



The Effect of Thermal Pressure on Collisionless Magnetic Reconnection Rate

Xiaocan Li and Yi-Hsin Liu

Dartmouth College, Hanover, NH 03750, USA; Xiaocan.Li@dartmouth.edu

Received 2021 January 22; revised 2021 March 16; accepted 2021 March 31; published 2021 May 17

Abstract

Modeling collisionless magnetic reconnection rate is an outstanding challenge in basic plasma physics research. While the seemingly universal rate of an order $\mathcal{O}(0.1)$ is often reported in the low- β regime, it is not clear how reconnection rate scales with a higher plasma β . Due to the complexity of the pressure tensor, the available reconnection rate model is limited to the low plasma- β regime, where the thermal pressure is arguably negligible. However, the thermal pressure effect becomes important when $\beta \gtrsim \mathcal{O}(1)$. Using first-principle kinetic simulations, we show that both the reconnection rate and outflow speed drop as β gets larger. A simple analytical framework is derived to take account of the self-generated pressure anisotropy and pressure gradient in the force balance around the diffusion region, explaining the varying trend of key quantities and reconnection rates in these simulations with different β . The predicted scaling of the normalized reconnection rate is $\simeq \mathcal{O}(0.1/\sqrt{\beta_{i0}})$ in the high- β limit, where β_{i0} is the ion β of the inflow plasma.

Unified Astronomy Thesaurus concepts: Plasma astrophysics (1261); Plasma physics (2089); Heliosphere (711); Intergalactic medium (813); Galactic center (565)

1. Introduction

Magnetic reconnection is a ubiquitous fundamental plasma process that reorganizes the magnetic topology and releases the magnetic energy into plasma kinetic energies (Zweibel & Yamada 2009). It occurs in laboratory experiments (Yamada et al. 2006) and confined fusion devices (Yamada et al. 1994) and drives explosive magnetic energy release in space (Schindler 1974; Øieroset et al. 2002), solar (Masuda et al. 1994; Lin 2011), and astrophysical plasmas (Colgate et al. 2001; Zhang & Yan 2011). A long-standing problem in reconnection studies is how fast reconnection processes available magnetic flux, the so-called reconnection rate problem (see Cassak et al. 2017, and reference therein).

Many reconnection models have been constructed to explain the reconnection rate observed—about 0.1 in normalized units—in numerical simulations (e.g., Birn et al. 2001) and space (e.g., Wang et al. 2015) and solar plasmas (e.g., Yokoyama et al. 2001; Qiu et al. 2002). The most famous one is the Sweet–Parker model (Parker 1957, 1963; Sweet 1958), which is the first quantitative reconnection model but predicts a rate far too slow to explain, for example, solar flares. In this model, the current sheet is long and thin, limiting the inflow flux of plasmas, and therefore the reconnection rate is low. The following Petschek model (Petschek 1964) predicts a much shorter current sheet and a much higher reconnection rate. However, numerical simulations have demonstrated that it requires an ad hoc localized resistivity to be stable, and thus the origin of the localization is not captured in this model (Sato & Hayashi 1979; Biskamp 1986). Recent progress features collisionless physics in the diffusion region to be the key to producing fast reconnection, notably Hall physics (e.g., Shay et al. 2001) or secondary islands (Daughton & Karimabadi 2007; Liu et al. 2014). The up-to-date model by Liu et al. (2017) shows that the value of fast reconnection rate is insensitive to these diffusion-region-scale physics but is instead constrained by the mesoscale magnetic geometry and force balance. This model expresses the reconnection rate as a function of the exhaust opening angle; it predicts that the fast rate on the order of $\mathcal{O}(0.1)$ persists for a wide range of opening angles, and the maximum

plausible reconnection rate is bounded by $\simeq 0.2$ in the low- β regime.

The thermal pressure is often neglected in those models, which might be valid in the low- β regime and suitable for studying reconnection in solar flares, Earth’s magnetotail, or magnetically dominated astrophysical plasmas. However, thermal pressure could dominate the plasma dynamics with a higher $\beta \equiv P/(B^2/8\pi)$, as in the outer heliosphere (β up to 10; Drake et al. 2010; Schoeffler et al. 2011), in the hot intracluster medium (ICM) of galaxy clusters ($\beta \sim 10^2$ – 10^4 ; Carilli & Taylor 2002; Schekochihin & Cowley 2006), or at the Galactic center ($\beta \sim 10^1$ – 10^2 ; Marrone et al. 2007). Self-generated pressure anisotropy and/or pressure gradient upstream (Egedal et al. 2013) and downstream (Bessho & Bhattacharjee 2010; Liu et al. 2011, 2012; Haggerty et al. 2018) of the diffusion region could affect the force balance and reduce the outflow speed. Therefore, it is critical to include the thermal correction to a reconnection model for studying the reconnection rate in the high- β regime.

In this paper, we extend the reconnection model by Liu et al. (2017) to include the thermal correction. By including the pressure anisotropy and pressure gradient force in the inflow and outflow force balance equations, we get the magnetic field immediately upstream of the diffusion region and the outflow speed as a function of both the plasma β and the exhaust opening angle. A prediction of the normalized reconnection rate in a given β can be obtained by maximizing the rate with respect to the opening angle, which scales as $\simeq \mathcal{O}(0.1/\sqrt{\beta_{i0}})$ in the high- β limit, where $\beta_{i0} \equiv 8\pi P_{i0}/B_{x0}^2$ defines the ratio of the inflow ion thermal pressure (P_{i0}) and the magnetic pressure of the asymptotic magnetic field B_{x0} . In Section 2, we perform 2D kinetic simulations with different plasma β to show the scaling of the reconnection rate and related quantities with plasma β . In Section 3, we present the extended model and discuss its predictions in the low- β and high- β limits. In Section 4, we compare the model results with the simulations. In Section 5, we discuss the conclusions and implications based on our results.

2. Numerical Simulations

We carry out 2D kinetic simulations of magnetic reconnection in plasmas with $\beta=0.25, 1, 10$, and 40 , studying how the reconnection rate responds to this change. The simulations were performed using the VPIC particle-in-cell code (Bowers et al. 2008), which solves Maxwell's equations and the relativistic Vlasov equation. The simulation employs a Harris current sheet with the magnetic profile $\mathbf{B} = B_{x0} \tanh(z/\lambda) \hat{x}$, where B_{x0} is the reconnecting magnetic field and λ is the half-thickness of the current sheet. The simulation size is $L_x \times L_z = 76.8d_i \times 76.8d_i$, which spans domain $[0, L_x] \times [-L_z/2, L_z/2]$, where d_i is the ion inertial length. We chose a proton-to-electron mass ratio $m_i/m_e = 400$. The plasma consists of a Harris sheet component with a peak density n_0 and a background component with a uniform density n_b . Its density profile satisfies $n = n_0 \text{sech}^2(z/\lambda) + n_b$ to maintain the initial pressure balance. In all runs, we choose $\omega_{pe}/\Omega_{ce} = 2$, where the plasma frequency $\omega_{pe} = (4\pi n_0 e^2/m_e)^{1/2}$ and the electron gyrofrequency $\Omega_{ce} = eB_{x0}/m_e c$, resulting in an Alfvén speed $v_{A0} \equiv B_{x0}/(4\pi n_0 m_i)^{1/2} = 0.025c$. To have a similar dynamical time ($\sim L_x/v_{Ab}$, where $v_{Ab} \equiv B_{x0}/(4\pi n_b m_i)^{1/2}$) for different runs, we choose $n_b = n_0$ in all runs to fix $v_{Ab} = v_{A0}$. Electrons and ions have the same uniform temperature T_s in the sheet component. The initial pressure balance $2n_0 kT_s = B_{x0}^2/8\pi$ results in an electron thermal speed $v_{the} \equiv (kT_s/m_e)^{1/2} = (v_{A0}/2)(m_i/m_e)^{1/2} = 0.25c$ for the sheet component. Electrons and ions have the same uniform temperature T_0 in the background component. In the four runs, $T_0 = 0.25T_s, T_s, 10T_s$, and $40T_s$, resulting in a plasma $\beta_0 \equiv 16\pi n_b kT_0/B_{x0}^2 = 0.25, 1, 10$, and 40 , respectively. Note that for this paper we will use β_0 and β interchangeably when it does not cause confusion. Since electrons and ions have the same temperature, $\beta_{e0} = \beta_{i0} = \beta_0/2$ in this study. The grid numbers are $n_x \times n_z = 12,288 \times 12,288$ for the run with $\beta = 0.25$ and 6144×6144 for runs with higher plasma β . For electric and magnetic fields, we employ periodic boundaries along the x -direction and perfectly conducting boundaries along the z -direction. For particles, we employ periodic boundaries along the x -direction and reflecting boundaries along the z -direction. A localized initial perturbation is added to induce reconnection with a single X-line. The y -component of the vector potential of the perturbation is

$$\begin{aligned} \delta A_y = & -\frac{\delta B}{(2\pi/L_p)^2 + (1/L_c^2)} \\ & \times \left[\frac{1}{L_c} \sin\left(\frac{2\pi|x'|}{L_p}\right) + \frac{2\pi}{L_p} \cos\left(\frac{2\pi x'}{L_p}\right) \right] \\ & \times \exp\left(-\frac{|x'|}{L_c}\right) \cos\left(\frac{\pi z}{L_z}\right), \end{aligned} \quad (1)$$

where $x' = x - 0.5L_x$, δB controls the amplitude of the perturbation, L_p is the length scale of the sinusoidal perturbation along the x -direction, and L_c is the length scale that controls how fast the perturbation decays along the x -direction. We have chosen $\delta B = 0.165B_{x0}$, $L_p = L_x/12$, and $L_c = L_x$ in our simulations.

