

The host galaxy and persistent radio counterpart of FRB 20201124A

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ABSTRACT

The physical properties of fast radio burst (FRB) host galaxies provide important clues towards the nature of FRB sources. The 16 FRB hosts identified thus far span three orders of magnitude in mass and specific star formation rate, implicating a ubiquitously occurring progenitor object. FRBs localized with \sim arcsecond accuracy also enable effective searches for associated multiwavelength and multi-time-scale counterparts, such as the persistent radio source associated with FRB 20121102A. Here we present a localization of the repeating source FRB 20201124A, and its association with a host galaxy (SDSS J050803.48+260338.0, $z = 0.098$) and persistent radio source. The galaxy is massive ($\sim 3 \times 10^{10} M_{\odot}$), star-forming (few solar masses per year), and dusty. Very Large Array and Very Long Baseline Array observations of the persistent radio source measure a luminosity of $1.2 \times 10^{29} \text{ erg s}^{-1} \text{ Hz}^{-1}$, and show that it is extended on scales $\gtrsim 50$ mas. We associate this radio emission with the ongoing star formation activity in SDSS J050803.48+260338.0. Deeper, high-resolution optical observations are required to better utilize the milliarcsecond-scale localization of FRB 20201124A and determine the origin of the large dispersion measure (150–220 pc cm $^{-3}$) contributed by the host. SDSS J050803.48+260338.0 is an order of magnitude more massive than any galaxy or stellar system previously associated with a repeating FRB source, but is comparable to the hosts of so far non-repeating FRBs, further building the link between the two apparent populations.

Key words: galaxies: star formation – radio continuum: galaxies – radio continuum: transients – fast radio bursts.

1 INTRODUCTION

The handful of fast radio bursts (FRBs) localized to individual galaxies comprises emerging evidence for a diversity of progenitor environments. The repeating source FRB 20121102A is associated with a rapidly star-forming ($0.13\text{--}0.23 M_{\odot} \text{ yr}^{-1}$) dwarf (stellar mass of $\sim 10^8 M_{\odot}$) galaxy (Bassa et al. 2017; Tendulkar et al. 2017), and two other repeating FRBs are also associated with galaxies that form stars $>10 \times$ more rapidly than the Milky Way (Heintz et al. 2020; Marcote et al. 2020). FRB 20121102A is additionally associated with a luminous persistent radio source (Chatterjee et al. 2017; Marcote et al. 2017) of unknown origin. Although the repeating source FRB 20200120E has a comparable typical burst luminosity, repetition rate, and spectro-temporal characteristics to the repeating-FRB population (Bhardwaj et al. 2021), it is associated with an otherwise unremarkable globular cluster in the nearby galaxy M81

(Kirsten et al. 2022). FRB sources from which repetition has not been observed are located in galaxies with stellar masses ranging within $10^9\text{--}10^{11} M_{\odot}$, and star formation rates up to a few $M_{\odot} \text{ yr}^{-1}$ (Heintz et al. 2020). The FRB locations within the hosts are unremarkable (Mannings et al. 2021), although a selection of biases remain to be untangled in their interpretation (Bochenek, Ravi & Dong 2021).

The origins, life cycles, and ultimate fates of FRB progenitors remain unknown. The magnetospheres and immediate environments of neutron stars are the leading candidate FRB emission regions (e.g. Lu, Kumar & Zhang 2020; Margalit, Metzger & Sironi 2020). An FRB-like burst has been associated with an active Galactic magnetar (Bochenek et al. 2020; CHIME/FRB Collaboration et al. 2020). However, observed phenomena like a several-day FRB periodic activity cycle (CHIME/FRB Collaboration et al. 2020), the FRB source associated with a globular cluster (Kirsten et al. 2022), and the $\sim 10^{29} \text{ erg s}^{-1}$ compact persistent radio source (PRS) associated with FRB 20121102A (Chatterjee et al. 2017; Marcote et al. 2017) are difficult to reconcile with the known Galactic magnetar population (Kaspi & Beloborodov 2017). The physical properties

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of the FRB 20121102A PRS place interesting constraints on the nature of the source of the bursts (e.g. Chen, Ravi & Hallinan 2022). Attempts to jointly model FRB power sources and the emission mechanism typically invoke young magnetars with millisecond spin periods (e.g. Lu & Kumar 2018; Metzger, Margalit & Sironi 2019; Levin, Beloborodov & Bransgrove 2020; Lyubarsky 2020), or accretion from or interaction with a companion (e.g. Zhang 2018; Sridhar et al. 2021). The means by which such extreme systems form, and the lifetimes during which they produce the prodigious observed FRB rate (Ravi 2019), continue to remain mysterious. Only through the continued characterization of FRB host galaxies, and the positive or negative identification of multiwavelength and multi-time-scale counterparts, will this problem be addressed.

The repeating source FRB 20201124A was reported to be in an active state by CHIME/FRB Collaboration (2021) on 2021 March 21, with six bursts detected within five days. The dispersion measure (DM) of the source is $413.52 \pm 0.05 \text{ pc cm}^{-3}$, with a Galactic-disc contribution of between 76 (Cordes & Lazio 2002) and 109 pc cm^{-3} (Yao, Manchester & Wang 2017). Over the next few months, further repeat bursts were detected at frequencies between 700 MHz and 2 GHz by the Australian Square Kilometre Array Pathfinder (ASKAP; Kumar et al. 2021a,b), the Five hundred metre Aperture Spherical Telescope (FAST; Xu et al. 2021), the Karl G. Jansky Very Large Array *realfast* instrument (VLA/*realfast*; Law et al. 2021), the upgraded Giant Metrewave Radio Telescope (uGMRT; Wharton et al. 2021a,b), the Stockert Radio Telescope (Herrmann 2021), the Onsala Radio Telescope (Kirsten et al. 2021), and the Allen Telescope Array (Farah et al. 2021). Initial interferometric localizations with few-arcsecond accuracy were obtained within days to weeks of the CHIME/FRB announcement by the teams at ASKAP (Day et al. 2021a,b), VLA/*realfast* (Law et al. 2021), and uGMRT (Wharton et al. 2021b). A candidate host galaxy at a redshift $z = 0.098 \pm 0.002$ (Kilpatrick et al. 2021), SDSS J050803.48+260338.0, was associated with FRB 20201124A by Day et al. (2021a). A milliarcsecond-accuracy localization of FRB 20201124A by the European VLBI Network (EVN; Marcote et al. 2021), reported on 2021 May 5, established the association beyond doubt.

