

Shock cooling of a red-supergiant supernova at redshift 3 in lensed images

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The core-collapse supernova of a massive star rapidly brightens when a shock, produced following the collapse of its core, reaches the stellar surface. As the shock-heated star subsequently expands and cools, its early-time light curve should have a simple dependence on the size of the progenitor¹ and therefore final evolutionary state. Measurements of the radius of the progenitor from early light curves exist for only a small sample of nearby supernovae^{2–14}, and almost all lack constraining ultraviolet observations within a day of explosion. The several-day time delays and magnifying ability of galaxy-scale gravitational lenses, however, should provide a powerful tool for measuring the early light curves of distant supernovae, and thereby studying massive stellar populations at high redshift. Here we analyse individual rest-frame exposures in the ultraviolet to the optical taken with the Hubble Space Telescope, which simultaneously capture, in three separate gravitationally lensed images, the early phases of a supernova at redshift $z \approx 3$ beginning within 5.8 ± 3.1 hours of explosion. The supernova, seen at a lookback time of approximately 11.5 billion years, is strongly lensed by an early-type galaxy in the Abell 370 cluster. We constrain the pre-explosion radius to be 533^{+154}_{-119} solar radii, consistent with a red supergiant. Highly confined and massive circumstellar material at the same radius can also reproduce the light curve, but because no similar low-redshift examples are known, this is unlikely.

As part of a search for gravitationally lensed, transient events in archival Hubble Space Telescope (HST) imaging of the Hubble Frontier Fields (HFF)¹⁵, we found three images of a rapidly evolving, strongly lensed supernova (SN) in imaging of the Abell 370 galaxy cluster ($z = 0.375$)¹⁶ field obtained in 2010 December. Figure 1 shows these sources, which we label S1, S2 and S3 (as shown in Fig. 1d), and images of the underlying strongly lensed host galaxy, which is approximately 1' away from the centre of the galaxy cluster.

As shown in Extended Data Fig. 2g,h, we measure the photometric redshift of the images of the SN's host galaxy, and find a joint probability distribution of $z = 2.93^{+0.06}_{-0.05}$ using BPZ¹⁷ and $z = 2.94^{+0.06}_{-0.07}$ using EAZY¹⁸. The fluxes of the images are listed in Extended Data Table 1a.

We construct a lens model of the early-type galaxy and galaxy cluster using GLAFIC^{19–22}, which predicts that the early-type galaxy creates a total of four magnified images of the SN host galaxy in an Einstein cross configuration, as shown in Fig. 1e. One image, however, is insufficiently magnified to detect given the underlying galaxy. As listed in Extended Data Table 1b, the model predicts that image S1 has a time delay of 30.6 ± 5.6 days relative to S3 and of 9.6 ± 2.3 days relative to S2, and S2 has a time delay of 21.0 ± 3.4 days relative to S3. Taking into account the factor of four redshift time dilation of sources at $z = 3$, the SN as seen in image S1 is 7.7 ± 1.4 days younger than it appears in image S3, and image

S2 shows the SN as it was 5.3 ± 0.9 days younger than it appears in image S3. According to the lens model, images S1, S2, S3 and S4 have respective magnifications of 4.2 ± 1.1 , 7.8 ± 1.3 , 5.8 ± 0.8 and 1.2 ± 0.5 .

We next fit separate blackbody models to the spectral energy distributions of the three delayed images of the SN. Using the predicted magnification values from the lens model, we find effective blackbody temperatures of $5.5^{+13.8}_{-3.0} \times 10^4$ K, $2.0^{+0.2}_{-0.2} \times 10^4$ K and $1.1^{+0.1}_{-0.1} \times 10^4$ K for SN images S1, S2 and S3, respectively. We show, in Fig. 1e, a pseudocolour image of the SN constructed by assigning the F160W, F110W and F814W difference images to the red, green and blue (RGB) channels. Figure 1e shows the rapid change in the colour of the supernova as it cools from approximately 100,000 K to 10,000 K over about 8 days in the rest frame at $z = 3$. As described in Methods, we find that the optical depth for microlensing is very low, precluding chromatic microlensing.

Following shock breakout, the photosphere rapidly cools as the thermal energy of the material is converted to radiation and the kinetic energy of the expanding envelope. Emission escapes when the optical depth of the shock drops below c/v , where c and v are, respectively, the speed of light and the shock velocity^{23,24}. Therefore, a shell of circumstellar material (CSM) can dominate the emission associated with shock cooling only if its optical depth before explosion exceeds c/v (refs. ^{25,26}).

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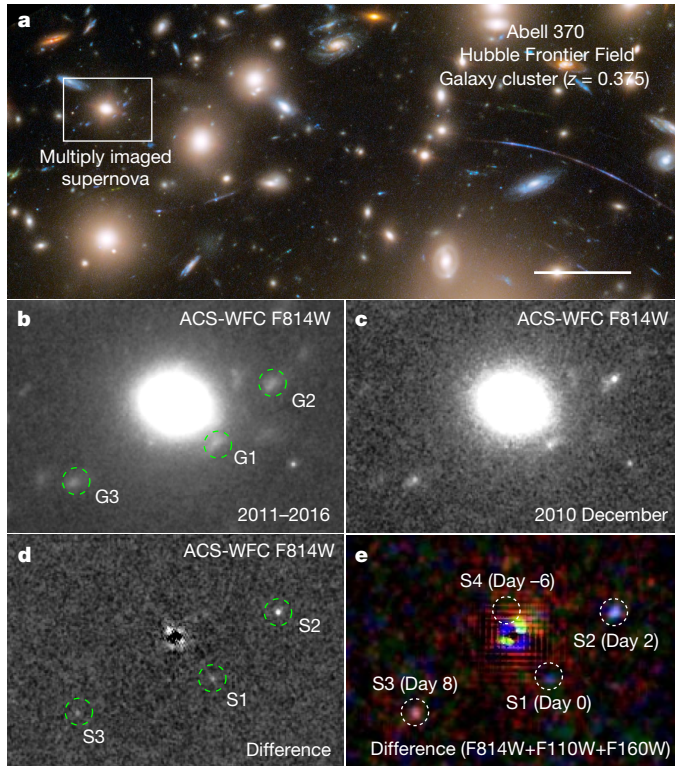


Fig. 1 | A multiply imaged SN in the archival HST imaging. **a**, A multiply imaged SN in a gravitationally lensed galaxy at $z \approx 3$ in the Abell 370 galaxy-cluster field. Scale bar, $10''$. **b**, An example deep template ACS-WFC F814W image of the field, where three images of the lensed galaxy or dwarf galaxy are marked by green dashed circles labelled as G1, G2 and G3. **c**, Image of the identified event in 2010 December from the same bandpass filter. **d**, The difference image; three SN images are marked as S1, S2 and S3 by green dashed circles. We have named the images according to the order in which the observed light was emitted, as predicted by our gravitational-lens model. **e**, Mixed difference imaging from the ACS-WFC F814W, WFC3-IR F110W and WFC3-IR F160W bandpasses. The de-noised difference images from F160W, F110W and F814W are assigned to the red, green and blue channels, respectively, providing a mixed image with pseudocolour. White dashed circles mark the predicted positions from our best-fit lens model. Background image in **a** by NASA, ESA/Hubble, HST Frontier Fields.

To calculate the predicted light curves of core-collapse supernovae (SNe) of supergiant stars where the CSM does not dominate the shock-cooling emission, we adopt an analytical model²⁷ with a planar-phase correction¹⁴. We label these models blue supergiant (BSG) and red supergiant (RSG), respectively. Modelling of the light curves of nearby type-IIP SNe shows that a large fraction of progenitors have CSM at radii greater than $1,000R_{\odot}$ (R_{\odot} , solar radii; refs. ^{28,29}), although almost all low-redshift datasets lack ultraviolet (UV) photometry within a day of explosion, which would provide additional constraints on the radius of optically thick material as we describe in the Methods. We therefore consider a pair of analytical models of a SN explosion surrounded by CSM which we refer to as CSM homologous¹⁰ and CSM planar³⁰.

Although our ten fitting parameters exceed the nine measured data points, the four parameters associated with the lens model are constrained by informative priors. We are able to use the Bayes factor to distinguish among potential models and marginalize fully over all parameters. We compute the Bayes factor for each SN model relative to the null hypothesis that the lensed source is a blackbody whose temperature and luminosity do not evolve with time. In Extended Data Table 1c, we list the natural logarithm of the ratio of the Bayes factors, and the ratios provide strong evidence against the null hypothesis. Given two hypothetical models A and B, $\log(B_{AB}) > 2$ is considered to

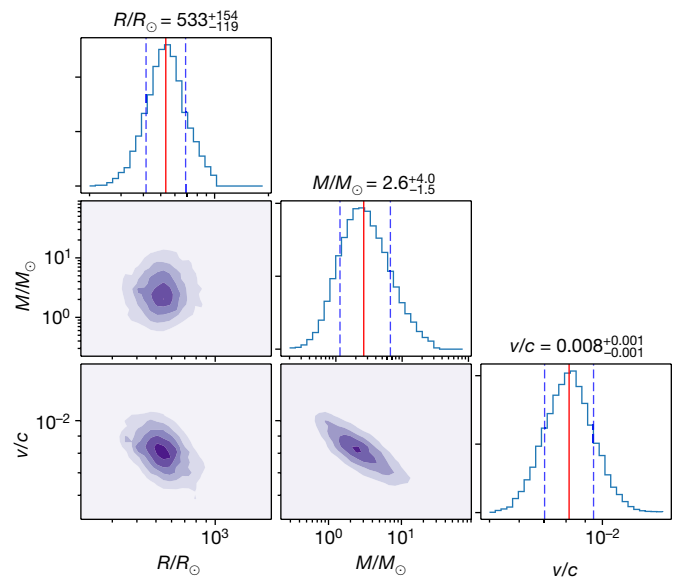


Fig. 2 | Posterior distributions of the progenitor radius R , envelope mass M and shock velocity v from the RSG model for $R_{\nu} = 3.1$. The red solid and the blue dashed lines overlaid on each histogram denote the median and the 68% confidence interval of each distribution, respectively. From the posterior distributions, we constrained the radius of the progenitor RSG to $533^{+154}_{-119}R_{\odot}$ and the envelope mass to $2.6^{+4.0}_{-1.5}M_{\odot}$. Distributions of all free parameters are shown in Extended Data Fig. 4.

be very strong evidence in favour of model A, whereas $\log(B_{AB}) > 0.5$ indicates little or no statistical evidence³¹. When we assume that the CSM has a mass within a range of $(0.001-0.1)M_{\odot}$ (M_{\odot} , solar mass) that is consistent with those used by modelling efforts^{10,30}, the RSG model is significantly favoured in comparison to both the BSG and CSM models (with $\Delta\log(B_{AB}) > 2$). In Fig. 2, we plot our constraints on the progenitor radius, envelope mass and shock velocity for the RSG model, and we plot the posterior probability distributions for all model parameters in Extended Data Fig. 4. The best-fit parameter values and their confidence intervals for all models are listed in Extended Data Table 2.

If we allow the CSM mass to become as large as $1M_{\odot}$, the Bayes factor for the CSM-homologous model becomes better than (but not significantly, with $\Delta\log(B_{AB}) = 0.7$) that of the RSG model, and the difference between the RSG model and the Bayes factor of the CSM-planar CSM model decreases. However, the CSM-homologous model yields a CSM radius of $481^{+157}_{-118}R_{\odot}$, and prefers a CSM mass of $1M_{\odot}$. The CSM radius is smaller than the greater than $1,000R_{\odot}$ inferred for a sample of SN IIP explosions of RSGs²⁹, and the dense, compact CSM-homologous model approximates the outer atmosphere of the RSG model.

The explosion of an RSG, the most common SN type in the local universe³², is the most probable explanation for the observations. As shown in Extended Data Fig. 7, the early light curve and the evolution of the temperature of this event are similar to those of the SN IIP explosions of RSG progenitor stars. Extended Data Fig. 7 shows that the observations are not well matched by those of well studied SNe IIP whose progenitors were inferred, from their light curves, to have CSM. Likewise, early available UV observations of SNe Ia, Ib/c, IIb or fast blue optical transients are not well matched to our photometry.

In Fig. 3, we show two illustrative examples of the reconstructed light curve of the SN. These plot the light curve when the images' magnification and times delays are set to their best-fitting values, and after they are shifted by 1σ (1 standard deviation) to smaller values; see Extended Data Table 1b for lens values. For each light curve, we show the best-fitting RSG model light curve, and the corresponding constraints on its pre-explosion radius, which are $615^{+86}_{-76}R_{\odot}$ and $509^{+74}_{-67}R_{\odot}$, respectively. These demonstrate the dependence of the inference about the size of

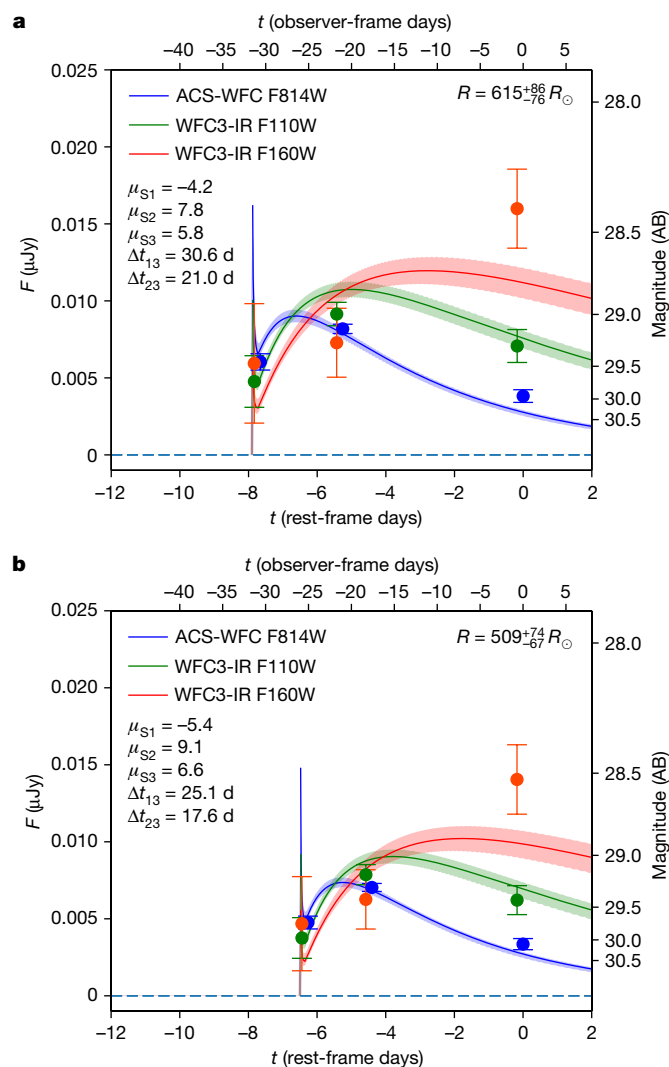


Fig. 3 | Two examples of the reconstructed light curve of the SN with fixed time delays and magnifications. **a**, Reconstructed light curves from the RSG model for $R_V = 3.1$ where the magnifications and time delays are set to the best-fit values from the lens model as listed in Extended Data Table 1b, where F and t are, respectively, the reconstructed flux density and time. Data points and error bars correspond to the measured values and $\pm 1\sigma$ uncertainties of the flux densities, respectively. Solid lines show the best-fit light curves and shaded regions are the 68% confidence intervals of the flux density. The radius of the progenitor is constrained to $615^{+86}_{-76} R_\odot$. **b**, Reconstructed light curves from the RSG model for $R_V = 3.1$, while we chose the magnifications to be 1σ larger, and the time delays to be 1σ smaller than their best-fit values. This infers a smaller progenitor radius as $509^{+74}_{-67} R_\odot$ compared with the result using best-fit magnifications and time delays.