Figure 1 shows the current layer near the primary X-line for runs with $\beta = 0.25$ and 10 . One single X-line forms in the reconnection layer under local perturbation, and the reconnection exhaust gradually opens up. The exhaust's opening angle is between 20° and 30° and can get over 40° when a magnetic

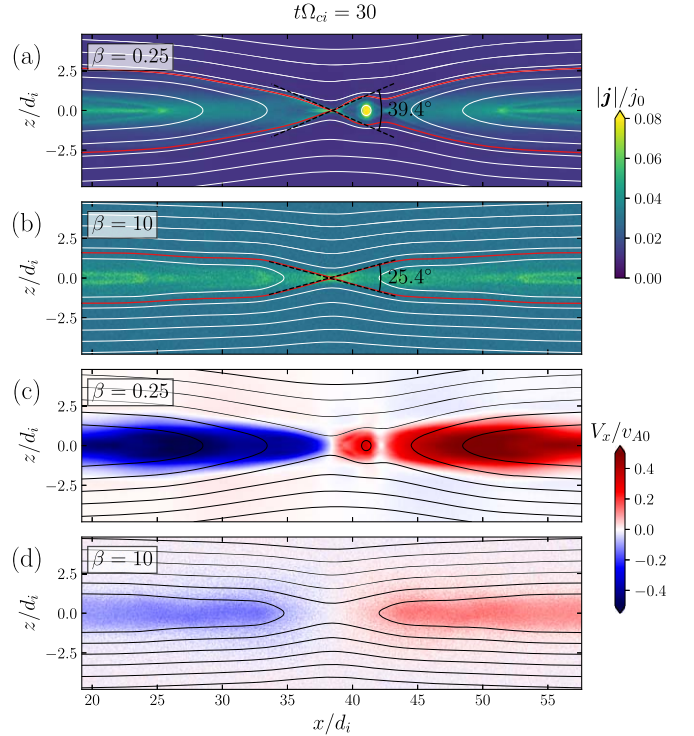


Figure 1. Reconnection layer near the main X-point. The top two panels show the current density for the runs with (a) $\beta = 0.25$ and (b) $\beta = 10$. The red lines indicate the separatrix. The point where the top line meets the bottom one indicates the X-point. Assuming that its coordinate is (x_0, z_0) , we then calculate the angle $\theta \equiv \arctan(|z - z_0|/|x - x_0|)$ for all the points (x, z) along the separatrix and get the maximum values for the four branches starting from the X-point ($\theta_1 - \theta_4$). The four maximums are similar except when a plasmoid is ejected and opens up the right side of the exhaust in the run with $\beta = 0.25$ (panel (a)). We evaluate $\theta_0 = \text{mean}(\theta_i)$. The exhaust opening angle is then about $2\theta_0$. The dashed lines with a slope $\pm \tan \theta_0$ indicate the boundaries for obtaining the opening angles. The bottom two panels show the outflow velocity V_x normalized by the upstream Alfvén speed v_{A0} (which is kept the same for all cases).

island forms in the case with $\beta = 0.25$. Figure 1(c) shows that the reconnection outflow V_x can reach $0.5v_{A0}$ in the $\beta = 0.25$ case, while Figure 1(d) shows that V_x is below $0.2v_{A0}$ in the $\beta = 10$ case. Given a similar exhaust opening angle, the reduction of the reconnection outflow at the high- β regime implies a weaker outflow motional electric field that potentially leads to a lower reconnection rate. We will model this reconnection outflow reduction in the next section.

Figure 2 shows the time evolution of several quantities critical to the determination of the reconnection rate. Figure 2(a) shows the exhaust opening angle. The reconnection exhaust opens up earlier in the low- β runs, corresponding to a faster reconnection onset in the low- β limit. The opening angles in the nonlinear stage fall between 20° and 35° in general. When a secondary plasmoid forms in the reconnection layer (Figure 1(a)), the opening angle can reach 45° in the $\beta = 0.25$ case, but this transient feature between time $20 < t\Omega_{ci} < 30$ is not our focus. As the exhaust opens up, the upstream-pointing magnetic tension force gets stronger in the inflow region. To maintain this field geometry, the upstream magnetic field strength needs to decrease toward the diffusion region so that the magnetic pressure gradient balances the tension force (Liu et al. 2017). The resulting reconnecting field immediately upstream of the diffusion region B_{xm} will thus be reduced; here the subscript “ m ” denotes the *microscopic* scale,

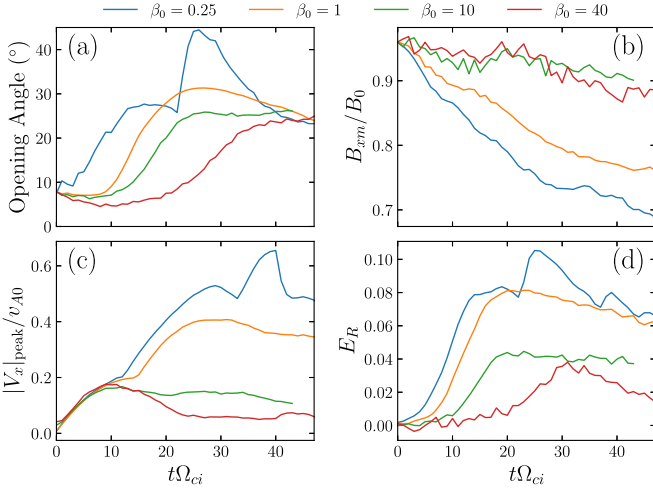


Figure 2. Time evolution of key physical quantities in the simulations. (a) Opening angle of the reconnection exhaust. (b) Magnetic field immediately upstream of the diffusion region. We obtain B_{xm} using the B_x profile along z and following a similar procedure in Liu et al. (2017) (see Appendix A for the detailed procedure). (c) Peak ion outflow velocity V_{ix} normalized by the upstream Alfvén speed v_{A0} . We pick the maximum V_{ix} near the X-line in regions between the X-point and $10d_i$ downstream of the X-point to estimate the V_{ix} downstream of the IDR. For lower- β runs, V_{ix} can still grow farther away from the X-point. (d) Normalized reconnection rate E_R .

which will be the ion inertial scale in electron–proton plasmas. Figure 2(b) shows the B_{xm} evolution in different runs. This B_{xm} gradually decreases as reconnection proceeds and saturates after $t\Omega_{ci}^{-1} > 30$ for all runs. Its value is larger when β is higher—about 0.73 when $\beta = 0.25$ and 0.9 when $\beta = 40$ at $t\Omega_{ci} = 30$. Based on this observation, one may expect the outflow speed V_{ix} in high- β runs to be higher since $V_{ix} \propto B_{xm}$, i.e., the outflow is driven by B_{xm} . However, as shown in Figure 2(c), the outflow speed V_x turns out to be lower in high- β runs. V_{ix} goes up to $0.5v_{A0}$ when $\beta = 0.25$ and decreases below $0.1v_{A0}$ when $\beta = 40$, suggesting that other factors besides B_{xm} play important roles in determining the outflow speed. Since the outflow is slower in the high- β limit, we expect a lower reconnection rate in this limit. Figure 2(d) shows that the reconnection rate indeed decreases with plasma β , as expected. When $\beta \leq 1$, the normalized reconnection rate $E_R \equiv cE_y/B_{x0}v_{A0}$ is around 0.08 and can be larger than 0.1 when the secondary plasmoids temporarily widen the reconnection exhaust in the $\beta = 0.25$ case, consistent with earlier simulations (e.g., Birn et al. 2001). In contrast, when $\beta = 40$, E_R is below 0.04. In the next section, we will develop a model to explain these simulated trends.

3. Rate Model with Thermal Correction

To explain the simulation results, we develop the framework to incorporate thermal pressure effects in the reconnection rate model. These effects are expected to be important when $\beta \gtrsim \mathcal{O}(1)$. Figure 3 is an overview of this model, which includes the thermal correction $-\nabla \cdot \mathbf{P}$ that participates in the force balance. As shown in the schematic, the thermal correction might change the inflow force balance and therefore B_{xm} , and the outflow speed can be slowed down by the back pressure and/or a weaker magnetic tension due to the pressure anisotropy. Following the approach in Liu et al. (2017), the force balance evaluated at point 1 in the inflow region will relate B_{xm} to the upstream asymptotic field B_{x0} . By matching

the upstream magnetic field opening angle θ to angle ϕ made by the reconnected field at the microscopic scale, we obtain the strength of the reconnected field $B_{zm} \simeq B_{xm}\Delta z/\Delta x$. On the other hand, the force balance evaluated at point 2 within the diffusion region will determine the outflow velocity $V_{out,m}$ immediately downstream of the diffusion region. With both B_{zm} and $V_{out,m}$, we can calculate the motional electric field adjacent to the diffusion region, which is essentially the reconnection electric field. An expression of the normalized reconnection rate as a function of the opening angle θ then can be derived to be $E_R \equiv cE_y/B_{x0}v_{A0} = B_{zm}V_{out,m}/(B_{x0}v_{A0})$.