Reports of a PRS associated with FRB 20201124A based on 650-MHz uGMRT observations (Wharton et al. 2021a), and 3- and 9-GHz observations at the VLA (Ricci et al. 2021), sparked particular excitement, given the heretofore fruitless hunt for FRB/PRS associations since the case of FRB 20121102A. The inferred radio luminosity of $\sim 10^{29} \text{ erg s}^{-1} \text{ Hz}^{-1}$ at 1.4 GHz is consistent with both the low end of the local luminosity function of radio AGN, and with typical galaxies in the local star-forming sequence (Mauch & Sadler 2007). An absence of milliarcsecond-scale radio emission from the PRS was reported by Marcote et al. (2021), suggesting an extended emission region.

In this paper, we present a comprehensive radio and optical study of FRB 20201124A and its host galaxy. We begin in Section 2 with a summary of observations by our group, including with the VLA, the Very Long Baseline Array (VLBA), and the Palomar 200-inch Hale Telescope (P200). We then extract critical parameters of the host galaxy and interpret the nature of the associated PRS in Section 3. A discussion of the implications for the source of FRB 20201124A is presented in Section 4, and we conclude in Section 5. Throughout, we adopt cosmological parameters from Planck Collaboration VI (2020), including a Hubble constant of $H_0 = 67.4 \text{ km s}^{-1} \text{ Mpc}^{-1}$, a matter density parameter of $\Omega_m = 0.315$, and a dark energy density parameter of $\Omega_\Lambda = 0.685$.

2 OBSERVATIONS

2.1 VLA/*realfast*

We used the VLA (program code 21A-387) to observe FRB 20201124A soon after CHIME/FRB Collaboration (2021) reported activity in the source. The FRB field was observed in 10, 52-min sessions from 2021 April 5–15. The first two observing blocks were in the *L* band (1–2 GHz) and all others were observed in the *C* band (4–8 GHz). The VLA antennas were arranged in the D configuration, which has baseline lengths up to 1 km and typical spatial resolution of 46 arcsec at 1.5 GHz and 12 arcsec at 5 GHz.

We recorded visibility data with a sampling time of 5 s while commensally streaming data with 10-ms sampling time into the *realfast* transient search system (Law et al. 2018). We used *realfast* to search for FRBs in real time with a typical 1σ sensitivity of 5 mJy in 10 ms. The search resampled the data to temporal widths of 10, 20, 40, and 80 ms in both bands and included DM trials up to 1500 and 3000 pc cm^{-3} at *L* and *C* bands, respectively.

The *realfast* system detected one burst in the first observing session on 2021 April 6 (burst MJD 59311.0129359, topocentric at 2.0 GHz). The burst had a signal-to-noise ratio (S/N) of 26, a DM of $420 \pm 10 \text{ pc cm}^{-3}$, and was unresolved within the 10-ms integration time. This significance was measured in a band from 1.3–1.5 GHz that included all of the burst emission. We calibrated the data with the VLA calibration pipeline (version 2020.1) using flux calibrator 3C 147. After applying these solutions, we measure a burst fluence of $2.4 \pm 0.1 \text{ Jy ms}$ and position (J2000 epoch) RA = $05^{\text{h}}08^{\text{m}}03^{\text{s}}.50$, Dec. = $26^{\circ}03'37.71$. The source size is similar to the synthesized beam size of $55 \times 50 \text{ arcsec}^2$ at a position angle of -2° at 1.4 GHz. The statistical position uncertainty is 0.8 arcsec, while the total error is 1.9 arcsec (systematics dominated, see below).

The standard (slow) visibility data were imaged to search for persistent emission at the FRB location and estimate systematic position errors. We imaged a single observing session of 31 min at *L* and *C* bands with usable bandwidths of approximately 400 and 1500 MHz, respectively. Systematic source position error is typically dominated by antenna phase calibration errors that shift the centroid for all sources in the image. We estimate this effect by cross-matching sources in the deep radio image to the PanSTARRS1 (PS1) catalogue (Chambers et al. 2016). At *L* band, the confusion limited image does not have enough sources to perform a useful cross-match. At the *C* band, we find five cross matched sources within 2 arcsec and estimate a systematic error of 0.5 arcsec. Assuming this error scales with synthesized beam size, we estimate an *L*-band systematic error of 1.7 arcsec.

At the *L* band, the naturally weighted image shows no source brighter than $500 \mu\text{Jy}$ at the FRB location. The noise level of the image is roughly $170 \mu\text{Jy}$, which is consistent with the expected confusion limit of this resolution at 1.4 GHz (Condon et al. 2012). The *C*-band image has a compact radio source with a peak flux density of $221 \pm 15 \mu\text{Jy}$ at position (J2000 epoch) RA = $05^{\text{h}}08^{\text{m}}03^{\text{s}}.45$, Dec. = $26^{\circ}03'38.00$. The statistical uncertainty in the position is 0.2 arcsec, but the total error is dominated by systematic effects and is 0.5 arcsec. The source size is $11.6 \times 10.8 \text{ arcsec}^2$ at a position angle of 0° at 5 GHz, which is consistent with the synthesized beam size. Hereafter we denote this source PRS 201124.

Fig. 1 shows the persistent *L*-band burst location and *C*-band emission overlaid on a PS1 image of the region. The burst and persistent emission are coincident with each other and the galaxy SDSS J050803.48+260338.0 (hereafter J0508+2603). No other optical source is consistent with the radio source locations. We use the