the progenitor on the lens-model uncertainties. The constraints in Fig. 3 are consistent with but shifted by a small amount from our primary estimate, because the SN model and photometry together favour larger magnifications (by less than or approximately equal to 1.2σ) and smaller time delays (by less than or approximately equal to 0.5σ). In the rest frame of the SN at $z \approx 3$, the central wavelength of the ACS-WFC814W (broad I) filter corresponds to 2,000 Å, WFC3-IR F110W (wide y) to 2,900 Å, and WFC3-IR (H) to 3,900 Å. Figure 3 shows that the luminosity measured in the F814W filter grew and then faded, whereas that in the reddest (F160W) filter only increased.

Discoveries of strongly gravitationally lensed SNe^{33–35} offer an avenue for probing the populations of massive stars present in the high-redshift universe^{36,37}. Two multiply imaged SNe (SN Refsdal³³ and this work) have

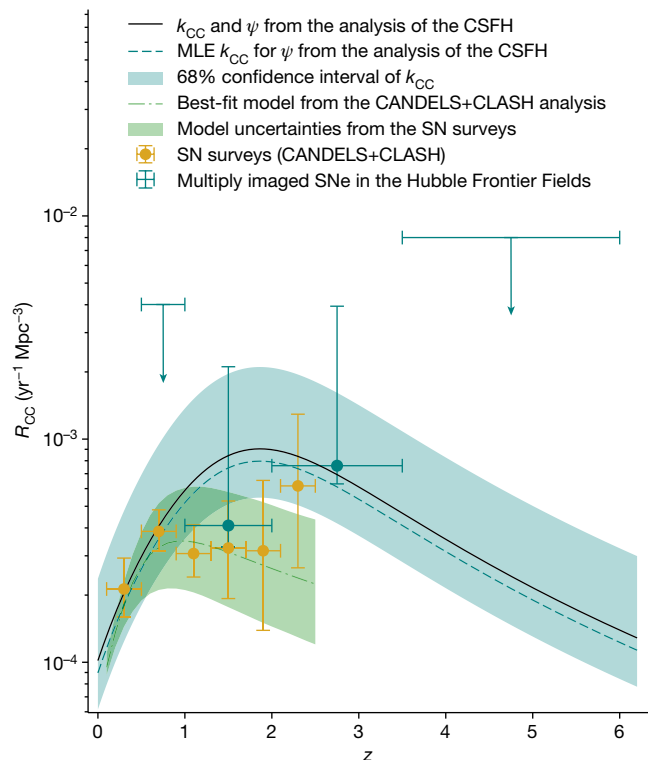


Fig. 4 | Volumetric CCSN rate (R_{cc}) as a function of redshift. Black solid curve is from the star-formation rate density (ψ)⁴⁶ and $k_{cc} = 0.0068 M_\odot^{-1}$ for a Salpeter IMF⁴⁷. Dark cyan data points are the MLE R_{cc} values estimated from the detection of multiply imaged SNe in the six Hubble Frontier Fields in the last decade. Horizontal error bars show the redshift bins, and vertical error bars give the 68% confidence intervals. Cyan shaded region is the range of the estimated R_{cc} based on ψ from the analysis of the CSFH⁴⁶ and the 68% confidence interval of k_{cc} from the detection of multiply imaged SNe. Golden data points, dash-dotted green line and green shaded region are R_{cc} constraints from SN surveys, the best-fit star-formation model (ψ and k_{cc}) and the error region of the model from the CANDELS+CLASH analysis⁴⁰, respectively.

been discovered in the six HFFs, and we use these to estimate the volumetric core-collapse SN (CCSN) rate to $z = 3$. Using our simulations of the CCSN light curves, we constrain the volumetric CCSN rate in four redshift bins from 0.5 to 6. Given these two magnified SNe, we obtain CCSN rates of $4.1^{+17.0}_{-0.8} \times 10^{-4} \text{ yr}^{-1} \text{ Mpc}^{-3}$ at $z = 1\text{--}2$ and $7.6^{+31.8}_{-1.3} \times 10^{-4} \text{ yr}^{-1} \text{ Mpc}^{-3}$ at $z = 2\text{--}3.5$, as shown in Fig. 4.

The principal existing constraints on the CCSN rates to $z \approx 2.5$ are from an analysis of the Cosmic Assembly Near-infrared Deep Extragalactic Legacy Survey (CANDELS)³⁸ and the Cluster Lensing and Supernova Survey with Hubble (CLASH)³⁹ (the CANDELS+CLASH analysis)⁴⁰. For a given survey, the expected number of CCSNe depends on both the intrinsic volumetric rate of CCSNe and an assumed prior on extinction due to dust. For a given number of detected CCSNe, greater average extinction will lead a greater inferred volumetric CCSN rate, because an increased fraction SNe will be too faint to detect. To draw comparison with the CANDELS+CLASH constraints, we must use both the same inferred volumetric CCSN rates and the extinction prior to predict the expected number of strongly lensed SNe in archival observations of the HFF cluster fields. We find a probability of only $P = 0.026$ of discovering at least two strongly lensed CCSN.

The exponential scale factor of 0.187 mag used by the CANDELS+CLASH analysis⁴⁰ to construct a dust-extinction prior corresponds to an extreme average $A_V = 5.35$ mag, which differs from the average of the sample ($A_V = 0.405$) used to construct it^{41,42}. For our primary analysis, we instead use a prior with mean $A_V = 1$ mag informed by measurements of low-redshift SNe^{43–45}, as shown in Extended Data Fig. 8a.

The rate of CCSNe per unit volume ($R_{\text{CC}}(z)$) can be expressed as a function of the volumetric star-formation density $\psi(z)$: $R_{\text{CC}}(z) = \psi(z)k_{\text{CC}}$. Given the redshift-dependent $\psi(z)$ inferred from observations of galaxies⁴⁶, and the two multiply imaged SN in the HFF fields, we obtain $k_{\text{CC}} = 0.0087^{+0.0071}_{-0.0046} M_{\odot}^{-1}$ and a maximum-likelihood estimate of $k_{\text{CC}} = 0.0060 M_{\odot}^{-1}$. In contrast to previous analysis of high-redshift CCSNe rates⁴⁰, these values are consistent with the theoretical expectation of $k_{\text{CC}} = 0.0068 M_{\odot}^{-1}$ for a Salpeter initial mass function (IMF)⁴⁷ and the approximate assumption that stars with initial masses between $8 M_{\odot}$ and $40 M_{\odot}$ successfully explode⁴⁶.

Online content

Any methods, additional references, Nature Research reporting summaries, source data, extended data, supplementary information, acknowledgements, peer review information; details of author contributions and competing interests; and statements of data and code availability are available at <https://doi.org/10.1038/s41586-022-05252-5>.

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Methods

Cosmology

In this paper, we assume a concordance cosmology described by the Λ cold dark matter model (Λ CDM) with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and a Hubble constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

HST imaging and transient searching

All HST images are aligned using the TweakReg tool and resampled to a scale of $0.03'' \text{ pixel}^{-1}$ using the AstroDrizzle package⁴⁸. We prepared our template images using the full set of the HST imaging from the WFC3 and ACS cameras. We grouped and coadded images from each HST visit and then created a difference image by subtracting a deep template image from the coaddition. We next identified all peaks in the resulting difference images with flux exceeding a 3σ signal-to-noise ratio, and classified the peaks by a convolutional neural network machine specifically developed for the transient search of the archival HST imaging. For photometry in the difference imaging, we used PythonPhot⁴⁹, which includes an implementation of point-spread-function (PSF) fitting photometry based on the DAOPHOT algorithm⁵⁰. Our search led us to identify this SN from HST images, as well as the already known SN Refsdal in the MACS J1149.5+2223 galaxy-cluster field³³. The SN images were detected in the HST imaging taken in 2010 December. Prior to the 2010 December HST visit, the field was observed most recently in 2009 September. The field was not observed again until 2014 October, and neither set of images from observations before 2009 and after 2014 show evidence of variability. The relative time delays and magnifications of the three strongly lensed images enable us to reconstruct the light curve of the SN beginning within a day after explosion, which we use to constrain the radius of its progenitor.

Photometry of the SN images

A difference image is constructed from a pair of coadditions of images acquired at two epochs. However, the roll angle of HST generally changes between visits. Consequently, the four diffraction spikes of the PSF will rotate, and a simple subtraction of the imaging contains a negative and positive set of residual spikes. In the case of the lens galaxy in the centre of the multiply imaged SN, as shown in Extended Data Fig. 1d, these residual spikes coincide with the SN images. To remove these spikes, we created a model of the lens galaxy and subtracted it from each set of images, after convolving the model with the PSF measured from the imaging acquired at each epoch.

We used GALFIT⁵¹ to model the lens galaxy and adjacent sources in deep imaging taken by the HFF survey. Extended Data Fig. 1a–c shows the GALFIT models. After convolution with the PSF measured for each epoch, we subtracted the best-fit GALFIT models from both the HFF coaddition, and the previous epoch of imaging. This method yielded much improved contamination of the difference images, as shown in Extended Data Fig. 1e. To generate a deep template for WFC3-IR F110W, we rescaled the WFC3-IR F105W and F125W templates to have zeropoints that match that of the WFC3-IR F110W imaging.

Fourth image of SN host galaxy

We are not able to detect G4 in coadded images, as expected given its small magnification and underlying early-type galaxy. According to the predicted magnification ratio of G2 to G4 and the photometry of G2, the apparent magnitude of G4 is expected to be 28.1 AB in the ACS-WFC F814W band and 27.7 AB in the WFC-IR F160W band, fainter than the 3σ noise levels, where the 3σ noise levels within an aperture for G2 are 28.0 AB in the ACS-WFC F814W band and 27.4 AB in the WFC-IR F160W band. The appearance of the fourth image of the SN (S4) should be the most delayed. As we will describe, we find that the SN exploded only ~ 1 day before we observed it in image S1. The lens model predicts that, in individual visits to the field, we see G4 as it was ~ 23 days before we

see G1. Consequently, the SN should not yet be visible in G4 when the 2010 data were taken, and we do not detect it.

Evaluating whether the colour differences of the SN images could arise from noise

We sought to ascertain whether the observed colour differences among images S1–S3 could be explained by flux-measurement uncertainties, given a lensed transient with an unchanging colour. As our null hypothesis, we consider a blackbody source with a constant temperature and luminosity, such that differences in flux arise only from magnification and uncertainties associated with the flux measurement. The alternative hypothesis is that the SN has distinct effective temperatures and luminosities across the three observed images. We placed a broad prior on the luminosity of the blackbody between the solar value ($3.828 \times 10^{26} \text{ W}$) and that of ASASSN-15lh⁵² ($2.2 \times 10^{38} \text{ W}$). The likelihood ratio test statistic is $-2\ln\Lambda \approx 43$, where Λ is the ratio between likelihoods of the null and alternative hypotheses. This test statistic asymptotically approaches the chi-squared (χ^2) distribution with degree of freedom equal to 4 (because the alternative hypothesis has four more free parameters than the null hypothesis), according to Wilks' theorem. Hence, the likelihood-ratio test gives a P value of $\sim 1.0 \times 10^{-8}$, equivalent to a $\sim 5.6\sigma$ significance for a normal-distributed variable. We constrained the effective temperature (T) and luminosity (L) using a Markov chain Monte Carlo (MCMC) analysis and uniform priors for $10^3 \text{ K} < T < 10^9 \text{ K}$ and $3.828 \times 10^{26} \text{ W} < L < 2.2 \times 10^{38} \text{ W}$. The posterior distributions are shown in Extended Data Fig. 1f–h.

Photometric redshift

We conducted photometry of the multiple galaxy images G1, G2 and G3 from the published coaddition of the HST imaging collected from the HFF survey¹⁵. Before the measurement, we used GALFIT to fit the bright galaxy in the centre of the system and subtracted the best-fit model from the field. We then fit photometry using BPZ¹⁷ and EAZY¹⁸ software to estimate the photometric redshift of the multiply imaged galaxy. As shown in Extended Data Fig. 2g,h, both software packages yield a photometric redshift consistent with $z \approx 3$. We performed a likelihood-ratio test between the hypothesis that the galaxy images have an identical redshift and the alternative hypothesis that the images have distinct redshifts. The test statistic is $-2\ln\Lambda \approx 3.5$, providing only a $\sim 0.9\sigma$ significance in favour of the hypothesis for unrelated galaxy images with independent redshifts. This indicates that the disagreement among the photometric redshift results from the three images is statistically insignificant. The 95% confidence interval of the photometric redshift from the joint probability is (2.85, 3.05) from BPZ and (2.81, 3.07) from EAZY. This result is also consistent with the photometric redshift from a previous study³³, in which the 95% confidence intervals of the G1, G2 and G3 redshifts are (3.122, 3.513), (2.888, 3.351) and (2.736, 3.298), respectively. The next most probable value of the photometric redshift is $z \approx 0.2$, but we can exclude this redshift range because it would lie in the foreground of the Abell 370 cluster at redshift $z = 0.375$. Fitting stellar population synthesis models⁵⁴ to the spectral energy distribution of the host galaxy, we measure a stellar mass of $\sim 2 \times 10^8 M_\odot$ and a star-formation rate of $\sim 0.2 M_\odot \text{ yr}^{-1}$. Extended Data Fig. 2i shows the distribution of CCSN host-galaxy stellar mass and star-formation rate⁵⁵, where the red star marks the host galaxy of the SN we report here.

