPT results are very similar as expected. The HMI results are largely consistent but with too small a peak amplitude predicted, and has some spurious data from some of the magnetograms considered. The DeRosa/LMSAL-based result shows significant departure from the data: the period before 2018 October 29 is predicted to connect to the positive side of the HCS, the 2018 November 4 bump which is smoothed out by time integration (Section 4.1) in the other models is still prevalent, and the peak positive amplitude is significantly overestimated. The variation between these models demonstrates that because *PSP* traversed generally very close to the HCS during E1, the predictions are quite sensitive and a small change in the modeled PIL can produce a sudden reversal in polarity measured at *PSP*. We err toward the GONG-based data due to this empirical observation.

In terms of the difference between GONG and ADAPT, the main noticeable change is that on time integration, the ADAPT prediction becomes "choppier," and subsequent 3 day intervals do not smoothly meet each other. This is likely due to the extra physical modeling in ADAPT, meaning flux variations are captured at higher time resolution than with pure GONG data. Nevertheless the major conclusions from Section 4.1 are unchanged from use of either of these magnetograms. Since these fluctuations do not immediately appear to correspond to data we infer, although they may be physical, they are likely smoothed out exterior to the source surface via processes not considered by PFSS. Beyond this, the choice of GONG versus ADAPT does not affect the outcome of this paper and hence we make the choice of the smoother predictions and proceed with our analysis using these.

For further comparison and to offer some insight into possible sources of discrepancy above, in Figure 10 we compare the same choices of magnetograms with a source surface height of $1.3\,R_\odot$ which, as discussed in Section 4.2.2, produces predictions of new small-scale polarity inversions prior to perihelion.

In this case we observe excellent consistency between all models including DeRosa/LMSAL. Since at this source surface height the polarity and field strength are much more related to the field strength radially below, this suggests that high latitude field is the dominant cause of disparity in Figure 9, for example differences in how the unobserved polar regions are modeled. Nevertheless, we again note conclusions based on the GONG prediction are unchanged with different magnetogram sources and we therefore work with these data in the main text.

