

DRAFT VERSION JANUARY 6, 2021

Typeset using L^AT_EX preprint style in AASTeX63

A New Method for Simulating Photoprocesses in Astrochemical Models

ELLA MULLIKIN,¹ HANNAH ANDERSON,¹ NATALIE O'HERN,¹ MEGAN FARRAH,¹
CHRISTOPHER R. ARUMAINAYAGAM ¹ EWINE F. VAN DISHOCK ^{2,3} PERRY A. GERAKINES ⁴,
LITON MAJUMDAR ⁵ PAOLA CASELLI ⁶ AND CHRISTOPHER N. SHINGLEDECKER ^{6,7,8}

¹*Department of Chemistry, Wellesley College, Wellesley, MA 02481, USA*

²*Leiden Observatory, Leiden University, P.O. Box 9513, NL-2300 RA Leiden, The Netherlands*

³*Max-Planck-Institut für extraterrestrische Physik, D-85748 Garching, Germany*

⁴*Astrochemistry Laboratory, NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA*

⁵*School of Earth and Planetary Sciences, National Institute of Science Education and Research, HBNI, Jatni 752050, Odisha, India*

⁶*Center for Astrochemical Studies Max Planck Institute for Extraterrestrial Physics Garching, Germany*

⁷*Institute for Theoretical Chemistry University of Stuttgart Pfaffenwaldring 55, 70569 Stuttgart, Germany*

⁸*Department of Physics & Astronomy, Benedictine College, Atchison, KS 66002, USA*

(Received –; Revised –; Accepted –)

Submitted to ApJ

ABSTRACT

We propose a new model for treating solid-phase photoprocesses in interstellar ice analogues. In this approach, photoionization and photoexcitation are included in more detail, and the production of electronically-excited (suprathermal) species is explicitly considered. In addition, we have included non-thermal, non-diffusive chemistry to account for the low-temperature characteristic of cold cores. As an initial test of our method, we have simulated two previous experimental studies involving the UV irradiation of pure solid O₂. In contrast to previous solid-state astrochemical model calculations which have used gas-phase photoabsorption cross-sections, we have employed solid-state cross-sections in our calculations. This method allows the model to be tested using well-constrained experiments rather than poorly constrained gas-phase abundances in ISM regions. Our results indicate that inclusion of non-thermal reactions and suprathermal species allows for reproduction of low-temperature solid-phase photoprocessing that simulate interstellar ices within cold (~ 10 K) dense cores such as TMC-1.

Keywords: astrochemistry — ISM, molecules — ISM, molecular processes, photoprocessing, astrochemical modeling

1. INTRODUCTION

Corresponding author: Christopher N. Shingledecker
cshingledecker@benedictine.edu

While gas-phase and surface reactions on bare carbonaceous or silicate dust grains contribute to cosmic chemistry, the energetic processing of cosmic ices within dark, dense molecular clouds via photochemistry (initiated by non-ionizing radiation) and radiation chemistry (initiated by ionizing radiation) is thought to be the dominant mechanism for the interstellar synthesis of prebiotic molecules (see, for example, review: (Arumainayagam et al. 2019)). Rate-equation based modeling treatments of UV-induced condensed-phase photochemistry have been moderately successful in reproducing the abundances of complex organic molecules (COMs) observed toward hot cores/corinos (Shingledecker et al. 2019b; Grassi et al. 2014; McElroy et al. 2013; Garrod 2013). However, recent detections of several COMs (e.g., methyl formate (HCOOCH_3) and dimethyl ether (CH_3OCH_3)) in cold (~ 10 K) dense cores (Vastel et al. 2014; Taquet et al. 2017; Scibelli & Shirley 2020; Bacmann et al. 2012; Jiménez-Serra et al. 2016; Öberg et al. 2010), albeit in smaller abundance than in hot cores, have led to the search for alternative mechanisms for complex molecule production through cold or non-thermal mechanisms (Shingledecker et al. 2018; Vasyunin et al. 2017). A recent radiolysis-related computational study (Shingledecker et al. 2019b) has provided an explanation for the unprecedented observations of chemical synthesis at temperatures as low as 10 K in starless and prestellar cores. In this modified bulk-chemistry method involving radiolysis by cosmic rays, radicals produced within the ice are considered to be trapped and attempt to react with a neighbor with approximately every vibration. In the study described herein, we use this non-diffusive mechanism to revise the treatment of solid-phase photoprocesses in astrochemical models to account for the complex organics observed in cold cores. A recent study by Jin & Garrod (2020) utilizes a non-diffusive rate-based model which demonstrates the dependence of COM production on non-diffusive reactions between radicals and ice species in cold astrochemical environments and achieves considerable success in reproducing observations toward prestellar core L1544. In contrast to that study, the model presented here incorporates (1) the detailed inclusion of photoionization and photoexcitation, and (2) explicit consideration of the production and reaction of electronically excited radicals (suprathermal species).

One of the main processing mechanisms of ices in molecular clouds is radiation chemistry, which involves ionization and the production of copious numbers of low-energy (< 15 eV) electrons, which are thought to be the dominant species involved in radiation chemistry (e.g., (Arumainayagam et al. 2010)). Ionizing radiation present in this environment include MeV to TeV cosmic rays ($\sim 85\%$ H^+ , $\sim 13\%$ He^{2+} , $\sim 1\%$ heavy bare nuclei, and $\sim 1\%$ electrons) and high-energy photons (e.g., vacuum ultraviolet photons with energies higher than ~ 10 eV, extreme ultraviolet, X-ray, and γ -ray).

Whereas high-energy photons contribute to radiation chemistry in dense molecular clouds, low-energy (< 10 eV) photons (e.g., far (deep)-UV (4.1 – 6.2 eV)) initiate photochemistry, a process that does not involve direct ionization, but instead involves reactions of electronically-excited species. The UV interstellar radiation field, consisting of radiation from nearby stars, is extinguished by dust well before reaching the interior of dark, dense molecular clouds where prebiotic molecules are synthesized. However, a local secondary UV field exists to initiate photo-processing of dust grain ices (Prasad & Tarafdar 1983). Cosmic rays excite gaseous molecular hydrogen, resulting in Lyman and Werner band emission with an estimated flux of $\sim 10^3 - 10^4$ photons $\text{cm}^{-2} \text{s}^{-1}$ (Gredel et al. 1989; Cruz-Diaz et al. 2014; Shen et al. 2004). Although this spectral distribution includes high-energy (10 to 13.6 eV) photons, over half of the secondary UV field consists of low-energy (< 10 eV) photons capable of photochemistry by exciting condensed species which may then react within the ice. Except during high photon-flux laser experiments which may involve multi-photon processes, photochemistry is

subject to the Bunsen-Roscoe law, which states that the photochemical yield is directly proportional to dose, irrespective of dose rate; this law allows for extrapolation from laboratory experiments to real astrochemical predictions, though even these low-flux experiments utilize fluxes much higher than those experienced by ice in dark, dense molecular clouds.

Astrochemical models provide a critical link between the fundamental chemical information revealed by laboratory experiments and predictions and observations of chemical abundances in the interstellar medium. Most models utilize a rate-based approach due to the convenience and speed of that method, though Monte Carlo models have been used to more accurately simulate processes such as the catastrophic impact of high energy radiation or multi-layer interactions (Cuppen et al. 2017; Öberg 2016). Because all such simulations involving reaction networks and rate-equations are extremely sensitive to input parameters, these models generally become more successful as parameters are better constrained by laboratory experiments. Abundances of several COMs in hot cores/corinos are well reproduced by modern rate-equation-based computational models, which include a coupled gas-phase and grain-surface chemistry or three-phase (gas, surface, and bulk) chemistry (Cuppen et al. 2017).