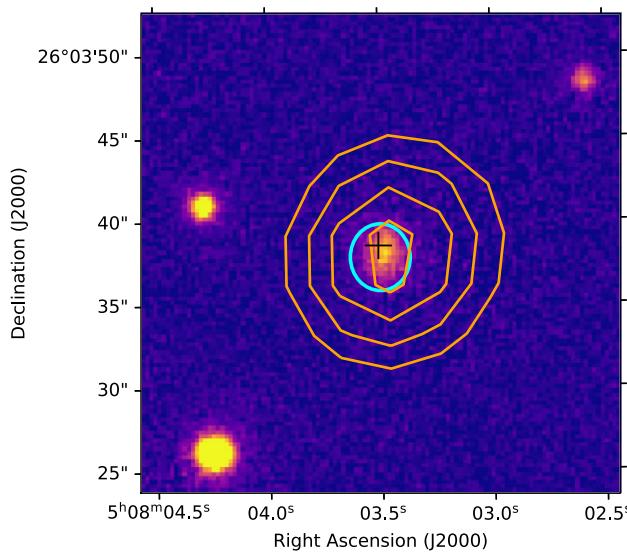


Figure 1. Localization of FRB 20201124A. Background: stacked *i*-band image from the PanSTARRS Data Release 1 (Chambers et al. 2016). Orange contours: VLA C-band image of the field, with contour levels of 6σ , 9σ , 12σ , and 15σ , where $\sigma = 13 \mu\text{Jy beam}^{-1}$ is the image noise level. Blue ellipse: VLA/realfast localization error region of a burst from FRB 20201124A. The dimensions represent the 1σ uncertainties in right ascension and declination. Black cross: representation of the EVN position of bursts from FRB 20201124A (Marcote et al. 2021). The position error is much smaller than the arms of the cross.

burst localization precision and PS1 catalogue to calculate the chance of association using the `astropath` Bayesian inference framework (Aggarwal et al. 2021). We estimated the association probability by assuming several built-in priors (e.g. host offset distributions of ‘exponential’, ‘core’, or ‘uniform’ and host magnitudes of ‘inverse’) and always find greater than 95 per cent chance of association. Consistent with previous analyses (Marcote et al. 2021; Day et al. 2021a), we conclude that J0508+2603 is the host galaxy of FRB 20201124A.

2.2 VLBA L-band continuum

We observed the location of FRB 20201124A with the VLBA for 6 h between 2021 April 08 20:00 and 2021 April 09 02:00 UTC. Data were recorded using the Digital Downconverter personalities of the ROACH Digital Backends with all 10 antennas in four 64-MHz-wide spectral windows centred on 1407.75, 1471.75, 1663.75, and 1727.75 MHz, each with 128 channels. Data were recorded in dual circular polarizations with a data rate of 2.048 Gbps at each station. The observations were phase referenced, with 71 scans of 210-s duration on the target interspersed with 45-s scans on the phase reference source J0500+2651 1°.88 distant, for which we adopt a position (International Celestial Reference Frame 3; Charlot et al. 2020) of RA = $05^{\text{h}}00^{\text{m}}27\text{s}87019156$, Dec. = $+26^{\circ}51'34''3393223$ (J2000). Observations of the bandpass calibrator 3C 84 (two 3-min scans) and a check source J0502+2516 (two 45-s scans) were also performed. The data were correlated at the NRAO Array Operations Center with the DIFX software correlator (Deller et al. 2011), with a target phase centre of RA = $05^{\text{h}}08^{\text{m}}03\text{s}.50$, Dec. = $+26^{\circ}03'37''.8$ (J2000).

We analysed the data using both CASA (version 5.6.1-8) and AIPS (Greisen 2003), and obtained similar results. After data editing to excise radio-frequency interference, we derived initial bandpass solutions using 3C 84, and complex gain solutions using global

Table 1. Summary of VLBA imaging results on FRB 20201124A.

uv taper (klambda)	Beam FWHM (mas)	RMS (mJy beam $^{-1}$)
None	10.2×4.9	0.04
5000	19.5×18.2	0.048
2000	41.4×39.2	0.062
1200	86.7×48.7	0.116

fringe fitting together with a single round of self-calibration on J0500+2651. We then applied the solutions to data on the target and the check source J0502+2516. Imaging was performed only in CASA using the `tclean` task, with the Högbom (1974) deconvolver and robust weighting with a robustness parameter of 0.5. For the target, we present results from images made with a phase centre corresponding to the EVN position of FRB 20201124A (Marcote et al. 2021), over a 1.024×1.024 arcsec 2 region.

No persistent emission was detected with the VLBA in any images made towards FRB 20201124A. Further, no emission was detected towards the centre of light of J0508+2603. This result is consistent with the previous EVN observations (Marcote et al. 2021), which concluded that PRS 201124 is extended beyond milliarcsecond scales. We made a series of images with different Gaussian tapers in the *uv* plane, as summarized in Table 1. The listed tapers correspond to the `tclean` ‘*uv taper*’ parameter, and we also list the full width at half-maximum (FWHM) of the synthesized beam and image noise rms in each case. The *uv taper* parameter controls the width of a multiplicative Gaussian taper applied to the gridded visibilities to downweight data on longer baselines. In all cases, the minimum projected baseline length was 210 km, corresponding to a largest angular scale of ~ 180 mas. In Fig. 2, we show images of the target, phase reference, and check source made with each of the *uv* tapers; note that all images were made with exactly the same calibrations applied to the data.

In summary, we place upper limits on the presence of a compact radio source associated with FRB 20201124A. The 3σ limit ranges from 0.12 and 0.2 mJy for angular scales from 5 to 50 mas (see Table 1 for more detail). These flux-density limits correspond to radio luminosities of $(3\text{--}9) \times 10^{28} \text{ erg s}^{-1} \text{ Hz}^{-1}$ at the distance of the FRB 20201124A host J0508+2603.

2.3 Palomar P200/DBSP

We observed the host galaxy of FRB 20201124A, J0508+2603, with the P200 Double Spectrograph (DBSP; Oke & Gunn 1982) on 2021 April 10 UT. Conditions were clear, with 1.2-arcsec seeing. We obtained a low-resolution ($R \sim 1000$) spectrum of J0508+2603 using the 600/4000 grating on the blue arm (central wavelength 4400 Å), and the 316/7500 grating on the red arm (central wavelength 7500 ÅÅ). A 1.5-arcsec slit was placed on the target, centred on the position RA = $05^{\text{h}}08^{\text{m}}03\text{s}.48$, Dec. = $+26^{\circ}03'38''.0$ (J2000) at a position angle of 304°. Two 1200-s exposures were obtained with the blue arm, and three 800-s exposures were obtained with the red arm, at a mean airmass of 1.5. The data were bias-subtracted, flat-fielded, cleaned of cosmic rays, wavelength calibrated using comparison-lamp spectra, sky-line subtracted, and optimally extracted using standard techniques implemented in a custom DBSP pipeline.¹ Flux calibration and telluric-line correction was performed with observations of the standard star Feige 34 at a comparable airmass.