Assuming that the pressure tensor can be approximated into the gyrotropic form, the momentum equation can be written as

$$nm_i \frac{\partial \mathbf{V}}{\partial t} + nm_i \mathbf{V} \cdot \nabla \mathbf{V} + \nabla \frac{B^2}{8\pi} + \nabla P_{\perp} \simeq \nabla \cdot \left(\varepsilon \frac{\mathbf{B}\mathbf{B}}{4\pi} \right), \quad (2)$$

where the anisotropy (firehose) parameter $\varepsilon \equiv 1 - 4\pi(P_{\parallel} - P_{\perp})/B^2$. P_{\parallel} and P_{\perp} are the pressure parallel and perpendicular to the local magnetic field, respectively. The plasma is firehose unstable (e.g., Wang et al. 2018, 2020) if $\varepsilon < 0$ since the magnetic tension on the right-hand side becomes negative. When $P_{\parallel} > P_{\perp}$, the pressure anisotropy will reduce the magnetic tension force. We seek a steady-state solution by neglecting the ∂t term. From simulations, we get the empirical relations $P_{e\parallel} \approx P_{e\perp} \approx P_{e0}n/n_0$, $P_{i\parallel} \approx P_{i0}n/n_0$, and $P_{i\perp} \approx P_{i0}nB_x/(n_0B_{x0})$ along the inflow symmetry line (see Appendix B for details, where the CGL (Chew et al. 1956) and Le & Egedal (Le et al. 2009; Egedal et al. 2013) closures are also explored). The simulations further suggest that the plasma number density $n \approx n_0[1 - C(1 - \bar{B}_x)]$, where the constant $C = 1/(\sqrt{2\beta_{i0}} + 1)$ (see Figure 7 for details). By discretizing Equation (2) along the inflow direction at point 1 and using these pressure closures, we get

$$\frac{B_{x0}^2 - B_{xm}^2}{8\pi\Delta z} + \frac{P_0 + P_{i0}(\sqrt{2\beta_{i0}} + B_{xm}/B_{x0})}{\sqrt{2\beta_{i0}} + 1} \frac{B_{x0} - B_{xm}}{B_{x0}\Delta z} = \frac{\varepsilon_1}{4\pi} B_{x1} \frac{2B_{zh}}{\Delta x}. \quad (3)$$

We made a reasonable approximation of $B_x \simeq B_x(z)$, $B_z \simeq B_z(x)$, and $\varepsilon \simeq \varepsilon(z)$ around the diffusion region. The first term is from the magnetic pressure gradient, the second term is from the perpendicular pressure gradient, and the term on the right is from the magnetic tension force modified by the pressure anisotropy. $B_{x1} = (B_{x0} + B_{xm})/2$ is the magnetic field at point 1, and the corresponding firehose parameter is

$$\varepsilon_1 \equiv \varepsilon(B_{x1}) = 1 + \frac{\beta_{i0}}{2} \left(\frac{1}{\bar{B}_{x1}} - \frac{1}{\bar{B}_{x1}^2} \right) \times \left(1 - \frac{1 - \bar{B}_{x1}}{\sqrt{2\beta_{i0}} + 1} \right), \quad (4)$$

where $\bar{B}_{x1} \equiv B_{x1}/B_{x0}$. Note that we can include the plasma inertia ($nm_i \mathbf{V} \cdot \nabla \mathbf{V}$) in the calculation, but the inflow contribution is negligible (i.e., can be ordered out); thus, we will not treat it here to avoid unnecessary complexity. Equation (3) can

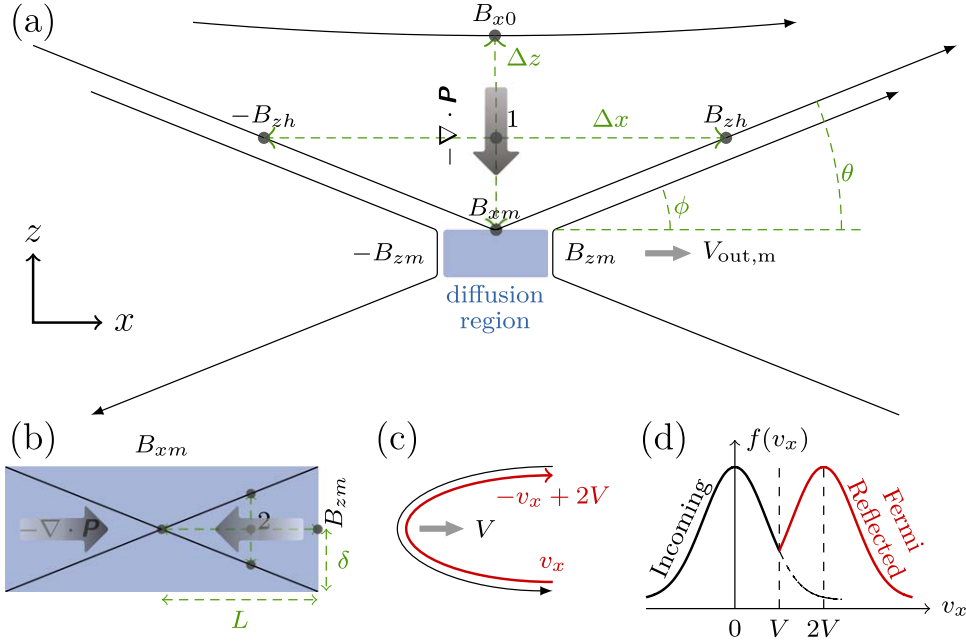


Figure 3. Reconnection rate model with thermal corrections. (a) Geometry of reconnection at the mesoscale. The large gray arrow indicates the force $-\nabla \cdot \mathbf{P}$ due to thermal correction. The force balance along the inflow will be evaluated at point 1. B_{x0} is the asymptotic magnetic field. B_{xm} is the magnetic field immediately upstream of the diffusion region. B_{zh} is evaluated near the separatrix. B_{zm} is the magnetic field immediately downstream of the diffusion region. Angles $\theta \equiv \tan^{-1}(\Delta z/\Delta x)$ and $\phi \equiv \tan^{-1}(B_{zm}/B_{xm})$. (b) Dimensions of the diffusion region $2L \times 2\delta$ at the microscale. The arrows indicate the forces toward the X-point arising from the thermal correction. The force balance along the outflow will be evaluated at point 2. (c) Plasma heating associated with the Fermi mechanism. The red curve indicates an example particle trajectory. Real particle trajectory might be more complicated (Drake et al. 2009). v_x is the particle's velocity along x when it enters the diffusion region. The flow velocity V increases from 0 at the X-point to $V_{\text{out},m}$ immediately out of the diffusion region. (d) Particle velocity distribution consisting of the incoming population and the Fermi reflected population.

be written in a dimensionless form (see Appendix B for details)

$$(1 - \bar{B}_{xm}^2) + \frac{\beta_0 + \beta_{i0}(\sqrt{2\beta_{i0}} + \bar{B}_{xm})}{\sqrt{2\beta_{i0}} + 1} \times (1 - \bar{B}_{xm}) = \varepsilon_1 \left(\frac{\Delta z}{\Delta x} \right)^2 (1 + \bar{B}_{xm})^2, \quad (5)$$

where $\bar{B}_{xm} \equiv B_{xm}/B_{x0}$ and the asymptotic upstream plasma $\beta_0 \equiv \beta_{i0} + \beta_{e0}$. In cases studied here $\beta_{i0} = \beta_{e0} = 8\pi n_0 kT_0/B_{x0}^2$. When the opening angle is small ($\Delta z/\Delta x \ll 1$), $\bar{B}_{xm} \rightarrow 1$ for any β_0 , consistent with that early in the simulations (Figure 2(b)). For an arbitrary opening angle, we can obtain \bar{B}_{xm} numerically finding the roots of Equation (5) (e.g., using Newton's method).

To obtain the outflow velocity, we follow a similar procedure to that in the inflow region but discretize the force balance along the x -direction at point 2 in Figure 3(b) and keep the plasma inertia ($nm_i \mathbf{V} \cdot \nabla \mathbf{V}$). We get

$$\frac{n_2 m_i V_{\text{out},m}^2}{2L} + \frac{B_{zm}^2}{8\pi L} + \frac{\Delta P_{xx,m}}{L} = \frac{1}{4\pi} \frac{B_{zm}}{2} \frac{\varepsilon_m B_{xm}/2}{\delta/2}, \quad (6)$$

where n_2 is the plasma density at point 2, $\Delta P_{xx,m}$ is the pressure increase from the X-point to immediately downstream of the ion diffusion region (IDR), and ε_m is the anisotropy parameter (similar to Equation (4)) immediately upstream of the IDR

$$\varepsilon_m \equiv \varepsilon(B_{xm}) = 1 + \frac{\beta_{i0}}{2} \left(\frac{1}{\bar{B}_{xm}} - \frac{1}{\bar{B}_{xm}^2} \right) \left(1 - \frac{1 - \bar{B}_{xm}}{\sqrt{2\beta_{i0}} + 1} \right), \quad (7)$$

which depends on β_{i0} and \bar{B}_{xm} and can be determined once we obtain \bar{B}_{xm} from Equation (5). To find the solutions of $V_{\text{out},m}$ from Equation (6), we need to model n_2 and $\Delta P_{xx,m}$. Unlike the inflow region, the CGL-like closure is not expected to work within the IDR. To estimate $\Delta P_{xx,m}$, we only need to know the difference between P_{xx} at the x -line and the edge of the IDR. While the particle heating mechanism can be complex inside the diffusion region (Hoshino et al. 2001; Shuster et al. 2015; Wang et al. 2016), it transitions to a simpler Fermi-type reflection outside the diffusion region. We thus will model the $P_{xx,m}$ using the Fermi mechanism, as illustrated in Figures 3(c) and (d); particle velocity changes from v_x to $2V - v_x$ during the reflection (Figure 3(c)), and the combination of the reflected/accelerated particles with the incoming population leads to the broadening of the distribution function (Figure 3(d); see Appendix C for details). Strictly speaking, the heating is primarily along the magnetic field, but particles can be scattered near the diffusion region, leading to an increase of the perpendicular pressure ΔP_{\perp} , which for ions is approximately

$$\Delta P_{\perp,ix}(V) = n_0 m_i \left[V^2 + \left(V^2 + \frac{\beta_{i0}}{2} v_{A0}^2 \right) \times \operatorname{erf} \left(\frac{V}{\sqrt{\beta_{i0}} v_{A0}} \right) + V v_{A0} \sqrt{\frac{\beta_{i0}}{\pi}} e^{-V^2/(\beta_{i0} v_{A0}^2)} \right]. \quad (8)$$