Abundances of COMs in cold (~ 10 K) cores, however, are generally underpredicted. Most astrochemical models that include bulk-phase processes require thermal diffusion before reaction (Cuppen et al. 2017); however, in the low-temperature conditions of starless and prestellar cores, this thermal motion within the bulk ice is not feasible. Several explanations for the formation of COMs in cold dense cores have been proposed, including: (1) photo-processing followed by reactive desorption (Watanabe & Kouchi 2002; Chuang et al. 2017; Aikawa et al. 2008; Herbst & van Dishoeck 2009; Jin & Garrod 2020), (2) gas-phase reactions (Balucani et al. 2015; Codella et al. 2020), and (3) astrophysical-shock-catalyzed chemistry (James et al. 2020), (4) methanol reactive desorption on CO-rich ices (within the catastrophic CO-freeze-out zone of pre-stellar cores), followed by gas-phase chemistry (Vasyunin et al. 2017). The model presented herein instead assumes that reactive species are trapped within the bulk ice, but have the possibility of reacting with neighboring molecules during each vibration. This work is an extension of a previous study (modeling the physicochemical effects of astrochemical O_2 and H_2O ice-analogue bombardment by energetic protons), which revealed the importance of considering fast non-thermal reactions in these systems (Shingledecker et al. 2019b). In what follows, we apply this assumption to the case of photon irradiation of cosmic-ice analogues. The model, as utilized in this work, includes only photon-initiated ice processing, including both excitation and ionization events. Cations produced via ionization are assumed to quickly recombine with a secondary electron, resulting in the electronically excited parent molecule, which can then dissociate into electronically-excited products. Models such as ours are essential for interpreting planetary and interstellar ice data generated by past, upcoming, and ongoing NASA missions such as Spitzer, Stratospheric Observatory for Infrared Astronomy (SOFIA), and the James Webb Space Telescope (JWST).

As a proof of concept, this new model is used to simulate two published laboratory studies that monitored the processing of O_2 ice by < 10.8 eV photons. Given that interstellar ice mantles are likely segregated into polar and apolar layers, it is useful to “tune” the model to simulate the photo-processing of a single species accurately, and then these species-specific models may be combined to simulate processing of the layers of more realistic cosmic ice analogues (Tielens et al. 1991; Pontoppidan 2006; Öberg et al. 2009, 2011). The first study (Gerakines et al. 1996) employed a

microwave-discharge hydrogen-flow lamp (MDHL) with a photon flux of 2.2×10^{14} photons $\text{cm}^{-2} \text{s}^{-1}$. The MDHL spectrum closely reproduces the calculated dark, dense molecular cloud secondary UV spectrum above 115 nm (below 10.8 eV). However, the fraction of Lyman alpha emission in a MDHL spectrum can change significantly based on the experimental settings such as microwave power and gas pressure (Ligterink et al. 2015). In Gerakines et al. (1996), oxygen ices (~ 100 nm in thickness) were deposited at ~ 10 K inside a vacuum chamber 10^{-7} Torr, mimicking conditions relevant to those of interstellar ices in cold cores. Two capping layers of argon precluded both contamination and significant desorption from the ices during photon irradiation. The oxygen ice was irradiated with photons for one hour corresponding to a maximum fluence of 7.9×10^{17} photons cm^{-2} , corresponding to approximately a million years of secondary interstellar UV irradiation. Production of O_3 during the irradiation was monitored via the 1043 cm^{-1} IR feature of O_3 .

In the second experiment (Raut et al. 2011), a pulsed ArF excimer laser (193 nm), defocused using a MgF_2 lens, with a flux of $\sim 2.3 \times 10^{14}$ photons $\text{cm}^{-2} \text{s}^{-1}$, was used to irradiate O_2 ices 80–84 nm thick. The use of 6.4 eV photons precludes radiation chemistry. The maximum fluence was 9.3×10^{18} photons cm^{-2} . A vacuum chamber with a base pressure of 10^{-9} Torr and an ice temperature of 22 K simulated interstellar-like conditions. The ozone column density as a function of fluence was monitored via the 1043 cm^{-1} and 2109 cm^{-1} IR features of O_3 .

The goal of the present work is to improve the current understanding of ice chemistry initiated by interstellar secondary UV radiation within dark, dense molecular clouds during the starless and prestellar stages well before the formation of protostars and planets. The successful reproduction of experimental results herein indicates that the inclusion of non-thermal reactions and suprathreshold species will allow for more accurate modeling of interstellar photoprocessing of ices in cold cores. Additional tuning of the model for other species, such as water, will render the model suitable for predicting cold-core COM abundances attributable to photo-processing of mixed ices.

2. METHODS

2.1. Theory

As in Shingledecker & Herbst (2018), the starting point of our proposed model is the assumption that the interaction between a UV photon and some target species, A , results in one of the following outcomes:

$$A \rightsquigarrow A^+ + e^- \quad (\text{P1})$$

$$A \rightsquigarrow A^+ + e^- \rightarrow A^* \rightarrow B^* + C^* \quad (\text{P2})$$

$$A \rightsquigarrow A^* \rightarrow B + C \quad (\text{P3})$$

$$A \rightsquigarrow A^* \quad (\text{P4})$$

Here, the curly arrow ($\rightsquigarrow$) represents the absorption of a photons, B and C are dissociation products, and $*$ indicates an electronically excited (suprathreshold) species. Of the four processes given above, (P1) and (P2) correspond to the photoionization of A to the cation A^+ — followed by the rapid recombination of the charged products in the case of (P2) — and are relevant in solids for $h\nu \gtrsim$

10 eV (Arumainayagam et al. 2019). Similarly, (P3) and (P4) account for photoexcitation to the excited state A^* , with (P3) leading to the photodissociation of A . One advantage of separating photoionization and photoexcitation is that the former can be enabled or disabled based on the energy of the incident photons.

The rate coefficients of photoionization and dissociation processes, k_{photo} , are usually calculated using

$$k_{\text{photo}} = \int \sigma(\lambda) I(\lambda) d\lambda \quad (1)$$

where here, $\sigma(\lambda)$ and $I(\lambda)$ are wavelength-dependent cross-section and photon flux, respectively. This formula can also be expressed as

$$k_{\text{photo}} = \frac{\int \sigma(\lambda) I(\lambda) d\lambda}{\int I(\lambda) d\lambda} \int I(\lambda) d\lambda = \bar{\sigma} \Phi \quad (2)$$

where $\bar{\sigma}$ is the average cross-section, and Φ is the integrated photon flux. Following Shingledecker & Herbst (2018), we can then express the rates of (P1) – (P4) in the following way:

$$k_{\text{P1}} = P_e \bar{\sigma}_{\text{ion}} \Phi \delta \quad (3)$$

$$k_{\text{P2}} = (1 - P_e) \bar{\sigma}_{\text{ion}} \Phi \delta \quad (4)$$

$$k_{\text{P3}} = P_{\text{dis}} \bar{\sigma}_{\text{exc}} \Phi \delta \quad (5)$$

$$k_{\text{P4}} = (1 - P_{\text{dis}}) \bar{\sigma}_{\text{exc}} \Phi \delta. \quad (6)$$