¹https://github.com/finagle29/DBSP_DRP.

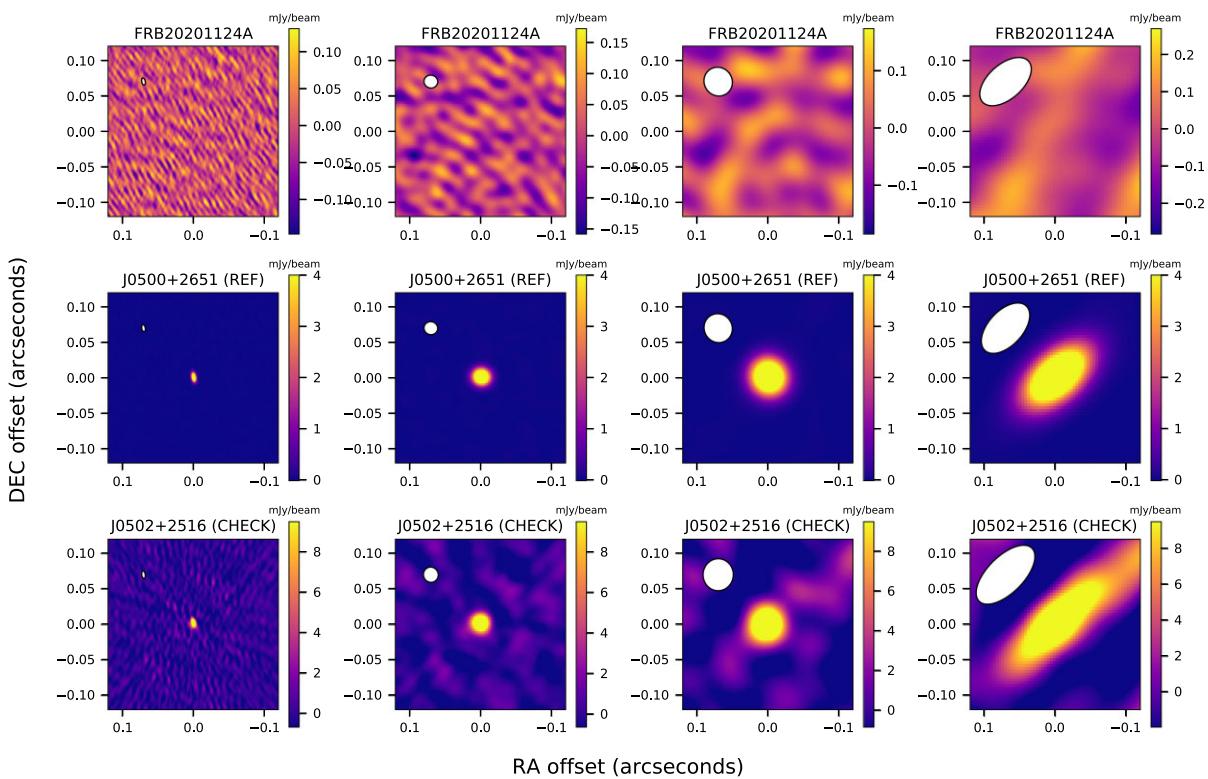


Figure 2. VLBA images of FRB 20201124A (top row) at different angular scales, with images of the phase-reference (middle row) and independent check (bottom row) sources shown for comparison. From the left- to right-hand side, images were made with no uv taper, and uv tapers of 5000, 2000, and 1200 klambda respectively (see Table 1 for more details). Note that only two 45-s scans were obtained on check source.

The spectrum of J0508+2603, shown in Fig. 3, exhibits a few clear emission-line features, including the $H\alpha/[N\,II]$ complex, and the $[S\,II]\lambda 6718, 6733$ doublet. A tentative indication of $H\beta$ is evident, although we do not claim a detection. After normalizing by the continuum and correcting for Galactic extinction of $A_V = 2.024$ (Schlafly & Finkbeiner 2011) assuming a Fitzpatrick & Massa (2007) extinction law, we measure line ratios of $\log([N\,II]/H\alpha) = -0.40 \pm 0.01$ and $\log([S\,II]/H\alpha) = -0.48 \pm 0.02$. Based on non-detections of the $H\beta$ and $[O\,II]\lambda 5007$ lines, we derive a 95 per cent confidence upper limit on their ratio of $\log([O\,III]/H\beta) < 0.6$. According to the diagnostics of Kewley et al. (2006), these ratios are consistent with a softer source of ionizing radiation corresponding to young stars, i.e. corresponding to ongoing star formation.

We caution that the above analysis of the P200/DBSP observations is not fully representative of the J0508+2603. The spectrum does not capture all the light from the galaxy, and we have applied no corrections for slit losses. We have also not attempted to model absorption features at the $H\alpha$ and $H\beta$ wavelengths corresponding to stellar photospheres. Finally, we do not spatially resolve the galaxy in the spectroscopic observations.

3 INTERPRETATION

3.1 The optical/infrared SED and spectrum of J0508+2603

We collated archival photometric observations of J0508+2603 to perform stellar population synthesis modelling of its spectral energy distribution (SED). J0508+2603 appears in all bands in the Sloan Digital Sky Survey (SDSS) Data Release 16 (DR16; Ahumada et al. 2020), the Two Micron All Sky Survey (2MASS; Skrutskie et al.

2006), and in the ALLWISE data release (Cutri et al. 2021) from the *Wide-field Infrared Survey Explorer* (WISE). These data, together with results from the modelling described below, are shown in Fig. 4. We discarded the u -band SDSS DR16 catalogue measurement of 23.09 ± 0.55 mag, as no source was evident in a visual inspection of the image. We also did not attempt to model thermal emission in the WISE $w3$ and $w4$ bands, given that model uncertainties in these emission bands are typically large and hard to quantify (e.g. Leja et al. 2017; Bellstedt et al. 2020).