The corresponding plasma density is

$$n(V) = n_0 + n_0 \operatorname{erf} \left(\frac{V}{\sqrt{\beta_{i0}} v_{A0}} \right), \quad (9)$$

which approaches n_0 when $\beta_{i0} \rightarrow \infty$ and $2n_0$ when $\beta_{i0} \rightarrow 0$, within the range predicted in Birn et al. (2010). The contribution from electrons through a single Fermi reflection is negligible because $\Delta P_{exx}/\Delta P_{icx} \sim \mathcal{O}(\sqrt{m_e/m_i}) \ll 1$ according to Figure 8. Even though the ion ΔP_{icx} may not totally account for ΔP_{xx} , we find that it remains the dominant term even in the large- β limit (shown in Section 4). Observationally, it has been shown that ion heating is stronger than electron heating during magnetic reconnection at Earth's magnetosphere, and $\Delta T_i/\Delta T_e \simeq 7$ (Phan et al. 2013, 2014). Encouraged by these observations, we will ignore electron heating ΔP_{exx} in the following analysis. Other potential heating mechanisms (Hoshino et al. 2001; Drake et al. 2005, 2006; Oka et al. 2010; Dahlin et al. 2014; Egedal et al. 2015; Li et al. 2015; Shuster et al. 2015; Wang et al. 2016) could be included in future work. In the following analysis, we will take $\Delta P_{xx,m} \simeq \Delta P_{icx,m} \equiv \Delta P_{icx}(V_{out,m})$, which is normalized to

$$\begin{aligned} \Delta \bar{P}_{icx,m} &= \frac{4\pi \Delta P_{icx,m}}{B_{x0}^2} \\ &= \bar{V}_{out,m}^2 + \left(\bar{V}_{out,m}^2 + \frac{\beta_{i0}}{2} \right) \text{erf} \left(\frac{\bar{V}_{out,m}}{\sqrt{\beta_{i0}}} \right) \\ &\quad + \bar{V}_{out,m} \sqrt{\frac{\beta_{i0}}{\pi}} e^{-\bar{V}_{out,m}^2/\beta_{i0}}, \end{aligned} \quad (10)$$

where $\bar{V}_{out,m} = V_{out,m}/v_{A0}$ and we have used $v_{A0} = B_{x0}/(4\pi n_0 m_i)^{1/2}$. We model n_2 as the average of the density at the X-point ($\simeq n_0$) and that immediately downstream of the IDR ($\simeq n(V_{out,m})$ evaluated from Equation (9)). Thus,

$$n_2 \simeq \frac{n_0 + n(V_{out,m})}{2} = n_0 + \frac{n_0}{2} \text{erf} \left(\frac{\bar{V}_{out,m}}{\sqrt{\beta_{i0}}} \right). \quad (11)$$

Dividing Equation (6) by $n_0 m_i v_{A0}^2$, we get the force balance at point 2 (see Appendix C for details)

$$\begin{aligned} \frac{1}{4} \left[2 + \text{erf} \left(\frac{\bar{V}_{out,m}}{\sqrt{\beta_{i0}}} \right) \right] \bar{V}_{out,m}^2 + \Delta \bar{P}_{icx,m} \\ + \frac{\bar{B}_{xm}^2}{2} \left[\left(\frac{\Delta z}{\Delta x} \right)^2 - \varepsilon_m \right] \simeq 0 \end{aligned} \quad (12)$$

in dimensionless form, where we have used $B_{zm}/B_{xm} \simeq \delta/L \simeq \Delta z/\Delta x$.

When $\beta_0 \ll 1$, then $\text{erf}(\bar{V}_{out,m}/\sqrt{\beta_{i0}}) \rightarrow 1$ and $\bar{V}_{out,m} \exp(-\bar{V}_{out,m}^2/\beta_{i0}) \rightarrow 0$. As a result, $\Delta \bar{P}_{icx,m} \rightarrow 2\bar{V}_{out,m}^2$ according to Equation (10). In the small opening angle limit, $\Delta z/\Delta x \ll 1$, $\bar{B}_{xm} \rightarrow 1$ according to Equation (5), and $\varepsilon_m \rightarrow 1$ according to Equation (7). Then,

$$V_{out,m} = \sqrt{\frac{2\varepsilon_m}{11}} v_{A0} \simeq 0.43 v_{A0}, \quad (13)$$

which is an Alfvénic outflow as expected in the low- β limit.

When $\beta_0 \gg 1$, then $\text{erf}(\bar{V}_{out,m}/\sqrt{\beta_{i0}}) \simeq (2/\sqrt{\pi}) \bar{V}_{out,m}/\sqrt{\beta_{i0}}$ and $\bar{V}_{out,m} \exp(-\bar{V}_{out,m}^2/\beta_{i0}) \simeq \bar{V}_{out,m}$. According to

Equation (10),

$$\Delta \bar{P}_{icx,m} \rightarrow \frac{2}{\sqrt{\pi\beta_{i0}}} \bar{V}_{out,m}^3 + \bar{V}_{out,m}^2 + 2\sqrt{\frac{\beta_{i0}}{\pi}} \bar{V}_{out,m}. \quad (14)$$

Then, Equation (12) gives

$$\frac{5}{\sqrt{\pi\beta_{i0}}} \bar{V}_{out,m}^3 + 3\bar{V}_{out,m}^2 + 4\sqrt{\frac{\beta_{i0}}{\pi}} \bar{V}_{out,m} = \varepsilon_m, \quad (15)$$

where $0 \leq \varepsilon_m \leq 1$. We seek a solution $\bar{V}_{out,m} > 0$ (flows moving away from the X-point). All the terms on the left-hand side of Equation (15) are positive, indicating that the third term $4\sqrt{\beta_{i0}/\pi} \bar{V}_{out,m} < \varepsilon_m$. Then, the first term $5\bar{V}_{out,m}^3/\sqrt{\pi\beta_{i0}} < (5\pi/(64\beta_{i0}^2))\varepsilon_m^3$ and the second term $3\bar{V}_{out,m}^2 < (3\pi/(16\beta_{i0}))\varepsilon_m^2$. Both terms will be much smaller than ε_m when $\beta_{i0} \gg 1$ and are therefore small corrections to the first-order linear equation $4\sqrt{\beta_{i0}/\pi} \bar{V}_{out,m} = \varepsilon_m$. Thus,

$$V_{out,m} \simeq \frac{\sqrt{\pi}}{4} \frac{\varepsilon_m v_{A0}}{\sqrt{\beta_{i0}}}. \quad (16)$$

A higher β will thus reduce the outflow speed. Equation (16) is almost identical to the expression obtained in Haggerty et al. (2018), which is $(\sqrt{2}/3)\varepsilon_m v_{A0}/\sqrt{\beta_{i0}}$ in our notation. They compared this expression against 81 kinetic simulations and 14 in situ observations that span a wide range of parameter regimes and showed an excellent agreement. However, the $\sqrt{2}/3$ factor is an empirical parameter in their model, while we derived it from Equation (15). It is also interesting to note that their prediction based on the 1D shock transition across the exhaust is consistent with our 2D model that accounts for the back pressure along the outflow direction. The single compact expression in Equation (12) explains the outflow speed in both the high- β and low- β limit reported in Haggerty et al. (2018). For an arbitrary plasma β and opening angle, we can find the roots of Equation (12) numerically. Combining Equations (5), (8), and (12) and $B_{zm} \simeq B_{xm} \Delta z/\Delta x$, the resulting reconnection rate $E_R = B_{zm} V_{out,m}/(B_{x0} v_{A0})$ can be derived for a general case.