Spectroscopic observations of the SN host galaxy

We acquired optical spectra of the host galaxy using the Blue Channel Spectrograph⁵⁶ on the MMT telescope (MMT-Blue; 6 October 2019 UT and 23 December 2019 UT), as well as near-infrared spectra using the Large Binocular Telescope Utility Camera in the Infrared⁵⁷ (LBT-LUCI; 30 October 2021 UT) and the Multi-Object Spectrometer For Infra-red Exploration at the Keck Observatory⁵⁸ (Keck-MOSFIRE; 10 January 2022 UT).

For the MMT-Blue observation, we set up a 2''-wide slit with an effective length of 150'' towards the image G2 with three different position angles of the slit, and obtained a total integration time of 4.22 h. The spectra were obtained using the 300 line mm⁻¹ grating, providing a wavelength coverage of 3,800–9,100 Å with a spectral resolution (full-width at half-maximum intensity, FWHM) of 6.47 Å. We centred two 1'' slits at the galaxy images G2 and G3 for our LBT-LUCI observation, and took a total exposure of 1.2 h on our targets. We obtained *H*- and *K*-band spectra of G2 and G3 from the LBT-LUCI observation, covering a wavelength range from 15,500 Å to 25,000 Å. For our Keck-MOSFIRE observation, we set up a 0.7'' slit going through the galaxy images G1 and G3, and obtained spectra from 14,500 Å to 18,000 Å using an exposure of 0.86 h in the *H* band. The MMT-Blue and LBT-LUCI data were reduced using Pypelt software^{59,60} and the Keck-MOSFIRE data were reduced using the MOSFIRE Data Reduction Pipeline⁶¹.

For the MMT-Blue spectroscopy, the spectrum of our target is expected to have two components: one from the bright foreground galaxy and one from the faint G2 image, where most of the continuum is contributed by the brighter component. We fit the continuum with FireFly software⁶² using stellar-population models⁵⁴. We find no $>3\sigma$ peaks detected from the residual spectrum. At $z = 3$, we expect the Lyman- α emission line to have an observer-frame wavelength of 4,863 Å, and we obtain a 3σ upper limit on its flux of 2.2×10^{-17} erg s⁻¹ cm⁻² within the 6.47-Å spectral bin, corresponding to an upper-limit equivalent width of ~ 200 Å in the observer's frame.

The near-infrared spectra were taken from Keck-MOSFIRE observations using two telescope pointings along the slit, and the difference spectroscopic imaging from the two positions was used to measure the spectra of our targets. There is no single peak with $>2\sigma$ significance in the *H* + *K* and *H* bands. A pair of low-significance peaks can be found at $-14,840$ Å from the coadded Keck-MOSFIRE spectrum, which may correspond to the [O II] doublet at 3,726.0 Å and 3,728.8 Å in the rest frame at $z = 2.983$. To calculate a 3σ upper limit for the equivalent width, we estimate the continuum from HST photometry. Our analysis yields a 3σ upper bound of ~ 350 Å for the [O II] doublet, and we compare this to the distribution measured by the MOSDEF survey⁶³. The host galaxy of the multiply imaged SN we have observed has a stellar mass of $\sim (2 \times 10^8) M_\odot$, which is less massive than a majority of the MOSDEF galaxies with detected [O II] emission lines. Therefore, an [O II] detection would not be expected, given the MOSDEF sample.

Modelling the gravitational lens

After the discovery of the transient, we first examined the predictions of the existing HFF lens models¹⁵. Because the multiple-image system is only several arcseconds in size, we considered three high-resolution published models with $\leq 0.06''$ pixel⁻¹: Keeton V4^{64–66}, GLAFIC V4^{19–21} and Sharon V4^{67,68}. Among these three models, the Keeton V4 model provides the best reconstruction of the multiple images given a source redshift of $z = 3$, as shown in Extended Data Fig. 3a. However, the mass of the early-type galaxy is not included as a free parameter for these models, because the multiply imaged host galaxy was not yet known. Consequently, we developed our own lens model using GLAFIC.

To explore geometric constraints on the SN's redshift independent of our photometric redshift using these 'blind' models, we next allow the redshift of the SN to be a free parameter. In Extended Data Fig. 3b, we plot the root-mean-square (r.m.s.) angular separation of the predicted images from the observed positions as a function of the source redshift. For all three published models, we are able to rule out low redshifts at $z \approx 1$, where the r.m.s. angular separation of the predicted images is much larger than that at higher redshifts.

In Extended Data Fig. 3b, we can see that, for redshifts exceeding 3, the models still yield relatively small r.m.s. residuals. This is because Einstein radii of the lens, for the published models, are slightly smaller than the true value obtained by allowing the mass of the galaxy to be a free parameter. The Einstein radius depends on the ratio of angular

diameter distances D_{ls}/D_{os} , where D_{ls} and D_{os} are angular diameter distances from the lens to the source and from the observer to the source, respectively. However, this ratio changes slowly for source redshifts beyond $z \approx 3$. Consequently, the models' underestimated mass for the early-type lens cannot be compensated for by increasing the redshift of the source.

To use the image positions as constraints and allow the parameters of the early-type lens to vary freely, we modelled the gravitational lens in the Abell 370 cluster using the GLAFIC modelling package^{21,22} and MCMC analysis. We explicitly included the observed image positions as part of the fit. We adopted a prior on each image's position within a standard deviation of 0.1'' as the standard deviation of the prior probability. We use a truncated singular isothermal ellipsoid model where the velocity dispersion, truncation radius, ellipticity and position angle are all free parameters independent of those of the other cluster members to model the early-type lens. We use the joint probability density we obtained from the photometric-redshift measurement as shown in Extended Data Fig. 2g,h to apply a Gaussian prior on the SN host galaxy's redshift at 2.95 with $\sigma = 0.05$. As shown in Extended Data Fig. 3c, our GLAFIC model can reproduce the multiple SN images with r.m.s. residuals of $< 0.02''$.

The original GLAFIC V4 model predicts that the early-type galaxy has a velocity dispersion of the lens of 152 km s⁻¹, whereas our revised model, which is constrained by the image positions and photometric redshift, yields a velocity dispersion of the lens of 172 km s⁻¹. The difference between these values is within the scatter of the Faber–Jackson relation⁶⁹. If we fix the SN redshift to $z = 1$, the velocity dispersion of the lens becomes much larger, ~ 268 km s⁻¹, and therefore its mass-to-light ratio deviates substantially from those of the other cluster member galaxies. This analysis provides independent evidence from gravitational lensing alone that excludes low redshifts at $z \approx 1$.

The resulting time delays (Δt) and magnifications (μ) of the multiple images are listed in Extended Data Table 1b. We note that the uncertainties associated with Δt and μ listed in Extended Data Table 1b are the standard deviations of each individual parameter. These parameters are highly correlated, as shown by the distribution of Δt and μ values from the MCMC samples of the lens model plotted in Extended Data Fig. 3d,e. The distribution of Δt and μ given by the model samples is used as the prior density for the following light curve fitting. The magnification contributed by the macro cluster-lens model at the position of the multiple images is ~ 8 , as given by the sum of the signed magnifications of the quadruple⁷⁰.

Light curve fitting

The breakout may take place when the shock reaches the surface of the progenitor star. In that case, the early-time emission from the shock is from the envelope of the star. The early light curve of SNe based on this mechanism has been modelled by previous studies^{23,24,27}. Alternatively, if an amount of CSM has been ejected prior to the SN explosion and its optical depth is larger than c/v , the breakout will occur at a large radius outside the star's surface. In such a case, the early-time SN emission is from the shocked CSM^{10,25,26,30}, in which a low-mass ($\sim (0.01\text{--}0.1) M_\odot$) and extended ($\sim 10^{13}$ cm) CSM is assumed. Such a scenario could also be used to explain the emission from fast blue optical transients^{30,71–73}.

We fit the observed data with shock-cooling models using our gravitational lens model. We considered two shock-cooling scenarios assuming shock breakouts in massive envelopes of supergiant stars²⁷ or in light-mass and dense CSM shells around progenitors¹⁰. Detailed descriptions of the models we use can be found in the literature^{10,14,30}. For all models, we assumed the opacity to be 0.34 cm² g⁻¹. The bandpass flux is computed using the pysynphot package (<https://pysynphot.readthedocs.io/en/latest/>) based on blackbody spectra.

Each of the SN models includes a set of four free parameters that comprise the radius R and mass M of the envelope or the CSM, the shock velocity v , and the initial time of the shock breakout t_0 , which are as

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follows: $R, M, v, t_0, \Delta t_{13}, \Delta t_{23}, \mu_{S1}, \mu_{S2}, \mu_{S3}$ and $E(B - V)$, where R and M are the radius and mass of the envelope or the CSM, v is the shock velocity, t_0 is the initial time of the shock breakout, Δt_{13} and Δt_{23} are, respectively, the relative time delays of S1 and S2 with respect to S3, and μ_{S1}, μ_{S2} , and μ_{S3} are the magnifications of S1, S2 and S3, respectively. R, M and v are on a logarithmic scale in the parameter space.

The reconstruction of the extinction-corrected rest-frame light curve from the observations involves six additional parameters that are constrained by the lens model. The lensing-related parameters are Δt_{13} and Δt_{23} , which are, respectively, the relative time delays of S1 and S2 relative to S3, and μ_{S1}, μ_{S2} and μ_{S3} , which are, respectively, the magnifications of S1, S2 and S3. Finally, $E(B - V)$ is the colour excess of the SN due to the host-galaxy dust extinction.

We next perform an MCMC analysis using the `emcee`⁷⁴ package to obtain the posterior probability distribution for the model parameters. We adopt the GLAFIC posterior probability density distributions as priors on $\Delta t_{13}, \Delta t_{23}, \mu_{S1}, \mu_{S2}$ and μ_{S3} . For all of the remaining parameters, we implement uniform priors within the parameter ranges listed in Extended Data Table 1c. The parameter values and their confidence intervals for all models are listed in Extended Data Table 2. In our fitting, epochs are defined relative to the rest-frame time when the F814W imaging of S3 was acquired.

The CSM-planar model has additional parameters, including the shock-velocity scale. We chose the same values as used in previous modelling efforts³⁰: $x_0 = R_0/7$, $v_0 = v_e$ and $\beta = 0.5$. We considered three options of the host-galaxy dust extinction with $R_V = 2.74, 3.1$ or 4 for the cases of the Small Magellanic Cloud bar, average Milky Way diffuse, and extragalactic starburst, respectively. We assumed that the colour excess from the host-galaxy extinction $E(B - V)$ is a free parameter. Corner plots in Extended Data Figs. 4, 5 show the distribution of free parameters from the MCMC samples for $R_V = 3.1$ for the two models that can reproduce our observations. Light curves from the best-fit models are shown Extended Data Fig. 6 for $R_V = 3.1$. MCMC samples, corner plots and best-fit light curves for all models with all our choices of R_V values are available in an online repository⁷⁵.

The first set of photometric measurements of the SN in its rest frame are those made from the image S1, which is closest to the early-type lens. This image has a greater background contribution from the galaxy's light than the other two images, especially in the two near-IR filters. There is no statistically significant detection of the S1 image from the F160W filter, where we measure a flux of 25.19 ± 16.45 nJy (nanojanskys). To explore the effect of the early measurement, we repeated light curve fitting using the RSG and CSM-homologous (with CSM mass $10^{-3} M_\odot < M < 1 M_\odot$) models after doubling the uncertainty associated with the F160W flux of S1. The constraints on the progenitor radii that result from these fits are $526^{+162}_{-118} R_\odot$ and $469^{+158}_{-118} R_\odot$, respectively. Even if the uncertainty of our photometry of the most contaminated image is underestimated by a factor of two, our constraints on the progenitor radius will not be greatly altered.

Bayes factors

To compare the results from fitting each light curve model, we evaluated the Bayes factor for each model against a constant flux model that describes a multiply imaged blackbody with an unvaried temperature that is differently magnified by the galaxy and cluster lenses. For a dataset D , the Bayes factor B_{10} of a hypothesis \mathcal{H}_1 against the null hypothesis \mathcal{H}_0 is given by

$$B_{10} = \frac{\mathcal{L}_B(D|\mathcal{H}_1)}{\mathcal{L}_B(D|\mathcal{H}_0)}, \quad (1)$$

in which the marginal likelihood \mathcal{L}_B is given by

$$\mathcal{L}_B(D|\mathcal{H}) = \int dx P(D|x, \mathcal{H}) \pi(x|\mathcal{H}), \quad (2)$$

where x is the set of model parameters, and P and π are the likelihood functions and prior densities (respectively). Prior densities of time-delay and magnification parameters are from the parameter distribution of the MCMC samples given by our lens modelling. We used the same prior densities for the other parameters as for the parameter estimation, whose upper and lower bounds are listed in Extended Data Table 1c. The integration in the high-dimensional parameter space is evaluated using the Monte Carlo integration method. In particular, we generated a number of Monte Carlo samples based on the prior densities. The marginal likelihood \mathcal{L}_B can be approximated by averaging the calculated likelihood over all the Monte Carlo samples, which is given by

$$\mathcal{L}_B \approx \frac{1}{n} \sum_i^n P(D|x_i, \mathcal{H}), x_i \sim \pi(x_i|\mathcal{H}) \quad (3)$$

where the \sim symbol denotes x_i drawn from the prior distribution π . In this work, we evaluated the marginal likelihood using a numerical integration based on 10^8 Monte Carlo samples for each model. We then calculated the normal logarithm Bayes factor for each shock-cooling model against the constant flux model. The results are listed in Extended Data Table 1c. Because the Bayes factor is the ratio of the marginal likelihoods of two competing statistical models, the difference between two logarithmic Bayes factors corresponds to the models' logarithmic Bayes factor.