Here, P_e is the electron escape probability (Elkomoss & Magee 1962), which we assume as a first approximation is equal to zero. A more comprehensive model will need to relax this approximation to account for the effects of low-energy secondary electrons thought to be the primary agents of radiation chemistry. All ionized molecules are assumed to quickly recombine to form an excited molecule which will subsequently dissociate, react, or be quenched. Quenching by the surrounding ice is assumed to be the dominant relaxation mechanism rather than radiative relaxation, and the attempt frequency is used as the first-order rate constant for this process (Shingledecker et al. 2019a). In reality, electronic excitations (excitons) may diffuse from the interior of the ice to the selvedge where they can drive the desorption of species into the gas (Thrower et al. 2011; Marchione et al. 2016). P_{dis} is the dissociation probability, which is ~ 1 in the gas, but in solids, we assume it to be 0.5 as a first approximation. This value was adjusted to account for spectral characteristics in later simulations. Dissociation products can recombine to reform the parent species, but this recombination is not assumed to occur preferentially to – or is calculated differently than – any other possible chemical reaction with other bulk species the fragments could undergo. The δ is a fitting factor that was introduced to account for assumptions of the model and absolute uncertainties in experimental data such as photon flux. More explicitly, δ is sensitive to, e.g., (a) the reactions and photoproducts included in the chemical network, (b) the associated branching fractions or cross sections, as well as (c) the methods for treating the underlying physical processes employed in the code. Thus, reasonable agreement between calculated

Table 1. Reactions comprising the chemical network used to model O₂ photo-processing and subsequent chemistry.

Photon-Induced Reactions	Non-Photon-Induced Reactions		
$\text{O} \xrightarrow{\text{ionization}} \text{O}^+ + \text{e}^- \rightarrow \text{O}^*$	$\text{O}_3^* + \text{O}_3 \rightarrow \text{O}_2 + \text{O}_2 + \text{O}_2$	$\text{O}_2^* + \text{O} \rightarrow \text{O}_3^*$	$\text{O}^* + \text{O}_3^* \rightarrow \text{O}_2 + \text{O}_2$
$\text{O}_2 \xrightarrow{\text{ionization}} \text{O}_2^+ + \text{e}^- \rightarrow \text{O}_2^* \rightarrow \text{O}^* + \text{O}^*$	$\text{O} + \text{O} \rightarrow \text{O}_2^*$	$\text{O}_2^* + \text{O} \rightarrow \text{O}_2 + \text{O}$	$\text{O}_2^* + \text{O}_2^* \rightarrow \text{O}_2 + \text{O}_2$
$\text{O}_3 \xrightarrow{\text{ionization}} \text{O}_3^+ + \text{e}^- \rightarrow \text{O}_3^* \rightarrow \text{O}_2^* + \text{O}^*$	$\text{O}^* + \text{O} \rightarrow \text{O} + \text{O}$	$\text{O}_2^* + \text{O}_2 \rightarrow \text{O}_2 + \text{O}_2$	$\text{O}_2^* + \text{O}_3^* \rightarrow \text{O}_2 + \text{O}_2 + \text{O}$
$\text{O} \xrightarrow{\text{excitation}} \text{O}^*$	$\text{O}^* + \text{O} \rightarrow \text{O}_2^*$	$\text{O}_2^* + \text{O}_3 \rightarrow \text{O}_2 + \text{O}_2 + \text{O}$	$\text{O}_3^* + \text{O}_3^* \rightarrow \text{O}_2 + \text{O}_2 + \text{O}_2$
$\text{O}_2 \xrightarrow{\text{excitation}} \text{O}_2^*$	$\text{O}^* + \text{O}_2 \rightarrow \text{O}_3^*$	$\text{O}^* + \text{O}^* \rightarrow \text{O} + \text{O}$	$\text{O} + \text{O}_3^* \rightarrow \text{O}_2 + \text{O}_2$
$\text{O}_3 \xrightarrow{\text{excitation}} \text{O}_3^*$	$\text{O}^* + \text{O}_2 \rightarrow \text{O} + \text{O}_2$	$\text{O}^* + \text{O}_2^* \rightarrow \text{O}_3^*$	$\text{O}_2 + \text{O}_3^* \rightarrow \text{O}_2 + \text{O}_2 + \text{O}$
$\text{O}_2 \xrightarrow{\text{excitation}} \text{O}_2^* \rightarrow \text{O} + \text{O}$	$\text{O} + \text{O}_3 \rightarrow \text{O}_2 + \text{O}_2$	$\text{O}^* + \text{O}_2^* \rightarrow \text{O} + \text{O}_2$	$\text{O} + \text{O}_2 \rightarrow \text{O}_3^*$
$\text{O}_3 \xrightarrow{\text{excitation}} \text{O}_3^* \rightarrow \text{O}_2 + \text{O}$	$\text{O}^* + \text{O}_3 \rightarrow \text{O}_2 + \text{O}_2$		

and experimental data obtained assuming $\delta \approx 1$ for all photoprocesses would suggest that (a), (b), and (c) capture the salient features of a given system. Conversely, shortcomings in (a), (b) or (c) can be compensated for to some degree by adjusting δ values to yield best agreement with experimental results. For photoprocesses occurring in the bulk of optically thick ices, an extinction factor, ϵ , can be included in Eqs. (3) – (6) to account for the reduced photon flux relative to the ice surface. Because the two experimental studies of interest used optically thin ices, the extinction factor was set equal to 1.

2.2. Model

In this work, we have utilized the MONACO model (Vasyunin et al. 2017), previously modified by us, to simulate ice radiation chemistry experiments (Shingledecker et al. 2019b). This code, written in Fortran 90, solves a system of coupled differential equations describing the evolution of the abundance of each species in our network. Unlike comparable astrochemical models, the model described herein accounts for electronically-excited suprathermal species produced during photo-processing of ices. Table 1 presents all photon-induced and non-photon-induced (reactions involving products of the initial photo-processing) reactions included in the model network for O₂. Reactions occurring in the seldedge, considered to be the top four monolayers of the ice (Vasyunin & Herbst 2013), are assumed to occur via the Langmuir-Hinshelwood mechanism, and rate coefficients are calculated using the standard formula for diffusive processes. For reactions in the bulk, rate coefficients are calculated using the non-diffusive formula of Shingledecker et al. (2019b),

$$k_{\text{fast}} = f_{\text{br}} \left[\frac{\nu_0^A + \nu_0^B}{N_{\text{bulk}}} \right] \exp \left(-\frac{E_{\text{act}}^{AB}}{T_{\text{ice}}} \right), \quad (7)$$

where f_{br} is the branching fraction, T_{ice} is the ice temperature, E_{act}^{AB} is the activation energy in Kelvins for reaction, N_{bulk} is the total number of bulk species in the simulated ice, and ν_0^A is the characteristic (hereafter, trial) vibrational frequency (Herbst & Millar 2008).

In our model, species in the seldedge can, in principle, desorb both thermally as well as following exothermic association reactions, with the latter being treated by the method of Garrod (2008) with a standard efficiency of 1%. To mimic the Gerakines experiments where such desorption would be

inhibited due to the presence of a capping noble gas layer, thermal, chemical, and photodesorption processes were disabled in our models. The ices studied by Raut and coworkers lacked such a noble gas cap, and thus some amount of desorption would have occurred during the course of the experiment. However, given the fairly low temperature (22 K) of the bulk ice, we assume that thermal desorption is negligible over the timescale of the experiment. Moreover, photodesorption for the Raut et al. experiments was also disabled, since the desorption rate is not well constrained (Fayolle et al. 2013; Bulak et al. 2020) and our focus for this study was, in any case, the chemistry occurring *within* the bulk ice. Chemical desorption is included for the simulation of the Raut et al. data but was found to have a negligible impact on the bulk chemistry we describe in detail below.