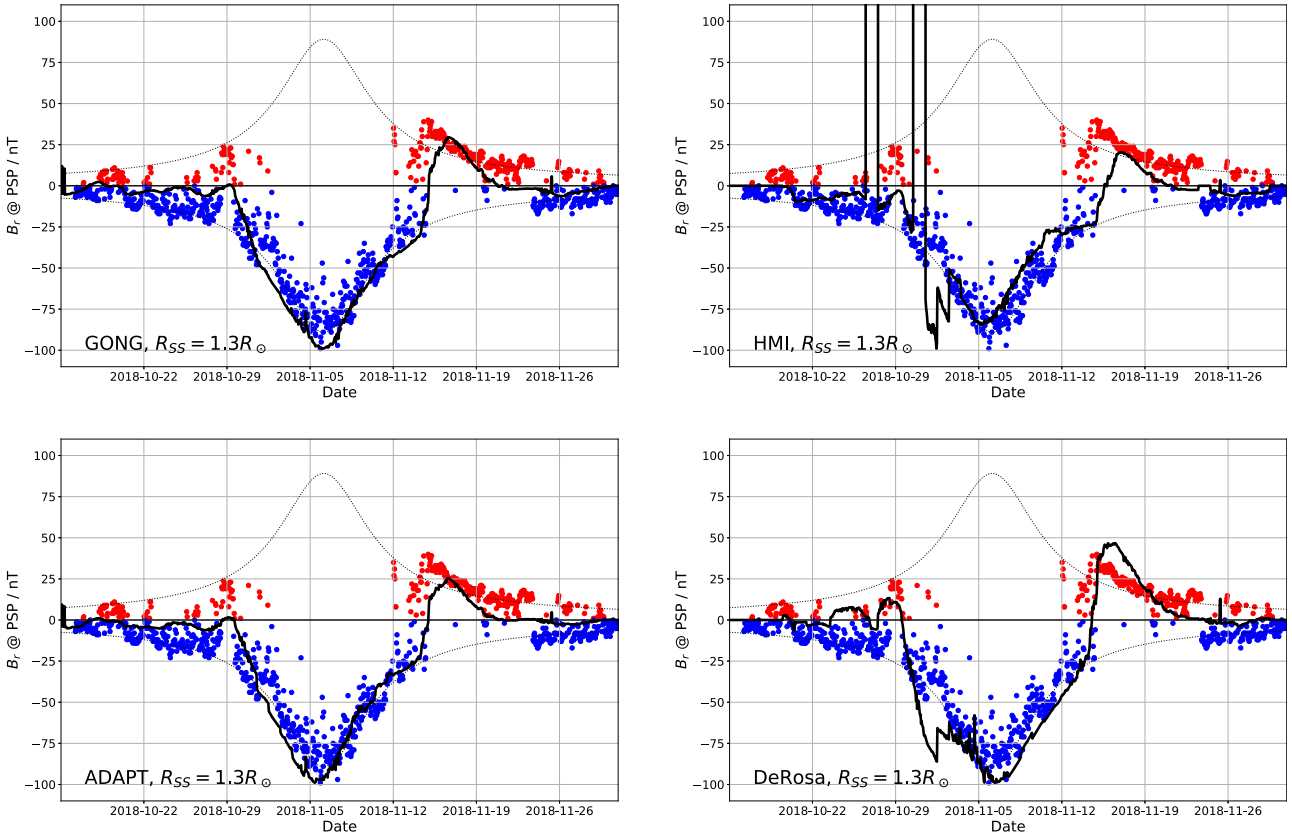

Figure 10. Comparison of time series predictions using different magnetograms at low source surface height. Compared to 2.0 $R_\odot$, here all the models are very consistent with each other, including LMSAL/DeRosa.

Appendix B
Cost Function for Comparing Source Surface Heights

Further supporting evidence for our general use source surface heights below the canonical value of 2.5 $R_\odot$ is shown in Figure 11. Here, we compute a least-squares cost function evaluating the relative similarity between time integrated models (see Section 4.1).

This cost function is given by

$$L(M, O) = \sqrt{\sum_{i=0}^{N} (M_i - O_i)^2} \quad (2)$$

where the model, M, and observations, O, are N-dimensional vectors. Since both the model and observations are expressed in nT this is also the unit of the cost function. In Figure 11 we show the least-squares result as a function of source surface height using the GONG (red) and ADAPT (gold) input magnetograms.

Both models show a distinct minimum (best fit) at a significantly lower source surface height than 2.5 $R_\odot$. The ADAPT "best" height is approximately 1.9 $R_\odot$ while for GONG it is approximately 2.2 $R_\odot$. However, these minima are both very broad and have overlapping full width at half maxima. In addition, as discussed in the main text, the concept of a single source surface height to fit all longitudes and for a six week long interval is likely not a good approximation. We settle on a value of 2.0 $R_\odot$ to discuss a global picture in the main text, but note here a range of $\pm 0.2\,R_\odot$ will have very little affect on the overall goodness of fit.

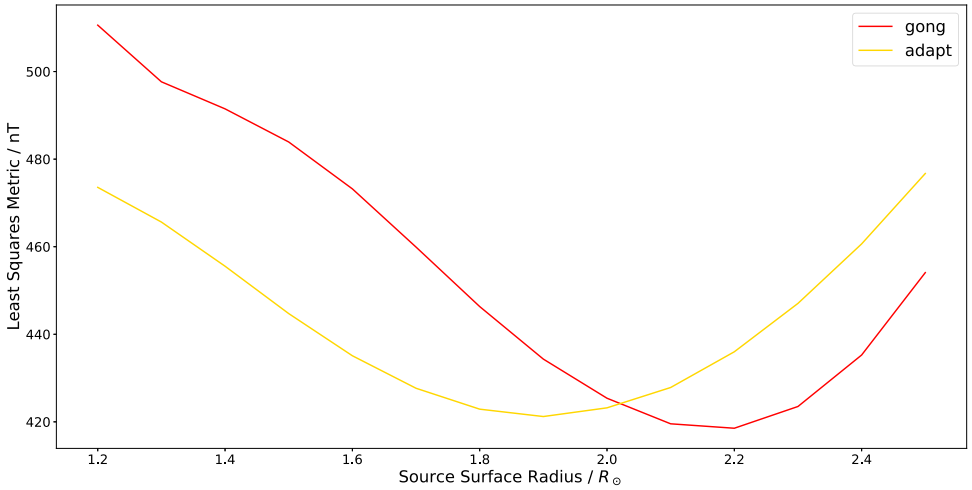

Figure 11. Least-squares metric computed between time-integrated model results and *PSP* data as a function of source surface height. Results for GONG and ADAPT are shown and exhibit a distinct minimum below the canonical 2.5 $R_\odot$ for both cases.

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