Microlensing optical depth

The colours of multiple images of a single source can become different if the caustic of a microlens, such as a star or compact object in the foreground lens, intersects the expanding SN photosphere. If the scale of the source-plane caustic from microlenses is comparable to the size of the source's photosphere, photospheric regions with different temperatures could be magnified differently, leading to a colour change in the microlensed images. The probability of microlensing can be estimated from the optical depth for microlensing. We measured the photometry of the intracluster light in the vicinity of the three images, and then fit it using the FAST++ software^{76,77} to obtain the stellar mass. We obtained a surface stellar mass density of $3.55 M_\odot \text{ pc}^{-2}$, $3.47 M_\odot \text{ pc}^{-2}$ and $2.09 M_\odot \text{ pc}^{-2}$ near G1, G2 and G3, respectively. The result is consistent with the surface mass density⁷⁸ at a few hundred kpc away from the centre of the Abell 370 cluster. The optical depth of microlensing⁷⁹ can be given by

$$\tau = \int_0^{D_L} \Omega_E n(D_L) dD_L, \quad (4)$$

where Ω_E is the solid angle covered by the Einstein ring from microlenses, and $n(D_L)$ is the number density of the microlenses at an angular-diameter distance D_L . To simplify the calculation, we adopt a thin-lens assumption that all mass along the line of sight is concentrated in the cluster and assume all microlenses have the same mass that generate an Einstein ring with radius θ_E . Equation (4) can then be simplified as

$$\tau \approx \mu \pi \theta_E^2 n_L, \quad (5)$$

where μ is the magnification of the macro lens from the cluster and n_L is the number density of the microlenses in the cluster lens. Plugging the surface star density (assuming solar-mass microlenses) and the magnification we obtained earlier, we have the microlensing optical depth of 0.010, 0.013 and 0.006 for S1, S2 and S3, respectively. The optical depth is very small, indicating that the probability that at least two images are significantly microlensed is negligible.

Comparison with UV light curves of low-redshift SNe

In Extended Data Fig. 7a, we compare the evolution of the effective blackbody temperature of the SN reported here with those of an RSG model, SN 2018fif¹⁴ (type IIP, with a $\sim 700R_{\odot}$ RSG progenitor)¹⁴, SN 2013ej⁶ (type IIP, with a $(400\text{--}600)R_{\odot}$ RSG progenitor⁶, or a $2,100R_{\odot}$ CSM shell²⁸), SN 2017eaw^{11,12} (type IIP, with a $\sim 600R_{\odot}$ RSG progenitor)¹², SN 1987A⁸⁰, and SN 2016gkg⁷⁻¹⁰ (type IIb, with the early-time emission driven by a $\sim 0.03M_{\odot}$ CSM shell¹⁰). We can see that the early-time temperature of this event is consistent with the examples of type-IIP SNe.

Modelling of the light curves of SNe IIP finds evidence that a large fraction of the progenitors of CCSNe have an extended, massive CSM²⁸. We note that the early-time light curve of an SN IIP is only sensitive to CSM that has an optical depth exceeding c/v . We next compare our event's light curve with those of two recently discovered SNe IIP/L, SN 2018fif¹⁴ and SN 2021yja⁸¹, whose early light curves allow only a small CSM mass ($\lesssim 0.001M_{\odot}$) around their progenitors^{14,81}. In the rest frame of our SN at $z=3$, the central wavelengths of the ACS-WFC F814W, WFC3-IR F110W, and WFC3-IR F160W filters correspond approximately to the Swift-UVOT's *B*, *UVW1* and *UVW2* bands, respectively. As shown in Extended Data Fig. 7b, the early light curve of our multiply imaged SN is consistent with the Swift-UVOT observations of SN 2018fif and SN 2021yja.

Analytical models show that very early UV observations at wavelengths shorter than $3,000\text{ \AA}$ are needed to constrain the SN's temperature, which is important to accurately infer the radius where the shock breaks out²⁷. As shown in Extended Data Fig. 7c,d, we construct SN IIP light curves using the analytical models with CSM at $R < 1,000R_{\odot}$ and $R > 1,000R_{\odot}$. Recent analysis finds evidence that a larger fraction RSG progenitors of SNe IIP may have CSM shells, but CSM with large masses ($M \gtrsim 0.05M_{\odot}$) are only found at $R \gtrsim 1,000R_{\odot}$ (ref. ²⁸). Our analytical light curves plotted in Extended Data Fig. 7 show that a photosphere or CSM (with optical depth $> c/v$) smaller than $\sim 1,000R_{\odot}$ causes the UV luminosity to fall more rapidly following shock breakout, with subsequent rebrightening, whereas the UV light curves for objects with more extended CSM exhibit much more modest rebrightening.

As Extended Data Fig. 7 demonstrates, detection of rebrightening requires UV observations within the first day in the rest frame. Although our dataset includes these very early UV observations, they are absent from almost all low-redshift SNe IIP. We show the light curves and CSM-homologous models of SN 2013fs⁸² and ASASSN-14gm⁸³, which have been inferred to have CSM shells around their progenitors at $R > 1,000R_{\odot}$ (refs. ^{28,29}). Overplotted CSM-homologous models for the CSM parameters for these events^{28,29} show that UV observations within the first day would be brighter than the measured luminosities for our SN. We note that, in those analyses of CCSNe with CSM^{28,29}, the UV data were not used to constrain the CSM properties.

In Extended Data Fig. 7c,d, we plot the observations of our event and the light curves of our best-fitting models (RSG and CSM-homologous). Comparison with the predicted light curves for the CSM-homologous model for the CSM radius and mass^{28,29}, for SN 2013fs and ASASSN-14gm, demonstrates the connection between the radius of optically dense material and the early UV light curve. When we rerun light curve fitting without the first rest-frame UV observation, our constraints become consistent with a much larger set of progenitor radii, $952^{+529}_{-398}R_{\odot}$ from the CSM-homologous model, and favour a value approximately twice as large ($\sim 480R_{\odot}$) as we obtained from the CSM-homologous model with the first rest-frame UV observation. In summary, the radius of the optically thick material for our progenitor system is strongly constrained by the earliest UV data points.

Extended Data Fig. 7e,f shows that the early light curve of our event is not well matched by those of SNe types Ic, IIb and Ia, and also of fast blue optical transients, with available early-time Swift photometry^{8,84-86}.

The earliest emission from an SN originates from the surface of the exploding progenitor. The differences between the gravitational potential of the photosphere at the observed epochs do not suffice to

produce a detectable colour change. For example, applying Birkhoff's theorem, for a BSG with $R = 10R_{\odot}$ and $M = 20M_{\odot}$, the gravitational redshift of light from the BSG surface is 4.2×10^{-6} for a distant observer. Even for a compact white-dwarf progenitor with $R = 0.01R_{\odot}$ and $M = 1M_{\odot}$, the gravitational redshift is only 2.1×10^{-4} for light from the object's surface. We also note that, because the light from the background SN travels into and then out of the potential, the SN's light should not experience any net gravitational redshift.

Core-collapse SN rate from multiply imaged SNe

To simulate the HST detection of multiply imaged SNe behind the six HFF galaxy clusters, we randomly generated CCSNe in comoving volumes behind the cluster lens within a $0.03^{\circ} \times 0.03^{\circ}$ field of view and simulate strong-lensing effects based on the GLAFIC lens models of the six clusters¹⁹⁻²². We used the *sncosmo* package (<https://sncosmo.readthedocs.io/en/stable/index.html>) to synthesize the multiband light curves for each SN. For each type of CCSN, we adopted the mean and standard deviation of its peak *B*-band absolute magnitude from a previous study⁸⁷. We estimated the significance of any detection of the SN using the actual HST exposures of archival imaging of the HFF clusters employing the HST Exposure Time Calculators (<https://etc.stsci.edu/>). In a redshift range from z_1 to z_2 , the expected number of detectable multiply imaged CCSNe in the SN rest-frame time t is given by

$$n_{\text{mcc}} = d_{\text{H}}^3 \theta^2 \int_{z_1}^{z_2} t(z) \beta(z) R_{\text{CC}}(z) \left(\int_0^z \frac{dz'}{E(z')} \right)^2 \frac{dz}{E(z)}, \quad (6)$$

where θ is the angular size of the searching window, $d_{\text{H}} = c/H_0$ is the Hubble distance, R_{CC} is the volumetric CCSN rate, and $E(z) = \sqrt{\Omega_m(1+z)^3 + \Omega_A}$.

For an observer-frame survey time t_{obs} , the rest-frame time follows as $t(z) = t_{\text{obs}}/(1+z)$. For a threshold of the signal-to-noise ratio (SNR) on a point-source detection, $\beta(z)$ is the fraction of the detectable multiply imaged SNe to the total number of the simulated SNe at redshift z . In this work, we assume that the detectable SN image has a point-source SNR larger than 5σ , which is the smallest SNR of the detected images of the known multiply imaged SNe. We also require that the detectable multiply imaged SN has at least one image brighter than the threshold.

In our simulation, we use an exponential A_V distribution that follows $P(A_V) \propto \exp(-\lambda_V A_V)$. The CANDELS+CLASH analysis⁴⁰ used $\lambda_V = 0.187$, although we are not able to reproduce this prior from the data used to construct it^{41,42}, and another study⁸⁸ gave $\lambda_V = 1/\tau_V \approx 2$ based on simulated CCSN host galaxies⁸⁹. For a few hundred CCSN host galaxies in the low-redshift universe⁴³, we obtained the best-fit $\lambda_V = 0.98 \pm 0.03$, as shown in Extended Data Fig. 8a. For $\lambda_V = 1$ and $R_V = 4.05$, the mean and median host-galaxy $E(B-V)$ are 0.25 and 0.17 mag, respectively, consistent with the mean and median $E(B-V)$ from previous studies^{44,45}. In this study, we used $\lambda_V = 1$ and the $R_V = 4.05$ (ref. ⁹⁰) as the primary values. We have also repeated our simulation with $R_V = 3.1$, as well as $\lambda_V = 0.187$ and $\lambda_V = 2$ as used in the two previous analyses^{40,88}, for the purpose of comparison.

Extended Data Fig. 8b shows the differential comoving volume for the strongly lensed sources as a function of redshift. Although the lensing volume remains large at high redshifts, we can see that the effective volume for $>5\sigma$ detection decreases with increasing redshift in the high-redshift regime. The resulting differential number of multiply imaged CCSNe with $>5\sigma$ detection significance by HST in the last decade is shown in Extended Data Fig. 8c. At high redshifts, the number of detectable multiply imaged SNe is small, not only because the effective lensing volume for detectable sources decreases with increasing redshift, but also because of the redshift time dilation. For a given survey period in the observer's frame, the corresponding rest-frame time at higher redshifts is much shorter than that at lower redshifts.

We chose four redshift bins: 0.5–1, 1–2, 2–3.5 and 3.5–6. The number of detected multiply imaged CCSNe (N_{mcc}) within each redshift bin

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follows a Poisson distribution as $\text{Poi}(N; n) = n^N e^{-n} / N!$, where $n = n_{\text{mcc}}$ and $N = N_{\text{mcc}}$. As shown in Fig. 4, we constrain the average R_{CC} for each redshift bin to the 68% confidence level based on the two discovered multiply imaged SNe, where upper limits are for redshift bins with nondetection.

The coefficient k_{CC} can be calculated to be $0.0068 M_{\odot}^{-1}$ (ref. ⁴⁶) for a Salpeter IMF⁴⁷ and an initial mass range of CCSN progenitors ($8\text{--}40 M_{\odot}$). Adopting the star-formation rate density from an analysis of the cosmic star-formation history (CSFH) inferred from observations of galaxies⁴⁶,

$$\psi(z) = \frac{A(1+z)^C}{((1+z)/B)^D + 1} M_{\odot} \text{ yr}^{-1} \text{ Mpc}^{-3}, \quad (7)$$

where $A = 0.015$, $B = 2.9$, $C = 2.7$ and $D = 5.6$, we reproduce the previously predicted CCSN rate⁴⁶, as shown by the solid black curve in Fig. 4 (the same as the curve shown in figure 10 of the CSFH analysis)⁴⁶.

On the other hand, for a given model of star-formation rate density (ψ), k_{CC} can also be constrained using the total number of detected multiply imaged CCSNe behind the six HFF clusters. Here we evaluate a posterior probability density as given by

$$P(k_{\text{CC}}) = \text{Poi}(N_{\text{mcc}}; n_{\text{mcc}}) \times \sqrt{1/n_{\text{mcc}}}, \quad (8)$$

where $\sqrt{1/n_{\text{mcc}}}$ is from the Jeffreys prior⁹¹ for Poisson distributed variables. The variable n_{mcc} is a function of R_{CC} , and therefore is a function of k_{CC} and ψ , which can be derived using equations (6), (7). Using equation (7) and the parameters from the analysis of the CSFH⁴⁶, we find $k_{\text{CC}} = 0.0087^{+0.0071}_{-0.0046} M_{\odot}^{-1}$ and a MLE value of $k_{\text{CC}} = 0.0060 M_{\odot}^{-1}$.

SN Refsdal, the first known example of a strongly lensed SN, and one of the most distant examples of an SN (at $z = 1.49$), was classified as SN 1987A-like⁹². The classification was surprising, because SN 1987A-like SNe account for only 1%–3% of CCSNe reported by low-redshift surveys^{93,94}. The Caltech Core-Collapse Project reported a somewhat greater percentage of 5%⁹⁴.

We used our simulations to estimate the probability that, of the two first discoveries of strongly lensed CCSNe (by a cluster lens), at least one was SN 1987A-like. The absolute magnitudes of the SN 1987A-like SNe in our simulations have a mean of -16 and standard deviation of unity, as an approximation of the distribution of absolute magnitudes of SN 1987A-like SNe⁹². As this figure shows, SN Refsdal would be a relatively luminous SN 1987A-like SN with an absolute B -band magnitude of approximately -16.5 given a magnification of 15 for images S1–S3. Assuming that SN 1987A-like SNe account for 5% of CCSNe to $z \approx 4$, we find that we should expect 1.7 strongly lensed CCSNe in the existing observations of the HFF cluster fields, but only 0.09 of these should be SN 1987A-like SNe. Consequently, the discovery of SN Refsdal, an SN 1987A-like SN⁹² at $z = 1.49$, is not highly probable. However, the P value exceeds $P = 0.05$, given the two known examples.

We have rerun our simulation with an extreme assumption that SN 1987A-like SNe account for 30% of CCSNe in the lensed volume, and find that we should expect 1.76 strongly lensed SNe in the existing observations of the HFF cluster fields. Consequently, inferences about the core-collapse rates are almost entirely unchanged (only $\sim 3.5\%$), indicating that our constraints will not be substantially affected by our assumptions about the rate of SN 1987A-like SNe at high redshift.

As shown in Extended Data Fig. 8e, our constraints on the volumetric rate of CCSNe obtained using $R_{\text{V}} = 3.1$ are not substantially different from the those derived with $R_{\text{V}} = 4.05$. However, our estimate of the CCSN rate changes significantly with the parameter λ_{V} of the exponentially distributed A_{V} . For $\lambda_{\text{V}} = 0.187$, the same as used in a previous analysis of high-redshift SN rates⁴⁰, we find $k_{\text{CC}} = 0.0216^{+0.0146}_{-0.0111} M_{\odot}^{-1}$ and the MLE $k_{\text{CC}} = 0.0162 M_{\odot}^{-1}$. After correction for an extraneous factor of $h^2 = 0.7^2$ (see discussion below), our reanalysis of the CANDELS+CLASH analysis⁴⁰ finds that their measurements yield $k_{\text{CC}} = (0.0045 \pm 0.0008) M_{\odot}^{-1}$.

The volumetric CCSN rate can be expressed as a function of the star-formation density: $R_{\text{CC}}(z) = \psi(z) k_{\text{CC}}$ (ref. ⁴⁶). The CANDELS+CLASH analysis introduced a factor of h^2 where $H_0 = 100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$ on the right-hand side of the equation⁴⁰, which we believe is not needed. Instead, the equation for the measured $\langle R_{\text{CC}} \rangle$ and $\langle \psi \rangle$ should be

$$\langle R_{\text{CC}} \rangle = k_{\text{CC}} \times \langle \psi \rangle \times \left(\frac{H_0^*}{H_0} \right)^2, \quad (9)$$

where H_0^* is the value of H_0 used for the star-formation density. However, both the volumetric SN rate from the CANDELS+CLASH analysis⁴⁰ and the star-formation density from the analysis of the CSFH⁴⁶ adopted $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. Consequently, the factor H_0^*/H_0 should be equal to 1. Following equation (9), we find that the k_{CC} value reported in the CANDELS+CLASH analysis⁴⁰ should be reduced by a factor of 0.7².

We next compute the probability of detecting the SN we have found at $z \approx 3$, given the constraints on the CCSN rate from the CANDELS+CLASH analysis⁴⁰. The authors used the functional form for the cosmic star-formation history⁴⁶, as shown in equation (7), but fit for new values of the coefficients: $A = 0.015$, $B = 1.5$, $C = 5.0$ and $D = 6.1$. We simulate multiply imaged SNe based on the cosmic star-formation history, the value of k_{CC} , and the extinction parameter $\lambda_{\text{V}} = 0.187$ (ref. ⁴⁰). Our simulations show that we expect to detect $n_{\text{mcc}} = 0.25$ multiply imaged SNe in the existing observations of the HFF cluster fields based on the parameters from the CANDELS+CLASH analysis⁴⁰. We find a small probability of detecting two or more multiply imaged SNe, $P(N_{\text{mcc}} \geq 2) = 0.026$. Moreover, the star-formation history from the CANDELS+CLASH analysis⁴⁰ declines rapidly in the range $z \approx 1\text{--}3$, and the probability of finding at least one multiply imaged SNe at $z \geq 3$ is very small ($P \approx 0.01$). The detection of our event is extremely unlikely given the constraints from the CANDELS+CLASH analysis⁴⁰. In Extended Data Fig. 8e, we compare our constraints on the volumetric CCSN rate with those from the CANDELS+CLASH analysis⁴⁰.

For $\lambda_{\text{V}} = 2$ (ref. ⁸⁸), the estimated CCSN rate is smaller than what we obtained from $\lambda_{\text{V}} = 1$ on the basis of the distribution of A_{V} measured for the host galaxies of nearby SNe⁴³, but the 68% confidence intervals for these two λ_{V} choices, as shown by the green-shaded region and the blue-shaded region in Extended Data Fig. 8e, highly overlap, indicating little statistical difference. For $\lambda_{\text{V}} = 2$, we obtained $k_{\text{CC}} = 0.0072^{+0.0059}_{-0.0038} M_{\odot}^{-1}$ and the MLE $k_{\text{CC}} = 0.0049 M_{\odot}^{-1}$ from observed multiply imaged SNe, on the basis of the cosmic star-formation history from analysis of the CSFH⁴⁶.

So far, we have only reanalysed archival imaging of the six HFF clusters, which led to the observation of the strongly lensed SN not identified by previous searches. In the earlier CLASH programme, imaging of 25 galaxy-cluster fields (including four of the six HFFs) was acquired with a total of 524 orbits. No multiply imaged SN has been publicly reported from the CLASH data in previous studies. If we assume that none was missed and an SN can be detected within a 0.5-yr time period bracketing each observation, our volumetric CCSN rate would decrease by a factor of ~ 2 , yielding $k_{\text{CC}} \approx 0.008 M_{\odot}^{-1}$ for the host-galaxy extinction prior from the CANDELS+CLASH analysis⁴⁰. Nonetheless, our CCSN rate and the k_{CC} value are still higher than the constraints obtained from blank-field HST surveys⁴⁰. If we assume that no multiply imaged SNe from the CLASH survey were missed, then the tension between our measurements and those of the CANDELS+CLASH analysis⁴⁰ (after correction) diminishes to $P = 0.088$, but the probability of finding at least one multiply imaged SN at $z \geq 3$ is still small ($P \approx 0.02$).

Data availability

The HST data used for this study can be retrieved from the NASA Mikulski Archive for Space Telescopes (<http://archive.stsci.edu>). The supernova is found in the HST imaging of the Abell 370 field acquired from programme GO-11591 (PIJ.-P. Kneib). The LBT spectroscopy data

are available from the LBT archive (<http://archive.lbt.org>). The Keck MOSFIRE data can be retrieved from the Keck Observatory Archive (<https://www2.keck.hawaii.edu/koa/public/koa.php>). The HFF data and models can be downloaded from <https://archive.stsci.edu/prepds/frontier/lensmodels/#modelsandinput>. Additional data including the MMT spectroscopic data, the HST coaddition and image differencing data, the GALFIT scripts and resulting models, the HST photometry data of the SN host galaxy, the SN light curve fitting script and resulting MCMC data, and our best-fit GLAFIC lens model are available from <https://doi.org/10.5281/zenodo.6725770>.

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Author contributions W.C. analysed the HST, Keck, LBT and MMT data, wrote the manuscript and developed simulations. P.L.K. aided the interpretation of the events, planned the Keck, LBT and MMT observations, edited the manuscript and cross-checked the results. M.O. constructed the gravitational-lensing model for this work. T.J.B., J.M.D. and A.Z. helped with the HST proposal and reviewed the manuscript. P.L.K., N.E. and A.Z. acquired the Keck observation. P.L.K. and N.E. acquired the LBT observation. A.Z. constructed independent lens models for the cluster to verify the time-delay results. A.V.F. obtained the Keck time, helped acquire the Keck MOSFIRE observations, and edited the HST proposal and this manuscript. T.L.T. contributed to the manuscript and interpretation.

Competing interests The authors declare no competing interests.

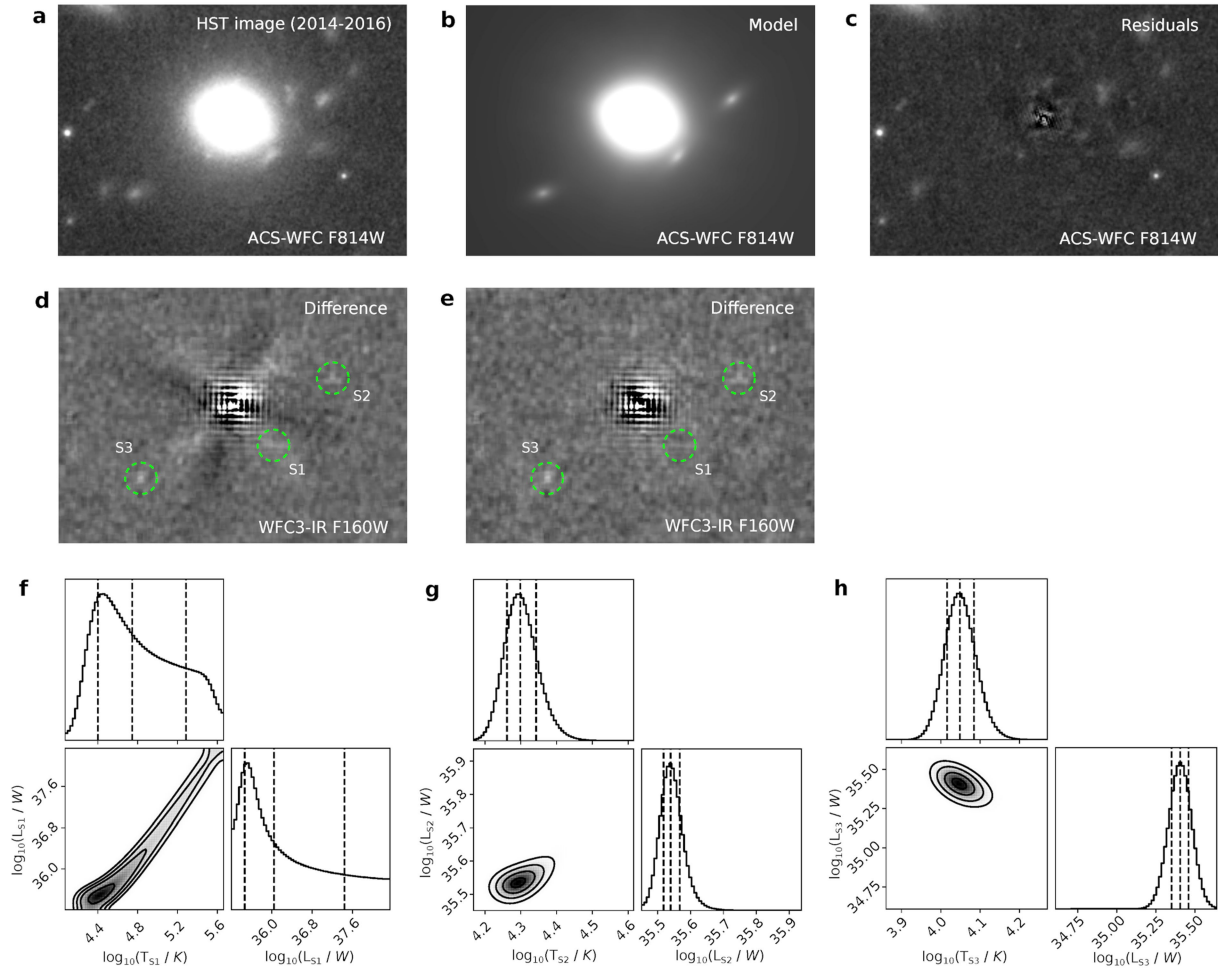
Additional information

Supplementary information The online version contains supplementary material available at <https://doi.org/10.1038/s41586-022-05252-5>.