We used the PROSPECTOR (Leja et al. 2017; Johnson et al. 2021) stellar population inference code to model the SED of J0508+2603. PROSPECTOR enables efficient sampling of the posterior distributions of model parameters that describe the stellar populations of galaxies. The forward model is built on the Flexible Stellar Population Synthesis code (Conroy, Gunn & White 2009). We used a standard ‘delay-tau’ parametric star formation history, and sampled from the posterior using EMCEE (Foreman-Mackey et al. 2013a). A model for dust attenuation and re-radiation was also included. Priors used included a log-uniform prior on the mass in formed stars (hereafter the stellar mass) of between 10^7 and $10^{13} M_\odot$, a top-hat prior on the internal dust extinction (A_V) of 0–3 mag, a top-hat prior on the age of the stellar population of between 0.001 and 13.8 Gyr, a log-uniform prior on the star formation time-scale of between 0.001–30 Gyr, and a top-hat prior on the ratio of the metallicity to the solar metallicity ($\log Z_{\text{sol}}$) of between –2 and 0.2. We found $\log(M/M_\odot) = 10.62^{+0.07}_{-0.06}$, $\log Z_{\text{sol}} = -0.9 \pm 0.2$, and $A_V = 1.5 \pm 0.2$ mag. This metallicity is remarkably low for such a massive galaxy, as evidenced by the position of this galaxy 0.8 dex below the mass-metallicity relation in the local Universe (Curti et al. 2020). Furthermore, the large ratio of $\log([N\,II]/H\alpha)$ is consistent with approximately solar metallicity (Pettini & Pagel 2004). We therefore

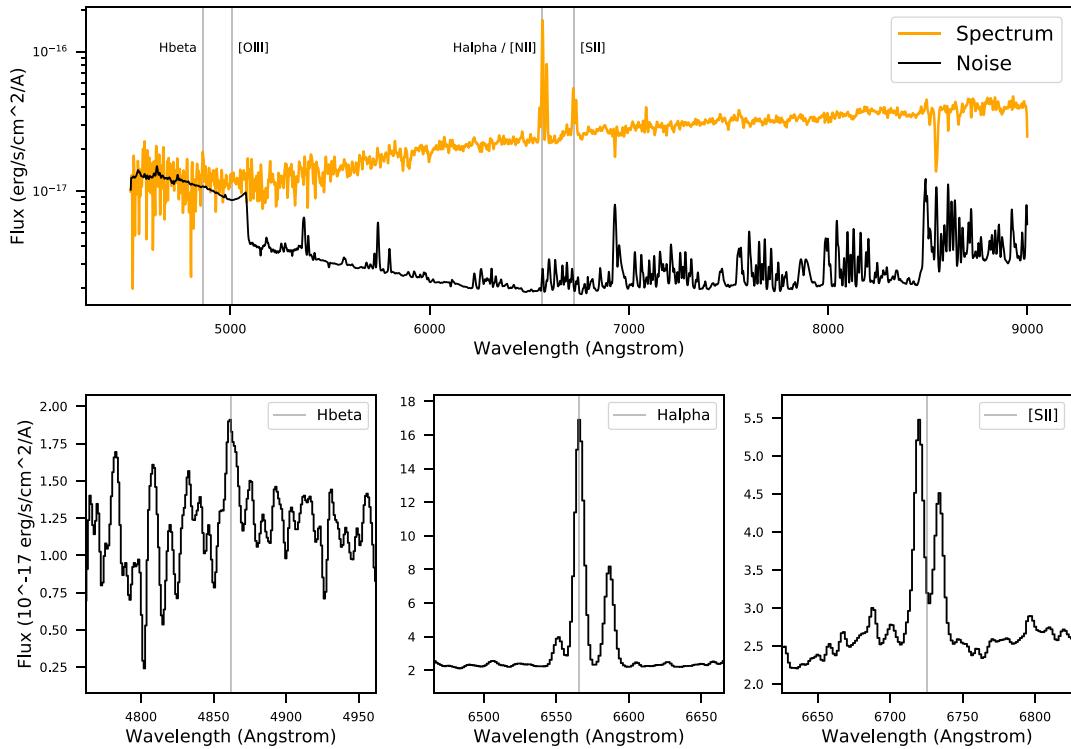


Figure 3. Palomar P200/DBSP spectrum of SDSS J050803.48+260338.0 (J0508+2603). Top panel: flux calibrated spectrum (gold) and 1σ noise spectrum (black), with the locations of a few important spectral lines shown. Bottom panels, from the left- to right-hand side: spectra centred on the $H\beta$, $H\alpha$, and $[SII]\lambda 6718, 6733$ lines. For better visualization, all spectra are smoothed by a Gaussian with a standard deviation of 1.5 pixels. No extinction correction is applied to the data in this figure.

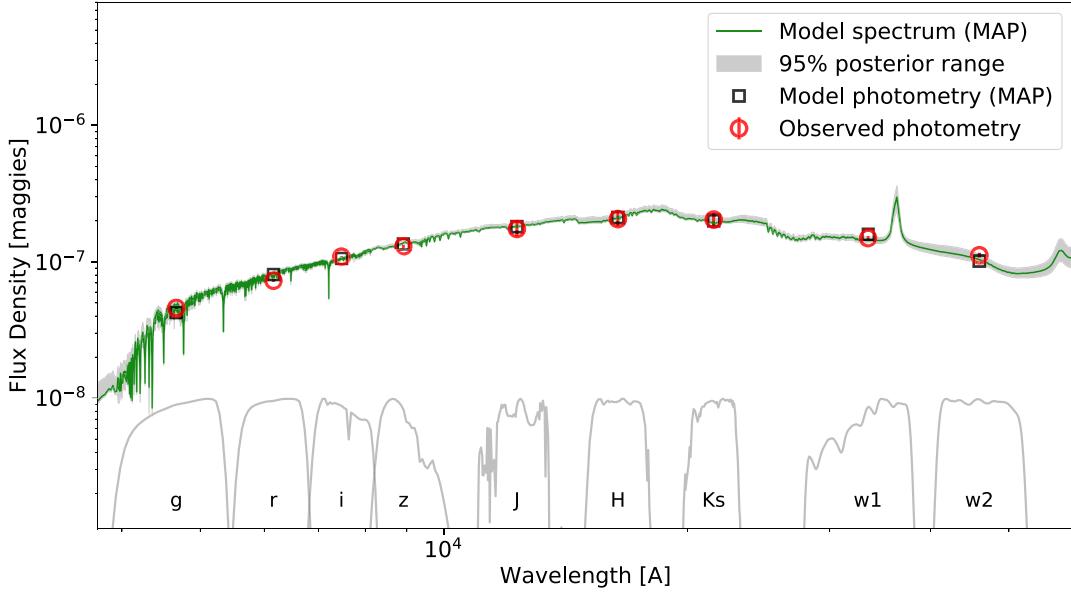


Figure 4. Observed and modelled spectral energy distribution (SED) of SDSS J050803.48+260338.0 (J0508+2603). Observed flux densities in various filters (representative transmission curves shown in grey), with 1σ error bars, are shown as red circles. The maximum a posteriori probability (MAP) model SED derived from PROSPECTOR is shown as a green curve, and the corresponding model photometry in all filters is shown as black squares. The grey shading indicates the 95 per cent confidence range in the modelled SED at every wavelength. Note that no nebular emission was included in the model. All points and curves are corrected for extinction in the Milky Way interstellar medium.