Rate constants for photon-induced reactions are dependent on the average cross-sections ($\bar{\sigma}$), photon flux (Φ), and the fitting factor (δ) as described in §2.1. Each product of photoionization or photoexcitation is treated as being trapped in a cage of neighboring bulk species molecules; reactions involving suprathermal species are assumed to be barrierless. Non-photon-induced reactions are assumed to occur non-diffusively, with rates of reaction between any two species being proportional to their abundances in the ice (Shingledecker et al. 2019b).

For the pure O₂ ice, we use the chemical network (Table 1) initially described in Shingledecker et al. (2017) and used in the microscopic Monte Carlo model, CIRIS, and later modified for use in rate-based kinetic codes in Shingledecker et al. (2019b). The choice of O₂ as the bulk species for this initial test of the model is appropriate given the relative simplicity of the products and subsequent possible reactions, especially compared to species such as water or methanol. The selvedge, which comprises the chemically distinct region near the top of the mantle, is considered to be the top four monolayers of the ice (Vasyunin & Herbst 2013).

Parameters relevant to the simulation of laboratory experiments include ice thickness, photon fluence (photon flux multiplied by irradiation time), and photon energy; these values were obtained directly from the manuscripts of the experiments chosen for simulation. The trial frequency, ν , parameterizes the vibrational frequency of a molecule, used as the pre-exponential factor in calculating bulk rate coefficients. The model assumes that with every vibration, there is a probability that a molecule will react with a neighboring molecule (Eq. (7)). For all simulations, the vibrational frequency was set to $1 \times 10^{15} \text{ s}^{-1}$, which is reasonable, assuming RRKM theory. Increasing the value by orders of magnitude has negligible impact on model simulations, while reducing the value below $1 \times 10^{15} \text{ s}^{-1}$ resulted in significant deviations from experimental data.

Specific to each photoprocess included in the chemical network are cross-sections and the δ fitting factor. Branching ratios for reactions with more than one product channel were assumed to occur with equal probability in early simulations, and later adjusted to match spectral characteristics. Cross-sections were obtained from various sources, as detailed below.

3. RESULTS AND DISCUSSION

We directly tested the validity of our new method by replicating experimental data of O₂ ice irradiation with UV photons ($< 10.8 \text{ eV}$). In contrast to simulations of poorly constrained gas-phase abundances in ISM regions, this method allows the model to be tested using well-constrained experiments. Simulated pure oxygen ice experiments shared the use of interstellar-like temperatures and pressures. The results of these simulations are described below.

3.1. Microwave discharge hydrogen flow lamp source

Table 2. Calculated average cross-sections and δ -values for pure O₂ ice irradiated by a MDHL and ArF laser.

Process	Type	$\bar{\sigma}_{\text{MDHL}}$ (cm ²)	$\bar{\sigma}_{\text{ArF}}$ (cm ²)	δ_{MDHL}	δ_{ArF}
O + $h\nu \rightarrow \text{O}^*$	(P2)	0	0	1.0	1.0
O + $h\nu \rightarrow \text{O}^*$	(P4)	0	0	1.0	1.0
O ₂ + $h\nu \rightarrow \text{O}^* + \text{O}^*$	(P2)	3.86×10^{-20}	0	1.0	1.0
O ₂ + $h\nu \rightarrow \text{O} + \text{O}$	(P3)	2.13×10^{-18}	2.10×10^{-19}	2.3	1.9 - 2.2
O ₂ + $h\nu \rightarrow \text{O}_2^*$	(P4)	2.13×10^{-18}	2.10×10^{-19}	1.0	0.25 - 0.35
O ₃ + $h\nu \rightarrow \text{O}_2^* + \text{O}^*$	(P2)	0	0	1.0	1.0
O ₃ + $h\nu \rightarrow \text{O}_2 + \text{O}$	(P3)	5.60×10^{-18}	2.15×10^{-18}	1.0	0.25 - 0.35
O ₃ + $h\nu \rightarrow \text{O}_3^*$	(P4)	5.60×10^{-18}	2.15×10^{-18}	1.0	0.25 - 0.35

NOTE—the values of δ_{MDHL} , were used to produce Figure 1. For δ_{ArF} , ranges of values are shown which were found to yield agreement with the data within experimental error.

Previous work by Gerakines et al. (1996) provides excellent data with which to quantify the validity of our approach. Their experiments on pure O₂ were carried out at 10 K. They utilized a microwave discharge hydrogen flow lamp (MDHL), two layers of inert argon, film thicknesses on the order of 0.1 μm , photon fluxes of $\sim 10^{14}$ photons cm⁻² s⁻¹, and an irradiation time of ~ 1 hr (Gerakines et al. 1996; Jenniskens et al. 1993). The O₃ production curve given in figure 8 of Gerakines et al. was digitized for comparison to the model output of O₃ abundance (Fig. 1). Listed in Table 2 are the effective cross sections, $\bar{\sigma}$, for this experiment. To calculate the effective solid-phase cross section for O₂ photoabsorption, first a reported spectrum of solid-phase O₂ absorption (but not cross section) as a function of wavelength was digitized (Lu et al. 2008). Next, this absorption data was scaled to the digitized solid phase cross section data of Mason et al. (2006) in order to obtain cross section values over a broader wavelength range corresponding to the spectrum of a MDHL. The spectrum of a MDHL (Jenniskens et al. 1993) was digitized, and intensity and cross section were multiplied together at each wavelength. Finally, the product was integrated over all wavelengths and divided by total flux (Eq. (2)). The effective solid-phase O₃ photoabsorption cross section was calculated by first scaling gas-phase absorption data from Sivaraman et al. (2014) to gas-phase cross section data from the Leiden database (Heays et al. 2017). Solid-phase absorption data taken in the same laboratory (Sivaraman et al. 2014) were multiplied by the same scaling factor. The resulting solid-phase photoabsorption cross section data as a function of wavelength was used to calculate effective solid-phase cross section. To obtain cross section values for the excited-state reaction path (e.g., (P4)) and the dissociation path (e.g., (P3)), total photoabsorption cross-sections were multiplied by the corresponding branching ratio. Average solid-phase cross sections for O₂ and O₃ photoionization were obtained by using gas phase cross section data from the Leiden database but shifting all data points by 1.5 eV for application to condensed species (Kahn 2015; Yu et al. 1975). After inserting the experimental parameters of ice thickness, photon flux, photon-irradiation time, and reaction cross sections, the δ fitting factors given in Table 2 were obtained by manually adjusting to maximize