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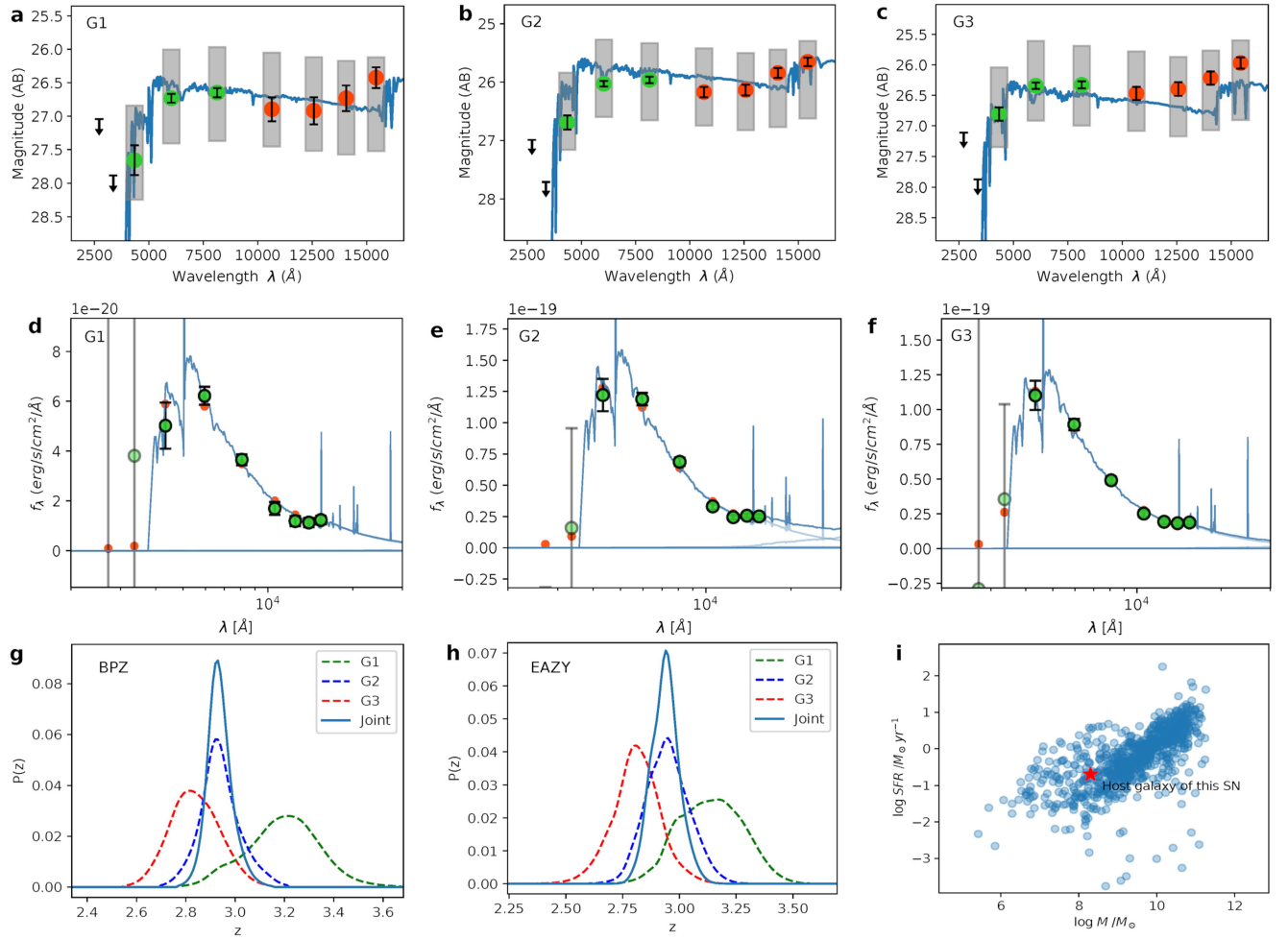
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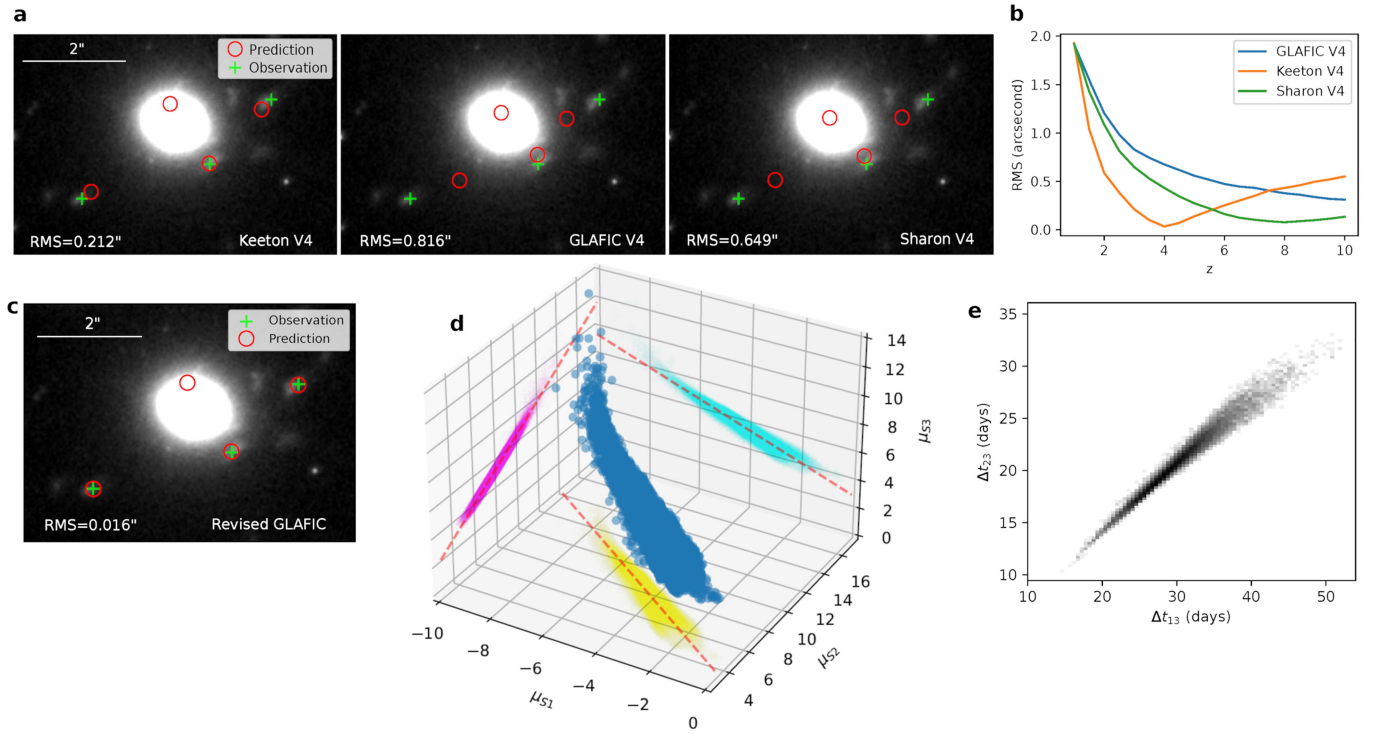
Extended Data Fig. 1 | Photometry of the multiply imaged SN: image differencing and the blackbody fitting. We used GALFIT⁵¹ to fit bright sources in the lensing system and subtracted the best-fit model from the field prior to measuring the flux. **a–c**, The template image (**a**), the best-fit GALFIT model (**b**), and the residual through the HST ACS-WFC F814W filter (**c**). **d**, The WFC3-IR F160W difference image by subtracting the template from the event image directly. **e**, The difference image using our GALFIT-based method, where

we subtracted the PSF-convolved best-fit GALFIT source models from the images, and then calculated the difference from GALFIT residuals between the coadded images and the event images. We can see that this method can reduce the significant residual from the mismatched PSFs around bright sources as shown in **d**. **f–h**, Distributions of effective temperature and luminosity from the MCMC samples from fitting the blackbody emission into the photometry of each SN image.



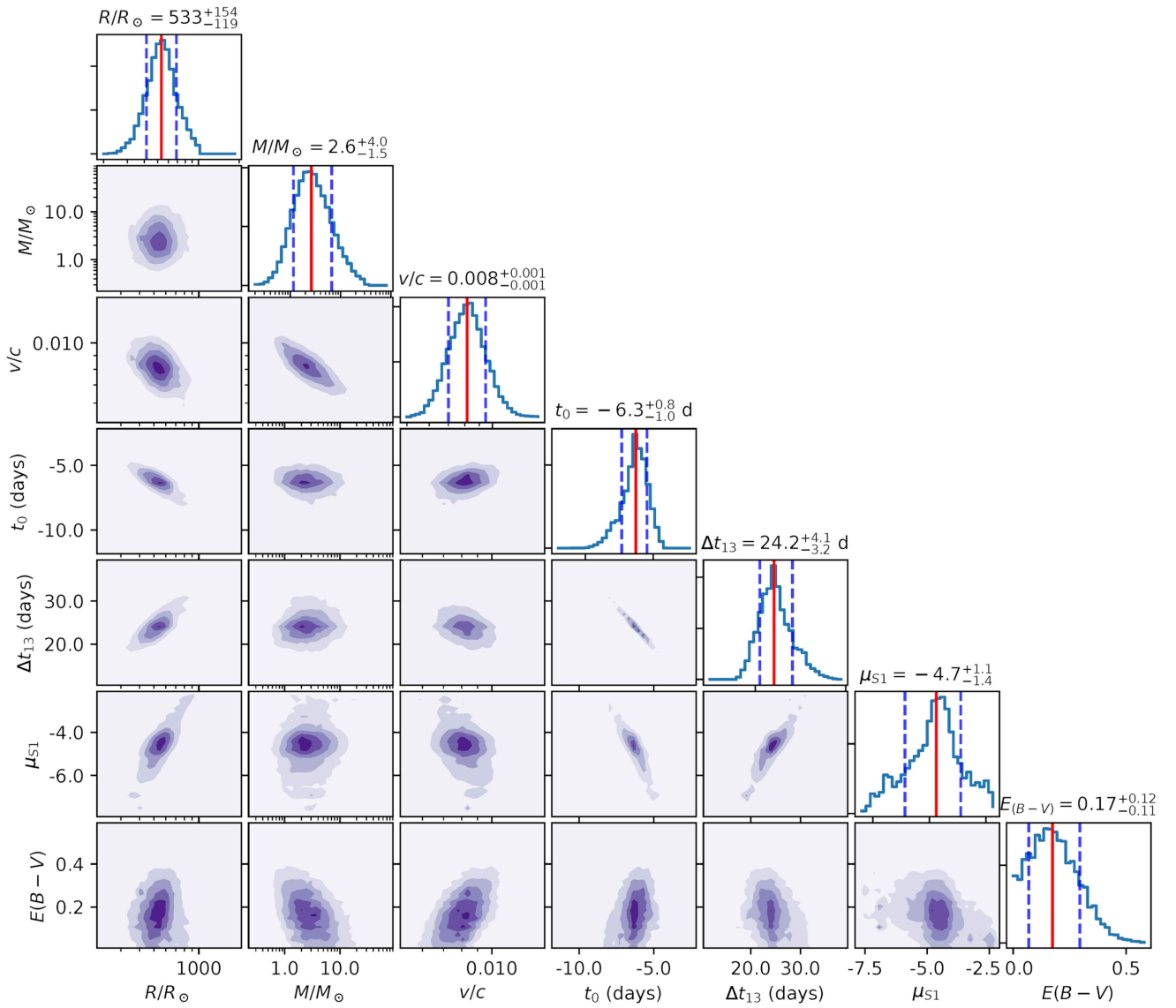
Extended Data Fig. 2 | Photometry and photometric redshift of the SN host galaxy. **a–c**, Broadband spectral energy distributions and best-fitting BPZ¹⁷ templates for G1, G2 and G3. The coloured circles designate the observed AB magnitudes with $\pm 1\sigma$ uncertainties measured from ACS-WFC (green) and WFC3-IR (red) images, and arrows correspond to 95% upper limits. Dark-blue curves plot the best-fit spectral templates. Grey rectangles mark the magnitudes calculated from the best-fitting BPZ templates (with approximate uncertainties) for those filters⁹⁵. **d–f**, Broadband spectral energy distributions and best-

fitting EAZY¹⁷ templates for G1, G2 and G3. The data points are the observed flux with $\pm 1\sigma$ uncertainties. Blue curves show the best-fit spectral template. The three panels on the right display the posterior probability distribution of the photometric redshift. **g, h**, Photometric redshift probability distributions for G1 (green dashed line), G2 (blue dashed line), G3 (red dashed line), and the joint analysis (solid line) derived from BPZ and EAZY fitting. **i**, The distribution of CCSN host-galaxy stellar mass (M) and star-formation rate (SFR)⁵⁵, where the red star marks M and SFR of the host galaxy of the SN reported in this work.



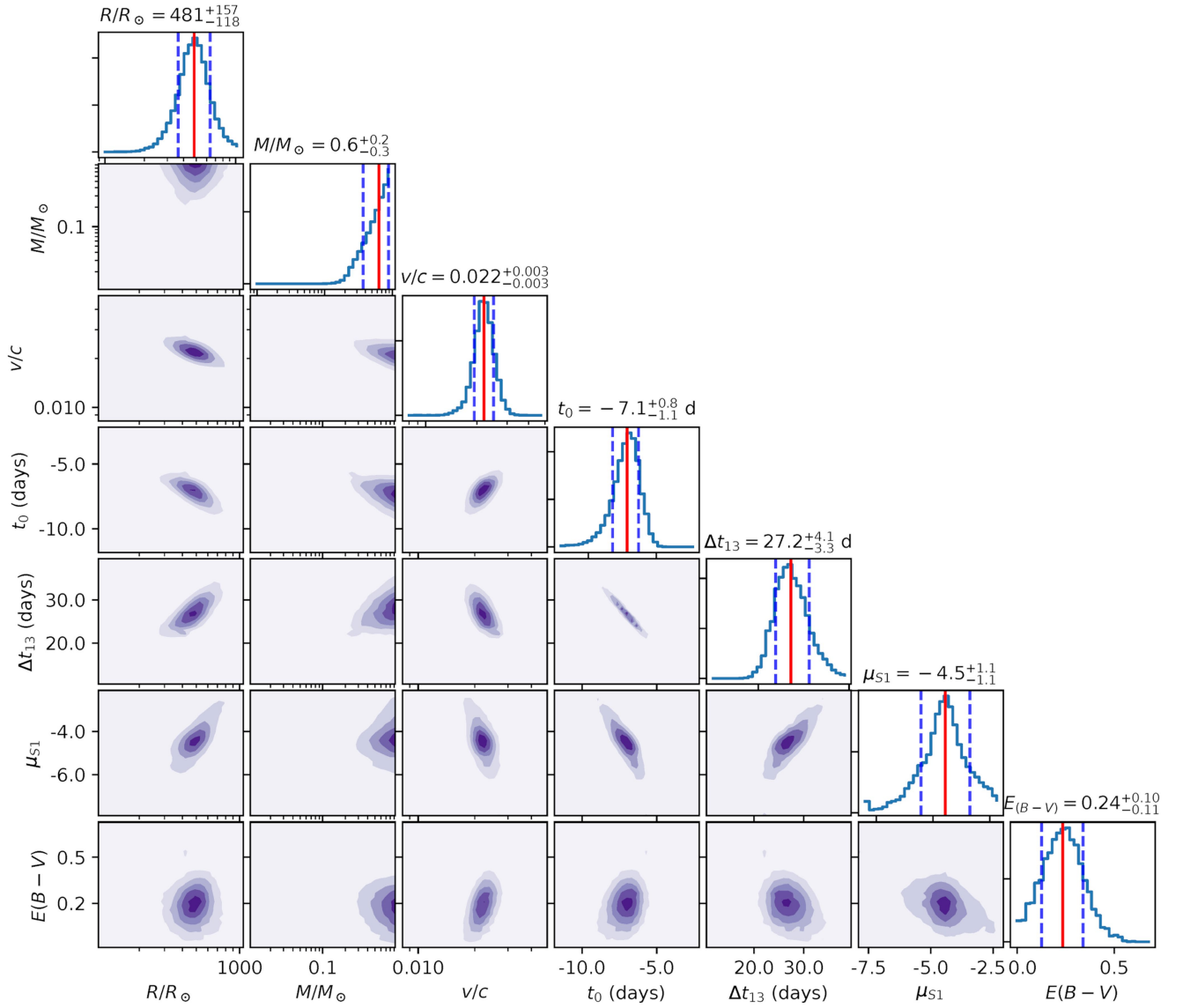
Extended Data Fig. 3 | Modelling the strongly lensed images. a, Best-fit multiple-image geometry from the published Keeton V4 (left), GLAFIC V4 (middle) and Sharon V4 (right) models, by minimizing the r.m.s. angular separation of the predicted images from the real images in the image plane, for $z = 3$. **b,** Best-fit r.m.s. angular separation of the predicted images as a function of z from the published Keeton V4, GLAFIC V4 and Sharon V4 models. **c,** Best-fit multiple-image geometry from our revised GLAFIC model. **d,** Distribution of

μ_{S1} , μ_{S2} and μ_{S3} from the MCMC samples given by our revised GLAFIC model, where μ_1 , μ_2 and μ_3 are predicted magnifications of images S1, S2 and S3, respectively. Cyan, yellow and magenta markers show the distribution projected on the $\mu_{S1}-\mu_{S3}$, $\mu_{S1}-\mu_{S2}$ and $\mu_{S2}-\mu_{S3}$ planes, where red-dashed lines are the best-fit lines from linear regressions on the projected distribution. **e,** Correlations between Δt_{13} and Δt_{23} from the MCMC samples, where Δt_{13} and Δt_{23} are the time delay of S1 and S2 relative to S3, respectively.