4. Model-Simulation Comparison

To illustrate the thermal effects in simulations, we plot the anisotropy (firehose) parameter ε and relevant components of the pressure tensors for the $\beta = 1$ case in Figure 4. We first examine quantities important to the force balance upstream of the IDR. The 2D map of ε near the IDR is shown in Figure 4(a), and its (vertical) cut across the x -line along the inflow direction is shown in Figure 4(b). The model based on the closure we choose (dotted black) captures the decreasing trend of ε toward the IDR, which is shaded in gray. Note that the region inside the gray area is not critical to the upstream force balance discussed here. This anisotropy parameter ε (solid black) is calculated using P_{\parallel} and P_{\perp} in the simulation and along the inflow symmetry line $P_{\parallel} \simeq P_{xx}$ and $P_{\perp} \simeq P_{zz}$. Therefore, in Figure 4(d) we show the vertical cuts of P_{icx} , P_{izz} , P_{exx} , and P_{ezz} . The perpendicular component P_{izz} follows well with the CGL closure based on the μ -conservation (solid orange). P_{icx} only decreases slightly and is closer to the Boltzmann closure (dashed orange). Both P_{exx} and P_{ezz} also follow the Boltzmann closure better, resulting in a much

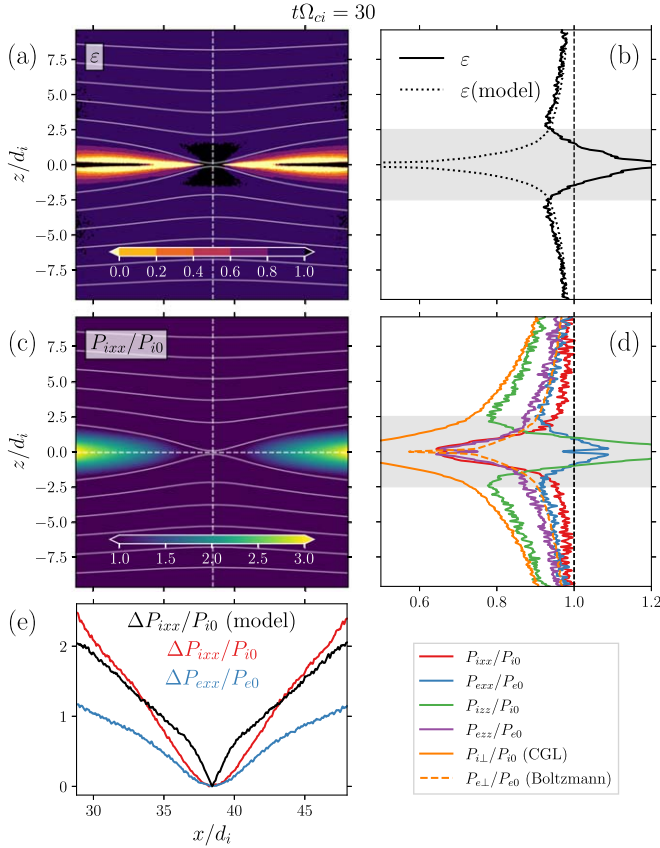


Figure 4. Thermal effects in the run with $\beta = 1$. (a) Firehose parameter ϵ . (b) ϵ along the vertical cut across the X-point (the dashed line in panel (a)). The dotted line is the modeled ϵ . The shaded region indicates the IDR. (c) P_{ixx} normalized by the background pressure. (d) Vertical cuts of the electron and ion pressure tensor components. The orange curves show the CGL (solid) and Boltzmann (dashed) scalings. The shaded region indicates the IDR. (e) Ion pressure enhancement along $z = 0$ (along the horizontal dashed line in panel (c)). The black line shows the predicted ΔP_{ixx} heating from the Fermi mechanism, which is evaluated from Equation (8) using measured V_{ix} along $z = 0$. For reference, the blue curve shows the enhancement of electron pressure ΔP_{exx} .

smaller pressure anisotropy than ions. In addition to the magnetic tension reduction due to ϵ , the pressure gradient $\partial_z P_{zz}$ resulting from the P_{zz} drop also plays an important role in counterbalancing the upstream-pointing tension force, mitigating the decrease of B_{xm} .

Along the outflow, the back pressure from the P_{xx} gradient is especially critical in the force balance. We thus plot the 2D map of the dominant component P_{ixx} in Figure 4(c) and its (horizontal) cut across the x -line in Figure 4(e). Our model ΔP_{ixx} (Equation (8) in black) based on the Fermi reflection reasonably captures the increasing trend toward the downstream region. In addition to this back pressure, ϵ upstream of the IDR (shown in panel (b)) can further reduce the magnetic tension force that drives the reconnection outflow. Note that this is consistent with the Walén test (Sonnerup et al. 1981) across exhausts, which indicates that the upstream ϵ (instead of the downstream ϵ) can affect the outflow velocity. This fact is also captured in the ϵ_m dependence in our model Equation (12). Figure 4(e) also shows a significant ΔP_{exx} , likely resulting from other heating mechanisms (Hoshino et al. 2001; Drake et al. 2005, 2006; Oka et al. 2010; Dahlin et al. 2014; Egedal et al. 2015; Li et al. 2015; Shuster et al. 2015; Wang et al. 2016) not considered here. However, simulations suggest that ion heating remains dominant (i.e., $\Delta P_{ixx} \gtrsim \Delta P_{exx}$) even in

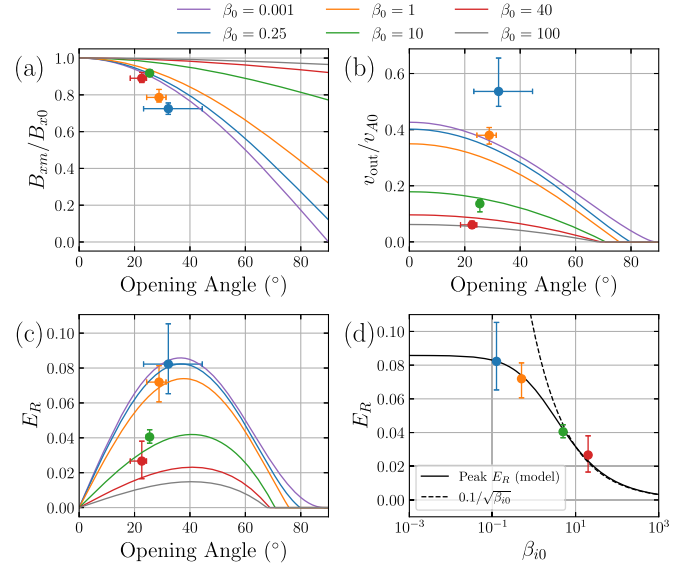


Figure 5. Model predictions using Equations (5), (8), and (12) and $E_R = B_{zm} V_{out,m} / (B_{x0} v_{A0})$. Here $\beta_0 = \beta_0/2$. The symbols indicate the simulation data points. We average the simulation results in Figure 2 after the reconnection rate peaks and use the minimum and maximum values in the same time range to determine the error bars. (a) Magnetic field upstream of the diffusion region. (b) Outflow velocity. (c) Normalized reconnection rate. (d) Scaling of the predicted maximum reconnection rate with β_0 .

the high- β limit, and the goal of this paper is to lay out a framework capable of incorporating the thermal correction on reconnection rates. A future more sophisticated and accurate heating model could be included in a similar manner.

These pressure models are included in the force balance along the inflow (Equation (5)) to obtain the strength of reconnecting field (B_{xm}) immediately upstream of the diffusion region and along the outflow (Equation (12)) to obtain the velocity ($V_{out,m}$) leaving the IDR. Figure 5(a) shows the modeled B_{xm} as a function of the opening angle θ for cases with a wide range of upstream plasma $\beta_0 = 10^{-3} - 10^2$. In the small opening angle limit, $B_{xm} \rightarrow B_{x0}$, as in an elongated Sweet-Parker reconnection layer. With the large opening angle, magnetic pressure $\partial_z B^2/8\pi$ is required to balance the upstream-pointing tension force, and this reduces B_{xm} ; this reduction can be mitigated in the large- β limit since (1) the thermal pressure gradient $\partial_z P_{zz}$ helps balance the magnetic tension and (2) the temperature anisotropy ϵ reduces the magnetic tension. This trend is observed in the simulation data ($\beta_0 = 0.25, 1, 10, 40$) overlaid in the same plot, although this model has overestimated B_{xm} .

Figure 5(b) shows the modeled $V_{out,m}$ as a function of the opening angle θ . With a large opening angle, the outflow speed can be reduced because of a weaker B_{xm} and larger (magnetic) back pressure $\partial_x B^2/8\pi$. The outflow velocity can be further reduced in the large- β limit because of (1) the thermal back pressure $\partial_x P_{xx}$ and (2) the reduced magnetic tension by the temperature anisotropy ϵ . There are critical angles above which the outflow speed vanishes; this cutoff behavior is caused by the complete loss of magnetic tension (i.e., the driver of reconnection) when $\epsilon_m < 0$ in the large opening angle limit. This trend of outflow reduction is also observed in the simulation data with different β , although this model has underestimated $V_{out,m}$ in the low- β limit.

The resulting reconnection rate is then $E_R = B_{zm} V_{out,m} / (B_{x0} v_{A0})$, and the reconnected field strength $B_{zm} \simeq B_{xm} \tan \theta$. Figure 5(c) shows the modeled rate as a function of the opening angle θ . The

model predicts a lower reconnection rate with a higher plasma β ; this trend is again observed in the simulation data. Note that the predicted peak E_R state has θ within $[30^\circ, 40^\circ]$ for a wide range of β , which is consistent with the β -insensitive opening angle observed in simulations. Figure 5(d) plots the predicted peak E_R as a function of upstream β_{i0} , and it captures the decreasing trend of simulated reconnection rate with a larger β_{i0} . In the $\beta_{i0} \gg \mathcal{O}(1)$ limit, the reduction of reconnection rate primarily correlates with the outflow speed reduction since B_{xm} is not reduced much in this limit (panel (a)). This leads to $\varepsilon_m \simeq 1$, and the outflow speed is thus $V_{out,m} \simeq (\sqrt{\pi}/4)(v_{A0}/\sqrt{\beta_{i0}})$ from Equation (16). Overall, the predicted reconnection rates show good agreement with simulations. In this high- β limit, the predicted scaling of the maximum plausible reconnection rate can be well approximated as $E_R \simeq 0.1/\sqrt{\beta_{i0}}$, as indicated by the dashed black curve in Figure 5(d).

