

MUSE-ALMA Haloes XI: gas flows in the circumgalactic medium

Simon Weng^{1b},^{1,2,3,4★} Céline Péroux^{1b},^{1,5} Arjun Karki,⁶ Ramona Augustin^{1b},⁷ Varsha P. Kulkarni,⁶ Aleksandra Hamanowicz^{1b},⁷ Martin Zwaan,¹ Elaine M. Sadler^{1b},^{2,3,4} Dylan Nelson^{1b},⁸ Matthew J. Hayes^{1b},⁹ Glenn G. Kacprzak^{1b},^{3,10} Andrew J. Fox,^{11,12} Victoria Bollo,¹ Benedetta Casavecchia¹³ and Roland Szakacs¹

¹European Southern Observatory, Karl-Schwarzschildstrasse 2, D-85748 Garching bei München, Germany

²Sydney Institute for Astronomy, School of Physics A28, University of Sydney, NSW 2006, Australia

³ARC Centre of Excellence for All Sky Astrophysics in 3 Dimensions (ASTRO 3D), Australia

⁴ATNF, CSIRO Space and Astronomy, PO Box 76, Epping, NSW 1710, Australia

⁵Aix Marseille Université, CNRS, LAM (Laboratoire d'Astrophysique de Marseille) UMR 7326, F-13388 Marseille, France

⁶Department of Physics and Astronomy, University of South Carolina, Columbia, SC 29208, USA

⁷Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA

⁸Universität Heidelberg, Zentrum für Astronomie, Institut für theoretische Astrophysik, Albert-Ueberle-Str 2, D-69120 Heidelberg, Germany

⁹Stockholm University, Department of Astronomy and Oskar Klein Centre for Cosmoparticle Physics, AlbaNova University Centre, SE-10691 Stockholm, Sweden

¹⁰Centre for Astrophysics and Supercomputing, Swinburne University of Technology, Hawthorn, Victoria 3122, Australia

¹¹AURA for ESA, Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA

¹²Department of Physics & Astronomy, Johns Hopkins University, 3400 N. Charles St., Baltimore, MD 21218, USA

¹³Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Strasse 1, D-85748 Garching b. München, Germany

Accepted 2023 May 11. Received 2023 May 11; in original form 2023 April 12

ABSTRACT

The flow of gas into and out of galaxies leaves traces in the circumgalactic medium which can then be studied using absorption lines towards background quasars. We analyse 27 $\log[N(\text{H I})/\text{cm}^{-2}] > 18.0$ H I absorbers at $z = 0.2$ to 1.4 from the MUSE-ALMA Haloes survey with at least one galaxy counterpart within a line of sight velocity of $\pm 500 \text{ km s}^{-1}$. We perform 3D kinematic forward modelling of these associated galaxies to examine the flow of dense, neutral gas in the circumgalactic medium. From the VLT/MUSE, *HST* broad-band imaging, and VLT/UVES and Keck/HIRES high-resolution UV quasar spectroscopy observations, we compare the impact parameters, star-formation rates, and stellar masses of the associated galaxies with the absorber properties. We find marginal evidence for a bimodal distribution in azimuthal angles for strong H I absorbers, similar to previous studies of the Mg II and O VI absorption lines. There is no clear metallicity dependence on azimuthal angle, and we suggest a larger sample of absorbers is required to fully test the relationship predicted by cosmological hydrodynamical simulations. A case-by-case study of the absorbers reveals that ten per cent of absorbers are consistent with gas accretion, up to 30 per cent trace outflows, and the remainder trace gas in the galaxy disc, the intragroup medium, and low-mass galaxies below the MUSE detection limit. Our results highlight that the baryon cycle directly affects the dense neutral gas required for star-formation and plays a critical role in galaxy evolution.

Key words: galaxies: evolution – galaxies: formation – galaxies: haloes – galaxies: kinematics and dynamics – quasars: absorption lines.

1 INTRODUCTION

Once known as ‘island universes’, galaxies are no longer considered isolated systems, especially when observing the cycle of baryons within these systems (Péroux & Howk 2020). The formation of stars is mediated by the delicate balance between outflowing and inflowing gas. Galactic winds driven by active galactic nuclei (AGN) or intense star formation remove gas from galaxies (Veilleux, Cecil & Bland-Hawthorn 2005). Some of the outflowing material condenses

back onto the galaxy in the form of fountains (Fraternali & Binney 2008; Marinacci et al. 2010; Fraternali 2017). At the same time, the accretion of cold gas from large-scale filaments and the cooling of hot halo gas replenish the gas reservoirs of galaxies (Kereš et al. 2005; Nelson et al. 2013; Hafen et al. 2022). Satellites embedded within the halo of galaxies are also an important source of cool gas (Wang 1993; Hafen et al. 2019). These phenomena leave their traces in the circumgalactic medium (CGM), the region that extends from the galactic disc to the intergalactic medium (IGM; Tumlinson, Peebles & Werk 2017; Faucher-Giguere & Oh 2023).

While the diffuse gas in the CGM can be studied in emission, it is typically detected in systems containing quasars, extreme starbursts,

★ E-mail: swen2649@uni.sydney.edu.au

and overdensities not representative of typical galaxies (e.g. Hayes et al. 2016; Epinat et al. 2018; Johnson et al. 2018; Chen et al. 2019; Burchett et al. 2021; Cameron et al. 2021; Helton et al. 2021). Stacking enables detections of the diffuse gas in emission, but it becomes difficult to link the detections to galaxy properties and the wider environment (Steidel et al. 2011; Momose et al. 2014; Chen et al. 2020b; Dutta et al. 2023). Typical galaxies on the main sequence require exceptional exposure times to obtain a detection in emission (Zabl et al. 2021; Leclercq et al. 2022; Bacon et al. 2023). Alternatively, analysing absorption lines in the spectra of bright background sources such as quasi-stellar objects (QSOs) allows us to probe the gas in the CGM to lower column densities. While these sightlines only sample the CGM on parsec scales, there has been an abundance of recent surveys examining the galaxies around H I (Lofthouse et al. 2020; Muzahid et al. 2020; Chen et al. 2020a; Berg et al. 2023, Karki et al., in preparation), Mg II (Bouché et al. 2016; Nielsen et al. 2020; Lundgren et al. 2021), and other gas phases.

Early studies of strong H I Ly- α absorbers associated them with the rotating discs of galaxies (Wolfe et al. 1986; Wolfe, Gawiser & Prochaska 2005). Damped Ly- α absorbers (DLAs) with column density $\log[N(\text{H I})/\text{cm}^{-2}] \geq 20.3$ are expected to originate from gas within ~ 20 kpc of the galaxy centre (Péroux et al. 2005; Zwaan et al. 2005; Stern et al. 2021). Since, the increasing number of galaxy-absorber systems brought by the advent of integral field spectroscopy suggest these absorbers can originate from beyond the disc. In particular, H I absorbers with lower column density found at larger impact parameters suggest different phenomena such as gas accretion and outflows are being traced.

The time galaxies take to deplete their molecular gas reservoirs is found to be shorter than the Hubble time at all redshifts, meaning that galaxies must have a way to maintain a gas supply through accretion (Daddi et al. 2010; Genzel et al. 2010; Péroux & Howk 2020; