



A Mystery in Chamaeleon: Serendipitous Discovery of a Galactic Symbiotic Nova

Lachlan Lancaster¹ , Jenny E. Greene¹, Yuan-Sen Ting^{1,2,3,4,10,11} , Sergey E. Koposov^{5,6,7} , Benjamin J. S. Pope^{8,9,12} , and Rachael L. Beaton^{1,3,10,11}

¹ Department of Astrophysical Sciences, Princeton University, 4 Ivy Lane, Princeton, NJ, 08544, USA; lachlanl@princeton.edu

² Institute for Advanced Study, Princeton, NJ 08540, USA

³ Observatories of the Carnegie Institution of Washington, 813 Santa Barbara Street, Pasadena, CA 91101, USA

⁴ Research School of Astronomy and Astrophysics, Mount Stromlo Observatory, Cotter Road, Weston Creek, ACT 2611, Australia

⁵ McWilliams Center for Cosmology, Carnegie Mellon University, 5000 Forbes Avenue, Pittsburgh, PA, 15213, USA

⁶ Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge, CB3 0HA, UK

⁷ Kavli Institute for Cosmology, University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK

⁸ Center for Cosmology and Particle Physics, Department of Physics, New York University, 726 Broadway, New York, NY 10003, USA

⁹ Center for Data Science, New York University, 60 Fifth Avenue, New York, NY 10011, USA

Received 2020 February 12; revised 2020 June 30; accepted 2020 July 7; published 2020 August 21

Abstract

We present the serendipitous discovery of a low optical-luminosity nova occurring in a D-type symbiotic binary star system in the Milky Way. We lay out the extensive archival data alongside new follow-up observations related to the stellar object CN Cha in the constellation of Chamaeleon. The object had long period (250 days), high amplitude (3 mag) optical variability in its recent past, preceding an increase in optical brightness by 8 magnitudes and a persistence at this brightness for about 3 yr, followed by a period of 1.4 mag yr^{-1} dimming. The object's current optical luminosity seems to be dominated by H α emission, which also exhibits blueshifted absorption (a P-Cygni-like profile). After consideration of a number of theories to explain these myriad observations, we determine that CN Cha is most likely a symbiotic (an evolved-star–white-dwarf binary) system that has undergone a long-duration, low optical brightness, nova, placing it squarely in the class of so-called “slow novae,” of which there are only a few known examples. The duration of the optical plateau in CN Cha would make it the shortest timescale plateau of any known slow symbiotic novae.

Unified Astronomy Thesaurus concepts: Symbiotic novae (1676); Novae (1127); Symbiotic binary stars (1674); Interacting binary stars (801)

1. Introduction

“One of the advantages to being disorganized is that one is always having surprising discoveries.”— A.A. Milne (Milne 1956)

The array of human-collected astronomical data is vast. Varied in both structure and content, there is no single archive that gathers all such data into a single, easily accessible place. Comparison between historical and contemporary data sets can reveal many previously overlooked astrophysical phenomena. In this paper, we present a dramatic outburst from the star CN Cha, whose enigmatic nature was only revealed through the broad time-domain and spectral coverage available across astronomical archives.

On a recent observing run, we chose bright and distant objects from the Gaia DR2 data (Gaia Collaboration et al. 2016, 2018) for spectroscopic follow-up. One such candidate was the star CN Cha in the constellation of Chamaeleon. As we will outline below, the observations that followed were quite intriguing. Yet the archival data that had already been gathered on this object were even more fascinating, including variations in brightness over a range of 10 magnitudes that occurred on a

variety of timescales. This variation is exemplified in Figure 1, which shows observations of CN Cha separated by 25 yr.

In Section 2, we present and summarize the set of archival data pertaining to this object. We then present the optical spectrum that we observed for CN Cha in Section 3. After having presented the data, we provide more detailed analysis of various parts of the data in Section 4. Using the net sum of the observational data, we then discuss the possible theories that could explain them in Section 5. Among these theories, we determine that it is likely that CN Cha is a symbiotic nova. In Section 6, we briefly discuss the implications of the interpretation of CN Cha as a symbiotic nova on the study of novae. Finally, we discuss possible follow-up observations that could be taken of this object and conclude in Section 7.

2. Archival Data

We present the various archival data that we have gathered on this object. We begin with a summary of the data as a whole. This is followed by a detailed description of each constituent data set, beginning with the Gaia data on the object, which first motivated our investigation. We then present the numerous archival photometric observations, roughly in chronological order, so as to give the reader a full explanation of the data underlying the photometric variability.

2.1. Archival Summary

CN Cha was first identified as a Mira variable in 1963 and was observed by multiple surveys throughout the 20th century that took measurements consistent with this initial conclusion.

¹⁰ Hubble Fellow.

¹¹ Carnegie-Princeton Fellow.

¹² NASA Sagan Fellow.



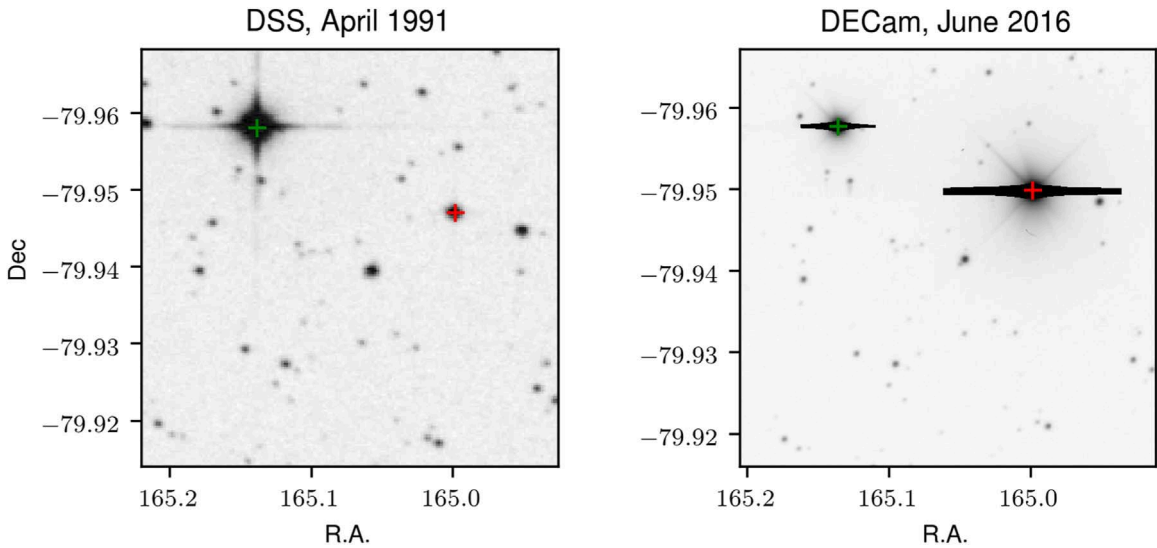


Figure 1. We show a comparison of two images of CN Cha, which is indicated by a red arrow in both panels. The left panel shows the DSS *R* band observation taken in 1991 April and the right shows the DECam *r* band observation taken in 2016 June. Despite the fact that the images are at different contrast levels and resolutions, and they are taken in slightly different bands, it is clear from comparison to the nearby star CD-79 452 (indicated with a blue arrow) that CN Cha has increased in brightness considerably.

A visual summary of the photometric data since 2000 is given in Figure 2. We gather multi-epoch photometry spanning nearly 20 yr ranging from the near-UV to the far-infrared. The information summarized in Figure 2 indicates that CN Cha experienced an outburst event at some point in 2013. We find that the pre-outburst SED is best described by the sum of two blackbodies, one cool (2000 K) bright component, and one hot (10,000 K) dim component. Observing the SEDs from before and after this event, it is clear that the object became much brighter in almost all bands for which we can make a reasonable comparison. This brightness difference is most extreme in the UV, where an observed increase in flux by a factor of 10,000 is seen between UVOT and SkyMapper measurements. It should be noted that these SkyMapper measurements in the UV were taken very soon after the outburst event, in 2014 March, so it is plausible to assume that the SED now is somewhat different than it was then, especially given the dimming that was observed later in the ASAS-SN light curve.

Finally, we show the TESS light curve in Figure 3, which we classify as “archival” despite the fact that it was taken after our own observation. The main conclusions from the TESS data are that (1) they indicate that the star is still steadily dimming at a rate consistent with the ASAS-SN observations and (2) they suggest the star has short-term variability on the timescale of several hours.

2.2. Gaia

The European Space Agency’s astrometric space mission Gaia, observed CN Cha during its first window of observation used in the mission’s second data release (DR2; DR2 Gaia Collaboration et al. 2016, 2018). This observation window was from 2014 July 24 to 2016 May 23. The mission observed the star to be peculiar in a number of ways that alerted the authors that the star may be interesting. The key observations that piqued our interest were:

1. The photometric observations gave $G = 7.41$, $G_{BP} = 7.72$, $G_{RP} = 6.98$ mag, making the star relatively bright.
2. The star was assigned a parallax of

$$\varpi = 0.3142 \pm 0.0249 \text{ mas} \quad (1)$$

putting the star’s distance at $3.18^{+0.27}_{-0.23}$ kpc, implying an absolute G band magnitude of -5.1 mag, making it unusually luminous.

3. The astrometric pipeline assigned zero astrometric excess noise (AEN) to this object, which indicates that the astrometry does not show obvious problems. The Renormalized Unit Weight Error (RUWE) is 1.19.
4. The star was not classified as variable.

The first three of the above items were enough to identify this object as one of interest for follow-up.

2.3. Identification

CN Cha was first identified as a Mira variable star in 1963 as reported by Hoffmeister (1963). At that time, the star was found to vary between magnitude 15 and 17,¹³ in the Johnson *I* band.

2.4. IRAS

The star was later identified by the Infrared Astronomical Satellite (IRAS) as a strong point source in all four bands (12, 25, 60, and 100 μm) as well as being given an 82% probability of being variable based on analysis of the 12 and 25 μm flux densities and their uncertainties (Neugebauer et al. 1984).

2.5. Hipparcos

Despite being very bright in the optical at the current epoch, it was not recorded as a source in any of the catalogs created from observations performed by the Hipparcos satellite (ESA 1997; Høg et al. 2000). The completeness of the

¹³ www.sai.msu.su/gcvs/cgi-bin/search.cgi?search=CN+Cha

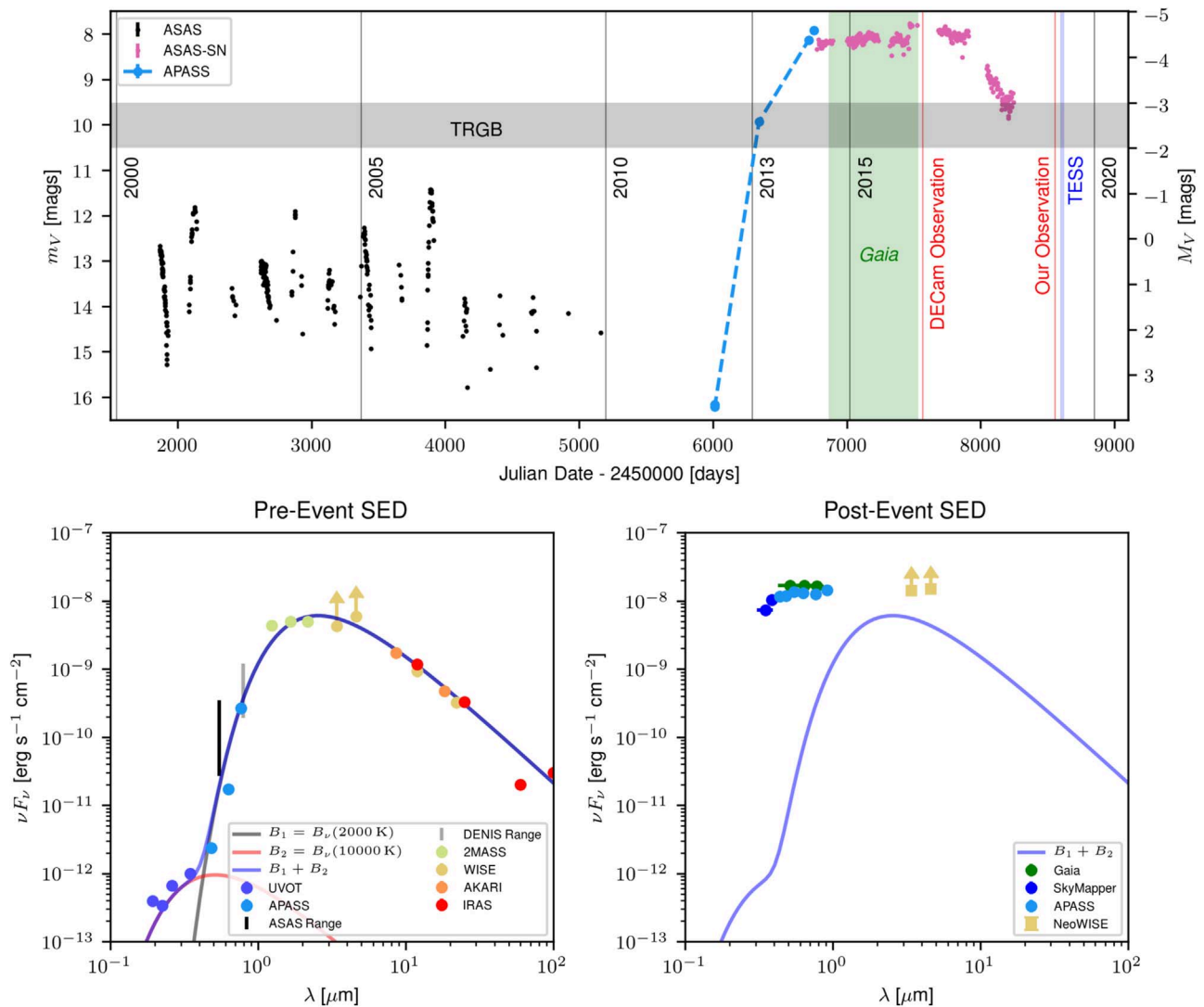


Figure 2. A summary of the photometric data. Top panel: we show a combination of light curves from the ASAS survey (black points), the APASS survey (light-blue points connected by a dashed line) and the ASAS-SN survey (pink points) all in approximations of the Johnson-Cousins V band; we make no correction for marginal differences in the definitions of these bands between surveys. The scale for the apparent (absolute, from Gaia parallax) magnitude is shown at left (right). We indicate intervals of 5 yr (along with 2013, the year we believe the outburst occurred) with vertical gray lines, the DECam and our spectroscopic observations as vertical red lines, and the period over which the Gaia mission used data for its DR2 as a green band. The small blue band indicates the TESS observations. We indicate the typical range in absolute magnitude spanned by the tip of the red giant branch (TRGB) as a horizontal gray band. Bottom left panel: we show an SED for the star for photometric data that is determined to have been taken “pre-event.” We compare the data with an SED, which is the sum of two blackbody spectra, one with a temperature of 10,000 K and one with a temperature of 2000 K for reference, the normalization of this spectrum was chosen to marginally follow the observed photometric data. We chose two blackbodies to match the UV excess observed in the data. As the star is observed to vary quite strongly in the “pre-event” phase we attempt to illustrate this variability by indicating the range of variability of the ASAS V band measurements as a black band and the range of the unsaturated DENIS measurements as gray bands. Due to the variability the deviation from the blackbody in the variable region should not be taken as significant. Bottom right panel: an analog of the bottom left panel but for “post-burst” photometric data. The sum of blackbodies has the same scale as the curve shown in the bottom left panel. The photometry is much more stable in the “post-outburst” phase so we can somewhat safely interpret these points as representing the true SED.

Tycho-2 catalog is 99% for stars brighter than 11 magnitude in the V band (Høg et al. 2000), suggesting that CN Cha was likely dimmer than 11th magnitude in the V band during Hipparcos’s observations from 1989 to 1993.

2.6. Digitized Sky Survey

The Digitized Sky Survey (DSS) consists of a series of photometric plate observations done on the Oschin and UK Schmidt Telescopes between 1983 and 2006 (Lasker et al. 1990; 1996). CN Cha was observed by the UK Schmidt Telescope using the RG610 filter (approximating the Johnson-Cousins R band) on 1991 April 12th. We show this observation

in the left hand panel of Figure 1. The star was reported to have magnitude $RG610 = 13.26$ mag in this band by the DSS (Lasker et al. 1990).

The star was also observed by the DSS in January of 1976 in a GG395 filter at a bluer wavelength and in March of 1996 in a RG715 filter at a redder wavelength. While these exposures are available online,¹⁴ we could not determine zero-point fluxes for stars in these bands or stars within the field with reference magnitudes in these bands, so we do not report any photometry for these images.

¹⁴ https://archive.stsci.edu/cgi-bin/dss_form

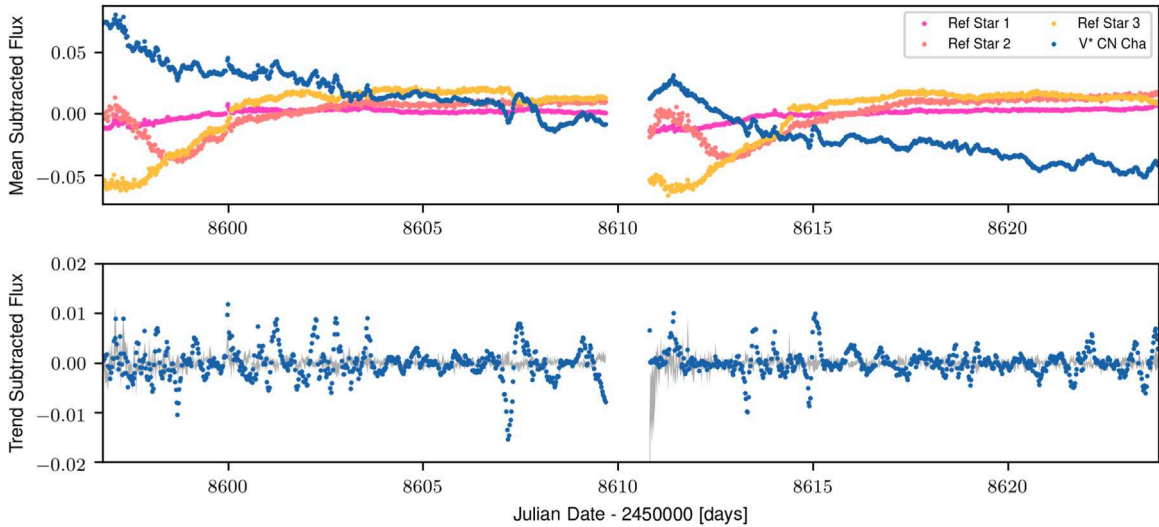


Figure 3. We show the light curve inferred from TESS photometry produced as described in Section 2.18. Top panel: the mean-subtracted flux variations (in units of the logarithm of electron counts on the CCD detector) of CN Cha (blue) are compared against three comparison stars (pink and yellow). While all four sources show variations over time, the behavior of CN Cha is distinct from that of the reference stars, whose variations are likely due to TESS systematics. It is also clear that CN Cha is dimming over the course of the TESS observation, this dimming is still significant over periods where the reference stars are relatively constant in flux. Bottom panel: we show the flux of CN Cha (blue) with a running mean on the scale of 10 hr subtracted out so as to make the short-period variations more visible. We additionally show the 1σ variations in trend-corrected flux from the “ensemble” of the three reference stars as the gray-shaded region. It is clear that CN Cha has significant variation on the timescale of hours.

2.7. Deep Near-Infrared Survey

The Deep Near-Infrared Survey (DENIS) was a near-infrared survey operated out of the 1 m European Southern Observatory (ESO) telescope at La Silla, Chile (Epchtein et al. 1994). DENIS also gathered multiple observations of 355 million objects in each of its three bands: I ($0.8\ \mu\text{m}$), J ($1.2\ \mu\text{m}$), and K_s ($2.1\ \mu\text{m}$) (Fouqué et al. 2000). The DENIS survey observed CN Cha on three occasions from early 1999 to January of 2000 (DENIS Consortium 2005). However all observations in the J and K_s bands were brighter than the saturation limits of the DENIS survey, as was the second observation in the I band (Epchtein et al. 1994). For this reason we only use the two DENIS observations in the I band that were unsaturated. The range of flux that these observations cover is indicated in Figure 2.

2.8. Two-Micron All-Sky Survey

The Two-Micron All-Sky Survey (2MASS) was a full-sky survey in the J , H ($1.65\ \mu\text{m}$), and K_s bands operated on two dedicated 1.3 m telescopes at Mount Hopkins, Arizona and Cerro Tololo, Chile between 1997 June and 2001 February. The survey was estimated to be complete at signal-to-noise greater than 10 to approximately 15th magnitude in the H band. CN Cha was observed by 2MASS on 2000 January 13th (Skrutskie et al. 2006).

2.9. All-Sky Automated Survey

The All-Sky Automated Survey (ASAS) was designed as one of the first major efforts to obtain time-variability information on a large fraction of stars in the night sky (Pojmanski 1997). The ASAS survey obtained observations of CN Cha from 2000 November 21 to 2009 November 27 (Pojmanski 2002a, 2003, 2004). The light curve for the star is reported in two separate online catalogs that we were able to find. We use both catalogs, as they seem consistent with one another, and discuss the relation between them in Appendix.

The ASAS survey used these observations to identify CN Cha as a Mira variable star with a period of 260 days. The light curves from the original ASAS survey were reanalyzed in 2016 by Vogt et al. (2016) with more rigorous analysis techniques. In this reanalysis, CN Cha has a period of 253.24 days.

2.10. Radial Velocity Experiment

The Radial Velocity Experiment (RAVE) obtained a spectrum of CN Cha on 2010 February 20 (Kordopatis et al. 2013). The RAVE team very generously provided us with a copy of the spectrum taken in 2010 in advance of its planned release in 2020. The star’s spectrum shows clearly very low T_{eff} with broad TiO absorption features, consistent with a cool giant star. Unfortunately, the RAVE pipeline fits seem to have hit the boundary of the stellar parameter grids and therefore extracted stellar model parameters are not reliable (RAVE Team 2020, private communication). We leave the analysis of this spectrum to future work.

2.11. AKARI

The AKARI satellite was a Japanese astronomical all-sky survey in six infrared bands between $9\ \mu\text{m}$ and $200\ \mu\text{m}$ (Murakami et al. 2007). This survey was carried out between 2006 May 6 and 2007 August 8, during which time AKARI observed CN Cha and reported measurements for its flux in band passes centered at $9\ \mu\text{m}$ and $18\ \mu\text{m}$ and non-detections in all longer wavelength bands at 65, 90, 140, and $160\ \mu\text{m}$ (Yamamura et al. 2010).

2.12. Swift/UVOT

The Neil Gehrels Swift Observatory is a space observatory mission directed by NASA and built for the main purpose of detecting Gamma Ray Bursts (GRBs).¹⁵ The Swift UV and Optical Telescope (UVOT) observes the sky in four broad

¹⁵ <https://swift.gsfc.nasa.gov>

bands *UVW2* (1928 Å), *UVM2* (2246 Å), *UVW1* (2600 Å), and *U* (3465 Å) spanning the far to near-UV (Kuin et al. 2015). Data from UVOT was used to create the Swift/UVOT Serendipitous Source Catalog (SUVOTSSC), which catalogs a number of point sources in these bands (Page et al. 2014; Yershov 2014). CN Cha was observed six times by UVOT and we display the photometry from 2008 January in Figure 2.

2.13. APASS

The American Association of Variable Star Observers (AAVSO) Photometric All-Sky Survey (APASS) is a survey performed primarily in five bands: Johnson-Cousins *B* and *V* as well as Sloan *g'*, *r'*, and *i'* (Henden et al. 2009; Henden & Munari 2014). With their latest data release (DR10), there have been observations in the Sloan *u'* and *z'* bands as well as the *Y* band (Henden et al. 2018).

The epoch photometry from the APASS survey is also available through the AAVSO's International Variable Star Index (VSX).¹⁶ The VSX data includes *g'*, *r'*, and *i'* values for five epochs spanning from 2012 March 26 to 2014 April 5. The last three of these five epochs additionally have photometric measurements in the *B*, *V*, and *z'* bands. This photometric information was essential in constraining the behavior of CN Cha in an otherwise unobserved portion of its evolution as shown in Figure 2, where we can see its rise in optical luminosity from magnitude 16 to magnitude 8.

As we have the most information on the long-term variability of CN Cha in the *V* band, in order to have the largest number of epochs for comparison we use the APASS *g'* and *r'* bands to calculate an effective *V* band magnitude for the two earliest observations where there is no directly measured *V* band magnitude. The calibration we use is

$$V = g' - 0.52(g' - r') - 0.03 \text{ mag}, \quad (2)$$

which comes from Jester et al. (2005). It should be noted that this relation is derived using relatively well-behaved stars, so our use of the relation here for an object with such an abnormal SED (as shown in Figure 2) may not be ideal.

2.14. WISE

The Wide-field Infrared Survey Explorer (WISE) observed CN Cha on 2010 February 20 and 27, using about 30 individual exposures for their total photometric measurement (Cutri et al. 2012;). CN Cha was observed as a strong point source in all four of the WISE photometric bands *W1*, *W2*, *W3*, and *W4*. More than 15% of the pixels were determined to be saturated in the *W1* and *W2* bands (Cutri et al. 2012), thus we quote the measured photometry as lower bounds of the true flux at these wavelengths.

Once the cryogenic cooling on the WISE satellite became non-operational, observations could no longer be performed in the *W3* and *W4* bands, but the mission continued to observe the infrared sky in the *W1* and *W2* bands as the NEOWISE mission (Mainzer et al. 2014). While the NEOWISE mission observed CN Cha on several occasions, it was highly saturated in all exposures, having derived mean magnitudes of *W1* = 3.22 and *W2* = 2.19 mag, which is much brighter than the objects that NEOWISE was designed to measure (Mainzer et al. 2014).

Regardless, we provide a further analysis of the light curve created by NEOWISE in Section 4.3.

2.15. SkyMapper

The SkyMapper Survey is a photometric survey of the southern sky performed on the 1.35 m telescope at the Siding Spring Observatory in Australia. The survey uses six photometric bands, which are *u*, *v*, *g*, *r*, *i*, and *z*, and they had their first data release in January of 2018 (Wolf et al. 2018). CN Cha was not included in the stellar catalog of SkyMapper DR1 despite being observed in all six bands on three separate nights in 2014 March, likely due to its saturation in many of the bands.

To make use of these data, we downloaded the images from the SkyMapper website¹⁷ and derived magnitudes based on reference to a nearby star, which is reported in the stellar catalog of SkyMapper with magnitudes in all six bands. The reference star we used has associated SkyMapper ID 285891100 and Gaia ID 5199012694392465792. We then derived fluxes for CN Cha in the SkyMapper bands by using the AB zero-point magnitudes reported in Wolf et al. (2018). CN Cha was saturated in all SkyMapper exposures in the *g*, *r*, *i*, and *z* bands, so we only report measurements for the *u* and *v* bands. We derive $u = 9.08 \pm 0.20$ and $v = 8.59 \pm 0.17$ mag (in AB system) for CN Cha.

2.16. ASAS-SN

The All-Sky Automated Survey for Supernovae (ASAS-SN) is a program aimed at searching for and monitoring bright supernovae in the variable sky (Shappee et al. 2014). The project currently consists of 24 telescopes that monitor a large fraction of the night sky for variable sources down to magnitudes $V < 18$.¹⁸ CN Cha was observed by ASAS-SN from 2016 December 28 to 2018 February 9 in the *V* band and was placed in the ASAS-SN variable star catalog as a Gamma Cassiopeia (GCAS) variable star (Jayasinghe et al. 2020). Based on communication with the ASAS-SN team, this classification is probably not correct, in part due to the fact that CN Cha is quite bright for the range of magnitudes that ASAS-SN observes (Kochanek et al. 2017).

It is clear from the ASAS-SN light curve, shown on the right side of the top panel of Figure 2, that CN Cha remained at an almost constant magnitude of $V \sim 8$ mag for a significant portion of the ASAS-SN monitoring before beginning to dim by 2017 September. One may be worried that the constant magnitude region is simply due to saturation of the ASAS-SN observations. Again, based on communication with the ASAS-SN team, along with the fact that this measured magnitude is consistent with the *G* band magnitude measured by Gaia, this flat portion being due to saturation seems unlikely, but is still plausible (private communication, Tharindu Jayasinghe, ASAS-SN Team). We further analyze these data in Section 4.2.

2.17. Dark Energy Camera

CN Cha was observed by the Dark Energy Camera (DECam), an instrument created as part of the Dark Energy Survey (DES, Abbott et al. 2018) and mounted on the Blanco

¹⁶ <https://www.aavso.org/vsx/>

¹⁷ <http://skymapper.anu.edu.au/image-cutout/>

¹⁸ <http://www.astronomy.ohio-state.edu/asasn/index.shtml>

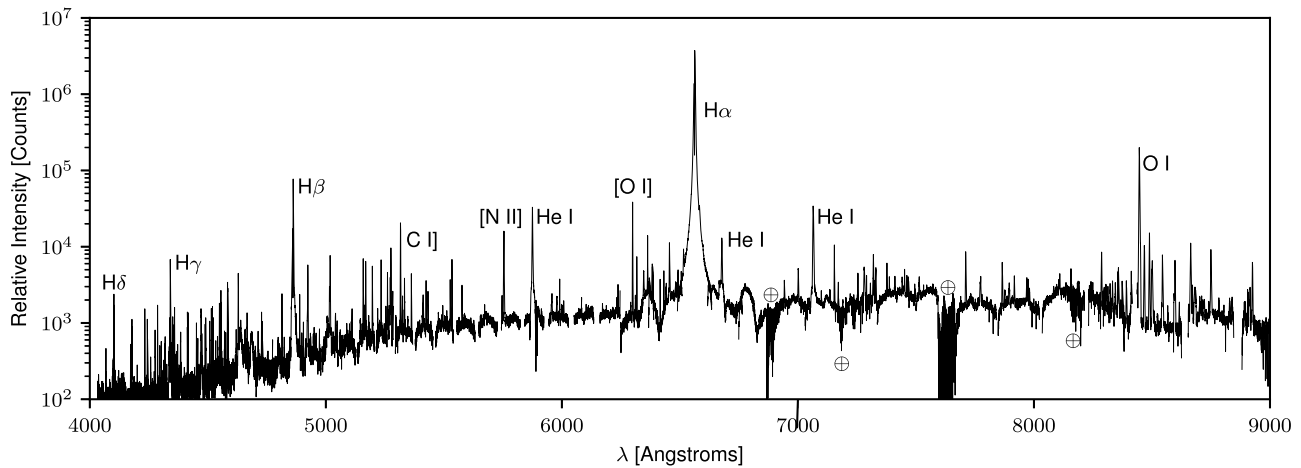


Figure 4. We show the optical spectrum of CN Cha observed using the Du Pont Echelle (DPE) instrument on the Irénée du Pont 2.5 m telescope at Las Campanas Observatory on 2019 March 12. This spectrum has been reduced using the CarPy software and the reduced spectrum has been divided by a normalized flat in order to account for the changing sensitivity as a function of wavelength. The vertical axis is then given at an arbitrary normalization that can roughly be interpreted as electron counts. We tentatively identify several of the emission lines, including several helium, nitrogen, and oxygen lines and three of the hydrogen Balmer lines. Regions of significant telluric absorption are indicated with the $\oplus$ sign.

4 m telescope at Cerro Tololo Inter-American Observatory (CTIO) in Chile. The observation was done as part of a search for substructure around the Magellanic clouds.¹⁹ One exposure in the DES r band, taken in 2016 June is shown in Figure 1. CN Cha is clearly saturated on that image, so we did not derive photometric measurements.

2.18. TESS

The Transiting Exoplanet Survey Satellite (TESS) is a photometric survey satellite launched on 2018 April 18 whose main scientific purpose is the discovery of exoplanets (Ricker et al. 2015). TESS observes single patches of the sky in 27.4 days segments, providing photometric measurements at a cadence of about 30 minutes. CN Cha was observed in TESS Sector 11 from 2019 April 22 to May 20.

To determine the light curve (LC) of the star, we have extracted the 100×100 pixels cutouts from the calibrated full-frame images. The star has ~ 1200 observations. Since CN Cha has a nearby bright star $\sim 30''$ (1.4 pixels) away, we decided to do a simultaneous PSF fit of the 16×16 pixel area around the star throughout all the epochs, while including three nearby stars in the model, and allowing for small rotation and translation between individual images. The PSF model was a linear combination of two Gaussian curves oversampled at five times the pixel size of the image, that was allowed to vary from frame to frame. From this modeling, we extract the light curve (LC) of CN Cha and of nearby stars. While the proper detrending and calibration of the LCs is beyond the scope of this paper, the extracted LCs of other stars in the field of CN Cha are quite constant with a typical median absolute deviation (MAD) of the LC of $< 0.5\%$, the MAD of the CN Cha LC is 2% and it shows an almost linear decline in flux by about 10% throughout the time of the TESS observations (27 days). That dimming corresponds to roughly a rate of 1.4 mag yr^{-1} (see Figure 3), which is consistent with the dimming starting at approximately Julian date 2458000 in the ASAS-SN light curve (see Figure 5).

3. Du Pont Echelle Spectroscopy

The Gaia observations outlined above motivated our spectroscopic follow-up with the Echelle Spectrograph on the 2.5 m Irénée du Pont telescope at Las Campanas Observatory.²⁰ This instrument provides coverage in the optical from 4000 to 9000 Å at $R \sim 40,000$. Our observations consisted of six exposures over 10 minutes from UTC 3:54 to UTC 4:08 on 2019 March 12 (Julian date 2,458,554.672). We performed two 90 s exposures, three 60 s exposures, and one 10 s exposure. The data was reduced using the Du Pont Echelle (DPE) reduction pipeline (publicly available online²¹). This reduction package makes use of the Carnegie Python Distribution (CarPy), which is also made available on the Carnegie Observatories Software Repository²² (Kelson et al. 2000; Kelson 2003).

The composite spectrum produced using all six exposures is shown in Figure 4. It is clear that the $H\alpha$ emission is quite broad and high equivalent width, making up the majority of the star’s luminosity in the optical. While there are several other emission lines, the continuum flux is only detected at a signal-to-noise greater than 10 redward of approximately $\lambda \sim 4800$ Å. Among the more obvious emission lines are the $H\beta$ and $H\gamma$ lines, which stand out clearly. We also tentatively identify several of the other prominent emission lines with oxygen ([O I] 6300 and 6360 along with O I 8446, which is the second strongest line), helium (He I 5875, 6678, and 7065), nitrogen ([N II] 5756 and 6585), and carbon (C I] 5317) and label these in Figure 4.

4. Further Data Analysis

Here we present some calculations performed on the archival data and Du Pont spectrum that go beyond simply presenting the data as provided by the various survey teams. We title our subsections by the key physical take-aways.

²⁰ www.lco.cl/telescopes-information/irenee-du-pont

²¹ code.obs.carnegiescience.edu/dpe-pipeline

²² code.obs.carnegiescience.edu/carnegie-python-distribution

¹⁹ <https://www.noao.edu/noaoprop/abstract.mpl?2016A-0366>

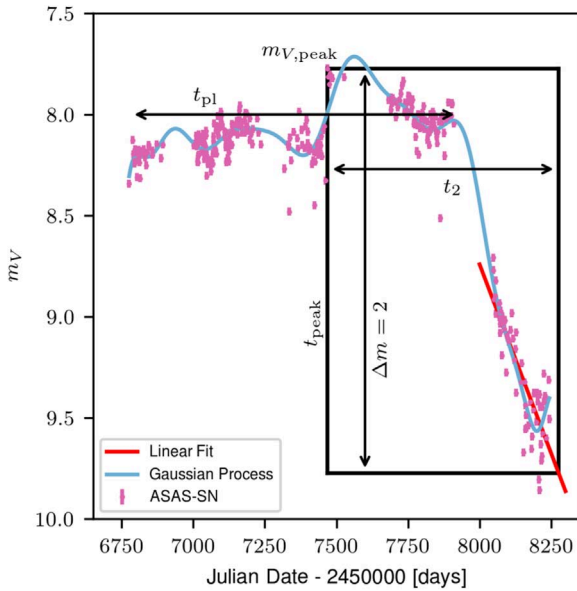


Figure 5. We present an in-depth look at the ASAS-SN light curve (shown in purple) for CN Cha. We indicate a timescale for the plateau phase of the light curve, in the context of slow novae, which spans the first four epochs on contiguous ASAS-SN observations of CN Cha. We indicate the peak apparent magnitude (top of black box) time at which the peak magnitude is reached (left side of box), two magnitudes above the peak magnitude (bottom of box) and the time at which this occurs (right side of box). To find the time at which the light curve has decayed by two magnitudes we fit a linear decay to the ASAS-SN light curve points taken after Julian Day 2,458,000 and find where this fit (shown in red) crosses 2 mag above the peak apparent magnitude. We additionally we show a Gaussian Process fit to the data in light blue.

4.1. Mira Past

It seems safe to conclude from several points of observation in the archival data laid out above that CN Cha was a Mira long-period variable star in its recent past. This is supported by its ASAS light curve shown at the left side of the top panel of Figure 2 and in detail in Figure A1. This observation is further supported by its archival identification as a Mira variable when it was first identified (Hoffmeister 1963), its identification as a likely variable by IRAS (Neugebauer et al. 1984), the NIR colors, e.g., in 2MASS $J - K = 1.69$ mag (Epchtein et al. 1994; Skrutskie et al. 2006), and the RAVE spectrum gathered in early 2010 and discussed briefly in Section 2.10 (Kordopatis et al. 2013).

4.2. Slowly Decaying Outburst

Here we perform a more detailed analysis of the ASAS-SN light curve that was presented in Section 2.16. In particular, we want to note the timescale of the plateau phase t_{pl} , which is typically used as a comparison for slow novae, and the peak absolute magnitude and decay time, which are observables typically used to compare classical novae events to one another. The latter quantities are usually quantified using the extinction-corrected, peak absolute magnitude in the V band ($M_{V,peak}^0$ mag) and the time it takes for the event to decay by two magnitudes from this peak (t_2). Our analysis is illustrated in Figure 5.

For the timescale of the plateau phase, we use a rough estimate spanning from the beginning of the ASAS-SN observations (which are roughly indicated as the beginning of the plateau phase by Figure 2) and the end of the fourth contiguous epoch of ASAS-SN observations, which again is

consistent with the onset of decay in the optical bands. This estimate provides $t_{pl} = 1153$ days, or just over 3 yr. We measure t_2 using a linear fit in magnitude space to the light curve for Julian Date greater than 2,458,000 days (the linear decay region). We then find the point at which this linear fit crosses two magnitudes decay from the peak observed apparent magnitude and define t_2 as the difference between that crossing time and the time that the peak magnitude was reached. We find $t_2 = 807$ days.

We find $M_{V,peak}^0$ using the peak apparent magnitude observed in the ASAS-SN light curve ($m_{V,peak} = 7.773$). We then get the peak absolute magnitude using the Gaia distance

$$M_{V,peak} = m_{V,peak} - 5 \log_{10} \left(\frac{d}{10 \text{ pc}} \right) = -4.74 \quad (3)$$

and correcting for extinction using $E(B - V) = 0.17$ mag from Schlegel et al. (1998) and $R_V = 3.1$ this gives us

$$M_{V,peak}^0 = M_{V,peak} - R_V E(B - V) = -5.27. \quad (4)$$

Finally, in order to get an estimate of the total amount of energy emitted during this outburst event, we calculate the total energy in the V band emitted over the course of the ASAS-SN observations. We do this by fitting a Gaussian Process (GP) regression, available through the `scikit learn` python package, to the extinction-corrected ($E(B - V) = 0.17$) ASAS-SN light curve (Pedregosa et al. 2011). This GP fit is shown (in observed magnitude space) in Figure 5. We then integrate the fit GP over the full course of the ASAS-SN observation. Using this integral in combination with the Gaia distance and an effective wavelength for the V band of $\lambda_V = 5450 \text{ \AA}$ we derive the total energy in the V band of $E_V = 2.8 \times 10^{45}$ ergs.

4.3. Continued Bright Emission in MIR

Though the observations made by the NEOWISE mission of CN Cha were saturated in all observations, we thought it useful to study the light curve if only to get a lower bound on the total amount of energy emitted in the W1 and W2 bands over the course of the NEOWISE observations. We restricted our analysis to photometric points from the NEOWISE catalog that were within $3''.6$ ($0''.001$) from CN Cha and included the original WISE observations. The NEOWISE data are grouped, with several observations occurring over a few days semi-annually (roughly 2014–2018 February and August). As these data are quite variable, we average all data points within each group to create a smoother overall light curve. We then integrate this light curve in a piecewise linear manner to get a total flux emitted over the observed period. We find 4.1×10^{45} ergs and 5.4×10^{45} ergs as lower bounds on the energy emitted in the W1 and W2 bands respectively. Additionally, there is no indication that the dimming that is seen in the V band is accompanied by dimming in the NEOWISE bands.

4.4. Balmer Lines: Massive Outflow and Self-absorption

It is clear from the spectrum that the star has an abnormally large Balmer decrement. In this subsection, we will explore these lines in some detail. We first give quantitative measurements of the Balmer decrement and then use these numbers to infer the extinction needed to explain these decrements solely through dust absorption. We give

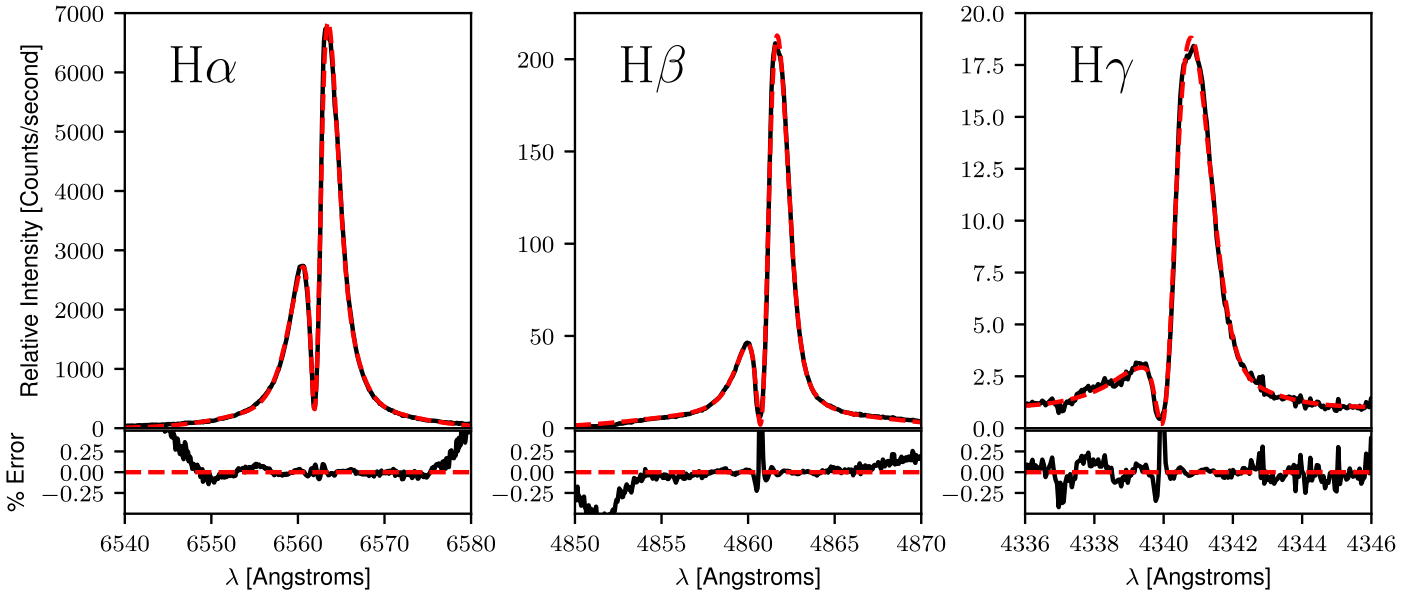


Figure 6. A zoom-in of the Balmer lines seen clearly in the spectrum shown in Figure 4. The data is shown in black while our five-component Gaussian fit to the lines is shown as a red dashed line. From left to right, we show the $H\alpha$, $H\beta$, and $H\gamma$ lines. All lines are shown on the same scale, though the $H\alpha$ line is reduced purely from a 10 s exposure, as it was saturated in all 60 and 90 s exposures. The bottom panels show the residuals of the data relative to the five-component Gaussian fit in terms of percent error. All lines are clearly quite broad and exhibit blueshifted absorption, P-Cygni like, profiles.

phenomenological fits to the lines and the physical interpretation of these fits. Finally, we also calculate the equivalent widths of the lines.