re-ran the PROSPECTOR model with a Gaussian prior on $\log Z_{\text{sol}}$ with mean 0.012 and standard deviation 0.205, based on the (Pettini & Pagel 2004) relation between $\log([N\text{II}]/H\alpha)$ and metallicity. This yielded $\log(M/M_{\odot}) = 10.48^{+0.03}_{-0.05}$, $\log Z_{\text{sol}} = -0.3^{+0.2}_{-0.3}$ and A_V

$= 1.3 \pm 0.2$. We adopt these latter values as representative of J0508+2603.

The star formation history was less well constrained by the data, with highly correlated posterior distributions of the delay and

time-scale parameters. We explored non-parametric models for the star formation history within PROSPECTOR (Leja et al. 2019), and found similar issues. However, it is clear that the galaxy is young, with a 90 per cent confidence upper limit on the delay parameter in the delay-tau model of 3 Gyr; this is also evidenced by the low metallicity (Gallazzi et al. 2005). The maximum a posteriori probability model, together with an indication of the range of possible models, is shown in Fig. 4.

With the measurement of the internal dust extinction, an estimate of the star formation rate (SFR) corresponding to the H α luminosity is possible. Combining Galactic and internal extinction, and adopting a luminosity distance of 467.3 Mpc, the P200/DBSP spectrum implies an H α luminosity of $L_{\text{H}\alpha} = (3.4 \pm 0.8) \times 10^{41} \text{ erg s}^{-1}$. Assuming a Balmer decrement of 3, the predicted H β flux is $4.3 \times 10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$, which is just approximately three times higher than the noise level in the P200/DBSP spectrum. This is consistent with the observed spectrum (Fig. 3). For consistency with the work of Heintz et al. (2020), we adopt their conversion between $L_{\text{H}\alpha}$ and SFR to derive an SFR of $\sim 1.7 \text{ M}_\odot \text{ yr}^{-1}$, with an uncertainty of ~ 35 per cent (including uncertainty in the $L_{\text{H}\alpha}$ –SFR relation). The specific SFR of J0508+2603 is thus approximately eight times that of the Milky Way. We emphasize that, as discussed in Section 2.3, our H α -based SFR is likely to be a lower limit.

A picture thus emerges of J0508+2603 as a young galaxy just a factor of a few less massive than the Milky Way. A high specific SFR is observed, together with significant internal dust extinction that partially attenuates the observed star formation. The internal extinction is indeed larger than in 90 per cent of SDSS galaxies with similar H α luminosities (Xiao et al. 2012). Better constraints on the metallicity would be derived with more complete line-ratio diagnostics from a deeper spectrum.

3.2 PRS 201124 represents star formation activity

The SED of PRS 201124, based on our VLA observations and those of Ricci et al. (2021) and Wharton et al. (2021a), is shown in Fig. 5. Our observations, together with those from the EVN (Marcote et al. 2021), clearly demonstrate that PRS 201124 is extended. For example, no emission is observed in our VLBA images on scales of $\lesssim 50$ mas (94 pc at the distance of J0508+2603), with a 3σ upper limit of 0.2 mJy or $5 \times 10^{28} \text{ erg s}^{-1} \text{ Hz}^{-1}$. Interpolating between the 650-MHz and 3-GHz observations of PRS 201124 (Ricci et al. 2021; Wharton et al. 2021a), the implied total flux density of PRS 201124 at 1567.75 MHz (the mid-point of the VLBA band) is ~ 0.46 mJy, or $1.2 \times 10^{29} \text{ erg s}^{-1} \text{ Hz}^{-1}$. For comparison, the only previously reported PRS associated with an FRB (FRB 20121102A) had a 1.77-GHz luminosity on milliarcsecond scales of $2 \times 10^{29} \text{ erg s}^{-1} \text{ Hz}^{-1}$ (Marcote et al. 2017). We conclude that a PRS, like that associated with FRB 20121102A, is not present at the location of FRB 20201124A.

On the other hand, the radio luminosity of PRS 201124 is consistent with the observed SFR of J0508+2603. The SFR derived from the H α luminosity implies a 1.4-GHz radio luminosity of $L_{1.4 \text{ GHz}} \sim 3 \times 10^{28} \text{ erg s}^{-1} \text{ Hz}^{-1}$ from the relation of Bell (2003). The relation of Murphy et al. (2011), which includes a factor that depends on the linear extent of the emission (D), implies $L_{1.4 \text{ GHz}} \sim 2 \times 10^{28} [D/(3 \text{ kpc})]^{0.75 \pm 0.11} \text{ erg s}^{-1} \text{ Hz}^{-1}$ with an uncertainty of ~ 1 dex. For the reasons discussed in Section 2.3, it is likely that the H α luminosity of J0508+2603 underestimates the total SFR. Furthermore, radio-derived SFRs are averaged over ~ 100 Myr (Condon 1992), whereas the H α luminosity represents more recent star formation. If the SFR is declining with time, the radio-derived

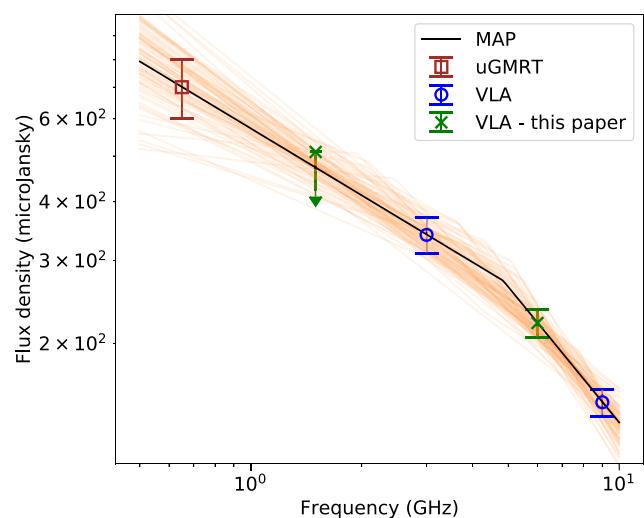


Figure 5. Radio spectrum of PRS 201124. The points indicate measurements presented here (green cross), and additional measurements from the uGMRT (brown square; Wharton et al. 2021a) and VLA (blue circles; Ricci et al. 2021). We fit a four-parameter broken power-law model to the four measurements using the EMCEE package (Foreman-Mackey et al. 2013b), and the MAP model is shown as a black line. Models corresponding to 100 random draws from the posterior are shown as orange curves.

