



The S-Web Origin of Composition Enhancement in the Slow-to-moderate Speed Solar Wind

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Abstract

Connecting the solar wind observed throughout the heliosphere to its origins in the solar corona is one of the central aims of heliophysics. The variability in the magnetic field, bulk plasma, and heavy ion composition properties of the slow wind are thought to result from magnetic reconnection processes in the solar corona. We identify regions of enhanced variability and composition in the solar wind from 2003 April 15 to May 13 (Carrington Rotation 2002), observed by the Wind and Advanced Composition Explorer spacecraft, and demonstrate their relationship to the separatrix–web (hereafter, S-Web) structures describing the corona’s large-scale magnetic topology. There are four pseudostreamer (PS) wind intervals and two helmet streamer (HS) heliospheric current sheet/plasma sheet crossings (and an interplanetary coronal mass ejection), which all exhibit enhanced alpha-to-proton ratios and/or elevated ionic charge states of carbon, oxygen, and iron. We apply the magnetic helicity–partial variance of increments (H_m –PVI) procedure to identify coherent magnetic structures and quantify their properties during each interval. The mean duration of these structures are ~ 1 hr in both the HS and PS wind. We find a modest enhancement above the power-law fit to the PVI waiting-time distribution in the HS-associated wind at the 1.5–2 hr timescales that is absent from the PS intervals. We discuss our results in the context of previous observations of the ~ 90 minutes periodic density structures in the slow solar wind, further development of the dynamic S-Web model, and future Parker Solar Probe and Solar Orbiter joint observational campaigns.

Unified Astronomy Thesaurus concepts: Slow solar wind (1873); Solar magnetic reconnection (1504); Space plasmas (1544); Interplanetary turbulence (830); Solar physics (1476); Solar magnetic fields (1503); Heliosphere (711)

1. Introduction

The global magnetic geometry of the solar corona directly determines the structure of the solar wind outflow (e.g., Zirker 1977; Axford et al. 1999; Antiochos et al. 2007, 2011; Cranmer 2012). Decades of in situ observations have shown that the heliospheric structure and solar wind properties reflect the coronal magnetic structure of its origin (Zurbuchen 2007; Zhao et al. 2014). During solar minimum, polar coronal holes are correlated with fast, tenuous solar wind (Geiss et al. 1995; McComas et al. 2002), while the helmet streamer (HS) belt and the heliospheric current sheet (HCS) are associated with slower, denser, and more variable solar wind (Gosling 1997; McComas et al. 1998b; Zurbuchen et al. 2002; Zhao et al. 2009). During solar maximum the HS belt is highly warped, and pseudostreamer (PS) coronal structures often make a significant contribution to the solar wind in the ecliptic plane (Riley & Luhmann 2012).

Whereas the large-scale closed-flux system of the HS belt separates open fields of opposite polarity, thus giving rise to the HCS, coronal PSs (sometimes called unipolar streamers) are closed-flux regions surrounded by open fields of a single polarity (e.g., Wang et al. 2007; Titov et al. 2012; Wang et al. 2012; Rachmeler et al. 2014; Wang & Panasenco 2019; Mason

et al. 2021). Solar wind originating from coronal PSs tends to be more similar to the dense, variable HS slow wind than to the fast wind from coronal holes (Crooker et al. 2012), but observations have established the existence of a continuum of states between the nominal fast and slow wind rather than a well-separated bimodal distribution (e.g., Stakhiv et al. 2015, 2016).

Connecting the solar wind to its source region of origin has become one of the central aims of heliophysics in order to test and constrain different theories of solar wind formation (Viall & Borovsky 2020). Additionally, accurate space weather prediction requires an understanding of the different solar wind streams in the heliosphere and where they were formed, e.g., mesoscale structures are known to drive magnetospheric dynamics (Viall et al. 2021). Therefore, establishing this solar–heliospheric connection is one of the fundamental science objectives of the Parker Solar Probe (PSP; Fox et al. 2016) and Solar Orbiter (Müller et al. 2020) missions.

White-light coronagraph and heliospheric imaging data have shown that the solar wind originating from the HS stalks includes a continual, intermittent outflow of intensity enhancements, called “streamer blobs,” that trace the bulk outflow of the slow solar wind (Sheeley et al. 1997, 1999, 2009; Rouillard et al. 2010a, 2010b; Sanchez-Diaz et al. 2017b). While the basic theory of steady-state, slow solar wind from the vicinity of coronal streamers and PSs is well-established (e.g., Arge & Pizzo 2000; Lepri et al. 2008; Riley & Luhmann 2012, and references therein), this steady-state picture is difficult to



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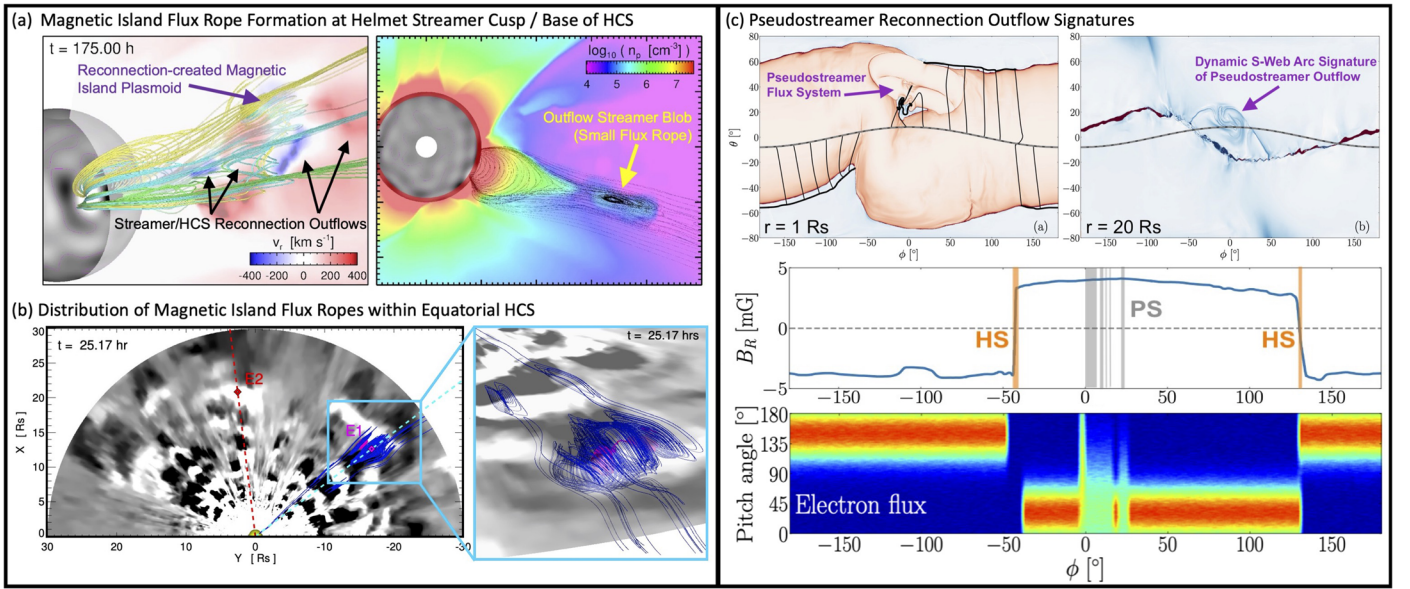


Figure 1. Reconnection mechanisms for generating intermittent outflow of dense, closed-field plasma in the slow-to-moderate speed solar wind from helmet streamers (HSs) and pseudostreamers (PSs). (a) ARMS simulation of HS blob pinch-off reconnection (adapted from Lynch 2020), and (b) the small flux rope/reconnection plasmoid structures of the heliospheric current sheet (adapted from Higginson & Lynch 2018). (c) ARMS simulation of interchange reconnection outflow from a PS and a synthetic proxy for suprathermal electron pitch angle based on magnetic connectivity (adapted from Aslanyan et al. 2022).

reconcile with the observed slow wind variability in both remote-sensing and in situ observations that likely require a time-varying magnetic reconnection component.

Demonstrating another example of solar wind variability, Kepko et al. (2020) analyzed 25 yr of solar wind data, expanding on the initial study of Viall et al. (2008), finding that intermittent, periodic density structures that range in size from 70 to 900 Mm are a ubiquitous feature of the slow solar wind, occurring a majority ($\gtrsim 60\%$) of the time. Furthermore, Viall et al. (2010) and Viall & Vourlidis (2015) examined the Solar Terrestrial Relations Observatory’s (Kaiser et al. 2008) SECCHI (Howard et al. 2008) HI1 and COR2 white-light imaging data and showed there were clear signatures of ~ 90 minutes variability in the intensity variations of coronal streamer outflow, confirming that many of the periodic density structures are the result of solar wind formation processes.

In the in situ slow solar wind, especially near the HCS, magnetic structures with timescales of several hours have been identified and linked to magnetic reconnection (Crooker et al. 1996, 2004; Suess et al. 2009). High-cadence composition data have revealed the presence of cyclic 0.5–3 hr solar wind structures with signatures in helium, oxygen, and carbon densities, and heavy ion charge states (Viall et al. 2009; Kepko et al. 2016). In situ elemental and ionic composition measurements are routinely used as proxies for solar wind formation processes and the “freeze-in” coronal electron temperatures in the low-to-middle corona; when the characteristic bulk solar wind expansion timescale exceeds the ionization and recombination timescales of various ion species, the ionic charge states can be considered frozen-in to the solar wind outflow (e.g., Hundhausen et al. 1968; Owocki et al. 1983; Ko et al. 1997; Landi et al. 2012b, 2012c; Landi & Lepri 2015). In fact, Kepko et al. (2016) showed that these density and compositional variations also often correspond to regions of coherent magnetic field signatures and periods of bidirectional electron streaming, suggestive of a succession of small magnetic flux ropes or flux-rope-like periods. There is

some preliminary indication that in situ small flux ropes can be coincident with periods of enhanced ionic composition (Foullon et al. 2011; Feng & Wang 2015; Kepko et al. 2016; Yu et al. 2016).

The separatrix–web (hereafter, S-Web) model for the origin of slow solar wind (Antiochos et al. 2011) is based on the magnetic geometry of the solar corona and predicts that the topological separatrix surfaces of the magnetic field are regions where interchange reconnection—the mechanism for releasing closed-flux coronal plasma onto adjacent open field lines—is most likely to occur.

The dynamic S-Web model extends previous observational and theoretical considerations of reconnection at coronal hole boundaries (Madjarska et al. 2004; Edmondson et al. 2009, 2010; Linker et al. 2011; Rappazzo et al. 2012; Brooks et al. 2015; Pontin & Wyper 2015; Scott et al. 2021) and solar wind outflows at the periphery of active regions (e.g., Sakao et al. 2007; Harra et al. 2008; Baker et al. 2009; Brooks & Warren 2011; Edwards et al. 2016), and aims to address a number of outstanding issues related to the slow solar wind, including its larger-than-expected latitudinal extent (Crooker et al. 2012) and the reconnection component seemingly required by the variability of the in situ measurements of slow wind plasma, field, and composition (e.g., Viall et al. 2009; Zhao et al. 2009; Lepri et al. 2013, 2014; Zhao et al. 2014; Kepko et al. 2016; Sanchez-Diaz et al. 2017a; Zhao et al. 2017; Di Matteo et al. 2019; Sanchez-Diaz et al. 2019; Réville et al. 2022).