Integrating the flux in each line and comparing we get $H\alpha/H\beta = 90.8$ (as opposed to the expected value of 3) and $H\gamma/H\beta \approx 0.08$ (typical value 0.45, Draine 2011). We note that the $H\alpha$ emission was saturated in all exposures except for the 10 s exposure, so we use this exposure alone to measure the $H\alpha/H\beta$ ratio.

This abnormal Balmer decrement would usually be indicative of extensive reddening, however, as is noted in Williams et al. (2017), nebulae commonly exhibit strong Balmer decrements, which can be due in large part to self-absorption in the higher order lines. We will come back to this in Section 4.5. As a nebula is only one of the explanations we explore below, we will nonetheless calculate the extinction that would be needed to explain the Balmer decrement that we see if there were no self-absorption. Estimating the extinction between the $H\alpha$ and $H\beta$ lines as

$$E(\beta - \alpha) = 2.5 \log_{10} \left(\frac{1}{3} \frac{F_{H\alpha, \text{obs}}}{F_{H\beta, \text{obs}}} \right), \quad (5)$$

gives $E(\beta - \alpha) = 3.70$ mag. Similarly, we estimate $E(\gamma - \beta) = 1.86$ mag. Using the $R_V = 3.1$ extinction curve from Weingartner & Draine (2001), we can infer that this extinction would imply a visual extinction of $E(B - V)_{\text{Ba}} = 3.07$ mag, where we have used the subscript Ba to indicate that this extinction measurement comes from the Balmer decrement. Using this extinction measurement to correct the absolute Gaia G band magnitude calculated in Section 2.2 would give an intrinsic magnitude of $G_0 = -14.4$ mag.

Even in the case that $E(B - V)_{\text{Ba}}$ is an accurate estimate of the true extinction, the above correction is not necessarily accurate given the abnormal distribution of light in the G band, as indicated by the Du Pont spectrum in Figure 4. However, no matter the SED, CN Cha would be intrinsically very bright if

$E(B - V)_{\text{Ba}}$ were an accurate estimate of the extinction. For this reason, it is reasonable to conclude that either $E(B - V)_{\text{Ba}}$ is not an accurate measure of the extinction or CN Cha is intrinsically as bright as a galaxy. While there have been transients with similar lifetimes and peak magnitudes as this, such as those found in the SPIRITS survey Kasliwal et al. (2017), we determine the former to be the more likely scenario.

In Figure 6, we show a zoom-in of the Balmer emission clearly visible in Figure 4. Note that while the lines are all displayed on the same scale, the $H\alpha$ line is processed from the 10 s exposure alone (since it is saturated in all other exposures) while the other lines are simply cutouts of Figure 4. It is clear from inspecting the Balmer lines that they all exhibit strong blueshifted absorption in P-Cygni-like profiles. These profiles are typically indicative of a massive outflow of gas from the star. Each line is also very broad with widths on the order of 120 km s^{-1} (see below).

Attempts to fit the lines with either single Gaussian profiles or single Voigt profiles for the emission and absorption (over the wavelength regions shown in Figure 6) were unsuccessful. In an attempt to interpret these profiles shapes, we fit a five-component Gaussian model to each line. The model is defined as

$$F_{\text{mod}}(\nu) = F_{\text{emit}}(\nu) - F_{\text{abs}}(\nu), \quad (6)$$

where $F_{\text{mod}}(\nu)$ is the model prediction as a function of frequency, $F_{\text{emit}}(\nu)$ is the component dedicated to capturing the emission profile and $F_{\text{abs}}(\nu)$ is the component meant to capture the absorption profile. We devote three Gaussian curves to the emission profile ($F_{\text{emit}}(\nu)$) and two Gaussian curves to the absorption profile ($F_{\text{abs}}(\nu)$). Within the emission or absorption profiles, respectively, all constituent Gaussian curves are enforced to have the same mean frequency, while all Gaussian curves are allowed to have different amplitudes and widths. The parameters of our model are then the central frequency of the emission profile (ν_e), the central frequency of the absorption

Table 1
The Physically Relevant Parameters of Our Five-component Gaussian Fit to the Profiles of the Balmer Lines Shown in Figure 6

Line	RV_{emit} (km s ⁻¹)	RV_{abs} (km s ⁻¹)	$(\Delta v)_{\text{FWHM, emit}}$ (km s ⁻¹)	$(\Delta v)_{\text{FWHM, abs}}$ (km s ⁻¹)	W	Absorption/Emission
H α	-14.04 ± 0.46	-34.51 ± 0.19	166.74	89.03	-1.81	0.44
H β	-9.26 ± 1.12	-29.56 ± 0.18	121.34	72.65	-0.12	0.48
H γ	3.19 ± 2.67	-28.25 ± 0.28	118.35	59.26	-0.02	0.33

Note. The first column specifies the line, columns 2 and 3 give the velocity blueshifts that were fit to each of the lines in emission and absorption, while columns 4 and 5 give the Full-Width at Half-Maximum (FWHM) of the profiles for the lines in both emission and absorption. Column 6 gives the equivalent width of the line and column 7 gives the ratio of absorption to emission in the line, both calculated as explained in the text.

profile (ν_a), the amplitudes of the Gaussian curves for the emission profile ($A_{e,1}, A_{e,2}, A_{e,3}$), the widths of the Gaussian curves for the emission profile ($\sigma_{e,1}, \sigma_{e,2}, \sigma_{e,3}$), the amplitudes of the Gaussian curves for the absorption profile ($A_{a,4}, A_{a,5}$), and the widths of the Gaussian curves for the absorption profile ($\sigma_{a,4}, \sigma_{a,5}$). This leaves us with 12 free parameters.

The physically relevant parameters of these fits are given in Table 1. The radial velocity shifts $RV_{\text{emit/abs}}$ are derived from the central frequencies of the fit profiles $\nu_{a/e}$ and the velocity Full-Width at Half-Maximum (FWHM) of the lines is derived by taking the frequency difference of the intersection of the profiles with their half-maximums. The dimensionless equivalent widths W , defined as

$$W = \int \frac{d\nu}{\nu_0} \left(1 - \frac{F_{\nu, \text{obs}}}{F_{\nu, 0}} \right), \quad (7)$$

where ν_0 is taken to be the rest frame wavelength of the particular line in air, $F_{\nu, \text{obs}}$ is the observed flux, and $F_{\nu, 0}$ is the background flux that would have been observed if the line were not present. We assume $F_{\nu, 0}$ to be a constant and go about estimating it by taking the two Echelle orders on either side of the line of interest, sigma-clipping them to remove smaller lines that may be present, and taking the mean of the resulting set of data. This method may not be entirely accurate for the H γ line, where the continuum is not detected at high signal-to-noise. Finally, the absorption-to-emission ratio given in Table 1 is estimated by integrating both the fitted emission and absorption profiles over the same wavelength regions used for the equivalent width estimates and taking the ratio of these integrals. We find the absorption to be between one-half and one-third of the emission across all lines.

4.5. N II Ratio Indicates High Density

We identify the emission lines [N II] 5756 and [N II] 6585, the latter of which sits on top of the wings of the H α line. The ratio of these emission lines can be used as a temperature diagnostic (Draine 2011). After subtracting out the background by fitting polynomials of degree 2 and 6 to the regions surrounding the [N II] 5756 and [N II] 6585 lines respectively, we sum their fluxes and find a ratio of 2.385. If we additionally account for the extinction that occurs between these two lines based on the $E(B - V)$ from Schlegel et al. (1998) the line ratio becomes:

$$\frac{[\text{N II}] 5756}{[\text{N II}] 6585} = 2.583. \quad (8)$$

While, as just mentioned, this ratio is usually used as a temperature diagnostic, at such large values it is more useful as a density diagnostic (Draine 2011). Assuming a gas temperature of $T = 10^4$ K, this ratio is indicative of an electron density of $n_e \gtrsim 10^7$ cm⁻³ (Draine 2011). These high densities would be consistent with self-absorption in the hydrogen recombination lines that we posited in the last section (Netzer 1975; Drake & Ulrich 1980).

4.6. Dynamics of a Thick-Disk Star

Here we review the dynamics of CN Cha within the Galaxy that are derived from the astrometric measurement from Gaia Collaboration et al. (2018) and a radial velocity of ~ 0 km s⁻¹, motivated by our fits to the Balmer lines given in Section 4.4. This is laid out in Figure 8.

Although CN Cha is spatially coincident on the sky with the Chamaeleon Complex, a star-forming molecular cloud complex lying between the Cha I ($\sim 2.5^\circ$ separation) and Cha III ($\sim 4^\circ$ separation) clouds, its distance as derived from its Gaia parallax ($3.18^{+0.27}_{-0.23}$ kpc) places it $20\times$ farther than the Chamaeleon Cloud (~ 180 pc, Zucker et al. 2019). Furthermore, the proper motion of CN Cha (a quantity that is in general much better measured by Gaia) is also inconsistent with the proper motions of stars known to be associated with the Chamaeleon Complex. To be precise, the measured proper motion of CN Cha is $(\mu_{\alpha*}, \mu_{\delta}) = (-5.42, 7.02)$ mas yr⁻¹ while the proper motions of HD 97300 ($(\mu_{\alpha*}, \mu_{\delta}) = (-21.01, -0.61)$ mas yr⁻¹) and HD 97048 ($(\mu_{\alpha*}, \mu_{\delta}) = (-22.44, 1.31)$ mas yr⁻¹), both YSOs known to be associated with the Cha I cloud, are significantly different with the motion of CN Cha (Luhman 2008; Gaia Collaboration et al. 2018).

In Figure 8, we compare the past orbit of CN Cha with a star known to be in the Chamaeleon Cloud Complex (HD 97048) and the orbit of the Sun. It seems quite clear that CN Cha has dynamics that are consistent with a thick-disk star, since the maximum vertical extent of its orbit is nearly 2 kpc above/below the disk plane and it is moving in a prograde fashion. More quantitatively, if we relate its vertical action $J_z = 75.3$ kpc km s⁻¹ to its age using the thin disk vertical heating study of Ting & Rix (2019), the probability that this star is from the thin disk is negligible. This demonstrates that the star is most likely from the thick disk and therefore likely more than 8 Gyr old.

4.7. Stochastic Optical Variability

We perform a power-spectrum analysis of the TESS light curve presented in Section 2.18 and displayed in Figure 3. We apply the ASTROPY Lomb-Scargle periodogram implementation

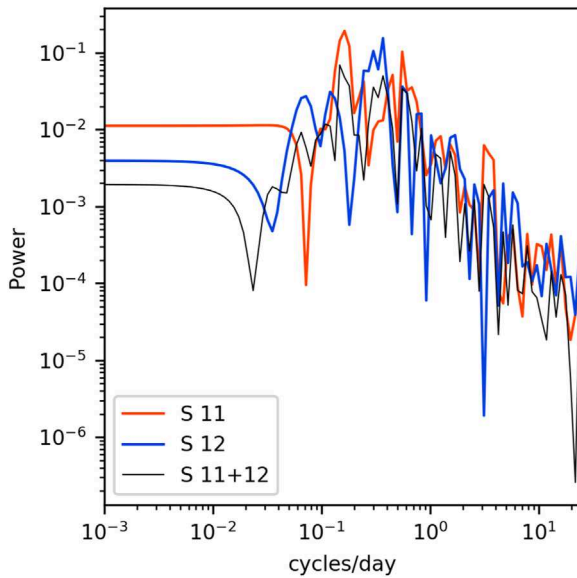


Figure 7. A power spectrum of the light curve taken by TESS, which is illustrated in Figure 7. The power spectrum of the data taken in Sector 11 of TESS is shown in red while that for Sector 12 is shown in blue. The joint power spectrum across both sectors is shown in black. All power spectra show red noise, which is indicative of stochastic variability.