SFR will exceed that derived from $L_{\text{H}\alpha}$. It is therefore not surprising that the observed luminosity of PRS 201124 mildly exceeds the predicted luminosity from the H α -derived SFR. The luminosity of PRS 201124 may indeed represent a more accurate estimate of the SFR, i.e. of $\sim 7 \text{ M}_\odot \text{ yr}^{-1}$.

The radio SED of PRS 201124 is also consistent with typical star-forming radio galaxies (Klein, Lisenfeld & Verley 2018; Tisanić et al. 2019). We model the SED (Fig. 5) using a broken power-law function (equation 3 of Tisanić et al. 2019), with distinct spectral indices α and β , respectively, below and above a break frequency ν_b :

$$\log F_\nu = \begin{cases} -\alpha \log(\nu/\nu_b) + F_0, & \text{if } \nu \leq \nu_b \\ -\beta \log(\nu/\nu_b) + F_0, & \text{if } \nu > \nu_b \end{cases} \quad (1)$$

Here, F_ν is the flux density, ν is the frequency, and F_0 is a reference flux density. Broken power laws are expected in the case of significant cooling of the synchrotron-emitting electrons. We use EMCEE to fit this four-parameter model to the four data points, and find $\alpha = 0.48^{+0.1}_{-0.18}$, $\beta = 0.96^{+0.29}_{-0.16}$, and $\nu_b = 5.0^{+2.2}_{-2.4} \text{ GHz}$.² Broken power-law spectra like this are observed in several nearby star-forming galaxies (Klein et al. 2018). Using a large sample of galaxies with SFRs at redshifts $0.3 < z < 4$ in excess of $100 \text{ M}_\odot \text{ yr}^{-1}$ from the VLA-COSMOS survey, Tisanić et al. (2019) find $\alpha = 0.53 \pm 0.04$, $\beta = 0.94 \pm 0.06$, and $\nu_b = 4.3 \pm 0.6 \text{ GHz}$. Our measurements are in good agreement with this sample of highly star-forming galaxies.

The extended morphology, luminosity and SED of PRS 201124 all correspond to the star formation activity inferred in J0508+2603. As discussed in Section 2.3, the emission-line ratios in our P200/DBSP spectrum of J0508+2603 show no evidence for an AGN, further establishing this correspondence. At the highest angular resolution of our VLBA observations (see Table 1), we place a 3σ limit on the 1567.75-MHz luminosity of a compact PRS associated with FRB 20201124A of $3 \times 10^{28} \text{ erg s}^{-1} \text{ Hz}^{-1}$, nearly an order of

²Although the ‘fit’ is deterministic, the use of EMCEE was to explore the uncertainties in the parameter estimates.

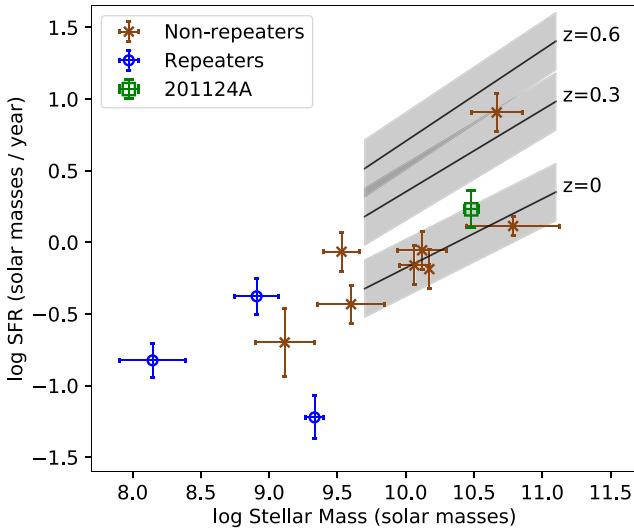


Figure 6. Stellar masses and SFRs of FRB host galaxies. The background sample of repeating and so far non-repeating FRB host data are from Heintz et al. (2020) (see the text for details on excluded FRBs), with the exception of FRB 190523 for which we use the [O II]-derived SFR (Ravi et al. 2019) of $1.3 \pm 0.2 \text{ M}_\odot \text{ yr}^{-1}$. The three black lines indicate fits to the star-forming main sequence of galaxies at different redshifts; the shaded regions indicate the 1 σ intrinsic scatter (Speagle et al. 2014).

magnitude below the 1.77-GHz luminosity of the PRS associated with FRB 20121102A (Marcote et al. 2017).

4 DISCUSSION

The host galaxy of FRB 20201124A, J0508+2603, is unremarkable within the diverse range of FRB host galaxies (Heintz et al. 2020). Among the sample in hand, stellar masses range from 10^8 – 10^{11} M_\odot , and star formation is evident in all but two cases (Ravi et al. 2019; Kirsten et al. 2022). Specific SFRs range from an order of magnitude below that of the Milky Way, to three orders of magnitude above that of the Milky Way. Metallicities range from sub-solar in the case of FRB 20121102A ($\log Z_{\text{sol}} < -0.58$; Tendulkar et al. 2017), to approximately solar in the remainder of the sample in hand. J0508+2603 is among the more massive FRB host galaxies, and the most massive host of an FRB that is observed to repeat (Fig. 6).³ We estimate a DM contributed by the host of 150 – 220 pc cm^{-3} , assuming a fiducial fraction of cosmic baryons in the intergalactic medium of 0.7 (Shull & Danforth 2018), a Milky Way halo contribution of between 50 and 80 pc cm^{-3} (Prochaska & Zheng 2019), and a range of Milky Way disc DM contributions between the Cordes & Lazio (2002) and Yao et al. (2017) models. This is larger than is observed in most FRBs localized to host galaxies (Macquart et al. 2020), and consistent with the host DM inferred for FRB 20121102A (Tendulkar et al. 2017). Further interpretation of this result will require high spatial resolution optical/infrared (OIR) observations of J0508+2603 to estimate the possible path-length of the FRB through the host interstellar medium, which will enable any egregious local DM contributions to be identified.