Higginson et al. (2017a, 2017b) presented simulation results showing that interchange magnetic reconnection is ubiquitous and most likely responsible for releasing much of the slow solar wind, in particular along S-Web topological features. Since that work there have been a number of significant developments in the modeling reconnection-generated slow solar wind structure and the interchange reconnection processes associated with dynamic S-Web outflows, summarized in Figure 1. Figure 1(a) presents the 3D structure of the pinch-off reconnection that

forms streamer blobs (i.e. flux rope plasmoids) in the simulation by Lynch (2020). These simulation results showed qualitative agreement with both the morphology and the kinematics of coronal inflows and streamer blob outflows in synthetic white-light coronagraph observations, as have other recent modeling efforts (e.g., Réville et al. 2020). Figure 1(b) presents simulation results from Higginson & Lynch (2018), who showed that the continual formation of flux rope/plasmoid structures essentially filled the entire HCS. Figure 1(c) shows the simulation results by Aslanyan et al. (2022) in which they examined interchange reconnection occurring in a 3D PS configuration and developed a synthetic suprathermal electron pitch-angle proxy based on the simulation’s instantaneous magnetic connectivity.

Previously, Zhao et al. (2017) used solar wind data from the Advanced Composition Explorer (ACE; Stone et al. 1998) during Carrington Rotation (CR) 2002 to develop a source-region classification scheme based on heliospheric back-mapping and potential field source surface (PFSS) modeling of observer-connected magnetic field lines and the pixel brightness in synoptic maps of extreme-UV (EUV) 195 Å emission in the vicinity of the field line foot point. Applying their EUV brightness-based source-region classifications (“Coronal Hole,” “Coronal Hole Boundary,” “Quiet Sun,” “Active Region Boundary,” “Active Region,” and “Helmet Streamer”) to in situ data from 1998–2011 resulted in a statistical ordering of the distributions of O^{7+}/O^{6+} by distance from coronal holes, representing a relatively smooth increase in some combination of coronal electron temperature, mass density, and/or outflow velocities.

A number of other solar wind classification schemes have been developed to identify specific solar wind “types” for the purpose of trying to uncover the physical relationships between different plasma, field, and composition signatures within and between different solar wind types (which are generally a proxy for coronal source-region classifications). For example, Xu & Borovsky (2015) constructed a “four-plasma” classification scheme based, in part, on the proton specific entropy, $S_p = T_p/n_p^{2/3}$, and showed this had a significant correlation with O^{7+}/O^{6+} and C^{6+}/C^{5+} signatures and a relatively clear separation in the Alfvén-speed-specific entropy ($v_A - S_p$) space between their “Ejecta,” “Coronal Hole,” “Streamer Belt,” and “Sector Reversal” (HCS/HPS crossing) types. Ko et al. (2018) examined the perpendicular velocity fluctuations (δv_T , δv_N in RTN coordinates) and presented superposed epoch trends in HS and PS intervals (low δv) for a variety of solar wind properties including magnetic field fluctuations, Alfvénicity, width of the suprathermal electron strahl, proton specific entropy S_p , helium abundance, the carbon, oxygen and iron charge states, and Fe/O composition. Bloch et al. (2020) investigated a couple of machine-learning techniques to identify “Streamer Belt” and “Coronal Hole” solar wind type clusters in the $S_p - O^{7+}/O^{6+}$ parameter space from Ulysses and ACE data. Roberts et al. (2020) used k -means clustering based on a number of solar wind variables including oxygen and iron charge states and the Fe/O ratio, which resulted in a mixture of some clearly separated solar wind types and some significantly overlapping solar wind types when visualized in the cross helicity (σ_c) and residual energy (σ_r) parameter space commonly used in turbulence studies.

In this paper, we extend the CR 2002 analysis of Zhao et al. (2017) to the magnetic complexity of the source region and

examine the relationship between measures of solar wind variability in plasma, field, and composition with the large-scale geometric S-Web configurations of the associated source regions. In Section 2, we present in situ solar wind observations from the Wind and ACE spacecraft during CR 2002 and define several slow-to-moderate speed intervals of enhanced variability in proton and alpha densities. We then show that each of these intervals correspond to enhancements in the ionic composition signatures of carbon, oxygen, and iron. In Section 3, we perform the heliospheric back-mapping procedure to map the in situ time series at 1 au to Carrington longitude at the PFSS at $2.5 R_\odot$ (Section 3.1) and show these intervals of enhanced variability and composition map back to the S-Web topological structures associated with the HS belt and coronal PSs (Section 3.2). In Section 4, we present the magnetic helicity–partial variance of increments (H_m –PVI) analysis during the enhanced variability intervals and quantify the similarities and differences between the HS (Section 4.2) and PS (Section 4.3) slow wind, and perform some statistical analyses on these time series (Section 4.4). Finally, in Section 5, we discuss the implications of our results for theory and modeling the origin of the slow solar wind and avenues for future progress with complementary PSP and Solar Orbiter observations.

2. Intervals of Enhanced Variability

2.1. Proton Density and the Alpha-to-proton Ratio

The slow solar wind shows considerably more variation in proton and helium densities (and their relative abundance ratio) than in the fast wind. The mean alpha particle (He^{2+}) to proton (H^+) ratio $A_{He} \equiv n_\alpha/n_p \times 100$ (or α/p , interchangeably) in both the fast and slow solar wind are on the order of 3%–5% but the relative variation in the fast solar wind is $\sim 10\%$ while in the slow solar wind it can be as high as $\sim 40\%$ (Gosling 1997; Schwenn 2006). Helium enhancements have long been associated with in situ observations of coronal mass ejection material (e.g., Borrini et al. 1982; Richardson & Cane 2004; Zurbuchen et al. 2016; Lepri & Rivera 2021), but recent analyses have also made significant progress quantifying the helium variability during ambient solar wind intervals (Kasper et al. 2007; Suess et al. 2009; Wang 2016a; Sanchez-Diaz et al. 2019). For example, Kasper et al. (2007, 2012) have shown the solar wind α/p ratio exhibits both a dependence on solar wind speed and the phase of the solar activity cycle, with the A_{He} in the slowest-speed solar wind intervals showing the most variation with sunspot number, in support of multiple sources and/or mechanisms for the solar wind’s helium component (Schwenn et al. 2006). Viall et al. (2009) and Kepko et al. (2016) and others have shown that the solar wind helium abundance (and the associated increase in the variance of the helium abundance) are often coincident with periodic proton density structures (and their increased variance), as well as periods of increased ionic and elemental composition (see also Kasper et al. 2012).

Figure 2 shows a plot of the Wind/Three-Dimensional Plasma and Energetic Particle Investigation (3DP; Lin et al. 1995) and Wind/Solar Wind Experiment (SWE; Ogilvie et al. 1995) data at 1 au for CR 2002 (from 2003 April 15 21:35 UT through 2003 May 13 03:24 UT). From top to bottom, we plot the bulk radial velocity V_r , proton number density n_p , alpha number density n_α , the A_{He} ratio, and its variance,

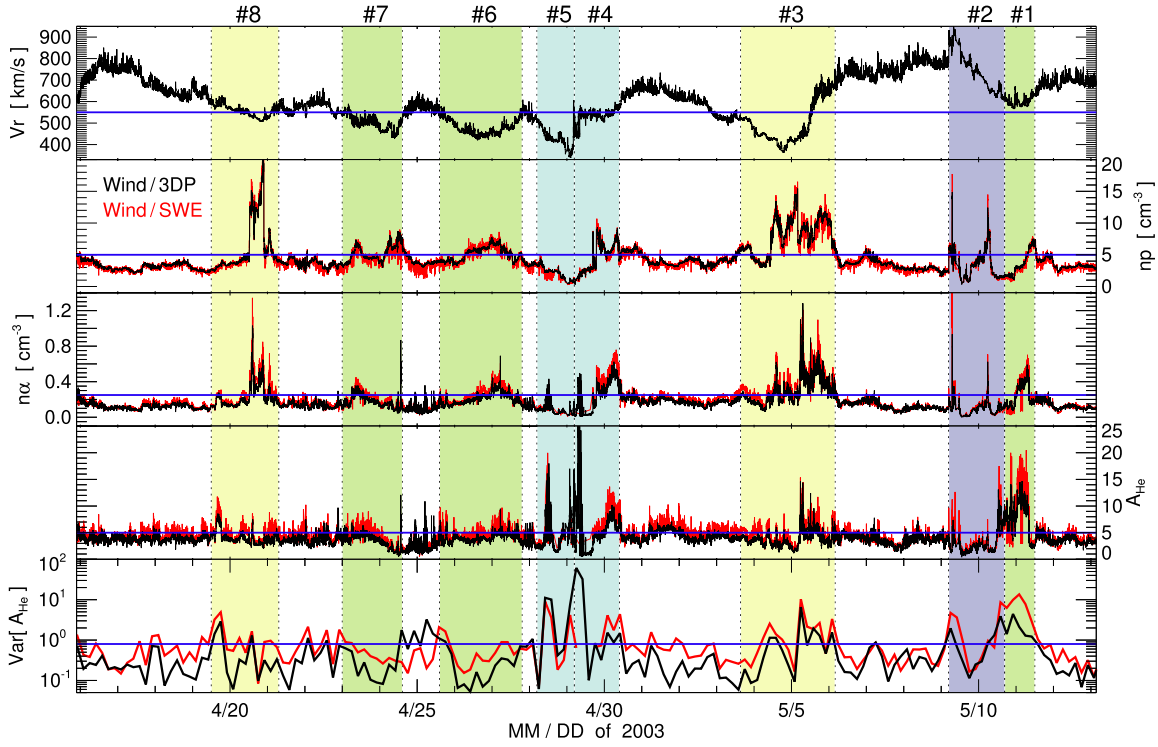


Figure 2. In situ solar wind data from the Wind spacecraft during Carrington Rotation (CR) 2002. Plotted, from top to bottom, are the bulk radial velocity V_r , proton number density n_p , alpha number density n_α , the $\text{He}^{2+}/\text{H}^+$ ratio, $A_{\text{He}} \equiv n_\alpha/n_p \times 100$, and its variance $\sigma_{\alpha/p}^2$ in 6 hr bins. The black curves are Wind/3DP 1 min data and the red curves are Wind/SWE 97 s data. The eight intervals, labeled #1–8 along the top axis, represent the different large-scale coronal source-region classifications (HS, yellow; PS, green, teal; ICME, purple). The interval properties are summarized in Table 1.