(VanderPlas 2018) to the light curves from TESS Sectors 11 and 12 considered separately and together. Since strong TESS systematics occur on the timescale of its orbit (13.5 days), these are filtered out at the calibration stage of our analysis and therefore longer-timescale astrophysical variability is also removed.

The resulting power spectra are displayed in Figure 7 and the results are similar for the sectors treated independently. The spectrum decays with a pink noise or “flicker” profile above about 0.1 cycles per day; systematics on this timescale have been filtered out. The “flicker” is consistent with stochastic variability, such as the turbulent motion of a convective photosphere or an accretion disk.

5. Possible Explanations

Here we present possible explanations for the observations of this object that were outlined in Sections 2–4. We present a simple comparison summary of how these theories explain the observations in Table 2. The theories explained below are ordered from least to most likely.

5.1. Young Stellar Object

Many characteristics of CN Cha seem to suggest we have caught a Young Stellar Object in outburst. In this case, the main YSO candidates would be either a T Tauri star (Appenzeller & Mundt 1989; Miller et al. 2011), a Herbig Ae/Be star (Hillenbrand et al. 1992; Waters & Waelkens 1998), or an FU Orionis type star (Hartmann & Kenyon 1985, 1996). Indeed many new Young Stellar Objects have been discovered using Gaia in combination with archival data (Hillenbrand et al. 2018, 2019). In particular, the timescale and relatively consistent brightness in the light curve post-outburst shown in the top panel of Figure 2 seem to be quite consistent with a period of increased accretion that is commonly seen in FU Ori type stars (Hartmann & Kenyon 1996).

There are two key observations that lead us to believe that a YSO is not a sufficient explanation of the data:

1. CN Cha appears to have been a Mira-type long-period variable star from 1963 to as late as 2010, as laid out in Section 4.1. As far as we are aware, there is no YSO object that has photometric variability characteristic of a Mira.
2. The dynamics of the star, suggested by Gaia proper motions, indicate that the star is quite old. This is explained above in Section 4.6.

This evidence strongly indicates that CN Cha is not some kind of YSO. Though it is true that young stars appear everywhere in our galaxy (Price-Whelan et al. 2019), they occur with much lower frequency outside of the disk plane (Binney & Tremaine 2008; Belokurov et al. 2020). The fact that, based on Gaia dynamics, the last time that CN Cha crossed the midplane of the disk was 52.5 Myr ago is a particularly bad sign for the YSO theory given our understanding of star formation mostly occurring in dense molecular clouds that are most common in the disk midplane (McKee & Ostriker 2007).

5.2. Protoplanetary Nebula

Another possible explanation for the phenomena we have outlined above would be the onset of the protoplanetary nebula (PPN) phase at the end of an evolved AGB cycle (Kwok 1993). While the observations of a massive outflow, combined with CN Cha being luminous in the infrared, and the presence of strong hydrogen recombination lines, are consistent with a planetary nebula, the evolutionary timescales exhibited by CN Cha are simply too short to be associated with a planetary nebula (Marigo et al. 2001, 2004). In particular, PPN theoretically should not show significant dimming from their peak magnitudes until several thousand years after the onset of the PPN phase (Marigo et al. 2001).

Additionally, our spectra of CN Cha exhibit no signs of the [O III]5007 Å line that is ubiquitously present in all known planetary nebulae (Méndez 2017). This would tend to indicate that our system is much cooler than the typical protoplanetary nebula.

5.3. Symbiotic Nova

A symbiotic binary star is a binary star system consisting of an evolved star, usually a red giant branch (RGB) star or asymptotic giant branch (AGB) star, and a hot, compact companion, usually a white dwarf (Podsiadlowski & Mohamed 2007; Mohamed & Podsiadlowski 2012). Those symbiotic binaries with RGB primaries are known as S-type systems while those with AGB companions, which are typically Mira variables and are surrounded by a dusty envelope, are known as D-type systems. A systematic, all-sky search for these objects was recently carried out by Akas et al. (2019), characterizing the relative proportion of these types. Binary systems with one component consisting of a white dwarf are thought to be the progenitors of classical novae, wherein hydrogen-rich material that is accreted from the binary companion onto the white dwarf triggers a thermonuclear runaway (TNR, Schatzman 1949, 1951; Gurevitch & Lebedinsky 1957; Cameron 1959; Starrfield et al. 1972, 2016). Though most classical novae are thought to have late-type main-sequence star companions (Darnley & Henze 2020), there are a few nova

Table 2
A Simple Comparison between the Key Observations We Have Described and the Theories We Have Proposed for Explaining Them

Theory	Mira Variability	UV Excess	Outburst Timescale	Balmer Decrement	Dynamics/ Age	Balmer Line Strength
YSO	X	✓	✓	X	X	?
pPN	✓	?	X	?	✓	?
Symbiotic Nova	✓	✓	✓	?	✓	?

Note. A ✓ indicates that the theory explains the observation, an X indicates that the theory does not explain the observation, and a ? indicates that it is uncertain whether or not the theory explains the observation.

systems known to contain evolved-star companions, and they particularly seem to make up the majority of known recurrent novae (RN), which have repeating outbursts on human-measurable timescales (Patat et al. 2011; Darnley et al. 2012).

There are many characteristics that make the observations of CN Cha consistent with a nova occurring in a symbiotic binary system. The star clearly was classified as a Mira-type variable (an evolved AGB star; Catelan & Smith 2015) for decades, explained in Section 4.1. The symbiotic binary theory would also explain the UV excess exhibited by the UVOT observations, visible in the bottom left panel of Figure 2, and fit by a 10,000 K blackbody. The P-Cygni-like absorption profiles exhibited in Section 4.4 would also be well explained by the dusty envelopes that are often observed in D-type symbiotic binaries. This would be consistent with the dynamics of the star outlined in Figure 8, which indicate that it is an old thick-disk star. The stochastic variability observed in the TESS light curve and analyzed in detail in Section 4.7 is also well explained by an accretion disk around the white-dwarf companion.

The plateau optical luminosity for the light curve of CN Cha as presented by the ASAS-SN data is also typical for symbiotic novae, especially for the class of symbiotic novae known as “slow novae” of which PU Vul is one prototypical examples along with a handful of others Mikolajewska (2010). On the other hand, CN Cha has quite a short plateau timescale compared with other members of this class. We discuss this further below.

6. Discussion

Of the explanations presented in Section 5, we believe that a symbiotic “slow” novae is most consistent with the wide range of observations we have gathered on CN Cha. In this section, we briefly discuss the implications for the interpretation of CN Cha as a symbiotic nova.

One of the key empirical insights that has resulted from the study of classical novae is the relationship between their peak magnitudes and the timescale over which they decay, known as the Maximum Magnitude Rate of Decline (MMRD) relation (Downes & Duerbeck 2000). This empirical relationship shows that brighter novae decay more quickly to their pre-nova magnitudes. The mass of the white dwarf in the system is thought to drive the MMRD (e.g., Livio 1992).

Though there are far fewer symbiotic novae known than classical novae, attempts have also been made at studying the relationship between peak magnitude and duration of outburst for the slow symbiotic novae. If we take the peak absolute magnitude in the *V* band of CN Cha to be $M_{V,\text{peak}}^0 = -5.27$ mag and its timescale to decay from this peak by two magnitudes to be $t_2 = 807$ days (calculations outlined in

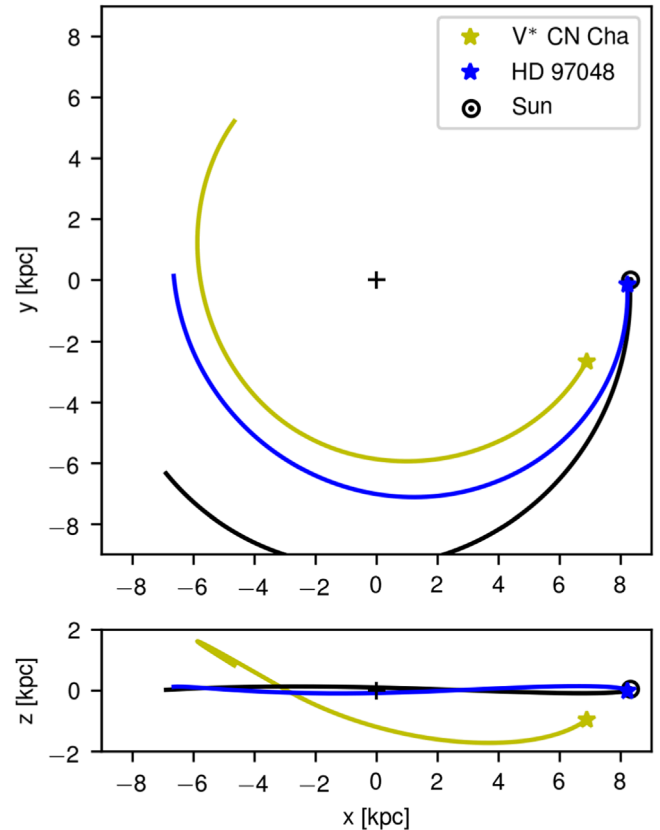


Figure 8. We show the galactic orbit of the star CN Cha (shown as a yellow star) along with the orbits of HD 97048 (a YSO known to be associated with the Chamaeleon Complex) in blue and the Sun in black. The top panel looks down up the disk plane, face-on while the bottom panel looks edge-on. The Galaxy is centered on the center of both panels. The orbits were integrated 100 Myr back in time using *galpy* and the 2014 Milky Way Potential from Bovy (2015).

Section 4.2), the outburst we observe would be one of the longest decays currently measured for classical novae. Symbiotic novae, on the other hand are typically compared on the timescale of their “plateau.” The plateau in CN Cha and its timescale of about 3 yr would be shorter than any plateau measured for other symbiotic novae by a factor of 3 (Mikolajewska 2010). In terms of peak brightness, the *V* band measurements that we obtained from the archive make CN Cha quite consistent with the plateau luminosity of PU Vul in the *V* band, which itself lies squarely in the middle of the population of slow symbiotic novae (Mikolajewska 2010).

Symbiotic binary systems are also thought to be possible progenitors for SNe Ia (Hachisu et al. 1999; Patat et al. 2011;

Mikołajewska 2013). Recent work has even suggested that mass accretion onto the compact companion via so-called Wind Roche–Lobe Overflow (WRLOF) could make these binary systems even more likely to create supernovae (Mohamed & Podsiadlowski 2007; Ikiewicz et al. 2019). In this context, CN Cha may be able to provide interesting insights into the viability of this progenitor channel for SNe Ia.

7. Summary and Follow-up

We have presented and organized the various archival data on the stellar object CN Cha in the constellation of Chamaeleon. We used these observations to show that the optical light from CN Cha was dominated by a Mira long-period variable star from 1963 until 2013, at which point the star seemed to enter an outburst phase. After remaining at roughly a constant optical brightness for about 3 yr, the object then entered a dimming phase, during which time it dimmed by about 1.4 magnitudes per year.