5. Conclusion and Discussion

In this paper, we derive an analytical framework to incorporate thermal effects in the reconnection rate model (Liu et al. 2017). These thermal effects are manifested as the pressure anisotropy and pressure gradient force, which can modify the force balance at both the inflow and outflow regions. In the large- β limit, we find that the reconnection rate decreases primarily because of the reduction of the outflow speed, hindered by the pressure gradient force. The spatial variation of thermal pressure is modeled using a combination of CGL and Boltzmann closures in the upstream and the kinetic heating by Fermi reflections in the downstream. The pressure gradient force derived from the Fermi heating only depends on the upstream plasma parameters and the outflow speed, and it can be easily included in the rate model. 2D kinetic simulations compare favorably with the heating mechanism and the β -dependency of various key quantities needed to model the outflow speed and reconnection rate.

While the force balance constraint laid out here is general in determining the reconnection rate, there are opportunities for improvement. First, the present model does not include a guide field, which could change the plasma heating (Dahlin et al. 2014; Li et al. 2017) and scattering processes, and therefore the pressure anisotropy and pressure gradient force. As shown by Haggerty et al. (2018), the outflow velocity in the reconnection exhaust tends to get closer to v_{A0} as the guide field increases. This is also expected from our model; a background guide field could inhibit the upstream pressure variation due to the μ -conservation since guide-field strength does not change much while convecting into the diffusion region; it may also reduce the acceleration rate of Fermi mechanism for P_{xx} heating since the field-line curvature (κ) and thus the heating rate $e_{||}(\mathbf{v}_E \cdot \kappa)$ is reduced (Dahlin et al. 2014; Li et al. 2017). Second, when the ion gyroradius is much larger than the current layer thickness, the gyrotopic approximation is expected to break, and the full pressure tensor needs to be considered. Third, the model only includes ion heating in the diffusion region, while additional electron heating can be as strong in some regimes (Dahlin et al. 2014; Haggerty et al. 2015; Li et al. 2015; Shuster et al. 2015; Wang et al. 2016), and the upstream electric potential may be important (Le et al. 2009; Egedal et al. 2013, 2015; Shuster et al. 2015). Fourth, pressure anisotropy instabilities (e.g., mirror instability) may arise in high- β plasmas and were suggested to

distort the field geometry and the current sheet (Alt & Kunz 2019). Finally, we do not take 3D physics into account, for example, self-generated turbulence (Daughton et al. 2011; Liu et al. 2013; Li et al. 2019) or localized reconnection layer along the third dimension (Liu et al. 2019; Huang et al. 2020). Nevertheless, this present model extends the reconnection rate model to the high- β regime and provides new insights into the reconnection rate problem in high- β plasmas, which can be applicable to the outer heliosphere, the hot ICM of galaxy clusters, and the Galactic center.

We thank the anonymous referee for very helpful and constructive reviews. We acknowledge support by the National Science Foundation grant PHY-1902867 through the NSF/DOE Partnership in Basic Plasma Science and Engineering and NASA MMS 80NSSC18K0289. Simulations were performed at National Energy Research Scientific Computing Center (NERSC), at the Texas Advanced Computing Center (TACC) at the University of Texas at Austin, and with Los Alamos National Laboratory (LANL) institutional computing.

Appendix A

Determining the Reconnecting Magnetic Field Upstream of the Diffusion Region

To determine the reconnecting magnetic field upstream of the IDR (B_{xm}), we need to locate the diffusion region first. In the low- β regime, one could locate the IDR by checking where the total electric field \mathbf{E} starts to deviate from the ideal electric field $-\mathbf{V} \times \mathbf{B}$ in the reconnection inflow region. However, this method does not work well in the high- β simulations, where both \mathbf{E} and \mathbf{V} are very noisy. In this study, we use only the B_x z -profile through the X-line to determine B_{xm} . Figure 6 shows the procedure. We fit the B_x profile with five piecewise linear segments and determine a pair of the breakpoints as the boundaries of the IDR. Then, we calculate $B_{xm} = (B_{x1} - B_{x2})/2$, where $B_{x2} < 0$.

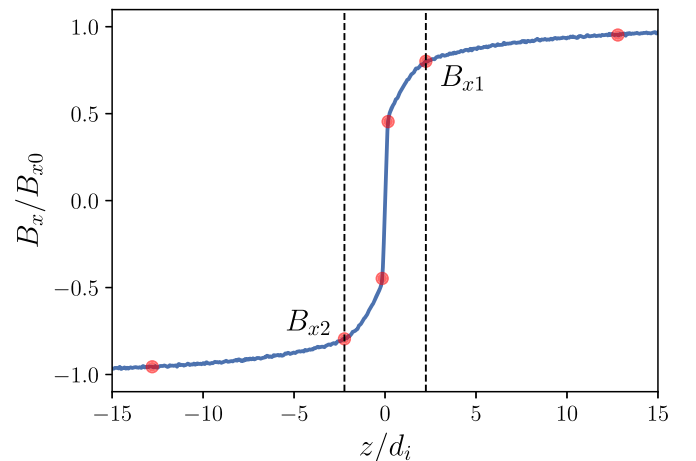


Figure 6. B_x profile through the X-line at $t\Omega_{ci} = 30$ in the run with $\beta = 1$. The red circles are the breakpoints automatically determined from piecewise linear fitting of the B_x profile using the `pwlf` Python library (Jekel & Venter 2019). The dashed lines indicate the boundaries of the IDR. B_{x1} and B_{x2} indicate the magnetic fields immediately upstream of the IDR. Note that we have applied a Gaussian filter to the B_x profile to reduce the grid-scale noises.

Appendix B

Inflow Force Balance

Starting from the single-fluid momentum equation

$$nm_i \frac{d\mathbf{V}}{dt} = -\nabla \cdot \mathbf{P} + \rho \mathbf{E} + \frac{\mathbf{j} \times \mathbf{B}}{c}, \quad (\text{B1})$$

we seek a steady-state solution by neglecting the time derivative of the plasma inertia. The electric force from charge separation is negligible in the nonrelativistic limit. In a well-magnetized plasma, $\mathbf{P} = P_\perp \mathbf{I} + (P_\parallel - P_\perp) \hat{\mathbf{b}} \hat{\mathbf{b}}$, where P_\parallel and P_\perp are pressures parallel and perpendicular to the local magnetic field, respectively, \mathbf{I} is the unit dyadic, and $\hat{\mathbf{b}}$ is the unit vector along the local magnetic field. Then,

$$\nabla \left(P_\perp + \frac{B^2}{8\pi} \right) = \nabla \cdot \left(\varepsilon \frac{\mathbf{B}\mathbf{B}}{4\pi} \right), \quad (\text{B2})$$

where $\varepsilon = 1 - 4\pi(P_\parallel - P_\perp)/B^2$ is the anisotropy (firehose) parameter. P_\parallel and P_\perp will be determined by the local plasma density and magnetic field strength through certain fluid closure. Note that the flux tube tends to expand while it is convected toward the IDR (i.e., as the result of exhaust opening), leading to a lower plasma density. However, particles can be redistributed along the flux tubes; thus, plasma density can change in a slower rate along the inflow symmetry line than the magnetic field does. Figure 7 shows the relation between the density change and the B_x change along the inflow in our simulations. When β is low and plasma is cold, the redistribution along the field line is less effective, resulting in $\bar{n} \sim \bar{B}_x$. When β is high, the thermal conduction can be large; thus, the redistribution is more effective, resulting in $\bar{n} \rightarrow 1$. Simple fitting shows a linear relation $\bar{n} \approx 1 - C(1 - \bar{B}_x)$ with a constant slope $C = 1/(\sqrt{2\beta_{i0}} + 1)$. Since β_{i0} is varied by changing the thermal temperature in our simulations, this expression indicates that the much heavier ions control the density variation.

The simulation results (e.g., Figure 4) show that the electron pressure anisotropy is much smaller than the ion pressure anisotropy. By comparing the results with Boltzmann ($P_\parallel = P_\perp \sim n$), CGL (Chew et al. 1956), and Le & Egedal (Le et al. 2009; Egedal et al. 2013) closures, we find that the pressure terms along the inflow symmetry line are best modeled as

$$P_{e\parallel} \approx P_{e\perp} = P_{e0} \bar{n}, \quad (\text{B3})$$

$$P_{i\parallel} = P_{i0} \bar{n}, \quad P_{i\perp} = P_{i0} \bar{n} \bar{B}_x. \quad (\text{B4})$$

Then, the total parallel and perpendicular pressures are

$$P_\parallel = P_0 \left(1 - \frac{1 - \bar{B}_x}{\sqrt{2\beta_{i0}} + 1} \right) \quad (\text{B5})$$

$$P_\perp = (P_{e0} + P_{i0} \bar{B}_x) \left(1 - \frac{1 - \bar{B}_x}{\sqrt{2\beta_{i0}} + 1} \right), \quad (\text{B6})$$

where $P_0 = P_{e0} + P_{i0}$ is the total pressure of the background plasma. By approximating $B_x \simeq B_x(z)$, $B_z \simeq B_z(x)$, and $\varepsilon \simeq \varepsilon(z)$, then we can write

$$\left(\frac{B^2}{8\pi} \right)_z = \frac{\partial}{\partial z} \frac{B^2}{8\pi} \simeq \frac{B_{x0}^2 - B_{xm}^2}{8\pi \Delta z} \quad (\text{B7})$$