Table 1
Average Properties of Each Slow-to-moderate Speed, Composition-enhanced Solar Wind Interval During CR 2002

#	Start Time DD/MM HH:MM (UT)	End Time DD/MM HH:MM (UT)	Source Region	$\langle V_r \rangle$ (km s $^{-1}$)	$\langle n_p \rangle$ (cm $^{-3}$)	$\langle A_{\text{He}} \rangle$ (%)	$\langle Q_C \rangle$ (4–6+)	$\langle Q_O \rangle$ (5–8+)	$\langle Q_{\text{Fe}} \rangle$ (6–20+)	$\langle \text{Fe}/\text{O} \rangle$
8	04/19 13:58	04/21 05:24	HS (Y)	556 \pm 26	6.1 \pm 4.1	5.1 \pm 1.6	5.07 \pm 0.15	6.14 \pm 0.08	10.82 \pm 0.96	0.11 \pm 0.02
7	04/23 00:00	04/24 14:12	PS (G)	498 \pm 32	5.2 \pm 1.3	4.4 \pm 1.6	5.19 \pm 0.08	6.08 \pm 0.02	10.12 \pm 0.52	0.13 \pm 0.02
6	04/25 14:24	04/27 19:12	PS (G)	478 \pm 34	5.2 \pm 1.1	5.0 \pm 1.1	5.24 \pm 0.16	6.09 \pm 0.05	9.85 \pm 0.32	0.16 \pm 0.03
5	04/28 04:48	04/29 04:47	PS (T)	432 \pm 45	2.2 \pm 1.0	3.9 \pm 1.8	5.52 \pm 0.18	6.32 \pm 0.16	11.83 \pm 1.29	0.19 \pm 0.09
4	04/29 04:48	04/30 09:36	PS (T)	534 \pm 32	4.5 \pm 2.3	7.6 \pm 2.4	5.35 \pm 0.15	6.29 \pm 0.14	11.07 \pm 0.51	0.31 \pm 0.28
3	05/03 15:27	05/06 03:56	HS (Y)	496 \pm 95	7.9 \pm 2.8	4.6 \pm 1.7	5.24 \pm 0.24	6.20 \pm 0.15	11.01 \pm 0.79	0.14 \pm 0.07
2	05/09 04:48	05/10 16:48	ICME (P)	738 \pm 89	3.5 \pm 2.3	3.2 \pm 2.1	5.06 \pm 0.15	6.12 \pm 0.05	11.10 \pm 1.14	0.21 \pm 0.11
1	05/10 16:48	05/11 12:00	PS (G)	601 \pm 20	3.6 \pm 1.8	9.4 \pm 5.3	5.19 \pm 0.08	6.22 \pm 0.08	10.09 \pm 0.25	0.20 \pm 0.07
Noninterval CR 2002 averages				637 \pm 93	4.0 \pm 1.9	4.5 \pm 1.2	5.01 \pm 0.17	6.04 \pm 0.09	10.37 \pm 0.55	0.11 \pm 0.03

Note. Interval-averaged solar wind properties V_r and n_p are from Wind/3DP, A_{He} is from Wind/SWE, and Q_C , Q_O , Q_{Fe} , and Fe/O are from ACE/SWICS. The interval source region shading is also indicated (Y, yellow; G, green; T, teal; and P, purple). Boldface values are slower and/or more enhanced than the noninterval averages over the remainder of CR 2002.

$\text{Var}[A_{\text{He}}] \equiv \sigma_{\alpha/p}^2$, calculated over 6 hr bins. The 3DP data are shown in black and the SWE data are shown in red. Based on visual inspection of the Figure 2 time series, we have identified eight distinct intervals during CR 2002 that can be considered slow-to-moderate speed solar wind ($V_r \lesssim 550 \text{ km s}^{-1}$) with one or more of the following: enhanced proton density ($n_p \geq 5 \text{ cm}^{-3}$), enhanced alpha density ($n_\alpha \geq 0.25 \text{ cm}^{-3}$), enhanced A_{He} ($\geq 5\%$), or enhanced $\sigma_{\alpha/p}^2$ (≥ 0.80). Each of the intervals are labeled above the top x-axis as #1–8 and shaded as yellow, green, teal, or purple. The colors were selected to represent different large-scale coronal source-region configurations, as will be discussed in Section 3.2. The one exception to our slow-to-moderate speed criteria is interval #2 (shaded purple), which is clearly identified as a fast interplanetary coronal mass

ejection (ICME) and cataloged as such by Richardson & Cane (2010). The start and end times of each Figure 2 interval are listed in Table 1 along with a synopsis of the relevant interval-averaged quantities.

2.2. Ionic and Elemental Composition Enhancement

Figure 3 shows ACE measurements for the CR 2002 solar wind. From top to bottom, we plot the Solar Wind Electron Proton Alpha Monitor (SWEPAM; McComas et al. 1998a) measurements of the bulk solar wind speed V_r , the normalized 272 eV suprathermal electron pitch-angle distribution (PAD), the Magnetometer Instrument (MAG; Smith et al. 1998) measurements of B in RTN coordinates, and the magnetic

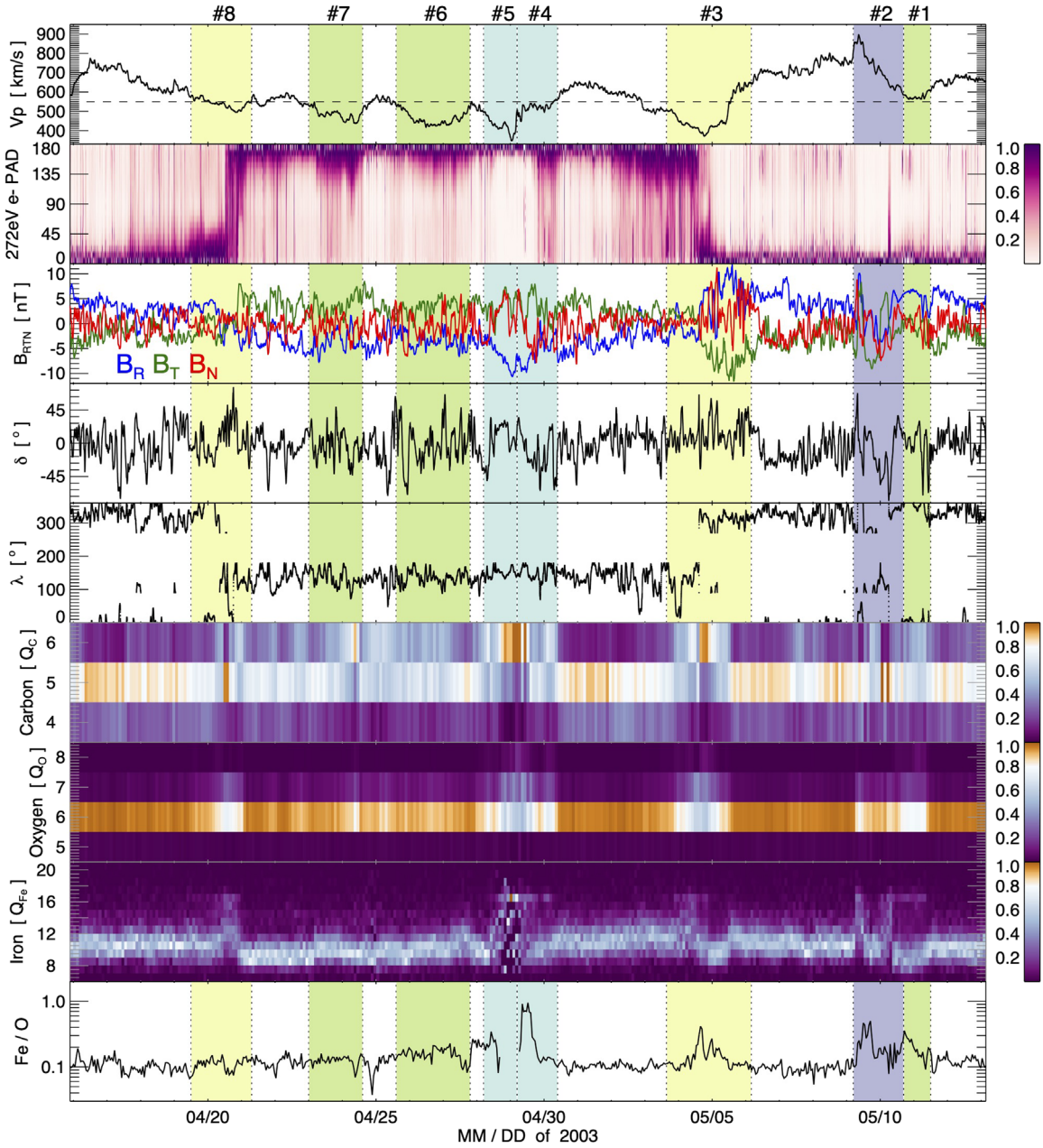


Figure 3. Solar wind and ionic and elemental composition properties from ACE/SWEPAM, ACE/MAG, and ACE/SWICS for CR 2002. From top to bottom, we plot proton V_p , the 272 eV suprathermal electron PAD, the magnetic field components of \mathbf{B}_{RTN} , the magnetic field elevation and azimuthal angles (δ , λ), the distributions of C^{4-6+} , O^{5-8+} , and Fe^{6-20+} , and the Fe/O ratio. The slow-to-moderate speed intervals from Figure 2 and Table 1 are also shown.

field orientation angles (δ is the elevation angle above/below the RT plane; λ is the azimuth angle within the RT plane), and the Solar Wind Ion Composition Spectrometer (SWICS; Gloeckler et al. 1998) composition measurements of select ion charge states of carbon (Q_C : 4–6+), oxygen (Q_O : 5–8+), and iron (Q_{Fe} : 6–20+), as well as the Fe/O abundance ratio. Here, the solar wind speed and magnetic field values are 1 hr averages whereas the SWICS composition measurements are 2 hr averages.

Figure 3 also shows each of the slow-to-moderate speed solar wind intervals associated with enhanced n_p , n_α , or A_{He} variability that were identified in the Wind data of Figure 2. With the inclusion of the magnetic field and suprathermal electron PAD, the intervals corresponding to sector boundaries

and heliospheric current sheet/plasma sheet (HCS/HPS) crossings are immediately apparent as #8 and #3, both shaded light yellow.

Another particularly noteworthy feature of Figure 3 is that each of the remaining slow-to-moderate speed intervals are coincident with broader suprathermal electron PADs and/or elevated charge states in carbon, oxygen, and iron. While recent analyses by Borovsky (2020, 2021) have shown that changes in the suprathermal electron strahl intensities often occur with simultaneous changes in other plasma and/or composition properties, here we note that the broader PADs of intervals #7, #6, and #4 exhibit remarkable, qualitative agreement with the synthetic PAD distribution constructed by Aslanyan et al. (2022) from their MHD simulation of

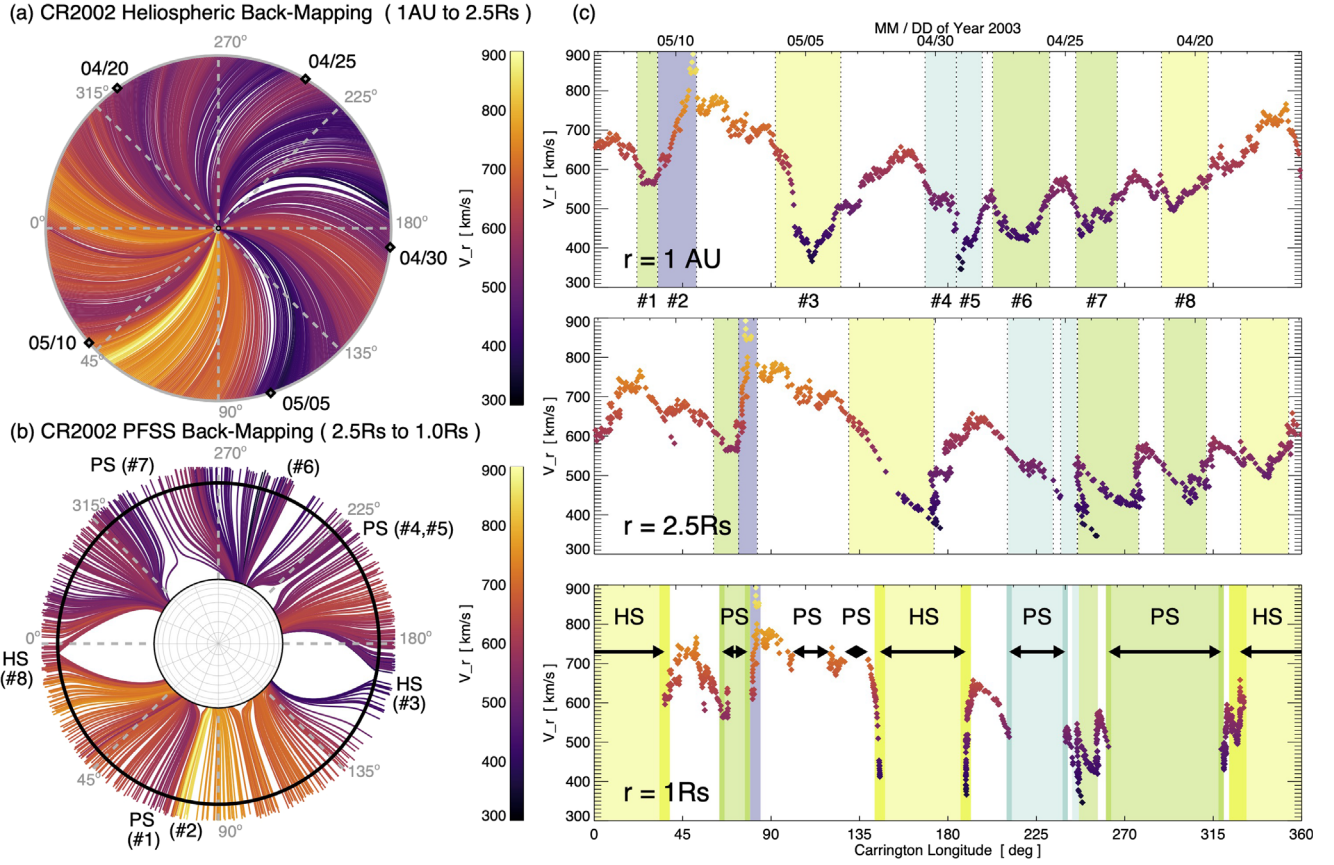


Figure 4. Heliospheric back-mapping for CR2002. (a) Ecliptic plane streamlines color-coded by 1 au radial velocity value to the $r = 2.5 R_\odot$ source surface. (b) Continuation of the back-mapping from R_{ss} to $1 R_\odot$ with the PFSS magnetic field. The view of the ecliptic plane is from the solar north pole. (c) The mapping of the time series of 1 hr ACE/SWEPAM radial velocity in Carrington longitude at 1 au (top panel) to the source surface (middle panel) and then to the solar surface (bottom panel). The intervals of high α/p from Section 2 are also shown in each location.

interchange reconnection PS outflow (lower panels of Figure 1(c)). We will show in the next section these intervals do, in fact, map to coronal PS source regions.

The charge state and elemental composition enhancements during each of the identified intervals have the following properties. The presence of increased C^{6+} and decreased C^{5+} will result in a substantial increase in the C^{6+}/C^{5+} ratio, which has similar properties to the O^{7+}/O^{6+} ratio commonly used to identify periods of increased coronal electron temperatures (e.g., Landi et al. 2012a; Kepko et al. 2016). Additionally, every interval except #6 and #7 also shows a significant increase in O^{7+} along with a corresponding decrease in O^{6+} , providing local maxima of the well-known O^{7+}/O^{6+} ratio (e.g., Zhao et al. 2009; Wang 2016b). During intervals #3, #4–5, and #8, there are also enhanced levels of the higher iron charge states, $Fe^{\geq 12+}$, including some traditionally “hot” signatures of $Fe^{\geq 16+}$ (Lepri & Zurbuchen 2004). Finally, the elemental composition ratio Fe/O shows clear enhancements during intervals #1–5 but less obvious enhancement during intervals #6–7. From Table 1, only interval #8 does not exceed the noninterval Fe/O average.

3. Solar–Heliospheric Connectivity to the Coronal S-Web

3.1. Heliospheric Ballistic Back-mapping

Here we follow the standard procedure for heliospheric back-mapping described by Parenti et al. (2021) and references therein. The in situ observations of solar wind at 1 au are

ballistically mapped from the spacecraft back to the Sun along the Parker (1958) spiral streamlines, assuming constant V_r values equal to the 1 hr averages measured by ACE. Figure 4(a) plots the heliospheric representation of Parker spiral streamlines colored by radial velocity.

In order to map the in situ solar wind observations back to their coronal source regions on the solar surface, we use the standard PFSS model (e.g., Altschuler & Newkirk 1969; Wang & Sheeley 1992) to approximate the large-scale geometry of the solar corona. We calculate the PFSS extrapolation from the line-of-sight observations of the photospheric magnetic field taken by the Michelson Doppler Imager (MDI; Scherrer et al. 1995) on board the Solar and Heliospheric Observatory (Domingo et al. 1995). The line-of-sight fields are transformed into radial fields via the $B_r = B_{los}/\sin \theta$ relation. Starting with the original high-resolution (3600×1080) MDI synoptic map for CR 2002 with the Sun et al. (2011) interpolation for the polar field values, we rebin it to 720×360 and calculate the PFSS spherical harmonics through order $l_{max} = 16$ with a source surface height of $R_{ss} = 2.5 R_\odot$. Figure 4(b) plots the magnetic field line mapping from R_{ss} to the lower boundary at $r = 1 R_\odot$. Here, the large-scale, closed-field coronal HS and PS structures are labeled with their corresponding intervals.

The top panel of Figure 4(c) plots the radial velocity as a function of Carrington longitude at 1 au (note time now runs from right to left, as indicated by the upper x-axis label). We have also drawn the corresponding intervals of enhanced variability identified in Section 2.1. The middle panel of

Figure 4(c) plots the 1 hr ACE velocity measurements as a function of Carrington longitude at R_{ss} , while the bottom panel of Figure 4(c) shows the Carrington longitude of the streamline foot points at $1 R_{\odot}$.

This ballistic mapping method has been widely used to estimate the coronal source regions of in situ solar wind measurements (e.g., Neugebauer et al. 2002, 2004; Gibson et al. 2011; Zhao et al. 2013, 2017), including with the new PSP (e.g. Badman et al. 2020; Panasenco et al. 2020; Griton et al. 2021) and Solar Orbiter data (e.g., Telloni et al. 2021). We note that, while the numerical errors associated with integrating velocity streamlines or magnetic field lines, e.g., from a PFSS extrapolation, are quite small (Stansby & Verscharen 2022), the overall “uncertainty” in the position of the foot point of the magnetic field lines as mapped by these techniques are typically within approximately 10° (Neugebauer et al. 2002; Leamon & McIntosh 2009), largely due to the assumptions and simplifications inherent in the models themselves, such as the current-free approximation in the corona and the unperturbed Parker spiral structure that does not account for the interaction between fast and slow solar wind streams, etc.

3.2. Separatrix–Web Source Region Configurations

The static representation of the S-Web’s topological structures is based on the Q -map, which is defined as the logarithmic “squashing factor,” $\log Q$. The Q -map quantifies the magnetic field’s geometric connectivity (Titov 2007), i.e., separatrix and quasi-separatrix surfaces are regions of high Q (e.g., Titov et al. 2011; Antiochos et al. 2012; Scott et al. 2018). We have calculated the Q -value from the CR 2002 PFSS magnetic field extrapolation via the formulation in Titov (2007), where $Q = N^2/|\Delta|$:

$$N^2 = \left(\frac{\partial Y}{\partial y}\right)^2 + \left(\frac{\partial Y}{\partial z}\right)^2 + \left(\frac{\partial Z}{\partial y}\right)^2 + \left(\frac{\partial Z}{\partial z}\right)^2, \quad (1)$$

and $|\Delta| = |B_x/B_x^*|$. While the full derivation (in arbitrary coordinates) is described in Titov (2007), the expression in spherical coordinates is straightforward to obtain: $B_x/B_x^* \rightarrow B_r/B_r^*$, the starting and ending field line positions become $(x_0, y_0, z_0) \rightarrow (r_0, \theta_0, \phi_0)$, $(X, Y, Z) \rightarrow (R, \Theta, \Phi)$, and the differentials become changes in arc length $\partial y \rightarrow r_0 \partial \theta$, $\partial Y \rightarrow R \partial \Theta$, $\partial z \rightarrow r_0 \sin \theta_0 \partial \phi$, and $\partial Z \rightarrow R \sin \Theta \partial \Phi$. We note that when the starting and ending radial surfaces are set to $r_0 = R_{\odot}$ and $R = R_*$, one arrives at the exact spherical definition of N^2 , given as Equation (22) in Titov (2007). We calculate the field connectivity from a grid of 1536×768 field lines starting at the desired radial distance r_0 uniformly spaced in (θ, ϕ) . As in the Q -map calculation of Wyper et al. (2016), we use a fourth-order central difference scheme for the derivatives.

Figure 5 summarizes the coronal portion of our heliospheric back-mapping procedure to illustrate the connectivity of our composition-enhanced intervals to their coronal S-Web structures of origin. Figure 5(a) repeats the 1 hr average V_r points mapped to R_{ss} (from Figure 4(c)) and plots the longitudinal extent of our back-mapped intervals of interest with their boundaries indicated in every subsequent panel. Figure 5(b) shows the MDI magnetogram for CR 2002. The positive (negative) open field regions calculated from the PFSS are

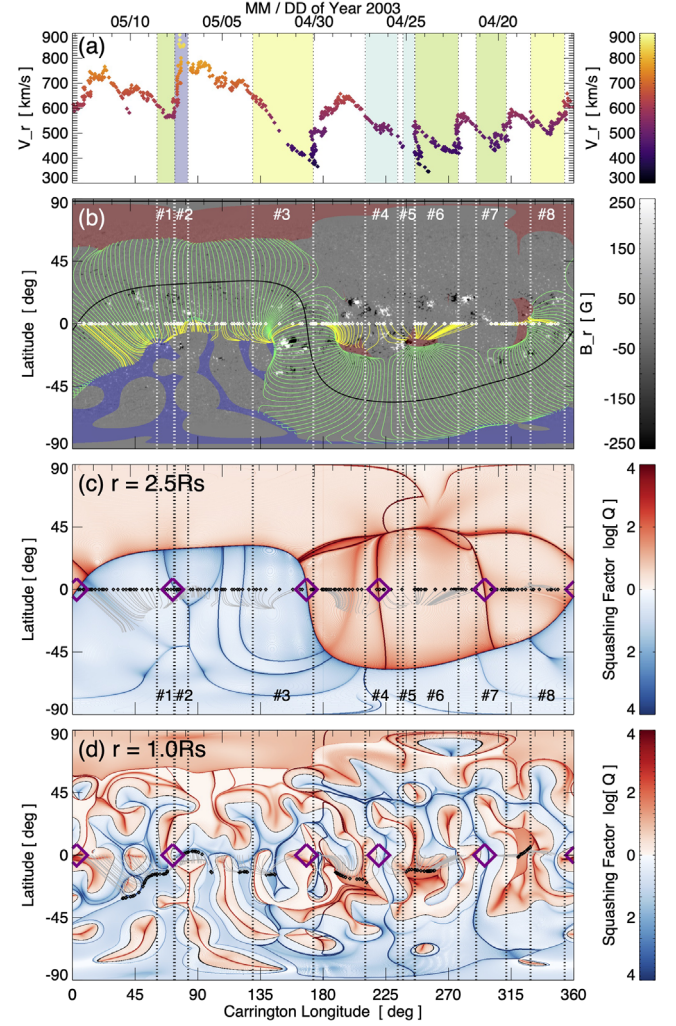


Figure 5. Magnetic structure of the PFSS extrapolation for CR2002 with back-mapped in situ intervals of slow-to-moderate speed solar wind. (a) The back-mapped V_r time series and intervals #1–8 at $R_{ss} = 2.5 R_{\odot}$ from Figures 4(c). (b) Synoptic map of the open field regions (blue positive polarity, red negative polarity). The configuration of the HS belt is shown as green field lines traced from the $B_r = 0$ contour at R_{ss} , representing the location of the HCS. (c) Q -map at R_{ss} showing the characteristic arcs of the S-Web structure. The $\log Q$ values are shaded blue (red) for positive (negative) polarity. (d) Q -map at $1 R_{\odot}$ showing the equatorial field line foot point locations and the low-latitude open field regions between PSs for intervals #4–7.