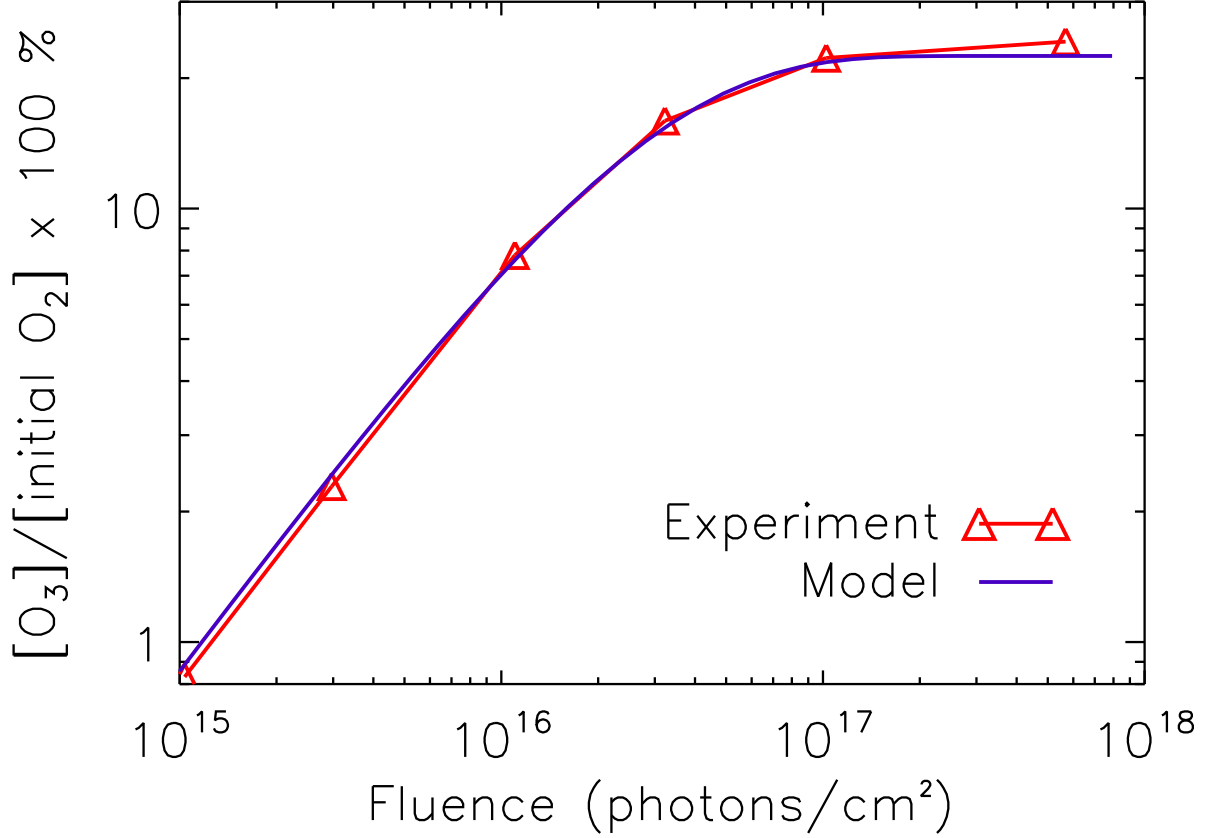


Figure 1. Calculated abundances of O₃ vs. photon fluence in a UV-irradiated pure O₂ ice (shown in blue), with corresponding experimental data from Gerakines et al. (1996) (shown in red). The values shown in Table 2 were used for branching ratios and cross sections. No error bars were reported in this study

the agreement with experimental data. The optimized steady-state O₃ abundance agrees with the experimental value to within 5%.

Only the δ fitting factor for O₂ photodissociation varied from 1.0 for agreement with the experimental data. Given the broad absorption cross-section peak even in the solid-phase spectrum and the high probability of dissociation following photoabsorption for gaseous O₂, a simulation was run with a branching ratio of 99:1 for O₂ dissociation to O₂ excited-state reaction, the results of which are shown in Fig. 2. All δ values could then be set to 1.0 for similar agreement to experimental data as when the dissociation and excited-state reactions were assumed to be equally likely but a δ value of 2.3 was required for the dissociation channel. Thus, the original deviation of δ from 1.0 was necessary to account for the high O₂ photodissociation probability when it was not otherwise included in the model.

3.2. Pulsed 193 nm ArF excimer laser source

A different condensed-phase O₂ experiment (Raut et al. 2011), which utilized a pulsed laser UV source, was also simulated to test the validity of our model. These experiments were conducted at

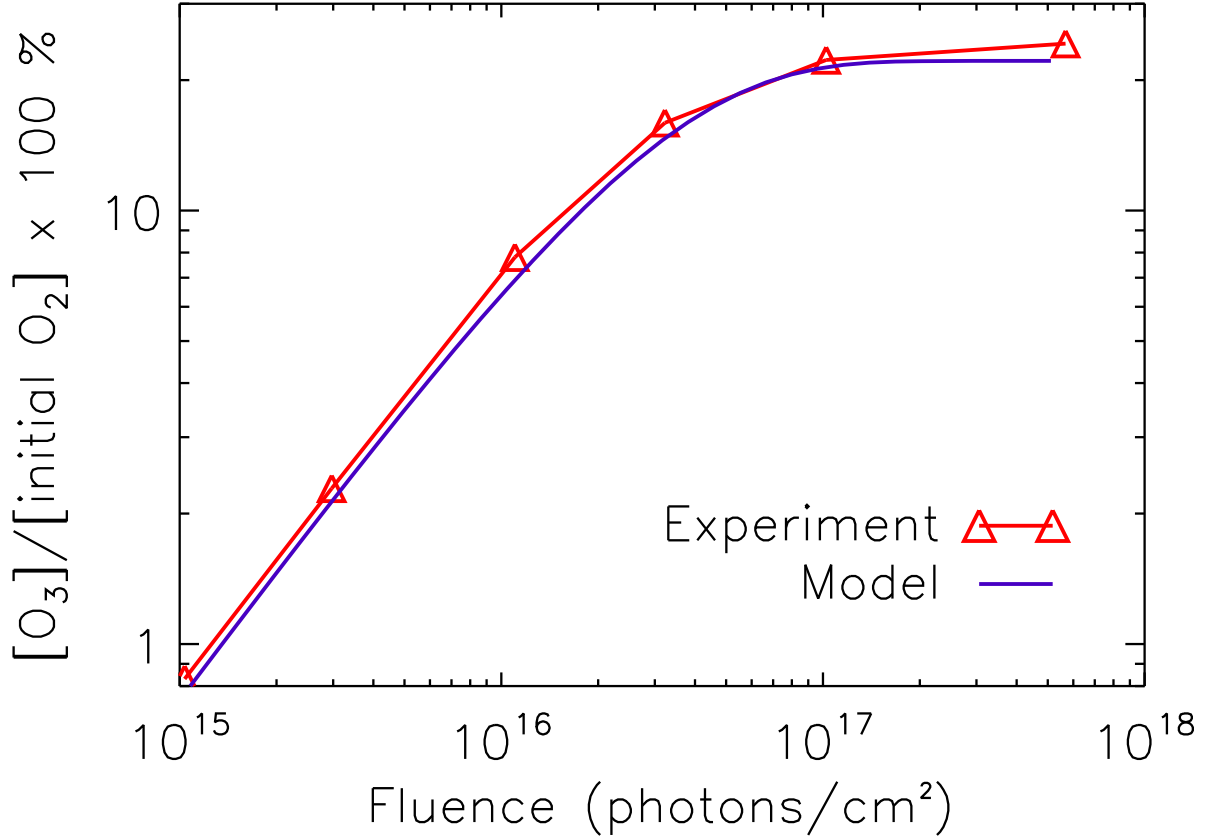


Figure 2. Calculated abundances of O₃ vs. photon fluence in a UV-photodissociated pure O₂ ice (shown in blue), with corresponding experimental data from Gerakines et al. (1996) (shown in red). For this simulation, the branching ratio of O₂ photodissociation to O₂ photoexcitement was set to 0.99, and all δ -values were set to 1.0.