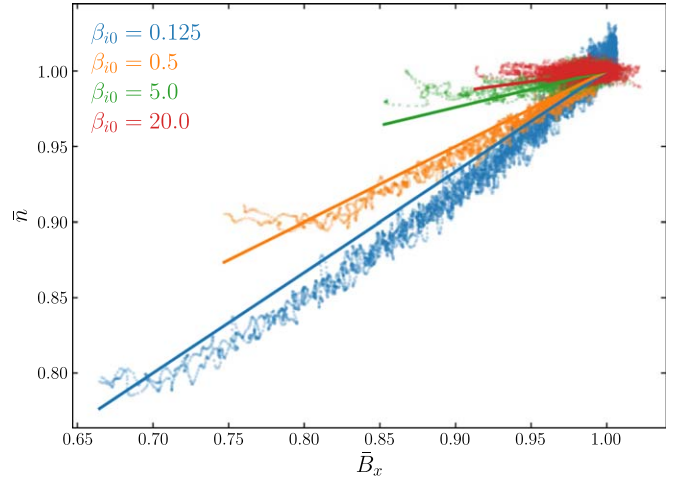


Figure 7. Plasma density vs. the magnetic field strength along the inflow symmetry line for the four runs at $t\Omega_{ci} = 30$. $\bar{B}_x = B_x/B_0$, where $B_0 = B_{x0}$ is the asymptotic magnetic field. $\bar{n} = n/n_0$. The straight lines show the linear relation $\bar{n} = 1 - C(1 - \bar{B}_x)$, where the slope $C = 1/(\sqrt{2\beta_{i0}} + 1)$.

$$(\nabla P_\perp)_z = \frac{\partial P_\perp}{\partial z} \simeq \frac{P_0 + P_{i0}(\sqrt{2\beta_{i0}} + \bar{B}_{xm})}{\sqrt{2\beta_{i0}} + 1} \frac{1 - \bar{B}_{xm}}{\Delta z} \quad (\text{B8})$$

$$\begin{aligned} \left[\nabla \cdot \left(\varepsilon \frac{\mathbf{B}\mathbf{B}}{4\pi} \right) \right]_z &= \frac{\partial}{\partial x} \left(\frac{\varepsilon B_x B_z}{4\pi} \right) + \frac{\partial}{\partial z} \left(\frac{\varepsilon B_z B_z}{4\pi} \right) \\ &\simeq \frac{\varepsilon_1}{4\pi} \frac{B_{x0} + B_{xm}}{2} \frac{2B_{zh}}{\Delta x}, \end{aligned} \quad (\text{B9})$$

where we have used $B_z = 0$ along the inflow symmetry line in Figure 3(a), and B_{zh} is evaluated at the separatrix. The anisotropy parameter at point 1 is

$$\begin{aligned} \varepsilon_1 &= 1 + \frac{4\pi P_{i0}(\bar{B}_{x1} - 1)}{B_{x1}^2} \left(1 - \frac{1 - \bar{B}_{x1}}{\sqrt{2\beta_{i0}} + 1} \right) \\ &= 1 + \frac{\beta_{i0}}{2} \left(\frac{1}{\bar{B}_{x1}} - \frac{1}{\bar{B}_{x1}^2} \right) \left(1 - \frac{1 - \bar{B}_{x1}}{\sqrt{2\beta_{i0}} + 1} \right), \end{aligned} \quad (\text{B10})$$

where

$$\bar{B}_{x1} = \frac{B_{x1}}{B_{x0}} = \frac{1 + \bar{B}_{xm}}{2}. \quad (\text{B11})$$

Since $\Delta z/\Delta x \simeq B_{zh}/B_{xh} \simeq B_{zh}/B_{x1} \simeq B_{zm}/B_{xm}$, then

$$\begin{aligned} B_{zh} &\simeq \frac{B_{zm} B_{xh}}{B_{xm}} \simeq \frac{B_{zm} B_{x0} + B_{xm}}{B_{xm}} \frac{1}{2} \\ &\simeq \frac{\Delta z}{\Delta x} \frac{B_{x0} + B_{xm}}{2}. \end{aligned} \quad (\text{B12})$$

The inflow force balance becomes

$$\begin{aligned} 1 - \bar{B}_{xm}^2 + \frac{\beta_0 + \beta_{i0}(\sqrt{2\beta_{i0}} + \bar{B}_{xm})}{\sqrt{2\beta_{i0}} + 1} (1 - \bar{B}_{xm}) \\ = \varepsilon_1 \left(\frac{\Delta z}{\Delta x} \right)^2 (1 + \bar{B}_{xm})^2, \end{aligned} \quad (\text{B13})$$

where $\beta_0 \equiv 8\pi P_0/B_{x0}^2$. We can numerically solve this equation to get \bar{B}_{xm} for a given slope $\Delta z/\Delta x$ (i.e., the opening angle $\theta \equiv \tan^{-1}(\Delta z/\Delta x)$) and plasma β_0 .

Appendix C

Outflow Force Balance

According to the momentum equation (Equation (B1)), the force balance equation that describes the outflow is

$$nm_i \mathbf{V} \cdot \nabla \mathbf{V} + \nabla \frac{B^2}{8\pi} + \nabla P_\perp \simeq \nabla \cdot \left(\varepsilon \frac{\mathbf{B}\mathbf{B}}{4\pi} \right). \quad (\text{C1})$$

At point 2 along the outflow direction (see Figure 3(b)),

$$\begin{aligned} \frac{n_2 m_i}{2} \frac{\partial V_x^2}{\partial x} + \frac{\partial}{\partial x} \frac{B^2}{8\pi} + \frac{\partial P_\perp}{\partial x} \\ = \frac{\partial}{\partial x} \left(\frac{\varepsilon B_x B_x}{4\pi} \right) + \frac{\partial}{\partial z} \left(\frac{\varepsilon B_z B_x}{4\pi} \right), \end{aligned} \quad (\text{C2})$$

where n_2 is the plasma density at point 2 in Figure 3(b). We again approximate $B_x \simeq B_x(z)$, $B_z \simeq B_z(x)$, and $\varepsilon \simeq \varepsilon(z)$ to simplify the problem. Since the magnetic field is primarily along the z -direction near the midplane ($z = 0$), $\partial P_\perp / \partial x \simeq \partial P_{xx} / \partial x$.

When ions are reflected by the outflow, they gain energy through the Fermi mechanism. Only ions that move toward the X-point, or those that move away from the X-point with a speed lower than the outflow speed V , will be reflected. As shown in Figure 3(c), particle velocity changes from v_x to $2V - v_x$ during the reflection. These particles will interpenetrate with incoming particles that have not been reflected yet and have v_x from $-\infty$ to V , i.e., particles with $v_x > V$ have escaped from the diffusion region (Figure 3(d)). The average velocity of these two populations is the outflow velocity V , as expected. This mixture results in a plasma density depending on V ,

$$n(V) = 2 \int_{-\infty}^V f(v_x) dv_x = n_0 + n_0 \operatorname{erf} \left(\sqrt{\frac{m_i}{2kT_0}} V \right), \quad (\text{C3})$$

where we use $f(v_x) = n_0 \sqrt{m_i / (2\pi kT_0)} \exp(-m_i v_x^2 / (2kT_0))$, a one-dimensional Maxwellian distribution with a temperature T_0 and a density n_0 . Here k is the Boltzmann constant. Then, the x -component of the diagonal part of the ion pressure tensor is

$$\begin{aligned} P_{ixx} &= 2m_i \int_{-\infty}^V (V - v_x)^2 f(v_x) dv_x \\ &= n_0 m_i \left[\left(V^2 + \frac{kT_0}{m_i} \right) \left(1 + \operatorname{erf} \left(\sqrt{\frac{m_i}{2kT_0}} V \right) \right) \right. \\ &\quad \left. + V \sqrt{\frac{2kT_0}{\pi m_i}} e^{-m_i V^2 / 2kT_0} \right]. \end{aligned} \quad (\text{C4})$$

Since the pressure P_{ixx} right at the x -line depleted by the inflowing plasma to near the background value, the increase of the ion pressure along the outflow can be modeled as

$$\begin{aligned} \Delta P_{ixx} &= P_{ixx} - n_0 kT_0 \\ &= n_0 m_i \left[V^2 + \left(V^2 + \frac{kT_0}{m_i} \right) \operatorname{erf} \left(\sqrt{\frac{m_i}{2kT_0}} V \right) \right. \\ &\quad \left. + V \sqrt{\frac{2kT_0}{\pi m_i}} e^{-m_i V^2 / 2kT_0} \right]. \end{aligned} \quad (\text{C5})$$

Inspired by in situ observations (and our simulations) that show a stronger ion heating than electron heating in reconnection exhausts (Phan et al. 2013, 2014), we ignore the electron heating in the following analysis. The outflow force balance

equation becomes

$$\begin{aligned} \frac{n_0 + n_m}{2} \frac{m_i V_{\text{out},m}^2}{2L} + \frac{B_{zm}^2}{8\pi L} + \frac{\Delta P_{ixx,m}}{L} \\ = \frac{1}{4\pi} \frac{B_{zm}}{2} \frac{(\varepsilon_m B_{xm}) / 2}{\delta / 2}, \end{aligned} \quad (\text{C6})$$

where we approximate $n_2 = (n_0 + n_m)/2$, n_m is evaluated from Equation (C3) using $V = V_{\text{out},m}$, $\Delta P_{ixx,m}$ is evaluated from Equation (C5) using $V = V_{\text{out},m}$, and $\varepsilon_m = \varepsilon(B_{xm})$ is the anisotropy parameter immediately upstream of the diffusion region,

$$\varepsilon_m = 1 + \frac{\beta_{i0}}{2} \left(\frac{1}{\bar{B}_{xm}} - \frac{1}{\bar{B}_{xm}^2} \right) \left(1 - \frac{1 - \bar{B}_{xm}}{\sqrt{2\beta_{i0} + 1}} \right). \quad (\text{C7})$$