solution shaded in blue (red), the structure of the HS belt is given with representative green field lines, and the HCS location ($B_r = 0$ at R_{ss}) is shown as the black contour. The green field lines are traced along the HCS location at a radial distance just below R_{ss} , and therefore represent the largest closed flux tubes belonging to the HS belt and illustrate the boundary between the large-scale open and closed coronal flux systems. The back-mapped intervals are labeled along the top axis of the plot.

Figure 5(c) plots the Q -map at R_{ss} . The values of $\log Q$ are also shaded blue and red to indicate B_r polarity. The position of the HCS current sheet is immediately identified as where the polarities change sign. The darker arcs contained within each polarity correspond to S-Web arcs. These S-Web arcs indicate the PFSS field line mapping of the outer spine and fan structures of PS flux systems (Scott et al. 2018) and/or the presence of narrow channels of open field (Antiochos et al. 2007). The purple diamonds indicate the S-Web features

associated with their corresponding back-mapped, composition-enhanced intervals. The in situ intervals that contain the HCS crossing (#3, #8) are clearly associated with the HS belt and the intersection of the HCS with the ecliptic plane, despite the spatial extent of interval #8 at R_{ss} (329° – 353°) missing the PFSS location of the HCS (3°) by $\sim 10^\circ$. This discrepancy is typical of the uncertainties associated with our simplified back-mapping (as mentioned above) but given that PFSS HS width beneath the HCS–ecliptic plane intersection spans $\sim 67^\circ$ in Carrington longitude (330° – 37°) at R_\odot , the association between the solar wind during interval #8 and its origin from this portion of the HS is evident. Intervals #1, #4, and #7 each include a well-defined, PS S-Web arc in their longitudinal range. Interval #5 is directly adjacent to the S-Web arc of interval #4, and interval #6 appears to straddle the midpoint between the #4 and #7 S-Web arcs.

Figure 5(d) plots the Q -map at $r = 1 R_\odot$ and shows the foot points of the PFSS magnetic field lines traced from R_{ss} . The positive polarity (blue) open field foot points map to the southern boundary of the HS belt/northern boundary of the polar coronal hole extensions. The negative polarity (red) open field foot points map to a series of low-latitude coronal holes sandwiched between the northern boundary of the HS belt and the southern boundaries of a series of large PSs above the AR magnetic fields between Carrington longitudes 180° – 315° . While intervals #5 and #6 were not associated with a distinct S-Web arc at R_{ss} , their field line foot points map to the vicinity of the open–closed flux boundaries between the low-latitude coronal holes and the large equatorial PSs.

4. Coherent Magnetic Structure in Composition-enhanced Solar Wind

4.1. H_m –PVI Analysis Procedure

We have implemented the Pecora et al. (2021) magnetic helicity–partial variance of increments (H_m –PVI) procedure to identify coherent magnetic structures within our intervals of composition-enhanced solar wind originating from coronal HS and PS source regions. Here, we briefly review the methodology for the identification of small-scale flux ropes and/or coherent flux tubes, while in Sections 4.2 and 4.3, we present the results from applying this technique to HS intervals (#3, #8) and PS intervals (#1, #4–#7), respectively. In Section 4.4, we compare and contrast properties of the H_m and PVI time series in each interval.

Quite generally, the magnetic helicity can be written as

$$H_m = H_m^+(\ell) + H_m^-(\ell), \quad (2)$$

where the temporal or spatial scale, ℓ , is used to define the magnetic helicity contained in scales greater than ℓ as $H_m^+(\ell)$ and less than ℓ as $H_m^-(\ell)$. Since we are interested in the local coherence, we will ignore the H_m^+ term and follow the Pecora et al. (2021) prescription for the local estimate of H_m^- using the two-point correlation function $C_{jk} = \langle B_j(\mathbf{r})B_k(\mathbf{r} + \mathbf{s}) \rangle$ (e.g., Matthaeus & Goldstein 1982). We take the spatial lag $\mathbf{s} = s \hat{\mathbf{e}}_i$ to be in the $\hat{\mathbf{r}}$ direction, so indices j, k are the tangent and normal directions of the spacecraft’s RTN coordinates.

We calculate the spatial average of the two-point correlation function over an interval of width $W = 2\ell$ centered at position

x via

$$C_{jk}(x, s) = \frac{1}{W} \int_{x-\frac{W}{2}}^{x+\frac{W}{2}} d\xi [B_j(\xi + s)B_k(\xi) - B_j(\xi)B_k(\xi + s)]. \quad (3)$$

Following Pecora et al. (2021), we apply a smooth windowing function to $C_{jk}(x, s)$ of the form $f(s) = \frac{1}{2} + \frac{1}{2} \cos(2\pi s/W)$ to obtain the local helicity estimate:

$$H_m(x, \ell) = \int_0^\ell ds C_{jk}(x, s) f(s). \quad (4)$$

The spatial domain quantities (x, s) can be converted to the temporal domain (t, τ) with the usual Taylor approximation of $x(t) = \int d\tau V_r(\tau)$.

In our implementation of $H_m(t, \ell)$, we use a spatial scale of $\ell_H = 4.3 \times 10^6$ km ($4300 R_\odot$, $\sim 6.2 R_\odot$); for a solar wind speed of 500 km s^{-1} , this corresponds to a temporal scale of 2.4 hr (i.e., ~ 135 data points at 64 s cadence), which is the mean correlation timescale of the vector magnetic field over our eight intervals (2.37 ± 1.83 hr). However, we note that the correlation timescales during the HCS/HPS intervals were significantly larger (4.58 ± 1.0 hr) than the PS intervals (0.98 ± 0.52 hr), which agree with previous estimates (e.g., Matthaeus et al. 2005). Typically, one decides that a given peak in $H_m(t, \ell)$ is significant if it exceeds a $\pm 1\sigma$ threshold. In the following sections, this standard deviation is calculated from the H_m curves over the entire interval of interest, i.e., those defined in Section 2 (and illustrated in Figures 2–5).

The PVI measure (e.g., Pecora et al. 2019, 2021) is defined as

$$\text{PVI}(t, \ell) = \frac{|\Delta \mathbf{B}(t, \ell)|}{\sqrt{\langle |\Delta \mathbf{B}(t, \ell)|^2 \rangle}}, \quad (5)$$

in which $|\Delta \mathbf{B}(t, \ell)| \equiv |\mathbf{B}(t + \ell) - \mathbf{B}(t)|$, the (temporal or spatial) averaging is over an appropriate interval, and ℓ represents the scale size of the increments. The PVI technique has been widely used to identify discontinuities, reconnecting current sheets, and as a measure of turbulence structures (e.g., Greco et al. 2009a, 2009b; Osman et al. 2014; Greco et al. 2018; Pecora et al. 2019). Since we aim to use PVI to identify the sharp magnetic boundaries of coherent flux tubes and/or small flux-rope plasmoids, we choose a temporal scale of $\ell_{\text{PVI}} = 2.13$ minutes and an averaging window of 24 hr (10 times the magnetic field’s mean correlation timescale above). Again, one makes a determination of the significance of any given PVI peak via thresholding, where some authors have used $\text{PVI} > 2$ (Pecora et al. 2021), > 2.4 (Greco et al. 2008), > 3 (Kilpua et al. 2022), or even larger thresholds of > 4 – 6 (e.g., Servidio et al. 2011; Zhou et al. 2019). Here, we use $\text{PVI} > 3.0$ during each of our CR 2002 solar wind intervals for ease of comparison between the HS and PS PVI statistics. The magnitude of the PVI peaks has been shown to be related to different types of boundaries or discontinuities in the solar wind. For example, the $\text{PVI} \gtrsim 3$ threshold has been interpreted as representing discontinuities that are actual physical boundaries of coherent magnetic structures rather than random statistical fluctuations, whereas PVI values $\gtrsim 5$ have been associated with reconnection events (Servidio et al. 2011).

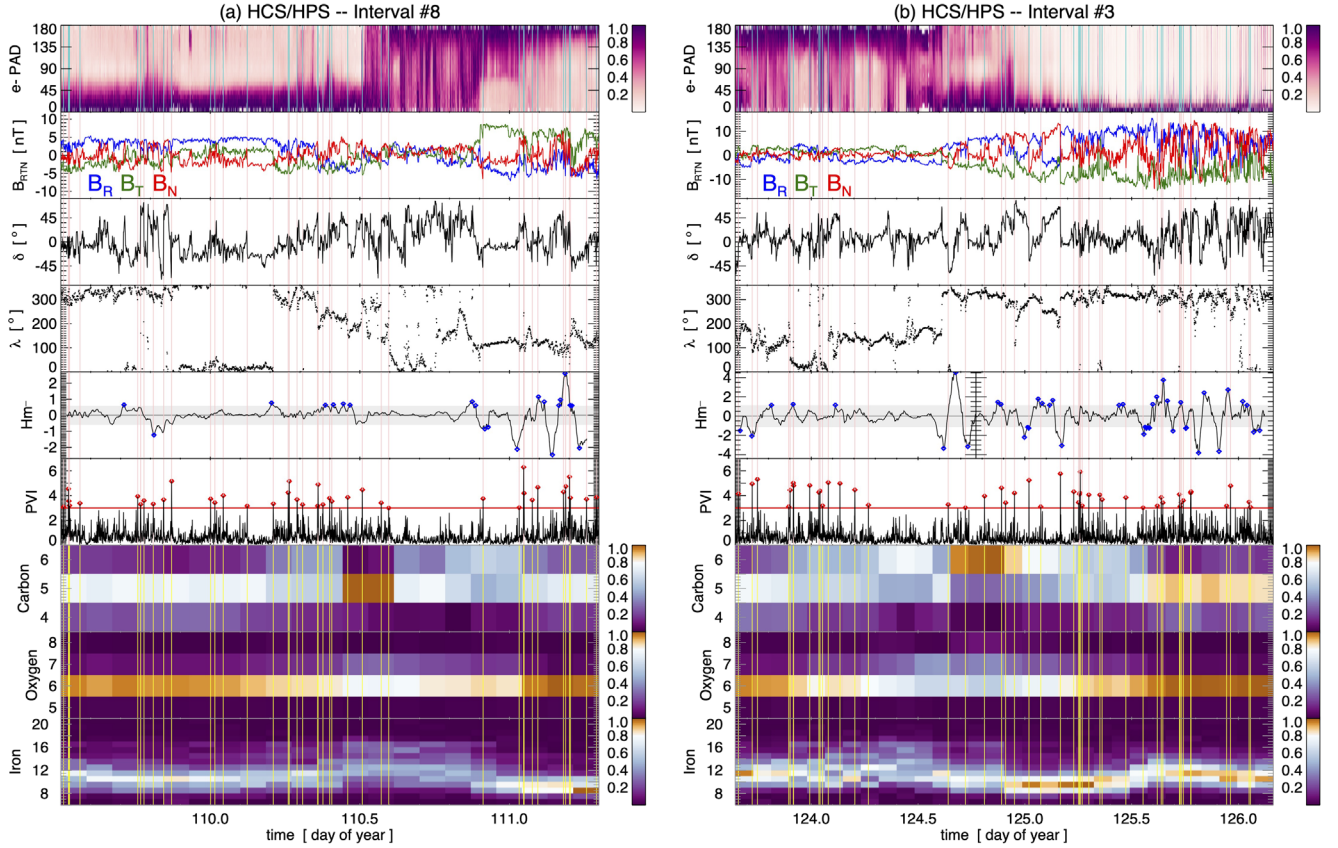


Figure 6. Intervals of the HS wind that include HCS/PS crossings. (a) HS #8 from DOY 109.582 to 111.267. (b) HS #3 from DOY 123.644 to 126.164.