22 K with a film thickness of ~ 80 nm, a photon flux averaging $\sim 10^{14}$ photons cm⁻² s⁻¹, and a total fluence of $\sim 10^{19}$ photons cm⁻² with 193 nm (6.4 eV) photons. The O₃ abundance data in Figure 5 of Raut et al. was digitized for comparison to the model output. Because the O₃ abundance was reported as column density, the model output was scaled to account for the thickness of the ice. In this case, cross-sections were provided in the experiment manuscript: the solid phase 193 nm photoabsorption cross section for O₂ and O₃ are reported as 4.2×10^{-19} cm² and 4.3×10^{-19} cm², respectively. To confirm these reported cross sections, two independently reported spectra of solid state O₂ absorption (but not cross-section) as a function of wavelength were digitized (Cruz-Diaz et al. 2014; Lu et al. 2008). Each included data at 193 nm. Next, a published spectrum of solid-phase O₂ cross section as a function of wavelength was also digitized (Mason et al. 2006); this spectrum did not include data at 193 nm. The absorption spectra were then scaled to the cross-section spectrum. The values found in the scaled data at 193 nm matched the cross-section values used by Raut et al. to within 25%.

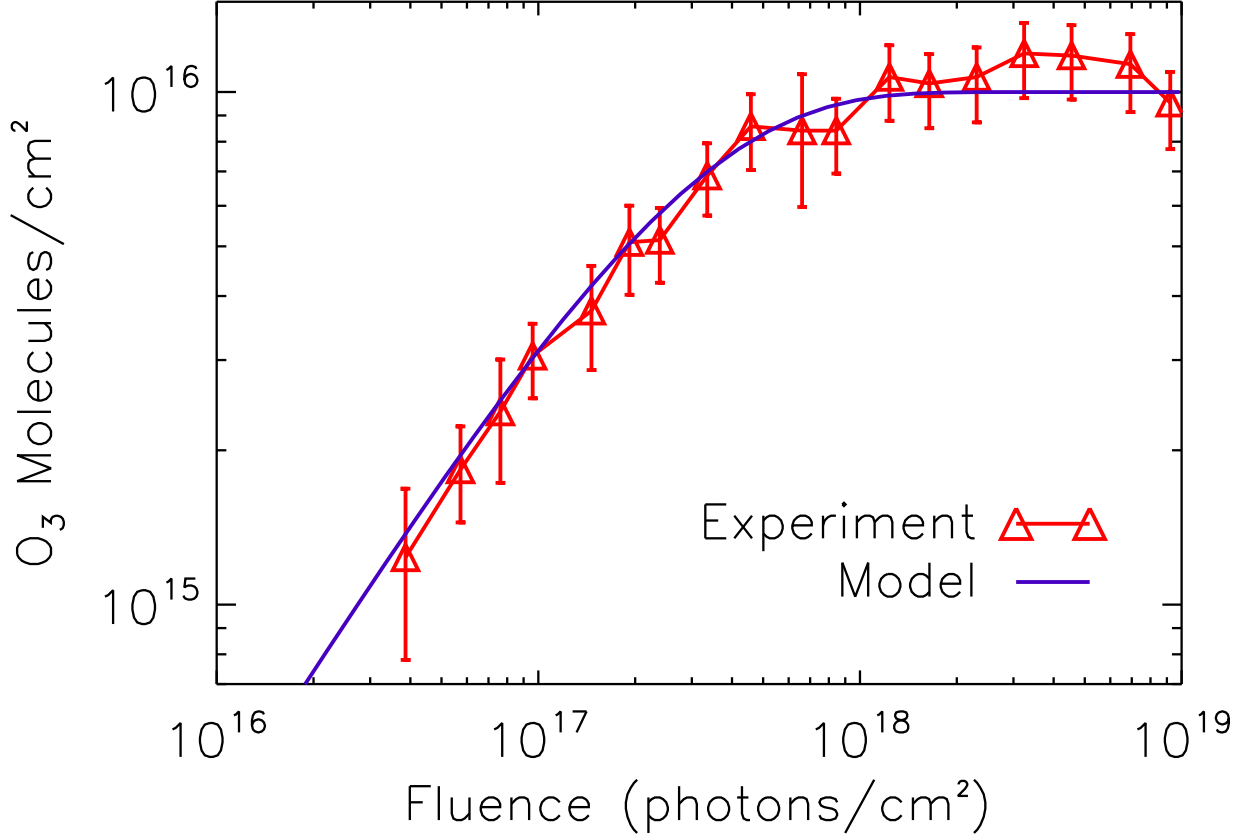


Figure 3. Calculated abundances of O_3 vs. photon fluence in a UV-irradiated pure O_2 ice (shown in blue), with corresponding experimental data from [Raut et al. \(2011\)](#) (shown in red). δ -values used in this simulation: $\text{O}_2 + h\nu \rightarrow \text{O} + \text{O}$ [$\delta = 2.0$]; $\text{O}_2 + h\nu \rightarrow \text{O}_2^*$ [$\delta = 0.3$]; $\text{O}_3 + h\nu \rightarrow \text{O}_2 + \text{O}$ [$\delta = 0.3$]; $\text{O}_3 + h\nu \rightarrow \text{O}_3^*$ [$\delta = 0.3$]; $\delta = 1.0$ for all other processes.

Since Raut et al. included error values in their experimental data, a python script was used to find optimized δ fitting factors which would maximize agreement with experimental data. “Maximal agreement” was considered as any model output which fell entirely within upper and lower error. Our script iteratively ran the model over a range of δ values for each process and indicated which combinations of δ values resulted in model outputs which agreed with the data within experimental error; these are given in Table 2. Because there is a range of outputs which may fall within upper and lower error, there are correspondingly ranges of δ values for the most influential processes. Simulation results using δ values within this range are shown in Fig. 3.

As displayed in Table 2, the δ fitting factors, although close to unity, vary somewhat between the simulations of the two experiments. When branching ratios are adjusted from the initial assumption of equal likelihood to more realistic values, all fitting factors could be set to 1.0 for the Gerakines experiment. As noted in §2.1, the δ fitting factors should be interpreted as effectively accounting for other factors (e.g., absolute uncertainties in the experimental data) not explicitly considered in the code. The fact that all δ values given in Table 2 are close to unity indicates that the overall

contribution of such unknown effects is likely small and are unlikely to be significant sources of uncertainty in astrophysical simulations.

In our models, it was found that our O_3 abundances were most sensitive to variations in the δ value for the $\text{O}_2 \rightsquigarrow 2\text{O}$ process, thereby revealing the importance of the $\text{O} + \text{O}_2$ reaction on the overall abundance of ozone. This finding reveals another useful role for the δ values, namely, that of highlighting key reactions for a given species based on how sensitive the calculated abundance is to variations in the assumed values of δ .

While the current model provides reasonable agreement with the findings of the considered experiments, a number of areas could be addressed in future studies that could increase both agreement with empirical data as well as the underlying physical realism of the simulation. As mentioned, the focus of this work has been on processes occurring in the mantles of thin ice films similar to those that coat interstellar dust grains, all of which were optically thin, and thus, we have not considered the effects of extinction that would be of particular importance in optically thick ices. To investigate this effect in more detail, a multi-layer model which more explicitly treats the vertical structure of ice mantles, such as the macroscopic Monte Carlo code described in [Vasyunin & Herbst \(2013\)](#), would be more appropriate.

Moreover, given the focus of this study on bulk chemistry, we have not considered photodesorption processes occurring in the top several monolayers of the ice. Current values used in models for these kinds of processes are not well constrained, however, our method of directly simulating laboratory experiments using astrochemical codes represents a promising means by which suitable values could be obtained.

Additionally, it is known that the absorption of photons of different energies will change both the efficiency and products of photo-processes ([Fayolle et al. 2013, 2011](#); [Fillion et al. 2014](#)). Absent theoretical/experimental cross-sections for photoprocesses as a function of energy, we have as a first approximation assumed that, for example, the photodissociation of O_2 produces with equal probability 2O^* and 2O , but not $\text{O}^* + \text{O}$.

