



Erratum: “The TREX Survey: Kinematical Complexity Throughout M33’s Stellar Disk and Evidence for a Stellar Halo” (2022, ApJ, 924, 116)

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The mean color uncertainties of the spectroscopic targets as a function of magnitude for the first three of the four panels in Figure 2 of the published article were incorrectly drawn due to an error in the plotting routine. The precision of the stellar colors as a function of magnitude, measured from the photometric data sets, is significantly better than shown in the original figure panels. Figure 2 of this erratum shows the correct color uncertainties as a function of magnitude. This error was isolated to the plotting of the figure panels, and does not impact any of the analysis.



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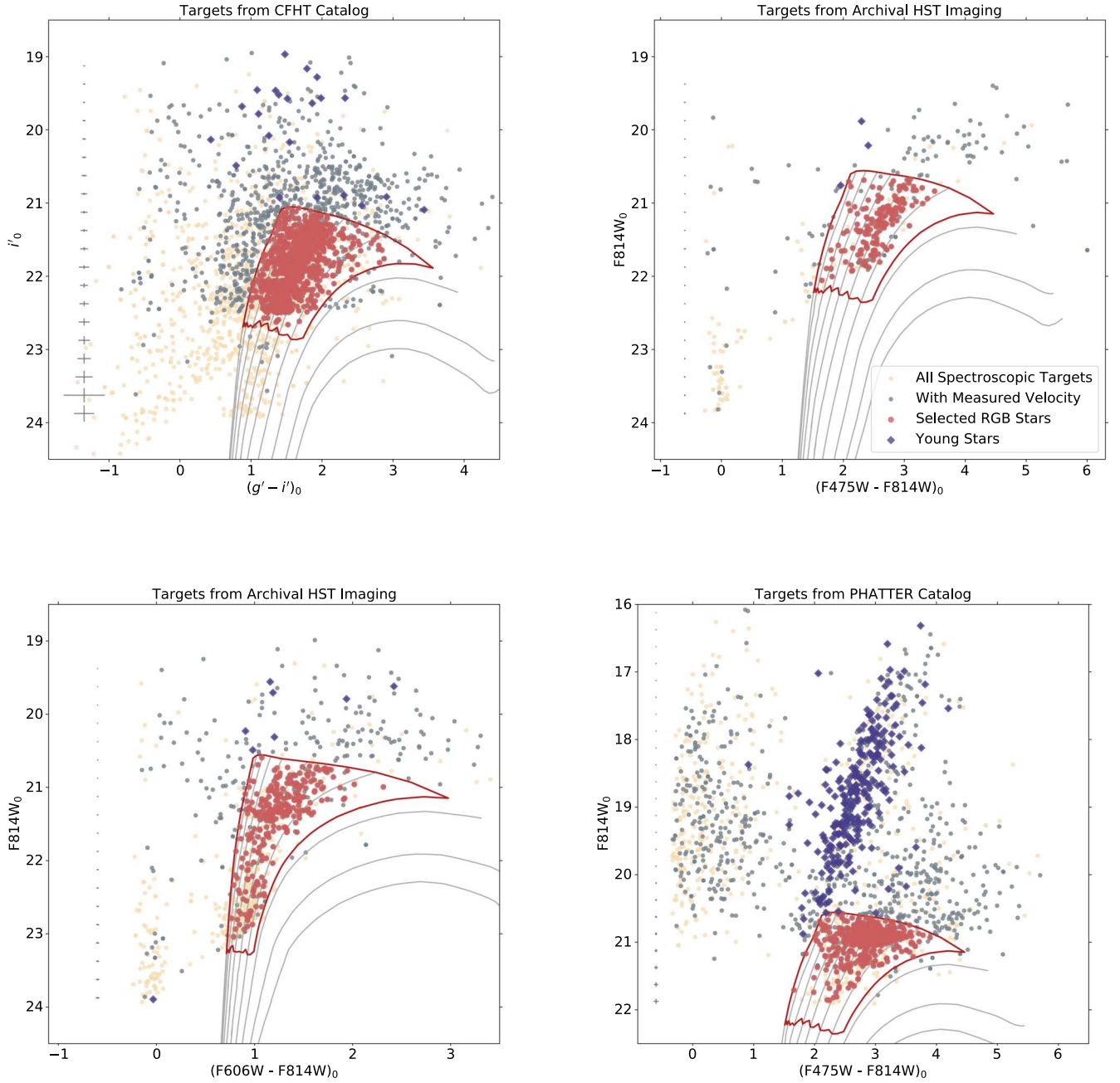


Figure 2. Color–magnitude diagrams for all spectroscopic targets (beige circles), as well as those with measured velocities (gray circles), and those selected for the RGB sample (larger red circles) and young red star comparison sample (blue diamonds; Section 2). Average photometric uncertainties as a function of magnitude are shown on the left-hand side of each panel. A consistent RGB sample selection region was identified across the various filter combinations using 10 Gyr, $[\alpha/\text{Fe}] = 0$ Padova isochrones (Marigo et al. 2017); the RGB selection box as well as select isochrone tracks are shown in each panel. The metal-rich extent of the RGB selection box shown here was selected to fully encompass the final RGB sample, and was used in estimating the MW contamination within the RGB sample region (Appendix A).

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References

Marigo, P., Girardi, L., Bressan, A., et al. 2017, *ApJ*, **835**, 77