The strength of the H_m –PVI procedure is that for a magnetic island or a coherent flux-rope-like structure there is a local $H_m(t)$ maximum somewhere within the flux rope and $PVI(t)$ yields local maxima at the flux-rope boundaries. For a given time series, local peaks in H_m or PVI can each occur for a variety of independent features, but the combination of two PVI peaks bracketing a local H_m maximum appears to be a fairly robust identification criteria (Pecora et al. 2021).

For completeness, we note there are a number of complementary methods to identify coherent intervals and/or solar wind flux-tube boundaries based on either statistical plasma properties or turbulence measures. For example, rapid changes in the magnetic field orientation (i.e., tangential discontinuities) can be characterized with $\Delta\theta_B$ (e.g., Borovsky 2008, and references therein), and these have recently been shown to coincide with abrupt changes in the suprathermal electron strahl width and/or intensity (Borovsky 2020, 2021; Borovsky et al. 2021).

4.2. Intervals of Helmet Streamer Wind

Figure 6 shows our two composition-enhanced intervals associated with HS wind and the in situ HCS/HPS and interplanetary magnetic field (IMF) sector boundary crossings. Figure 6(a) shows interval #8, which is from DOY 109.582 to 111.267 (40.44 hr total duration), and Figure 6(b) shows interval #3, from DOY 123.644 to 126.164 (60.48 hr total duration). In each plot the top panel shows the (normalized) 272 eV suprathermal electron PAD. The next three panels show the 64 s vector magnetic field in RTN coordinates (B_R , blue; B_T , green; B_N , red) along with its local orientation angles: latitude $\delta \in [-90^\circ, 90^\circ]$ and longitude $\lambda \in [0^\circ, 360^\circ]$. The fourth and

fifth panels show the magnetic helicity measure, H_m^- , and the PVI profiles. The remaining three panels show the 2 hr ionic composition measurements from ACE/SWICS of C^{4-6+} , O^{5-8+} , and Fe^{6-20+} .

The PVI panels show the $PVI \geq 3.0$ threshold as a solid red line and the local maxima are indicated with red diamond symbols. Vertical lines associated with the locations of the PVI peaks are drawn over each panel. The H_m^- profile in Figure 6(a) is normalized by a value of 1.0×10^8 and has a standard deviation of $\sigma = 0.582$. The H_m^- profile in Figure 6(b) has two separate normalizations, as indicated by the additional y-axis at DOY 124.764 due to the magnitude of B increasing after the HCS/HPS crossing. For $t < 124.764$ the normalization value is 1.0×10^8 , whereas for $t > 124.764$ we normalize by 7.68×10^8 so that $\sigma = 1.124$ for both sides. In each of the H_m^- panels the $\pm 1\sigma$ range is shaded light gray. The local H_m^- magnitude maxima larger than 1σ within each PVI interval are indicated with blue diamond plot symbols.

The H_m –PVI procedure identifies a number of coherent magnetic structures throughout each interval. The occurrence of significant PVI peaks tends to be clustered in local patches and each interval’s HCS/HPS crossing (the $\sim 180^\circ$ transition in λ coincident with the bidirectional/broadening of the suprathermal PADs) is bracketed by a cluster of PVI peaks.

In HS interval #8 there are four main clusters of PVI peaks: DOY 109.7–109.9, 110.0–110.15, 110.2–110.6, and 110.9–111.3. The two largest clusters of PVI peaks occur on either side of the HCS/HPS crossing and contain the greatest number of significant H_m peaks. Once the suprathermal electron PAD transitions from a unidirectional (0°) strahl to a broader, more isotropic distribution (DOY 110.5 through 111.1) there is a train of three coherent north-

to-south magnetic field rotations (positive-to-negative profile in δ) at the beginning of the PAD transition and a number of larger H_m structures as the PAD transitions to oppositely directed (180°) strahl on the other side. Notably, the structure centered at DOY 111.0 corresponds to a 1.5 hr wide, relatively flat profile in both δ and λ . Finally, there is a slight increase in O^{7+} (and decrease in O^{6+}) for the duration of the HCS/HPS crossing during the broad electron PAD region which coincides with a slight shift to higher iron charge states during this same period. Likewise, there is a significant increase in C^{6+} at the beginning and end of the PAD transitions at the same time as a number of the HCS/HPS interval-related H_m -PVI structures.

HS interval #3 exhibits many similar features to those of interval #8. For example, significant H_m^- peaks occur on either side of the IMF sector boundary at DOY 124.6, including three consecutive structures between DOY 124.5–124.8, followed by three more, centered on DOY 125.0, 125.1, and 125.2. These H_m -PVI structures are also associated with coherent rotations in (δ, λ) , as well as a sharp local maximum in C^{6+} at the HCS superimposed on top of a broader region of enhanced C^{6+} from DOY 124.0–125.5. The O^{7+} signal shows a similar, but less pronounced, trend over a slightly narrower range (DOY 124.3–125.2). However, the enhanced high iron charge states (up to Fe^{16+}) tend to be shifted earlier (DOY 123.9–124.6) and return to being strongly peaked at Fe^{9-11+} for $t > 125.0$. The suprathermal electron PADs leading up to the HCS crossing are more patchy, alternating between the 180° strahl and broader, more isotropic (and even bidirectional) PADs before transitioning to more steady 0° strahl after DOY 125.0.

Throughout both intervals the PVI peaks occur almost exclusively at discontinuities in the magnetic field angles, and these are often also coincident with changes in the electron PADs. Thus, the conjecture that the PVI peaks select boundaries of distinct plasma intervals (either magnetic flux tubes, discrete solar wind flows, or magnetic island plasmoid/small flux ropes) appears to be supported by our results. Another feature of the H_m -PVI analysis in these intervals is that, even when the H_m profiles do not exceed the 1σ significance threshold, there are often still local peaks and coherent magnetic field signatures within the bracketing PVI peaks.

4.3. Intervals of Pseudostreamer Wind

Figure 7 shows the remaining composition-enhanced intervals associated with non-HS wind, i.e., from PS or PS-adjacent source regions, in the same format as Figure 6. Figure 7(a) shows interval #7, which is from DOY 113.0 to 114.6 (38.4 hr duration), Figure 7(b) shows interval #6, from DOY 115.6 to 117.8 (52.8 hr), Figure 7(c) shows the combined intervals of #4 and #5, from DOY 118.2 to 120.4 (52.8 hr), and, lastly, Figure 7(d) shows the ICME interval #2 from DOY 129.2 to 130.7 (36 hr) and the subsequent, brief PS interval #1, from DOY 130.7 to 131.5 (19.2 hr).

The qualitative features of the HS intervals described above are also present in each of the PS intervals. Specifically, the H_m -PVI analysis continues to identify magnetic field discontinuities and/or sudden changes in the electron PAD via the significant PVI peaks, the PVI peaks clearly show clustering, and these peaks often bracket significant local maxima in the H_m^- magnitude. The H_m normalization for intervals #7, #6 is 1.0×10^8 , which results in a standard deviation of $\sigma = 0.940$ for #7 and $\sigma = 0.768$ for #6.

In interval #7 (Figure 7(a)), there is a series of significant H_m peaks from DOY 113.1–113.55 that begin before the cluster of PVI > 3 events ranging from DOY 113.35–113.75 and another succession of H_m peaks coincident with the next large PVI cluster at $t \gtrsim 114.2$. During interval #6 (Figure 7(b)), the PVI clusters are more frequent and of shorter duration, whereas the significant H_m peaks are more spread out over the entire interval. The overlap between the two occur primarily for DOY 116.3–116.7 and for $t > 116.9$. Essentially the entire #7 interval has a moderate enhancement of C^{6+} , but almost no corresponding enhancement in the hotter charge states of oxygen or iron. Interval #6 is similar with perhaps a very slight enhancement in Fe^{10-12+} between DOY 116–117, and a more prominent C^{6+} region for $t > 117$. Additionally, there is one 2 hr data point centered at 117.5 that includes a slight increase in O^{7+} (and decrease in O^{6+}), coincident with a coherent magnetic structure interval. In general, PS intervals #7 and #6 can be considered to have fewer composition enhancements than our previous HS intervals, and, while there are still some discrete regions of broader electron PAD signatures, especially in #7, for the most part these PS intervals have less variation in the PAD profiles—as may be expected for unipolar PS solar wind.

PS intervals #5+4 and #2+1, shown in Figures 7(c) and 7(d), respectively, also show PVI clusters and sequential trains of H_m coherent structures bracketed by PVI peaks. The H_m normalization in interval #5+4 is 1.0×10^8 , resulting in a standard deviation of $\sigma = 1.608$. We use the same normalization (10^8) for the ICME interval (#2), which yields $\sigma = 2.902$, while for the trailing PS interval (#1) we use a normalization of 3.877×10^7 to obtain the matching σ value.

PS interval #5+4 has the most enhanced heavy ion charge states of our PS intervals, including a significant increase in C^{6+} and O^{7+} from DOY 118.7–119.6, coincident with H_m peaks at DOY 118.9, 119.1, 119.2, and 119.3. Interval #5+4 also contains highly variable and enhanced hot iron charge states, $Fe^{\geq 12+}$, including a Fe^{16+} component present throughout almost the entire time range, DOY 118.5–120.2. Additionally, there are (small) flux-rope-like rotations in the DOY 118.4–118.5 and 119.7–119.8 structures. Again, the PVI peaks representing coherent structure boundaries are seen to line up with discontinuities in magnetic field (δ, λ) angles.