Finally, because subionization UV photons dominate the MDHL lamp/ArF laser UV, it is likely that the role of low-energy electrons is not significant in this study. To simulate the effects of secondary UV radiation within dark, dense molecular clouds, this model must be modified to include secondary low-energy electron-induced process such as dissociative electron attachment that can occur at electron energies almost as low as 0 eV.

4. CONCLUSIONS

We have simulated the < 10.8 eV UV photodissociation of solid O_2 at 10-22 K by a microwave-discharge hydrogen flow lamp and an ArF excimer laser using a rate-based model. Our methodology incorporates: (a) non-diffusive bulk reactions for radicals and other reactive species and (b) a new theoretical method for simulating photoprocesses which, for the first time, distinguishes between photoexcitation and photoionization. We explicitly account for the production and reactivity of short-lived suprathermal photoproducts. In contrast to previous condensed phase astrochemical model calculations that have used gas-phase photoabsorption cross sections, we have employed solid-phase cross sections in our calculations. This method allows the model to be tested using well-constrained experiments rather than poorly constrained gas-phase abundances in regions of the ISM. The semi-quantitative agreement of the model with experimental O_3 abundances obtained in two different laboratories indicates that the methodology is promising for simulating interstellar ice

photoprocessing. This new computational method, focusing on non-diffusive reactions for radicals and suprathermal species, results in improved agreement with experimental data compared to techniques that rely on bulk thermal radical diffusion, an unlikely mechanism at the exceedingly low temperatures of cold cores. Ultimately it would be fruitful to incorporate these types of rate-based photoprocessing calculations into models that account for atom addition, gas-phase reactions, and cosmic-ray bombardment. Such models, together with observations and laboratory simulations, are necessary for a fundamental understanding of interstellar chemistry which is the likely source of prebiotic molecules in the universe.

ACKNOWLEDGMENTS

C.N.S. thanks the Alexander von Humboldt Stiftung/Foundation for their generous support. E.M. gratefully acknowledges funding from the Arnold and Mabel Beckman Foundation. The Massachusetts Space Grant Consortium supported the work of MF. CRA's work was supported by grants from the National Science Foundation (NSF grant number CHE-1955215), Wellesley College (Faculty Awards and Brachman Hoffman small grants).

Software: MONACO - Vasyunin et al. (2017), MATLAB

REFERENCES

- Aikawa, Y., Wakelam, V., Garrod, R. T., & Herbst, E. 2008, *The Astrophysical Journal*, 674, 984, doi: [10.1086/524096](https://doi.org/10.1086/524096)
- Arumainayagam, C. R., Lee, H.-L., Nelson, R. B., Haines, D. R., & Gunawardane, R. P. 2010, *Surface Science Reports*, 65, 1, doi: [10.1016/j.surfrep.2009.09.001](https://doi.org/10.1016/j.surfrep.2009.09.001)
- Arumainayagam, C. R., Garrod, R. T., Boyer, M. C., et al. 2019, *Chemical Society Reviews*, 48, 2293, doi: [10.1039/C7CS00443E](https://doi.org/10.1039/C7CS00443E)
- Bacmann, A., Taquet, V., Faure, A., Kahane, C., & Ceccarelli, C. 2012, *Astronomy & Astrophysics*, 541, L12, doi: [10.1051/0004-6361/201219207](https://doi.org/10.1051/0004-6361/201219207)
- Balucani, N., Ceccarelli, C., & Taquet, V. 2015, *Monthly Notices of the Royal Astronomical Society*, 449, L16, doi: [10.1093/mnrasl/slv009](https://doi.org/10.1093/mnrasl/slv009)
- Bulak, M., Paardekooper, D. M., Fedoseev, G., & Linnartz, H. 2020, *Astronomy & Astrophysics*, 636, A32, doi: [10.1051/0004-6361/201937298](https://doi.org/10.1051/0004-6361/201937298)
- Chuang, K.-J., Fedoseev, G., Qasim, D., et al. 2017, *Monthly Notices of the Royal Astronomical Society*, 467, 2552, doi: [10.1093/mnras/stx222](https://doi.org/10.1093/mnras/stx222)
- Codella, C., Ceccarelli, C., Bianchi, E., et al. 2020, *Astronomy & Astrophysics*, 635, A17, doi: [10.1051/0004-6361/201936725](https://doi.org/10.1051/0004-6361/201936725)
- Cruz-Diaz, G. A., Caro, G. M. M., Chen, Y.-J., & Yih, T.-S. 2014, *Astronomy & Astrophysics*, 562, A120, doi: [10.1051/0004-6361/201322621](https://doi.org/10.1051/0004-6361/201322621)
- Cuppen, H. M., Walsh, C., Lamberts, T., et al. 2017, *Space Science Reviews*, doi: [10.1007/s11214-016-0319-3](https://doi.org/10.1007/s11214-016-0319-3)
- Elkomoss, S. G., & Magee, J. L. 1962, *The Journal of Chemical Physics*, 36, 256, doi: [10.1063/1.1732308](https://doi.org/10.1063/1.1732308)
- Fayolle, E. C., Bertin, M., Romanzin, C., et al. 2011, *The Astrophysical Journal*, 739, L36, doi: [10.1088/2041-8205/739/2/L36](https://doi.org/10.1088/2041-8205/739/2/L36)
- . 2013, *Astronomy & Astrophysics*, 556, A122, doi: [10.1051/0004-6361/201321533](https://doi.org/10.1051/0004-6361/201321533)
- Fillion, J.-H., Fayolle, E. C., Michaut, X., et al. 2014, *Faraday Discussions*, 168, 533, doi: [10.1039/C3FD00129F](https://doi.org/10.1039/C3FD00129F)
- Garrod, R. T. 2008, *Astronomy and Astrophysics*, 491, 239, doi: [10.1051/0004-6361:200810518](https://doi.org/10.1051/0004-6361:200810518)
- . 2013, *The Astrophysical Journal*, 765, 60, doi: [10.1088/0004-637X/765/1/60](https://doi.org/10.1088/0004-637X/765/1/60)