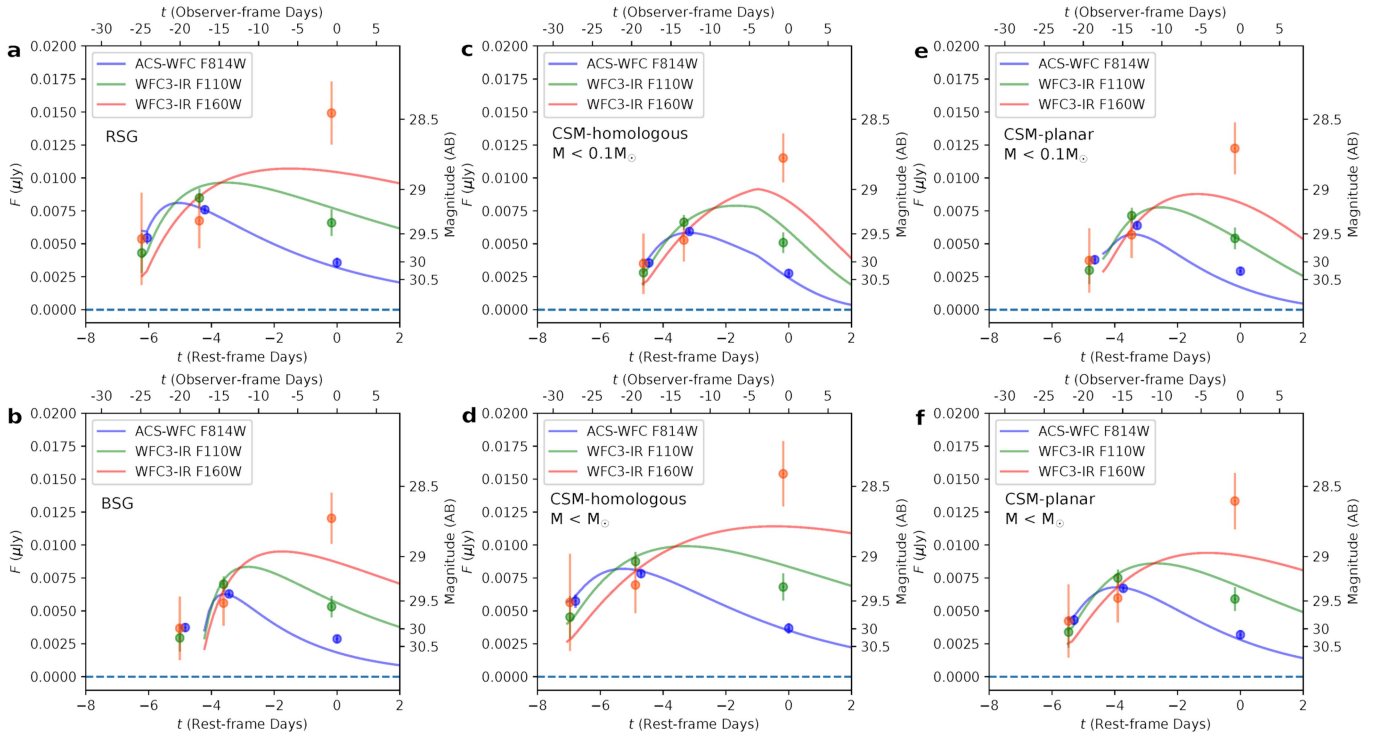
Extended Data Fig. 4 | Distribution of parameters from fitting the RSG model to the observations. Here R is the extended envelope radius, M is the envelope mass, v is the shock velocity, t_0 is the initial time, and Δt_{13} and μ_{S1} are, respectively, the time delay and magnification of the S1 image. We set $t = 0$ at

the observation time of the S3 image in the ACS-WFC F814W band. The time delay of each image is defined as the difference of its light-travel time from that of S3. The result is for $R_V = 3.1$.



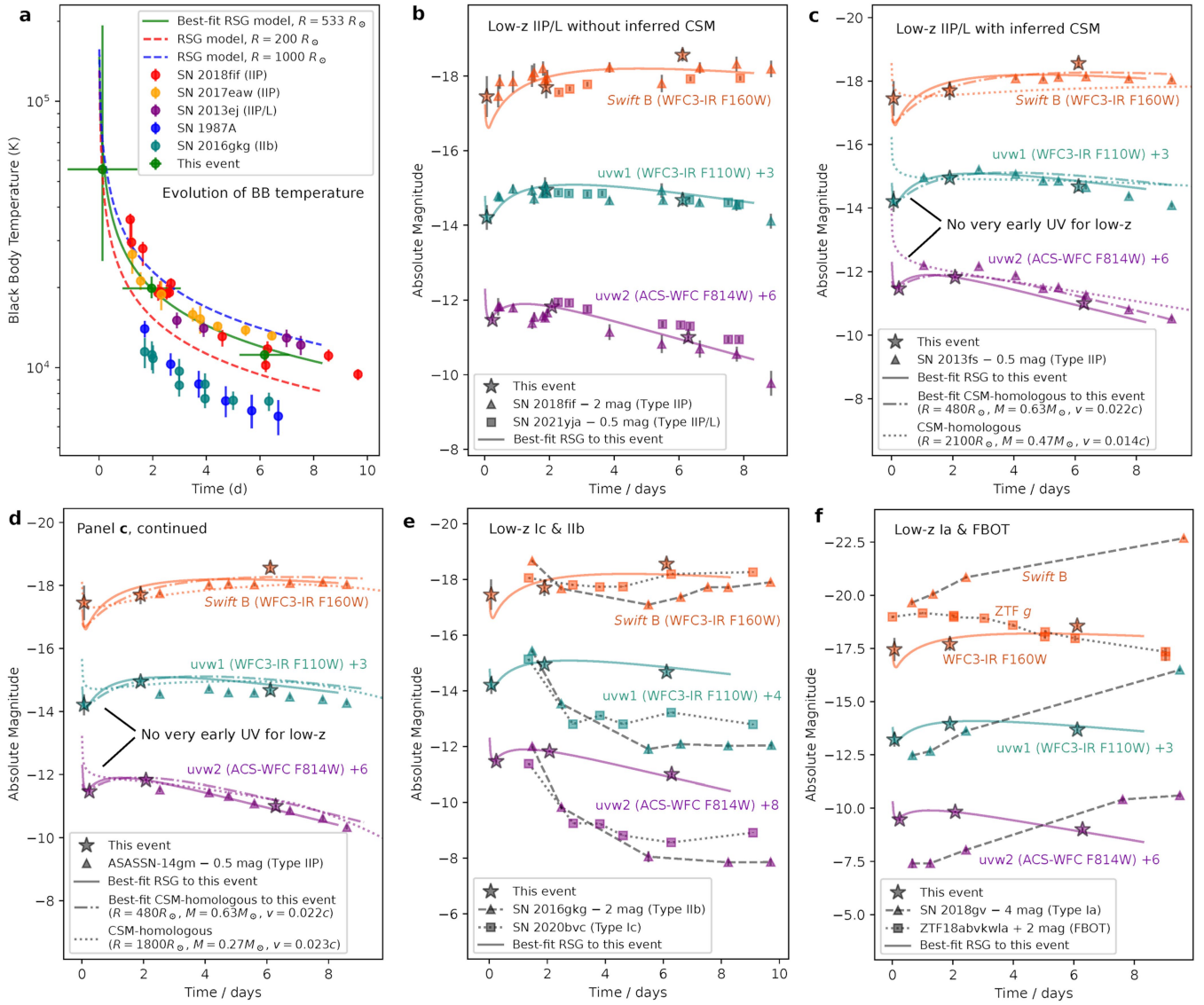
Extended Data Fig. 5 | Distribution of parameters from fitting the CSM-homologous model to the observations. The result is for a prior on the CSM mass of $M < 1M_{\odot}$. The parameters are the same as in Extended Data Fig. 5.

Data of the MCMC samples for all models with all our choices of R_V and priors are available from an online repository⁷⁵.



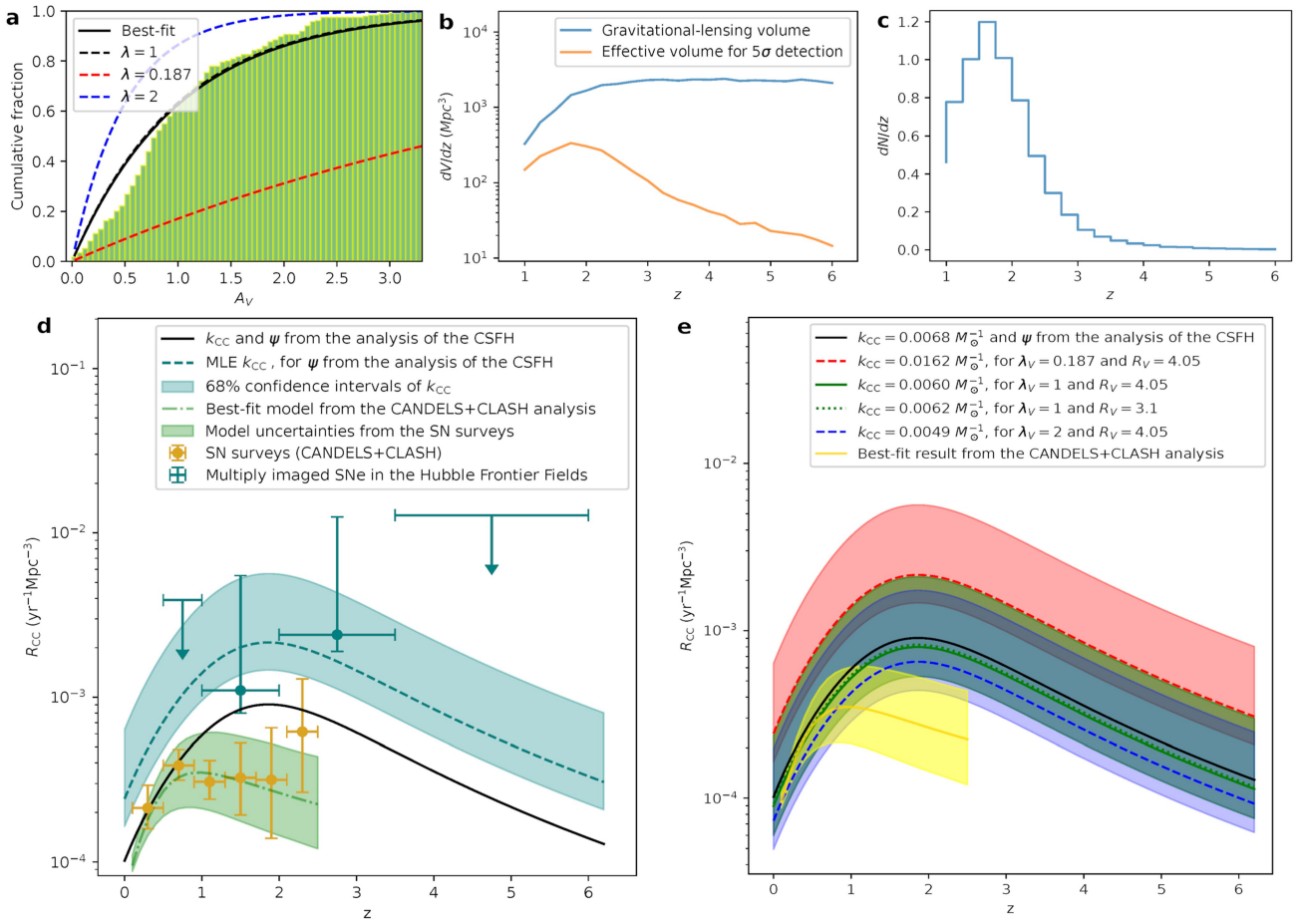
Extended Data Fig. 6 | Reconstructed light curves from the best-fit model parameters from all the light curve models for $R_v=3.1$. a–f. Data shown are from the RSG model (a), the BSG model (b), the CSM-homologous model for a prior on the CSM mass of $M < 0.1M_\odot$ (c), the CSM-homologous model for a prior on the CSM mass of $M < 1M_\odot$ (d), the CSM-planar model for a prior on the CSM

mass of $M < 0.1M_\odot$ (e), and the CSM-planar model for a prior on the CSM mass of $M < 1M_\odot$ (f). Data points and error bars correspond to the measured values and $\pm 1\sigma$ uncertainties of the flux densities, respectively. Reconstructed light curves with all our choices of R_v are available from an online repository⁷⁵.



Extended Data Fig. 7 | Comparisons of the reconstructed SN light curve with early UV light curves from other SNe. **a**, The early-time evolution of the effective blackbody temperature of the SN reported here (green data points, where error bars are 68% confidence intervals), where the effective temperatures of the SN images are obtained by independently fitting the blackbody emission into the photometry of each SN image. The green solid line is from the best-fit RSG model. Red and blue dashed lines are two examples from the RSG model with progenitor radius of $200 R_\odot$ and $1000 R_\odot$, respectively, for the same envelope mass and shock velocity from our best-fit RSG model for the SN reported here. Red, yellow, blue and cyan data points show the early-time evolution of SN 2018fif¹⁴ (type IIP), SN 2013ej⁶ (type IIP/L), SN 2017eaw^{11,12} (type IIP), SN 1987A⁸⁰ and SN 2016gkg⁷⁻¹⁰ (type IIB). **b-f**, Stars and solid curves show the absolute magnitude of the SN reported here and reconstructed light curves from the best-fit RSG model, respectively, with arbitrary magnitude offsets for better visualization. For this SN, orange, teal and purple points and lines are for ACS-WFC F814W, WFC3-IR F110W and WFC3-IR F160W, respectively.

In the rest frame of the SN, central wavelengths of these three filters fall into the *B*, *UVW1*, and *UVW2* bands of *Swift*-UVOT. In **b**, we compare the SN's early light curve to the *Swift*-UVOT observations of SN 2018fif¹⁴ and SN 2021yja⁸¹—two type-II SNe whose early light curves indicate only small amounts of CSM around their progenitors^{14,81}. **c,d**, Comparison to the two type-II SNe, SN 2013fs⁸² and ASASSN-14gm⁸³, which are believed to have dense CSM shells around their progenitors. In **c,d**, we plot the light curve from our best-fit CSM-homologous model for the SN of this work (dash-dotted lines) and light curves from the CSM-homologous model for the CSM radius and mass given by the analyses of CSM-rich CCSNe^{28,29} (dotted lines) for the two type-II SNe. **e,f**, Comparisons to the early UV light curves of SN 2016gkg⁸ (type IIB), SN 2020bvc⁸⁴ (type Ib/c), SN 2018gv⁸⁵ (type Ia), and ZTF18abvkwla⁸⁶ (fast blue optical transient). ZTF-g band light curve is shown for ZTF18abvkwla at $z = 0.27$, corresponding to the rest-frame wavelength of 3,820 Å. All magnitudes are AB magnitudes. BB, blackbody.



Extended Data Fig. 8 | Estimating the CCSN rate based on observations of the multiply imaged SNe. **a**, Cumulative distribution of CCSN host galaxies as a function of the host-galaxy extinction A_V ⁴³. The black solid line shows the cumulative distribution function of the exponential distribution $P(A_V) = \lambda_V \exp(-\lambda_V A_V)$ with best-fit $\lambda_V = 0.98$. Black, red and blue dashed lines show the cumulative distribution function of the exponential distribution with $\lambda_V = 1$, $\lambda_V = 0.187$ and $\lambda_V = 2$, respectively. **b**, Differential comoving volume (dV/dz) for the strongly lensed sources as a function of redshift z . Blue curve shows the lensed volume for all the sources in the six Hubble Frontier Fields within a $0.03^\circ \times 0.03^\circ$ search window for each cluster field. Orange curve shows the

effective lensing volume for sources that can be significantly ($>5\sigma$) detected by HST in the last decade. **c**, Differential number of detectable multiply imaged core-collapse supernovae (dN/dz) above the 5σ signal-to-noise-ratio level by HST in the last decade as a function of z . **d**, Volumetric CCSN rate (R_{CC}) as a function of redshift. Contents are the same as in Fig. 4, but for $\lambda_V = 0.187$. **e**, Volumetric CCSN rate (R_{CC}) as a function of redshift for different choices of λ_V and R_V . Red, green and blue shaded regions are for $\lambda_V = 0.187$, $\lambda_V = 1$ and $\lambda_V = 2$, respectively, for $R_V = 4.05$. Red, green and blue lines are from the MLE k_{CC} for the three λ_V options. The green dotted line is the same as the green solid line, but for $R_V = 3.1$.

Extended Data Table 1 | Results of photometry, gravitational-lensing modelling and the Bayesian test of light curve models

a

SN host galaxy image	G1	G2	G3
WFC3-UVIS F275W	27.04*	26.99*	27.11*
WFC3-UVIS F336W	27.89*	27.71*	27.88*
ACS-WFC F435W	27.66 (0.22)	26.69 (0.12)	26.81 (0.11)
ACS-WFC F606W	26.73 (0.07)	26.03 (0.05)	26.34 (0.05)
ACS-WFC F814W	26.65 (0.07)	25.96 (0.05)	26.33 (0.06)
WFC3-IR F105W	26.90 (0.18)	26.17 (0.09)	26.47 (0.11)
WFC3-IR F125W	26.92 (0.20)	26.14 (0.09)	26.40 (0.11)
WFC3-IR F140W	26.73 (0.19)	25.84 (0.08)	26.22 (0.11)
WFC3-IR F160W	26.42 (0.15)	25.65 (0.07)	25.97 (0.09)

*95% upper limits

b

Image	S1	S2	S3	S4
α	2 ^h 39 ^m 55.248 ^s	2 ^h 39 ^m 55.156 ^s	2 ^h 39 ^m 55.429 ^s	--
δ	-1°33'52.39"	-1°33'51.02"	-1°33'53.08"	--
$F_{V, F814W}$ / nJy	25.60 (2.24)	63.73 (2.35)	22.12 (2.38)	--
$F_{V, F110W}$ / nJy	20.20 (7.10)	71.24 (5.93)	40.95 (6.20)	--
$F_{V, F160W}$ / nJy	25.19 (16.45)	56.71 (17.43)	92.60 (14.86)	--
μ	-4.24 (1.14)	7.79 (1.27)	5.79 (0.80)	-1.21 (0.47)
Δt / days	30.62 (5.56)	21.03 (3.41)	0 (0)	53.49 (7.17)

c

Model	RSG	BSG	CSM-homologous 10 ⁻³ M _⊙ < M < 10 ⁻¹ M _⊙	CSM-homologous 10 ⁻³ M _⊙ < M < M _⊙	CSM-planar 10 ⁻³ M _⊙ < M < 10 ⁻¹ M _⊙	CSM-planar 10 ⁻³ M _⊙ < M < M _⊙
R / R _⊙	(10 ² , 10 ³)	(10 ² , 10 ^{2.5})	(10 ⁰ , 10 ³)	(10 ⁰ , 10 ⁴)	(10 ⁰ , 10 ³)	(10 ⁰ , 10 ⁴)
M / M _⊙	(10 ^{-0.5} , 10 ^{1.5})	(10 ⁰ , 10 ²)	(10 ⁻³ , 10 ⁻¹)	(10 ⁻³ , 10 ⁰)	(10 ⁻³ , 10 ⁻¹)	(10 ⁻³ , 10 ⁰)
v / (0.1c)	(10 ⁻² , 10 ¹)	(10 ⁻² , 10 ¹)	(10 ⁻² , 10 ¹)	(10 ⁻² , 10 ¹)	(10 ⁻² , 10 ¹)	(10 ⁻² , 10 ¹)
t ₀ / days	(-12, -4)	(-12, -4)	(-12, -4)	(-12, -4)	(-12, -4)	(-12, -4)
E(B - V)	(0, 3)	(0, 3)	(0, 3)	(0, 3)	(0, 3)	(0, 3)
logB _{i0}	13.6	6.3	11.1	14.1	5.5	12.6

a. Photometry of G1, G2 and G3 in units of AB magnitude. **b.** Positions (α, δ), flux density (f), magnification (μ) and time delay (Δt) of the multiple SN images. Flux density and time delay are in units of nanojanskys (nJy) and days, respectively, where the time delay is defined as the relative delay with respect to S3. Photometry has been corrected for foreground Milky Way extinction. A negative magnification indicates an opposite parity. Numbers in the parentheses are the standard deviations. **c.** Ranges of the prior density of physical parameters R , M , v , t_0 and $E(B - V)$ for the light curve models and logarithm Bayes factors (log B_{i0}) of the i th light curve model against the constant flux model (the null hypothesis) for $R_V=3.1$.

Extended Data Table 2 | Best-fit values of free model parameters from light curve models

Model	R_V	RSG	BSG	CSM-homologous $10^{-3} M_{\odot} < M < 10^{-1} M_{\odot}$	CSM-homologous $10^{-3} M_{\odot} < M < M_{\odot}$	CSM-planar $10^{-3} M_{\odot} < M < 10^{-1} M_{\odot}$	CSM-planar $10^{-3} M_{\odot} < M < M_{\odot}$
R / R_{\odot}	2.74	510 (402, 667)	295 (272, 308)	361 (305, 427)	464 (352, 638)	335 (279, 683)	283 (223, 363)
	3.1	533 (414, 688)	297 (277, 309)	407 (337, 503)	481 (363, 638)	399 (323, 811)	325 (250, 417)
	4.05	538 (419, 688)	294 (267, 308)	419 (345, 515)	495 (373, 653)	379 (319, 798)	340 (263, 435)
M / M_{\odot}	2.74	3.0 (1.3, 7.3)	16.91(7.2, 35.5)	0.09 (0.08, 0.10)	0.64 (0.38, 0.88)	0.09 (0.08, 0.10)	0.84 (0.64, 0.96)
	3.1	2.6 (1.2, 6.6)	16.0 (6.9, 39.9)	0.09 (0.08, 0.10)	0.63 (0.37, 0.88)	0.09 (0.08, 0.10)	0.83 (0.64, 0.95)
	4.05	2.6 (1.1, 6.6)	16.3 (6.8, 40.6)	0.09 (0.08, 0.10)	0.63 (0.37, 0.88)	0.09 (0.08, 0.10)	0.84 (0.63, 0.96)
$v / (0.1c)$	2.74	0.78 (0.67, 0.90)	0.87 (0.70, 1.19)	2.6 (2.4, 2.8)	2.1 (1.8, 2.4)	1.5 (1.2, 1.6)	1.0 (0.9, 1.2)
	3.1	0.82 (0.70, 0.95)	1.03 (0.75, 1.58)	2.8 (2.5, 3.1)	2.2 (1.9, 2.5)	1.6 (1.4, 1.8)	1.1 (1.0, 1.3)
	4.05	0.83 (0.71, 0.97)	1.11 (0.80, 1.75)	2.9 (2.6, 3.1)	2.2 (2.0, 2.5)	1.7 (1.5, 1.9)	1.2 (1.1, 1.3)
t_0 / days	2.74	-6.3 (-7.4, -5.5)	-4.4 (-4.7, -4.2)	-4.7 (-5.0, -4.4)	-7.2 (-8.4, -6.3)	-4.4 (-4.6, -4.2)	-5.5 (-6.1, -4.9)
	3.1	-6.3 (-7.3, -5.5)	-4.3 (-4.7, -4.1)	-4.7 (-5.0, -4.4)	-7.1 (-8.1, -6.3)	-4.4 (-4.8, -4.2)	-5.5 (-6.2, -5.0)
	4.05	-6.2 (-7.2, -5.4)	-4.3 (-4.7, -4.1)	-4.6 (-5.6, -4.4)	-7.2 (-8.2, -6.3)	-4.4 (-4.6, -4.2)	-5.6 (-6.3, -5.0)
$\Delta t_{13} / \text{days}$	2.74	24.3 (21.1, 28.6)	18.0 (16.7, 20.9)	17.8 (16.8, 19.1)	27.3 (24.0, 31.7)	17.6 (16.6, 22.8)	20.8 (18.6, 23.5)
	3.1	24.2 (21.0, 28.2)	19.3 (17.6, 21.8)	17.8 (16.7, 19.2)	27.2 (23.8, 31.2)	18.5 (16.6, 23.3)	21.2 (18.9, 23.8)
	4.05	24.0 (20.9, 27.7)	19.8 (17.3, 22.3)	17.7 (16.7, 19.0)	27.2 (24.0, 31.2)	17.3 (16.6, 22.7)	21.3 (19.0, 24.0)
$\Delta t_{13} / \text{days}$	2.74	16.8 (14.9, 19.4)	12.8 (11.8, 14.7)	12.7 (11.8, 13.5)	18.9 (16.7, 21.4)	12.8 (11.9, 15.9)	14.8 (13.2, 16.4)
	3.1	16.8 (14.8, 19.3)	13.8 (12.4, 15.2)	12.6 (11.8, 13.7)	18.8 (16.7, 21.2)	13.1 (12.0, 16.2)	14.9 (13.4, 16.6)
	4.05	16.7 (14.9, 19.0)	13.9 (12.3, 15.6)	12.6 (11.8, 13.5)	18.9 (16.7, 21.3)	12.4 (11.9, 15.9)	15.1 (13.4, 16.7)
μ_{S1}	2.74	-4.6 (-5.9, -3.4)	-7.2 (-7.6, -6.2)	-7.3 (-7.7, -6.4)	-4.4 (-5.4, -3.1)	-6.9 (-7.5, -5.6)	-6.2 (-7.2, -4.8)
	3.1	-4.7 (-6.1, -3.6)	-6.9 (-7.5, -5.8)	-7.2 (-7.6, -6.3)	-4.5 (-5.5, -3.4)	-6.8 (-7.5, -5.4)	-6.0 (-7.3, -4.7)
	4.05	-4.8 (-6.2, -3.8)	-6.8 (-7.6, -5.5)	-7.2 (-7.7, -6.3)	-4.5 (-5.5, -3.4)	-7.0 (-7.6, -5.7)	-6.0 (-7.2, -4.7)
μ_{S2}	2.74	8.3 (7.1, 9.5)	10.5 (9.4, 11.3)	11.0 (9.8, 11.7)	8.1 (6.8, 9.1)	10.4 (9.1, 11.2)	9.6 (8.4, 11.0)
	3.1	8.4 (7.3, 9.6)	10.1 (9.2, 11.1)	10.7 (9.7, 11.6)	8.1 (7.0, 9.2)	10.0 (8.9, 11.5)	9.5 (8.3, 10.9)
	4.05	8.5 (7.4, 9.8)	10.2 (8.8, 11.2)	10.8 (9.7, 11.7)	8.1 (7.0, 9.2)	10.4 (9.2, 11.4)	9.5 (8.4, 10.8)
μ_{S3}	2.74	6.2 (5.5, 6.9)	8.1 (7.0, 8.6)	8.2 (7.4, 8.8)	6.0 (5.3, 6.6)	7.8 (6.6, 8.3)	7.1 (6.2, 8.2)
	3.1	6.2 (5.5, 7.0)	7.7 (6.8, 8.4)	8.0 (7.2, 8.6)	6.0 (5.3, 6.7)	7.6 (6.6, 8.4)	6.9 (6.2, 8.1)
	4.05	6.3 (5.6, 7.1)	7.6 (6.6, 8.5)	8.1 (7.3, 8.8)	6.0 (5.3, 6.6)	7.8 (6.7, 8.3)	7.0 (6.2, 8.0)
$E(B - V)$	2.74	0.12 (0.04, 0.22)	0.18 (0.08, 0.47)	0.30 (0.20, 0.39)	0.19 (0.09, 0.28)	0.57 (0.47, 0.69)	0.47 (0.38, 0.56)
	3.1	0.17 (0.07, 0.30)	0.40 (0.13, 0.64)	0.37 (0.26, 0.48)	0.24 (0.13, 0.34)	0.66 (0.54, 0.78)	0.55 (0.45, 0.65)
	4.05	0.35 (0.15, 0.57)	0.77 (0.24, 1.15)	0.69 (0.50, 0.87)	0.44 (0.24, 0.61)	1.16 (0.97, 1.35)	1.00 (0.82, 1.16)

Numbers in parentheses are 68% confidence intervals.