While there is interesting, composition-enhanced internal magnetic structuring present within the ICME interval of Figure 7(d), including large ICME boundary enhancements in the iron distribution (e.g., DOY 129.3–129.7 and 130.0–130.4), in this work we will concentrate on PS interval #1. The Fe^{16+} component is also present for a large percentage of this interval, through DOY 131.3. There is an intriguing sequence of short, intermittent bursts of bidirectional electrons from DOY 130.6–131.0 which have corresponding H_m structures that do not exceed the 1σ threshold but occur toward the latter portion of an extended PVI peak cluster. The H_m peaks that do exceed the threshold occur toward the end of interval #1 and into the beginning of interval #2 (DOY 130.2–130.7) and the coherent magnetic structures at $t \gtrsim 131.3$ also show flux-rope-like rotation signatures.

4.4. Statistical Properties

Given the variation and “randomness” of the magnetic field structure(s) and fluctuations within our slow-to-moderate speed, composition-enhanced HS and PS solar wind intervals,

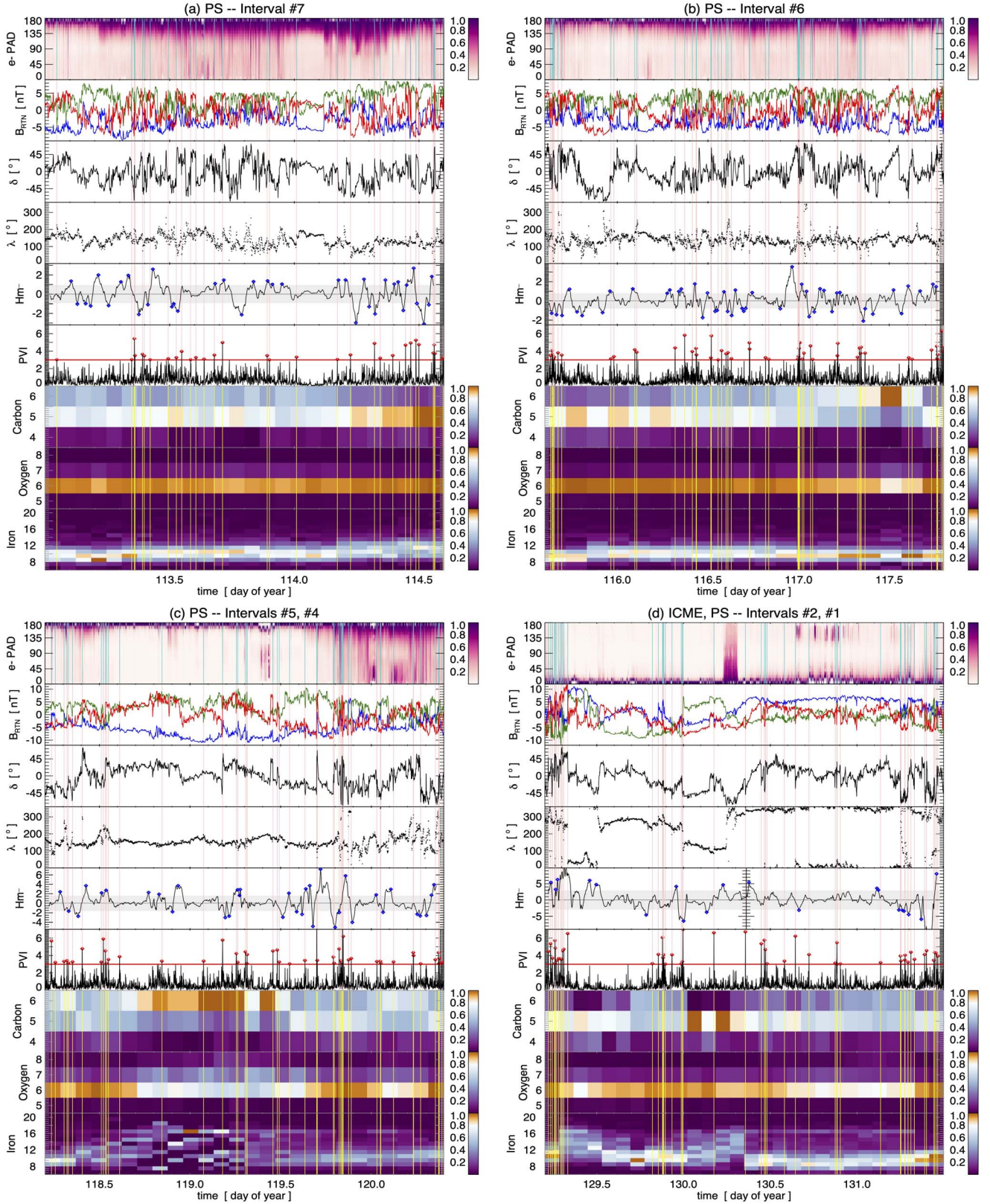


Figure 7. Non-HS, composition-enhanced intervals (primarily from PSs) in the same format as Figure 6. (a) PS (#7) from DOY 113.0 to 114.60. (b) PS (#6) from DOY 115.60 to 117.80. (c) PS (#5, #4) from DOY 118.20 to 120.40. (d) ICME (#2) and PS (#1) from DOY 130.70 to 131.50.

statistical methods are required to characterize various properties of the time series (e.g., Zurbuchen et al. 2000). A summation of these analyses is presented in Figure 8.

Figure 8(a) plots the autocorrelation functions, $A_{Hm}(\tau)$, of the $H_m^-(t)$ time series of the HS intervals (top row) and the PS intervals (bottom row). The average e -folding time, $\langle \tau_{1/e} \rangle$, for

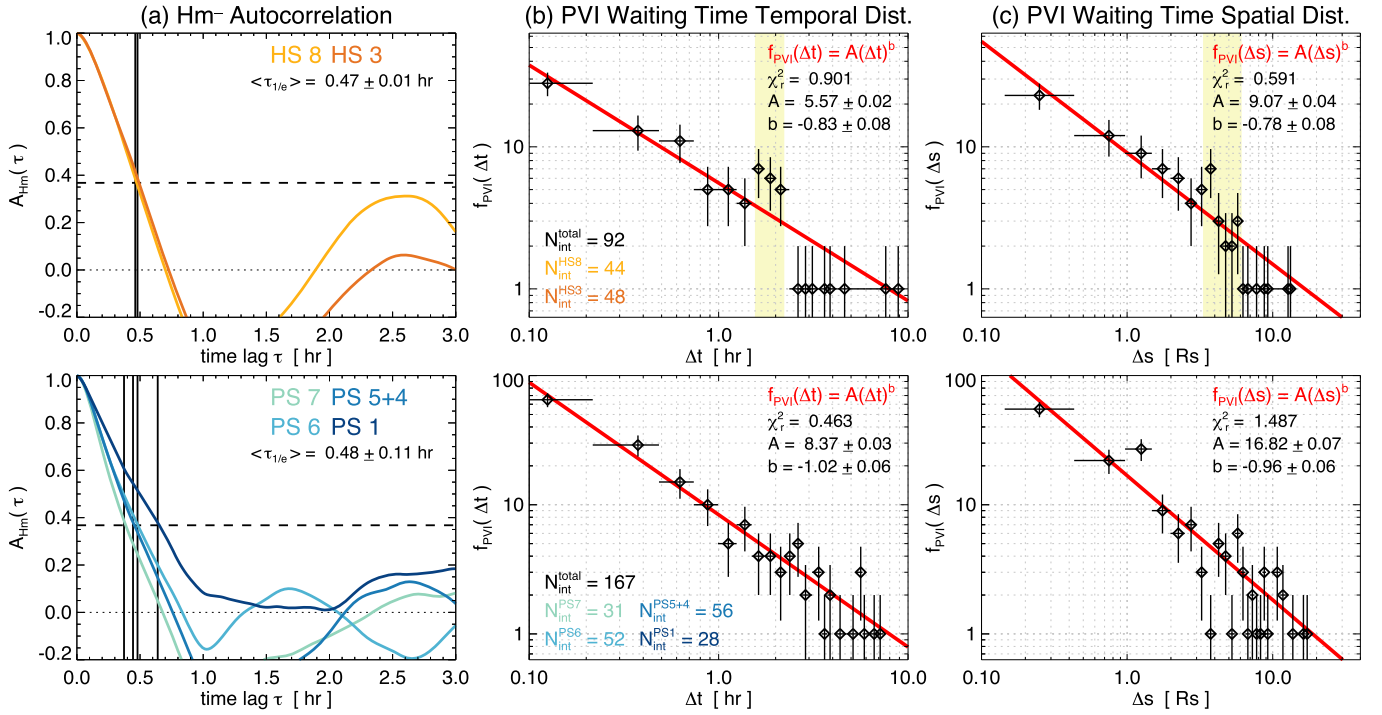


Figure 8. Statistical properties of the intermittency and coherent magnetic structures defined via the H_m -PVI analysis during our slow-to-moderate solar wind intervals with enhanced α/p and heavy ion charge states. (a) Autocorrelations of the $H_m(t)$ profiles from Figures 6 (top) and 7 (bottom). (b) Temporal waiting-time distribution of Δt between PVI peaks for the HS (#8, #3) and PS (#7, #6, #5+4, #1) intervals. (c) Spatial waiting-time distribution of Δs between PVI peaks for the HS and PS intervals. In columns (b) and (c), the red curves show power-law fits to the respective waiting-time distributions.

each set of curves is given in their respective panels. If one defines a characteristic width (duration) of the magnetic helicity-carrying structures as $w = 2\langle \tau_{1/e} \rangle$, then the mean HS interval width is $w_{HS} = 0.94 \pm 0.02$ hr and the mean PS interval width is $w_{PS} = 0.96 \pm 0.22$ hr. These values are consistent with, i.e., on the order of, the ~ 90 minutes periodicity found in solar wind proton density structures (e.g., Viall et al. 2010; Viall & Vourlidas 2015; Kepko et al. 2016; Di Matteo et al. 2019).

Figure 8(b) plots the temporal waiting-time histogram, $f_{PVI}(\Delta t)$, during the HS (top row) and PS (bottom row) intervals. We have fit a line to each of the distributions in log-log space using the IDL `linfit.pro` least-squares minimization procedure representing a $f(x) = Ax^b$ power-law form. The best-fit lines are also plotted in red in each panel and the fit parameters (and their 1σ uncertainties) are given in the plot. The HS and PS distributions have very similar slopes: $b = -0.83 \pm 0.08$ in the HS case and $b = -1.02 \pm 0.06$ for the PS case. If the PVI peaks represent boundaries of coherent magnetic structures, i.e., plasmoid flux ropes or individual flux tubes, then the Δt “waiting time” between PVI peaks should be roughly the flux structure’s diameter (with some variation due to the spacecraft’s relative impact parameter). The mean waiting times are $\langle \Delta t \rangle = 1.10$ hr and 1.01 hr for the HS and PS distributions, respectively. The vertical yellow bar in the HS waiting-time distribution highlights the bins centered at $\Delta t = 1.625$, 1.875 , and 2.125 hr. Each of these bins having counts $\gtrsim 1\sigma$ above the best-fit line may indicate the presence of additional coherent structure at these timescales, which is, again, remarkably consistent with the Viall et al. (2010) ~ 90 minutes timescales for periodic density structures. Interestingly, the PS waiting-time distribution does not appear to have a similar enhancement in the 1.5 – 2 hr scale range,

although the counts in the PS bins at $\Delta t = 2.875$ and 5.62 hr are also on the order of 1σ above the best-fit line.

Figure 8(c) plots the spatial waiting-time histogram, $f_{PVI}(\Delta s)$, in the same format as Figure 8(b). Here, we note the mean spatial lengths for the HS and PS intervals are, again, essentially identical at $\langle \Delta s \rangle = 2.44 R_\odot$ (1698 Mm) and $2.41 R_\odot$ (1677 Mm), respectively. An interesting feature is the “disappearance” of the small enhancement at the 1.5 – 2 hr scale range in the HS PVI waiting-time distribution when plotted as spatial scales. Since we used the radial velocity time series to integrate the distance between PVI peaks rather than a constant V_r value, the PVI Δs distribution is not merely a rescaled version of the Δt distribution. This means, at least in the case of HS slow-to-moderate speed solar wind, that it may be possible to miss a periodic or quasi-periodic signal associated with solar wind formation/source-region properties during its subsequent heliospheric evolution if one is focusing on the spatial domain. Conversely, the counts in the PS $\Delta s = 1.25 R_\odot$ bin are significantly above the power-law fit without an obvious corresponding enhancement in the PS Δt distribution. The average solar wind speed obtained from the first moment of the temporal and spatial times is $\langle V_r \rangle = \langle \Delta s \rangle / \langle \Delta t \rangle = 429$ km s $^{-1}$ for the HS intervals and $\langle V_r \rangle = 461$ km s $^{-1}$ for the PS intervals. These values appear to be slightly lower than the averages obtained directly from the $V_r(t)$ profiles during our composition-enhanced intervals (Table 1, Figures 2–4).

Our PVI waiting-time statistics seem compatible and consistent with previous applications of these analyses; at scales below the magnetic correlation scale, the PVI waiting-time distribution is well approximated by a power law, and at scales greater than the correlation scale the distribution takes on more of the classic Poisson waiting-time exponential form (Greco et al. 2009a, 2009b). The temporal/spatial plots in

Figures 8(b) and (c) show a consistent departure/rollover from the best-fit line for $\Delta t \gtrsim 2.4$ hr ($\Delta s \gtrsim 6 R_\odot$) and the first moments of the waiting times/length scales ($\langle \Delta t \rangle$, $\langle \Delta s \rangle$) are on the order of the associated correlation scales (see Section 4.1). In fact, the range of values we obtain for the power-law fit exponents (-1.02 to -0.78) are entirely consistent with those found by Greco et al. (2009b) in MHD turbulence simulation data (-0.92) and in the solar wind at 1 au (-1.29), and in PSP observations of the PVI > 3 magnetic field fluctuations at ~ 0.25 au (-1.29 to -0.83 ; Chhiber et al. 2020).

5. Summary and Discussion

It is well established that the in situ solar wind composition, its variation, and its associated plasma structures are all remnant signatures of the physical processes of solar wind formation and the coronal conditions of its origin. We have presented a comprehensive analysis of a set of slow-to-moderate speed, composition-enhanced solar wind intervals at 1 au during CR 2002. Our intervals were selected on the basis of solar wind speed and observed enhancements in some combination of n_p , n_α , A_{He} , or their variability. We have shown that each of these intervals correspond to solar wind flows with complex, broadened, or bidirectional suprathermal electron strahl, elevated (hot) ionic charge states of carbon, oxygen and iron, and an enhanced Fe/O ratio.

PSs are a prime location for interchange reconnection and they are thought to be responsible for a component of intermittent, slow solar wind outflow (e.g., Masson et al. 2012; Wang et al. 2012; Higginson et al. 2017b; Wang & Panasenco 2019). In general, energizing surface flows (e.g., translation or rotational shearing flows, flux emergence, and/or flux cancellation/tether cutting) will build up volumetric currents, stress magnetic null points, and develop strong current sheets at topological boundaries, thereby creating favorable conditions for magnetic reconnection (e.g., Antiochos et al. 2012; Rappazzo et al. 2012; Burkholder et al. 2019; Mason et al. 2021).

Lynch & Edmondson (2013) showed that 2.5D PS interchange reconnection (in the form of pre-eruption breakout reconnection) could result in bursty, quasi-steady signatures in density along the external spine and coronal dimming signatures near the stressed null point and current sheet (see also Kumar et al. 2021), while recent simulations from Aslanyan et al. (2021, 2022) have illustrated that the complex 3D interchange reconnection dynamics seen by Higginson et al. (2017a, 2017b) can also be produced at the open–closed boundaries of PS flux systems.

There is an implicit relationship between the Zhao et al. (2017) source-region categories and the large-scale coronal magnetic topology in the neighborhood of the PFSS field line foot points. For example, their “Quiet Sun,” “Active Region,” and “Active Region Boundary” classifications—typically thought of as closed-flux regions—are likely to be associated with structures giving rise to the S-Web, i.e., PSs and small/narrow open field regions such as low-latitude coronal holes. With the application of standard back-mapping techniques, we showed that the slow-to-moderate speed, composition-enhanced solar wind intervals at 1 au map to large-scale coronal features such as the HS belt and S-Web arcs. These are precisely the locations predicted by Q -map topological analysis to be sites favorable for interchange reconnection during the dynamic evolution of the solar corona’s open–closed flux

boundaries. Lastly, we note that the presence of relatively slow, highly structured, and composition-enhanced solar wind that originates from S-Web arcs far from the HCS is a crucial test of the S-Web theory (e.g., Higginson et al. 2017b; Di Matteo et al. 2019).

We have analyzed the properties of the in situ coherent magnetic structures within each composition-enhanced interval as determined by the Pecora et al. (2021) H_m –PVI procedure for the identification of helicity-carrying flux tubes and/or magnetic island plasmoids. The characteristic widths of these coherent magnetic structures (~ 1 hr from H_m , ~ 2 hr from PVI) are consistent with the ~ 90 minutes periodicities determined from either in situ proton density time series (Viall et al. 2010) or in the Thomson-scattered white-light coronagraph brightness fluctuations that are proportional to the line-of-sight integrated electron density n_e (Viall & Vourlidis 2015). There appears to be a 1.5–2 hr timescale signature above the expected power-law distribution of PVI waiting times in HS-associated solar wind that is either significantly less obvious or nonexistent in our PS intervals. There also appears to be an enhancement of the PS-associated waiting-time length scale $s \sim 1.25 R_\odot$ without a corresponding enhancement in the temporal distribution. One may expect different types of reconnection-generated magnetic structures at the boundaries of HS and PS regions due to the topological differences, for example as discussed by Edmondson & Lynch (2017) and Higginson & Lynch (2018), but further numerical modeling of their origin and heliospheric evolution will be needed.

This work complements previous statistical studies characterizing magnetic field and plasma properties within coherent intervals or by solar wind type (e.g., Ko et al. 2018; Borovsky et al. 2019; D’Amicis et al. 2019; Borovsky et al. 2021), as well as those studies of specific, small-scale structures (Khabarova et al. 2021; Gershkovich et al. 2022) such as small magnetic flux ropes (e.g., Feng et al. 2008; Yu et al. 2016; Murphy et al. 2020; Choi et al. 2021). Importantly, our attempt to relate various in situ properties of the structured variability in slow-to-moderate speed solar wind through an application of the H_m –PVI methodology represents a significant extension of previous work where coherent magnetic structures identified “by eye” were shown to be coincident with structure in the proton density and A_{He} observations (e.g., Kepko et al. 2016; Di Matteo et al. 2019). Given recent interest in the further refinement and development of sophisticated automated methods such as machine-learning/artificial intelligence neural networks, the H_m –PVI procedure appears to be a promising candidate for inclusion in the suite of tools being constructed to classify solar wind types and properties (e.g., as discussed in Section 1) and to identify and characterize coherent flux-rope intervals, ranging in spatiotemporal scales from ICMEs (e.g., Nguyen et al. 2019; dos Santos et al. 2020; Roberts et al. 2020; Narock et al. 2022) to small-scale flux ropes embedded in the slow solar wind and HCS/HPS crossings (e.g., Hu et al. 2018; Zhao et al. 2020).

The results presented herein open up a number of avenues for future research efforts: (i) extending the current analysis to in situ solar wind plasma, field, and composition measurements to many more CRs over different phases of the activity cycle; (ii) performing forward modeling of heavy ion charge states and elemental abundances associated with the spatial distribution of discrete, observer-connected solar wind flux tubes with varying solar wind outflow properties based on coronal

conditions of their foot point locations/source-region topologies; and (iii) further analysis of existing and future numerical MHD simulations of dynamic S-Web outflow and their derived observational signatures.

Since there has been recent progress integrating aspects of heavy ion composition forward modeling into steady-state MHD solar wind calculations (e.g., Oran et al. 2015; Shen et al. 2017; Lionello et al. 2019; Szente et al. 2022), it would be extremely interesting to perform these calculations on dynamic, time-dependent MHD modeling of the formation and evolution of coherent magnetic structures generated under different reconnection scenarios. For example, the Aslanyan et al. (2022) calculation of the synthetic suprathermal electron PAD “time series” associated with PS interchange reconnection outflows shows excellent qualitative agreement with the observed broadening of the strahl for some of our PS intervals (#4, #6, and #7 in particular). Lynch et al. (2014) showed the largest (i.e., low-frequency) $\delta B/\langle B \rangle$ signatures resulting from PS reconnection had characteristic length scales of 100–350 Mm (0.14–0.50 R_\odot) in the corona, which reflected the spatial scale of their PS flux system of origin, and Higginson & Lynch (2018) demonstrated that the MHD-simulation-derived, synthetic in situ magnetic field signatures of a similarly sized, nonlinear torsional Alfvén wave could resemble the coherent magnetic structure of small-scale magnetic flux ropes/streamer blob plasmoids typically associated with HS slow wind in the vicinity of the HCS/HPS.

On the largest scales (the tens of hours of our interval durations), there is a remarkably clear association between our HS and PS S-Web arc intervals and in situ composition enhancements. On the scales of the coherent magnetic structures depicted in Figures 6–7, there are some indications that the PVI boundaries are also associated with discrete changes in coronal freeze-in temperatures as inferred from the heavy ion charge states. The 2 hr cadence of the ACE/SWICS data used herein obviously limits our ability to resolve charge-state structure below the averaging window duration. Smaller-scale features have been observed and reported in Kepko et al. (2016) and Gershkovich et al. (2022) using periods of high-cadence (12 minutes native instrument resolution) ACE/SWICS data to argue that some of the discrete magnetic flux-tube intervals of interest did line up with sudden changes in various composition measures (i.e., helium, carbon and oxygen abundances, the C^{6+}/C^{5+} charge-state ratio, etc). Measurements from the Heavy Ion Sensor (with a native resolution of 30 s for heavy ions), part of the Solar Orbiter Solar Wind Analyser instrument suite (Owen et al. 2020), should enable the identification and characterization of smaller-scale associations of coherent magnetic structures with in situ composition enhancements.

Additionally, the scientific importance for multispacecraft measurements and remote -sensing and in situ quadrature observational geometries to establish the solar–heliospheric connection for specific plasma features of well-observed interchange reconnection events has been recently demonstrated by Telsoni et al. (2022), with the first direct imaging of a “switchback” with Solar Orbiter’s Metis coronagraph (Antonucci et al. 2020). This switchback event’s likely origin from the complex S-Web configuration of a small PS S-Web arc coming off the main HS belt/HCS and a null-point spine–fan curtain topology where the PS and HS flux systems intersect strongly motivates continued theoretical development,

data analysis, and numerical simulations of the dynamic S-Web model for the slow solar wind.

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