Dividing Equation (C6) by $n_0 m_i v_{A0}^2$, we get the normalized equation

$$\begin{aligned} \frac{1}{4} \left[2 + \operatorname{erf} \left(\frac{\bar{V}_{\text{out},m}}{\sqrt{\beta_{i0}}} \right) \right] \bar{V}_{\text{out},m}^2 + \Delta \bar{P}_{ixx,m} \\ + \frac{\bar{B}_{xm}^2}{2} \left[\left(\frac{\Delta z}{\Delta x} \right)^2 - \varepsilon_m \right] = 0, \end{aligned} \quad (\text{C8})$$

where we have used $B_{zm}/B_{xm} \simeq \delta/L \simeq \Delta z/\Delta x$, $\beta_{i0} = 8\pi P_{i0}/B_{x0}^2$, $\bar{V}_{\text{out},m} = V_{\text{out},m}/v_{A0}$, $\bar{B}_{xm} = B_{xm}/B_{x0}$, $\bar{B}_{zm} = B_{zm}/B_{x0}$, and

$$\begin{aligned} \Delta \bar{P}_{ixx,m} &= \frac{4\pi \Delta P_{ixx,m}}{B_{x0}^2} \\ &= \bar{V}_{\text{out},m}^2 + \left(\bar{V}_{\text{out},m}^2 + \frac{\beta_{i0}}{2} \right) \operatorname{erf} \left(\frac{\bar{V}_{\text{out},m}}{\sqrt{\beta_{i0}}} \right) \\ &\quad + \bar{V}_{\text{out},m} \sqrt{\frac{\beta_{i0}}{\pi}} e^{-\bar{V}_{\text{out},m}^2 / \beta_{i0}}. \end{aligned} \quad (\text{C9})$$

ORCID iDs

Xiaocan Li  <https://orcid.org/0000-0001-5278-8029>
Yi-Hsin Liu  <https://orcid.org/0000-0001-5880-2645>

References

- Alt, A., & Kunz, M. W. 2019, *JPhPh*, **85**, 764850101
Bessho, N., & Bhattacharjee, A. 2010, *PhPl*, **17**, 102104
Birn, J., Borovsky, J. E., Hesse, M., & Schindler, K. 2010, *PhPl*, **17**, 052108
Birn, J., Drake, J. F., Shay, M. A., et al. 2001, *JGR*, **106**, 3715
Biskamp, D. 1986, *PhFl*, **29**, 1520
Bowers, K. J., Albright, B. J., Yin, L., Bergen, B., & Kwan, T. J. T. 2008, *PhPl*, **15**, 055703
Carilli, C. L., & Taylor, G. B. 2002, *ARA&A*, **40**, 319
Cassak, P. A., Liu, Y. H., & Shay, M. A. 2017, *JPhPh*, **83**, 715830501
Chew, G. F., Goldberger, M. L., & Low, F. E. 1956, *RSPSA*, **236**, 112
Colgate, S. A., Li, H., & Pariev, V. 2001, *PhPl*, **8**, 2425
Dahlin, J. T., Drake, J. F., & Swisdak, M. 2014, *PhPl*, **21**, 092304
Daughton, W., & Karimabadi, H. 2007, *PhPl*, **14**, 072303
Daughton, W., Roytershteyn, V., Karimabadi, H., et al. 2011, *NatPh*, **7**, 539
Drake, J. F., Opher, M., Swisdak, M., & Chamoun, J. N. 2010, *ApJ*, **709**, 963
Drake, J. F., Shay, M. A., Thongthai, W., & Swisdak, M. 2005, *PhRvL*, **94**, 095001
Drake, J. F., Swisdak, M., Che, H., & Shay, M. A. 2006, *Natur*, **443**, 553
Drake, J. F., Swisdak, M., Phan, T. D., et al. 2009, *JGRA*, **114**, 5111
Egedal, J., Daughton, W., Le, A., & Borg, A. L. 2015, *PhPl*, **22**, 101208
Egedal, J., Le, A., & Daughton, W. 2013, *PhPl*, **20**, 061201
Haggerty, C. C., Shay, M. A., Chasapis, A., et al. 2018, *PhPl*, **25**, 102120

- Haggerty, C. C., Shay, M. A., Drake, J. F., Phan, T. D., & McHugh, C. T. 2015, *GeoRL*, **42**, 9657
- Hoshino, M., Mukai, T., Terasawa, T., & Shinohara, I. 2001, *JGR*, **106**, 25979
- Huang, K., Liu, Y.-H., Lu, Q., & Hesse, M. 2020, *GeoRL*, **47**, e88147
- Jekel, C. F., & Venter, G. 2019, pwlf: A Python Library for Fitting 1D Continuous Piecewise Linear Functions, https://github.com/cjekel/piecewise_linear_fit_py
- Le, A., Egedal, J., Daughton, W., Fox, W., & Katz, N. 2009, *PhRvL*, **102**, 085001
- Li, X., Guo, F., & Li, H. 2019, *ApJ*, **879**, 5
- Li, X., Guo, F., Li, H., & Li, G. 2015, *ApJL*, **811**, L24
- Li, X., Guo, F., Li, H., & Li, G. 2017, *ApJ*, **843**, 21
- Lin, R. P. 2011, *SSRv*, **159**, 421
- Liu, Y.-H., Daughton, W., Karimabadi, H., Li, H., & Peter Gary, S. 2014, *PhPI*, **21**, 022113
- Liu, Y.-H., Daughton, W., Karimabadi, H., Li, H., & Roytershteyn, V. 2013, *PhRvL*, **110**, 265004
- Liu, Y.-H., Drake, J. F., & Swisdak, M. 2011, *PhPI*, **18**, 092102
- Liu, Y.-H., Drake, J. F., & Swisdak, M. 2012, *PhPI*, **19**, 022110
- Liu, Y.-H., Hesse, M., Guo, F., et al. 2017, *PhRvL*, **118**, 085101
- Liu, Y.-H., Li, T. C., Hesse, M., et al. 2019, *JGRA*, **124**, 2819
- Marrone, D. P., Moran, J. M., Zhao, J.-H., & Rao, R. 2007, *ApJL*, **654**, L57
- Masuda, S., Kosugi, T., Hara, H., Tsuneta, S., & Ogawara, Y. 1994, *Natur*, **371**, 495
- Øieroset, M., Lin, R. P., Phan, T. D., Larson, D. E., & Bale, S. D. 2002, *PhRvL*, **89**, 195001
- Oka, M., Phan, T.-D., Krucker, S., Fujimoto, M., & Shinohara, I. 2010, *ApJ*, **714**, 915
- Parker, E. N. 1957, *JGR*, **62**, 509
- Parker, E. N. 1963, *ApJS*, **8**, 177
- Petschek, H. 1964, in AAS-NASA Symp. 50 The Physics of Solar Flares, ed. W. N. Hess (Washington, DC: NASA), 425
- Phan, T. D., Drake, J. F., Shay, M. A., et al. 2014, *GeoRL*, **41**, 7002
- Phan, T. D., Shay, M. A., Gosling, J. T., et al. 2013, *GeoRL*, **40**, 4475
- Qiu, J., Lee, J., Gary, D. E., & Wang, H. 2002, *ApJ*, **565**, 1335
- Sato, T., & Hayashi, T. 1979, *PhFI*, **22**, 1189
- Schekochihin, A. A., & Cowley, S. C. 2006, *PhPI*, **13**, 056501
- Schindler, K. 1974, *JGR*, **79**, 2803
- Schoeffler, K. M., Drake, J. F., & Swisdak, M. 2011, *ApJ*, **743**, 70
- Shay, M. A., Drake, J. F., Rogers, B. N., & Denton, R. E. 2001, *JGR*, **106**, 3759
- Shuster, J. R., Chen, L. J., Hesse, M., et al. 2015, *GeoRL*, **42**, 2586
- Sonnerup, B. U. O., Paschmann, G., Papamastorakis, I., et al. 1981, *JGR*, **86**, 10049
- Sweet, P. A. 1958, in IAU Symp. 6, Electromagnetic Phenomena in Cosmical Physics, ed. B. Lehnert (Cambridge: Cambridge Univ. Press), 123
- Wang, C.-P., Liu, Y.-H., Xing, X., et al. 2020, *JGRA*, **125**, e27406
- Wang, C.-P., Xing, X., Liu, Y.-H., & Runov, A. 2018, *JGRA*, **123**, 1822
- Wang, S., Chen, L.-J., Bessho, N., et al. 2016, *JGRA*, **121**, 2104
- Wang, S., Kistler, L. M., Mouikis, C. G., & Petrinec, S. M. 2015, *JGRA*, **120**, 6386
- Yamada, M., Levinton, F. M., Pomphrey, N., et al. 1994, *PhPI*, **1**, 3269
- Yamada, M., Ren, Y., Ji, H., et al. 2006, *PhPI*, **13**, 052119
- Yokoyama, T., Akita, K., Morimoto, T., Inoue, K., & Newmark, J. 2001, *ApJL*, **546**, L69
- Zhang, B., & Yan, H. 2011, *ApJ*, **726**, 90
- Zweibel, E. G., & Yamada, M. 2009, *ARA&A*, **47**, 291