- Gerakines, P. A., Schutte, W. A., & Ehrenfreund, P. 1996, *Astronomy and Astrophysics*, 312, 289, <http://adsabs.harvard.edu/abs/1996A%26A...312..289G>
- Grassi, T., Bovino, S., Schleicher, D. R. G., et al. 2014, *Monthly Notices of the Royal Astronomical Society*, 439, 2386, doi: [10.1093/mnras/stu114](https://doi.org/10.1093/mnras/stu114)
- Gredel, R., Lepp, S., Dalgarno, A., & Herbst, E. 1989, *The Astrophysical Journal*, 347, 289, doi: [10.1086/168117](https://doi.org/10.1086/168117)
- Heays, A. N., Bosman, A. D., & Dishoeck, E. F. v. 2017, *Astronomy & Astrophysics*, 602, A105, doi: [10.1051/0004-6361/201628742](https://doi.org/10.1051/0004-6361/201628742)
- Herbst, E., & Millar, T. J. 2008, in *Low Temperatures and Cold Molecules*, ed. I. W. M. Smith (London: Imperial College Press)
- Herbst, E., & van Dishoeck, E. F. 2009, *Annual Review of Astronomy and Astrophysics*, 47, 427, doi: [10.1146/annurev-astro-082708-101654](https://doi.org/10.1146/annurev-astro-082708-101654)
- James, T. A., Viti, S., Holdship, J., & Jiménez-Serra, I. 2020, *Astronomy & Astrophysics*, 634, A17, doi: [10.1051/0004-6361/201936536](https://doi.org/10.1051/0004-6361/201936536)
- Jenniskens, P., Baratta, G. A., Kouchi, A., et al. 1993, *Astronomy and Astrophysics*, 273, 583, <https://ui.adsabs.harvard.edu/abs/1993A%26A...273..583J/abstract>
- Jiménez-Serra, I., Vasyunin, A. I., Caselli, P., et al. 2016, *The Astrophysical Journal*, 830, L6, doi: [10.3847/2041-8205/830/1/L6](https://doi.org/10.3847/2041-8205/830/1/L6)
- Jin, M., & Garrod, R. T. 2020, arXiv:2006.11127 [astro-ph]. <http://arxiv.org/abs/2006.11127>
- Kahn, A. 2015, *Materials Horizons*, 3, 7, doi: [10.1039/C5MH00160A](https://doi.org/10.1039/C5MH00160A)
- Ligterink, N. F. W., Paardekooper, D. M., Chuang, K. J., et al. 2015, *Astronomy & Astrophysics*, 584, A56, doi: [10.1051/0004-6361/201526930](https://doi.org/10.1051/0004-6361/201526930)
- Lu, H.-C., Chen, H.-K., Cheng, B.-M., & Ogilvie, J. 2008, *Spectrochimica Acta Part A: Molecular and Biomolecular Spectroscopy*, 71, 1485, doi: [10.1016/j.saa.2008.05.007](https://doi.org/10.1016/j.saa.2008.05.007)
- Marchione, D., Thrower, J. D., & McCoustra, M. R. S. 2016, *Physical Chemistry Chemical Physics*, 18, 4026, doi: [10.1039/C5CP06537B](https://doi.org/10.1039/C5CP06537B)
- Mason, N. J., Dawes, A., Holtom, P. D., et al. 2006, *Faraday Discussions*, 133, 311, doi: [10.1039/B518088K](https://doi.org/10.1039/B518088K)
- McElroy, D., Walsh, C., Markwick, A. J., et al. 2013, *Astronomy & Astrophysics*, 550, A36, doi: [10.1051/0004-6361/201220465](https://doi.org/10.1051/0004-6361/201220465)
- Öberg, K. I. 2016, *Chemical Reviews*, 116, 9631, doi: [10.1021/acs.chemrev.5b00694](https://doi.org/10.1021/acs.chemrev.5b00694)
- Öberg, K. I., Boogert, A. C. A., Pontoppidan, K. M., et al. 2011, *The Astrophysical Journal*, 740, 109, doi: [10.1088/0004-637X/740/2/109](https://doi.org/10.1088/0004-637X/740/2/109)
- Öberg, K. I., Fayolle, E. C., Cuppen, H. M., van Dishoeck, E. F., & Linnartz, H. 2009, *Astronomy and Astrophysics*, 505, 183, doi: [10.1051/0004-6361/200912464](https://doi.org/10.1051/0004-6361/200912464)
- Öberg, K. I., Qi, C., Fogel, J. K. J., et al. 2010, *The Astrophysical Journal*, 720, 480, doi: [10.1088/0004-637X/720/1/480](https://doi.org/10.1088/0004-637X/720/1/480)
- Pontoppidan, K. M. 2006, *Astronomy & Astrophysics*, 453, L47, doi: [10.1051/0004-6361:20065569](https://doi.org/10.1051/0004-6361:20065569)
- Prasad, S. S., & Tarafdar, S. P. 1983, *The Astrophysical Journal*, 267, 603, doi: [10.1086/160896](https://doi.org/10.1086/160896)
- Raut, U., Loeffler, M. J., Famá, M., & Baragiola, R. A. 2011, *The Journal of Chemical Physics*, 134, 194501, doi: [10.1063/1.3589201](https://doi.org/10.1063/1.3589201)
- Scibelli, S., & Shirley, Y. 2020, *The Astrophysical Journal*, 891, 73, doi: [10.3847/1538-4357/ab7375](https://doi.org/10.3847/1538-4357/ab7375)
- Shen, C. J., Greenberg, J. M., Schutte, W. A., & van Dishoeck, E. F. 2004, *Astronomy and Astrophysics*, 415, 203, doi: [10.1051/0004-6361:20031669](https://doi.org/10.1051/0004-6361:20031669)
- Shingledecker, C. N., Gal, R. L., & Herbst, E. 2017, *Physical Chemistry Chemical Physics*, 19, 11043, doi: [10.1039/C7CP01472D](https://doi.org/10.1039/C7CP01472D)
- Shingledecker, C. N., & Herbst, E. 2018, *Physical Chemistry Chemical Physics*, 20, 5359, doi: [10.1039/C7CP05901A](https://doi.org/10.1039/C7CP05901A)
- Shingledecker, C. N., Álvarez Barcia, S., Korn, V. H., & Kästner, J. 2019a, *The Astrophysical Journal*, 878, 80, doi: [10.3847/1538-4357/ab1d4a](https://doi.org/10.3847/1538-4357/ab1d4a)
- Shingledecker, C. N., Tennis, J., Gal, R. L., & Herbst, E. 2018, *The Astrophysical Journal*, 861, 20, doi: [10.3847/1538-4357/aac5ee](https://doi.org/10.3847/1538-4357/aac5ee)
- Shingledecker, C. N., Vasyunin, A., Herbst, E., & Caselli, P. 2019b, *The Astrophysical Journal*, 876, 140, doi: [10.3847/1538-4357/ab16d5](https://doi.org/10.3847/1538-4357/ab16d5)
- Sivaraman, B., Nair, B. G., Raja Sekhar, B. N., et al. 2014, *Chemical Physics Letters*, 603, 33, doi: [10.1016/j.cplett.2014.04.021](https://doi.org/10.1016/j.cplett.2014.04.021)

- Taquet, V., Wirström, E. S., Charnley, S. B., et al. 2017, *Astronomy & Astrophysics*, 607, A20, doi: [10.1051/0004-6361/201630023](https://doi.org/10.1051/0004-6361/201630023)
- Thrower, J. D., Collings, M. P., Ruten, F. J. M., & McCoustra, M. R. S. 2011, *Chemical Physics Letters*, 505, 106, doi: [10.1016/j.cplett.2011.02.029](https://doi.org/10.1016/j.cplett.2011.02.029)
- Tielens, A. G. G. M., Tokunaga, A. T., Geballe, T. R., & Baas, F. 1991, *The Astrophysical Journal*, 381, 181, doi: [10.1086/170640](https://doi.org/10.1086/170640)
- Vastel, C., Ceccarelli, C., Lefloch, B., & Bachiller, R. 2014, *The Astrophysical Journal*, 795, L2, doi: [10.1088/2041-8205/795/1/L2](https://doi.org/10.1088/2041-8205/795/1/L2)
- Vasyunin, A. I., Caselli, P., Dulieu, F., & Jiménez-Serra, I. 2017, *The Astrophysical Journal*, 842, 33, doi: [10.3847/1538-4357/aa72ec](https://doi.org/10.3847/1538-4357/aa72ec)
- Vasyunin, A. I., & Herbst, E. 2013, *The Astrophysical Journal*, 762, 86, doi: [10.1088/0004-637X/762/2/86](https://doi.org/10.1088/0004-637X/762/2/86)
- Watanabe, N., & Kouchi, A. 2002, *The Astrophysical Journal Letters*, 571, L173, doi: [10.1086/341412](https://doi.org/10.1086/341412)
- Yu, K. Y., McMenamin, J. C., & Spicer, W. E. 1975, *Surface Science*, 50, 149, doi: [10.1016/0039-6028\(75\)90179-X](https://doi.org/10.1016/0039-6028(75)90179-X)