As just the fifth reported host system of a repeating FRB, our analysis of J0508+2603 further demonstrates the wide range of

³In making this statement, we do not consider M81 as the true host of FRB 20200120E, but rather associate it directly with its host globular cluster (Kirsten et al. 2022).

possible hosts of active FRB sources. Indeed, the J0508+2603 is at the more massive end of the star-forming main sequence of galaxies, consistent with the hosts of so far non-repeating FRBs but distinct from the remainder of the repeating-FRB host sample (Fig. 6). This empirical fact further strengthens the link between the sources of actively repeating and so far non-repeating FRBs (Ravi 2019; James et al. 2020). We caution, however, against detailed statistical inference from the data in Fig. 6. The star formation rates and stellar masses are measured with different methods, with for example different amounts of the host-galaxy light included in the measurements of H α luminosities. Sporadic attempts have been made to separate nuclear activity from star formation activity. Some FRBs have been excluded from Fig. 6 because of insecure host associations (FRB 190611; Heintz et al. 2020), incomplete data (FRB 190614; Law et al. 2020), and low luminosity (FRBs 200428 and 20200120E; Bochenek et al. 2020; Kirsten et al. 2022).

The stellar mass and SFR of J0508+2603 provide some insight into the source of FRB 20201124A. We can compare these properties to the samples of core-collapse supernovae (CCSNe), superluminous supernovae (SLSNe), and long gamma-ray bursts (LGRBs), assembled by Taggart & Perley (2021). These authors corrected the distributions of stellar mass and SFR for cosmic evolution, which is critical in comparing them to the local-Universe galaxy distribution (see also Bochenek et al. 2021). Only two SLSNe out of 53 (one of each of types I and II) and no LGRBs (out of 17) within the Taggart & Perley (2021) sample are found in galaxies as massive as J0508+2603. On the other hand, the properties of J0508+2603 place it between the 80th and 90th percentiles of the stellar-mass and SFR distributions of the hosts of CCSNe. Formation channels for the FRB 20201124A source associated with stellar evolutionary pathways associated with LGRBs and SLSNe are therefore not implicated by the properties of J0508+2603, consistent with previous findings based on the FRB host-galaxy population (Bochenek et al. 2021).

We anticipate that more detailed observations of J0508+2603 will yield significant further insights. In particular, space-based OIR observations with high angular resolution will enable a direct comparison with observations of the hosts of other repeating FRBs. Specifically, in combination with the EVN localization of FRB 20201124A (Marcote et al. 2021), this will enable the local environment of the FRB source to be placed in the context of the host-galaxy structure, and allow the host DM to be better interpreted (Tendulkar et al. 2021). A deeper, possibly spatially resolved OIR spectrum of J0508+2603 will enable more accurate metallicity and SFR measurements, in turn providing more robust measurements of the stellar mass and star formation history. Finally, radio observations of J0508+2603/PRS 201124 with 0.1–1 arcsec resolution will likely resolve the morphology of the PRS, better establishing its nature and link with FRB 20201124A.

5 CONCLUSIONS

We present a study of the host galaxy and persistent radio counterpart of the repeating FRB 20201124A. We conclude the following:

- (i) FRB 20201124A is associated with a galaxy (J0508+2603) at $z = 0.098$. J0508+2603 has a stellar mass of $\log(M/\text{M}_\odot) = 10.48^{+0.03}_{-0.05}$, a ratio between its metallicity and the solar metallicity of $\log Z_{\text{sol}} = -0.3^{+0.2}_{-0.3}$, and a significant internal dust extinction of $A_V = 1.3 \pm 0.2$. Based on an extinction-corrected H α luminosity of $L_{\text{H}\alpha} = (3.4 \pm 0.8) \times 10^{41} \text{ erg s}^{-1}$ (not corrected for slit losses or

the presence of absorption in the stellar continuum), we derive a star formation rate of $\sim 1.7 \text{ M}_\odot \text{ yr}^{-1}$ (35 per cent 1σ uncertainty).

(ii) We find that the persistent radio source (PRS 201124) associated with FRB 20201124A is extended on scales $\gtrsim 50 \text{ mas}$ (94 pc at the distance of J0508+2603). Considering the spatial extent, luminosity, and SED of PRS 201124, the persistent emission is consistent with non-thermal emission caused by the observed ongoing star formation activity. We place an upper bound on the luminosity of a compact ($\lesssim 10 \text{ pc}$) PRS at the location of FRB 20201124A of $3 \times 10^{28} \text{ erg s}^{-1} \text{ Hz}^{-1}$. Future searches for PRSs at the locations of FRBs should be careful to exclude radio sources originating in star formation, in addition to nuclear activity.

(iii) The host galaxy of FRB 20201124A, J0508+2603, is more massive (by an order of magnitude) than any of the three previous known host galaxy of a repeating FRB, but has a comparable stellar mass and SFR to known hosts of so far non-repeating FRBs. This provides further, weak, evidence for commonality between the sources of repeating and so far non-repeating FRBs.

(iv) The stellar mass of J0508+2603 is much larger than the typical host galaxies of SLSNe and LGRBs, but together with the SFR is consistent with CCSNe host galaxies.

More detailed studies of J0508+2603 and PRS 201124, with higher sensitivity and a wider range of angular resolutions in the radio and OIR bands, are required to better interpret the exquisite data in hand on FRB 20201124A. Larger samples of localized FRBs with systematic host-galaxy studies will continue to refine our understanding of the origins of FRBs.

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DATA AVAILABILITY

Original data presented herein include VLA, VLBA, and P200/DBSP observations. The VLA and VLBA primary data products are archived at the NRAO Data Archive (<https://archive.nrao.edu/archive/advquery.jsp>). The VLA secondary (*realfast*) data products are will soon be available through the NRAO archive, but are currently available on request. The raw data obtained during the P200/DBSP observations are available upon request. All data that shown in figures are available upon request.

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