We then presented the optical spectrum that motivated our initial interest in this system. The spectrum is dominated by emission in hydrogen recombination lines, predominantly from H α . The Balmer series lines exhibit blueshifted absorption, indicative of a massive outflow from the object. The spectrum additionally presents emission lines of [He I, C I], [N II], O I, and [O I], among many others that we did not explicitly identify. Many of these lines exhibit complicated profiles.

We estimate the electron density of the environment using the ratio of [N II] forbidden lines and we find a density of $n_e \gtrsim 10^7 \text{ cm}^{-3}$. We additionally concluded that the extreme Balmer decrement seen in our spectrum was most likely not due to extinction but due to self-absorption; this is consistent with the high densities indicated by the [N II] ratio. We made phenomenological fits to the profiles of the Balmer series lines, which indicated large velocity widths ($\sim 120 \text{ km s}^{-1}$) and outflow velocities from the absorption of $\sim 20 \text{ km s}^{-1}$. We additionally provide an analysis of the time and magnitude scales of the outburst, indicating a slow (long duration), constant, optical-luminosity outburst with a brightness scale that is consistent with other known slow symbiotic novae, but a timescale that is significantly shorter than most other members of this class of objects. Finally, we showed that the dynamics of the object, as inferred from the Gaia astrometry and a radial velocity, indicate an old star, perhaps a component of the thick disk (lying 1 kpc below the disk midplane).

After considering that CN Cha may be either a young stellar object or an early phase protoplanetary nebula, we used the evidence summarized above to argue that the data are best explained by a nova occurring in a symbiotic binary system. If this is indeed the case, CN Cha would be the latest addition to the handful of known slow symbiotic novae, and would have the lowest plateau duration timescale of any member of this population.

More concrete conclusions will be drawn about this object from an in-depth analysis of the spectral data that we have gathered, including follow-up high-cadence spectra at both high and moderate resolution that we gathered in late 2019 November/early December. We reserve detailed analyses of this data to future work.

We will be also able to learn a great deal more about this object’s evolution over the course of its outburst once the Gaia BP/RP spectra (which cover almost the entirety of the outburst phase) are released to the community; this is anticipated with

Gaia DR3 planned for late 2021. Supplementing our current data with observations in the radio and X-ray will be extremely informative in discerning between the different scenarios we have proposed here.

As we pointed out in the introduction to this manuscript, there is no single database that encapsulates the vast array of human-collected astronomical data. We believe that the discovery we have described in this work presents a concrete example of the consequences of not having such a resource. Though nearly all of the information that made this object interesting was freely available to the community, it was scattered between many catalogs, interfaces, and repositories. As a result, this object was overlooked and its discovery was serendipitous. This paper by no means constitutes the sole such example (Hillenbrand et al. 2018, 2019).

This problem will only become worse in the near future with large upcoming time-domain astronomical surveys such as the Vera C. Rubin Observatory (VRO, formerly LSST) and large spectroscopic surveys such as SDSS-V, WEAVE, and 4MOST (Bonifacio et al. 2016; Kollmeier et al. 2017; Ivezić et al. 2019; de Jong et al. 2019). The importance of well-thought-out and easily searchable/accessible data will be paramount in the coming astronomical age, and that doing this correctly should be considered a priority for the community.

The desire for a tool that would facilitate the access of data across surveys of widely varying structure is certainly not a new concept. In fact, this study has made use of several such tools intended for this purpose and developed over the past generation, including the “Whole Sky Database” created by Sergey Koposov and maintained at the University of Cambridge (Koposov & Bartunov 2006), the SIMBAD database created by the Centre de Données astronomique de Strasbourg (CDS, Wenger et al. 2000), the VizieR online database of catalogs (Ochsenbein et al. 2000), and the sky visualizer Aladin Bonnarel et al. (2000). There have been several other great efforts made by researchers along these lines including the National Virtual Observatory (Szalay 2001), the CDS overall (Egret & Albrecht 1995; Genova et al. 1996), and the Mikulski Archive for Space Telescopes (MAST) (Rafelski et al. 2019), along with too many others to give a complete list here.

Despite the largely successful efforts of these programs, this work demonstrates that there are still regions of discovery that are made less accessible by the lack of a perfectly effective cross-survey search tool. Though it is difficult to find and characterize anomalous objects, these objects have the potential to make us fundamentally reconsider our understanding of the world around us. We look forward to an age of astronomy in which these discoveries can be made commonplace.

The authors would like to thank Howard Bond, Matteo Cantiello, and Adam Jermyn for pointing us toward symbiotic binary systems as a possible explanation for this object and for further useful discussions on the nature of the object. We would additionally like to thank Michael Shara for recommending relevant literature references with regard to symbiotic novae and novae in general. We would like to thank Tharindu Jayasinghe and Krzysztof Stanek for consulting us on the ASAS–SN observations. We would like to thank Arne Henden and Sara Beck for consulting us on the APASS measurements. The authors would additionally like to thank Gaspar Bakos, Adam Burrows, Trevor David, Bruce Draine, Andy Goulding, Emily Levesque, Erin Kado-Fong, Eliot Quataert, Nathan

Smith, David Spergel, Joshua Winn, and Vasily Belokurov, Wyn Evans, and the other members of the Cambridge Streams group for useful conversations and references. The authors would additionally like to thank the anonymous referee for useful comments on the draft. L.L. would like to thank the Center for Computational Astrophysics at the Flatiron Institute in New York City for their hospitality while part of this work was completed.

This work was performed in part under contract with the Jet Propulsion Laboratory (JPL) funded by NASA through the Sagan Fellowship Program executed by the NASA Exoplanet Science Institute. Support for this work was provided by NASA through Hubble Fellowship grant #51386.01 awarded to R.L.B. and grant #51425.001 awarded to Y.S.T. by the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., for NASA, under contract NAS 5-26555. S.K. is partially supported by NSF grants AST-1813881, AST-1909584, and Heising-Simons foundation grant 2018–1030.

This publication makes use of data products from the Wide-field Infrared Survey Explorer, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration.

This work has made use of data from the European Space Agency (ESA) mission Gaia (<https://www.cosmos.esa.int/gaia>), processed by the Gaia Data Processing and Analysis Consortium (DPAC, <https://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for the DPAC has been provided by national institutions, in particular the institutions participating in the Gaia Multilateral Agreement.

This research made use of NASA’s Astrophysics Data System and the SIMBAD database, operated at CDS, Strasbourg, France. Some of the data presented in this paper were obtained from the Mikulski Archive for Space Telescopes (MAST). STScI is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS

5-26555. Support for MAST for non-Hubble space telescope data is provided by the NASA Office of Space Science via grant NNX13AC07G and by other grants and contracts.

This paper made use of the Whole Sky Database (wsdb) created by Sergey Koposov and maintained at the Institute of Astronomy, Cambridge by Sergey Koposov, Vasily Belokurov and Wyn Evans with financial support from the Science & Technology Facilities Council (STFC) and the European Research Council (ERC).

Appendix ASAS Catalogues

Here we briefly discuss the two separate photometric catalogs that we found describing the light curve of CN Cha based on the ASAS survey. The first, which included photometric data from 2003 October 11, was obtained from the online catalog *Vizier* (catalog 1, Pojmanski 2002b). The second catalog is the “ASAS All Star Catalogue” (catalog 2), which is hosted on servers at the University of Washington (Pojmanski 2010). This second catalog contains photometric information over the full range of observations, with magnitudes obtained through measurement in four separate apertures. When using the data we simply take the error-weighted mean of these four measurements.

When looking at the data in catalog 2, which was taken during the same time interval as is covered in catalog 1, one would expect these data to agree with one another. However, there seem to be several data points in catalog 2 that do not appear in catalog 1 and vice versa. Additionally, there are some minor inconsistencies in magnitudes between measurements in each catalog that are meant to be measured at the same time. As we were not able to resolve which of these measurements to trust more, we use the data given in both catalogs here. A comparison of the two catalogs, over the period covered by catalog 1, is given in Figure A1.

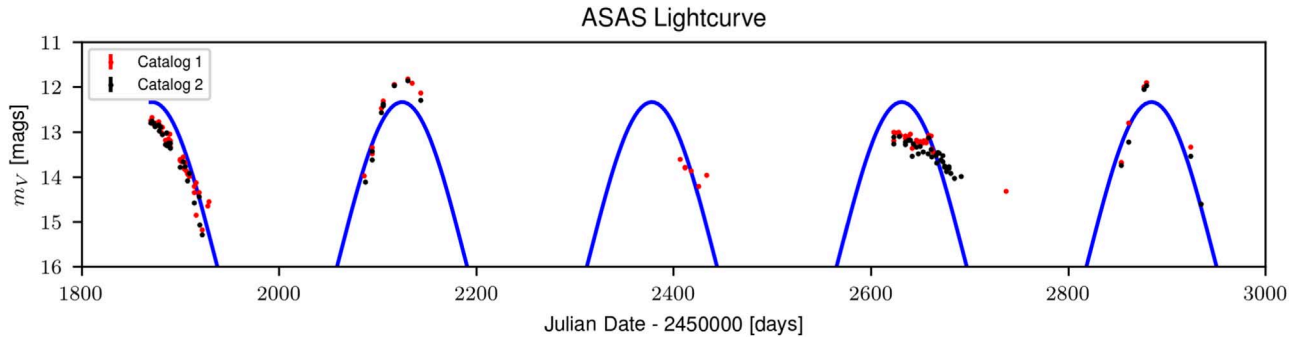


Figure A1. We show the ASAS light curve for CN Cha in the V band for both catalog 1 (red points) and catalog 2 (black points) as referred to in the text, though only over the period covered by catalog 1. Both catalogs have typical photometric errors of 0.02 magnitudes. To guide the eye, we also include a sine curve with a period of 253.24 days and a mean of 12.34 mags (as determined by the original ASAS team, blue curve) that we fit for the optimal phase shift and amplitude (3.43 mags).

ORCID iDs

Lachlan Lancaster  <https://orcid.org/0000-0002-0041-4356>
 Yuan-Sen Ting  <https://orcid.org/0000-0001-5082-9536>
 Sergey E. Koposov  <https://orcid.org/0000-0003-2644-135X>
 Benjamin J. S. Pope  <https://orcid.org/0000-0003-2595-9114>
 Rachael L. Beaton  <https://orcid.org/0000-0002-1691-8217>

References

- Abbott, T. M. C., Abdalla, F. B., Alarcon, A., et al. 2018, *PhRvD*, **98**, 043526
 Akras, S., Guzman-Ramirez, L., Leal-Ferreira, M. L., & Ramos-Larios, G. 2019, *ApJS*, **240**, 21
 Appenzeller, I., & Mundt, R. 1989, *A&A*, **1**, 291
 Belokurov, V., Sanders, J. L., Fattahi, A., et al. 2020, *MNRAS*, **494**, 3880
 Binney, J., & Tremaine, S. 2008, *Galactic Dynamics* (2nd ed.; Princeton, NJ: Princeton Univ. Press)
 Bonifacio, P., Dalton, G., Trager, S., et al. 2016, in *Proc. Annual Meeting of the French Society of Astronomy and Astrophysics*, ed. C. Reylé et al. (Paris: SAF), 267
 Bonnarel, F., Fernique, P., Bienaymé, O., et al. 2000, *A&AS*, **143**, 33
 Bovy, J. 2015, *ApJS*, **216**, 29
 Cameron, A. G. W. 1959, *ApJ*, **130**, 916
 Catelan, M., & Smith, H. A. 2015, *Pulsating Stars* (Hoboken, NJ: Wiley)
 Cutri, R. M., Wright, E. L., Conrow, T., et al. 2012, *yCat*, **2311**, 0
 Darnley, M. J., & Henze, M. 2020, *AdSpR*, **66**, 1147
 Darnley, M. J., Ribeiro, V. A. R. M., Bode, M. F., Hounsell, R. A., & Williams, R. P. 2012, *ApJ*, **746**, 61
 de Jong, R. S., Agertz, O., Berbel, A. A., et al. 2019, *Msngr*, **175**, 3
 DENIS Consortium 2005, *yCat*, **2263**, 0
 Downes, R. A., & Duerbeck, H. W. 2000, *AJ*, **120**, 2007
 Draine, B. T. 2011, *Physics of the Interstellar and Intergalactic Medium* (Princeton: Princeton Univ. Press)
 Drake, S. A., & Ulrich, R. K. 1980, *ApJS*, **42**, 351
 Egret, D., & Albrecht, M. A. 1995, *Information and on-line data in astronomy*, doi:10.1007/978-94-011-0397-8
 Epchtein, N., de Batz, B., Copet, E., et al. 1994, *Ap&SS*, **217**, 3
 ESA 1997, *The HIPPARCOS and TYCHO catalogs. Astrometric and photometric star catalogs derived from the ESA HIPPARCOS Space Astrometry Mission*, Vol. 1200 (Noordwijk: ESA)
 Fouqué, P., Chevallier, L., Cohen, M., et al. 2000, *A&AS*, **141**, 313
 Gaia Collaboration, Brown, A. G. A., Vallenari, A., et al. 2018, *A&A*, **616**, A1
 Gaia Collaboration, Prusti, T., de Bruijne, J. H. J., et al. 2016, *A&A*, **595**, A1
 Genova, F., Bartlett, J. G., Bienaymé, O., et al. 1996, *VA*, **40**, 429
 Gurevitch, L. E., & Lebedinsky, A. I. 1957, in *IAU Symposium*, Vol. 3, *Non-stable Stars*, ed. G. H. Herbig (Cambridge: Cambridge Univ. Press), 77
 Hachisu, I., Kato, M., & Nomoto, K. 1999, *ApJ*, **522**, 487
 Hartmann, L., & Kenyon, S. J. 1985, *ApJ*, **299**, 462
 Hartmann, L., & Kenyon, S. J. 1996, *ARA&A*, **34**, 207
 Henden, A., & Munari, U. 2014, *CoSkA*, **43**, 518
 Henden, A. A., Levine, S., Terrell, D., et al. 2018, *AAS Meeting*, **232**, 223.06
 Henden, A. A., Welch, D. L., Terrell, D., & Levine, S. E. 2009, *AAS Meeting*, **214**, 407.02
 Hillenbrand, L. A., Contreras Peña, C., Morrell, S., et al. 2018, *ApJ*, **869**, 146
 Hillenbrand, L. A., Reipurth, B., Connelley, M., Cutri, R. M., & Isaacson, H. 2019, *AJ*, **158**, 240
 Hillenbrand, L. A., Strom, S. E., Vrba, F. J., & Keene, J. 1992, *ApJ*, **397**, 613
 Hoffmeister, C. 1963, *VeSon*, **6**, 1
 Høg, E., Fabricius, C., Makarov, V. V., et al. 2000, *A&A*, **355**, L27
 Ilkiewicz, K., Mikołajewska, J., Belczyński, K., Wiktorowicz, G., & Karczmarek, P. 2019, *MNRAS*, **485**, 5468
 Jayasinghe, T., Stanek, K. Z., Kochanek, C. S., et al. 2020, *MNRAS*, **491**, 13
 Jester, S., Schneider, D. P., Richards, G. T., et al. 2005, *AJ*, **130**, 873
 Ivezić, Ž., Kahn, S. M., Tyson, J. A., et al. 2019, *ApJ*, **873**, 111
 Kasliwal, M. M., Bally, J., Masci, F., et al. 2017, *ApJ*, **839**, 88
 Kelson, D. D. 2003, *PASP*, **115**, 688
 Kelson, D. D., Illingworth, G. D., van Dokkum, P. G., & Franx, M. 2000, *ApJ*, **531**, 159
 Kochanek, C. S., Shappee, B. J., Stanek, K. Z., et al. 2017, *PASP*, **129**, 104502
 Kollmeier, J. A., Zasowski, G., Rix, H.-W., et al. 2017, arXiv:1711.03234
 Koposov, S., & Bartunov, O. 2006, in *ASP Conf. Ser. 351, Astronomical Data Analysis Software and Systems XV*, ed. C. Gabriel et al. (San Francisco, CA: ASP), 735
 Kordopatis, G., Gilmore, G., Steinmetz, M., et al. 2013, *AJ*, **146**, 134
 Kuin, N. P. M., Landsman, W., Breeveld, A. A., et al. 2015, *MNRAS*, **449**, 2514
 Kwok, S. 1993, *ARA&A*, **31**, 63
 Lasker, B. M., Doggett, J., McLean, B., et al. 1996, in *ASP Conf. Ser. 101, Astronomical Data Analysis Software and Systems V*, ed. G. H. Jacoby & J. Barnes (San Francisco, CA: ASP), 88
 Lasker, B. M., Sturch, C. R., McLean, B. J., et al. 1990, *AJ*, **99**, 2019
 Livio, M. 1992, in *ASP Conf. Ser. 29, White Dwarf Masses in Nova Systems and the Maximum-Magnitude vs. Rate-of-Decline Relation*, ed. N. Vogt (San Francisco, CA: ASP), 4
 Luhman, K. L. 2008, in *Handbook of Star Forming Regions*, ed. R. Reipurth, Vol. 5 (San Francisco, CA: ASP), 169
 Mainzer, A., Bauer, J., Cutri, R. M., et al. 2014, *ApJ*, **792**, 30
 Marigo, P., Girardi, L., Groenewegen, M. A. T., & Weiss, A. 2001, *A&A*, **378**, 958
 Marigo, P., Girardi, L., Weiss, A., Groenewegen, M. A. T., & Chiosi, C. 2004, *A&A*, **423**, 995
 McKee, C. F., & Ostriker, E. C. 2007, *ARA&A*, **45**, 565
 Méndez, R. H. 2017, in *IAU Symp. 323, Planetary Nebulae: Multi-Wavelength Probes of Stellar and Galactic Evolution*, ed. X. Liu, L. Stanghellini, & A. Karakas (Cambridge: Cambridge Univ. Press), 298
 Mikołajewska, J. 2010, arXiv:1011.5657
 Mikołajewska, J. 2013, in *IAU Symp. 281, Binary Paths to SNe Ia Explosions*, ed. R. Di Stefano, M. Orio, & M. Moe (Cambridge: Cambridge Univ. Press), 162
 Miller, A. A., Hillenbrand, L. A., Covey, K. R., et al. 2011, *ApJ*, **730**, 80
 Milne, A. A. 1956, *The Complete Tales of Winnie-the-Pooh* (New York: Dutton Children's Books)
 Mohamed, S., & Podsiadlowski, P. 2007, in *ASP Conf. Ser. 372, Wind Roche-Lobe Over ow: a New Mass-Transfer Mode for Wide Binaries*, ed. R. Napiwotzki & M. R. Burleigh (San Francisco, CA: ASP), 397
 Mohamed, S., & Podsiadlowski, P. 2012, *BaltA*, **21**, 88
 Murakami, H., Baba, H., Barthel, P., et al. 2007, *PASJ*, **59**, S369
 Netzer, H. 1975, *MNRAS*, **171**, 395
 Neugebauer, G., Habing, H. J., van Duinen, R., et al. 1984, *ApJL*, **278**, L1
 Ochsenbein, F., Bauer, P., & Marcout, J. 2000, *A&AS*, **143**, 23
 Page, M. J., Yershov, V., Breeveld, A., et al. 2014, in *Swift: 10 Years of Discovery*, ed. P. Caraveo et al. (Trieste: Sissa), 37, <https://pos.sissa.it/233/037/pdf>
 Patat, F., Chugai, N. N., Podsiadlowski, P., et al. 2011, *A&A*, **530**, A63
 Pedregosa, F., Varoquaux, G., Gramfort, A., et al. 2011, *JMLR*, **12**, 2825
 Podsiadlowski, P., & Mohamed, S. 2007, *BaltA*, **16**, 26
 Pojmanski, G. 1997, *AcA*, **47**, 467
 Pojmanski, G. 2002a, *AcA*, **52**, 397
 Pojmanski, G. 2002b, *ASAS Variable Stars in Southern hemisphere*, <http://vizier.u-strasbg.fr/viz-bin/VizieR-3?-source=II/264/var>
 Pojmanski, G. 2003, *AcA*, **53**, 341
 Pojmanski, G. 2004, arXiv:astro-ph/0401125
 Pojmanski, G. 2010, PhD thesis, University of Washington, <http://www.astrouw.edu.pl/asas/?page=aasc>
 Price-Whelan, A. M., Nidever, D. L., Choi, Y., et al. 2019, *ApJ*, **887**, 19
 Rafelski, M., Cherinka, B., Smith, A. M., Peek, J., & Hargis, J. 2019, *AAS Meeting*, **233**, 157.30
 Ricker, G. R., Winn, J. N., Vanderspek, R., et al. 2015, *JATIS*, **1**, 014003
 Schatzman, E. 1949, *AnAp*, **12**, 281
 Schatzman, E. 1951, *AnAp*, **14**, 294
 Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, **500**, 525
 Shappee, B. J., Prieto, J. L., Grupe, D., et al. 2014, *ApJ*, **788**, 48
 Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, **131**, 1163
 Starrfield, S., Iliadis, C., & Hix, W. R. 2016, *PASP*, **128**, 051001
 Starrfield, S., Truran, J. W., Sparks, W. M., & Kutter, G. S. 1972, *ApJ*, **176**, 169
 Szalay, A. S. 2001, in *ASP Conf. Ser. 238, Astronomical Data Analysis Software and Systems X*, ed. J. Harnden, F. A. Primini, & H. E. Payne (San Francisco, CA: ASP), 3
 Ting, Y.-S., & Rix, H.-W. 2019, *ApJ*, **878**, 21
 VanderPlas, J. T. 2018, *ApJS*, **236**, 16
 Vogt, N., Contreras-Quijada, A., Fuentes-Morales, I., et al. 2016, *ApJS*, **227**, 6
 Waters, L. B. F. M., & Waelkens, C. 1998, *ARA&A*, **36**, 233
 Weingartner, J. C., & Draine, B. T. 2001, *ApJ*, **548**, 296
 Wenger, M., Ochsenbein, F., Egret, D., et al. 2000, *A&AS*, **143**, 9
 Williams, S. C., Darnley, M. J., & Henze, M. 2017, *MNRAS*, **472**, 1300
 Wolf, C., Onken, C. A., Luvaul, L. C., et al. 2018, *PASA*, **35**, e010
 Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, **140**, 1868
 Yamamura, I., Makiuti, S., Ikeda, N., et al. 2010, *yCat*, **2298**, 0
 Yershov, V. N. 2014, *Ap&SS*, **354**, 97
 Zucker, C., Speagle, J. S., Schlafly, E. F., et al. 2019, *ApJ*, **879**, 125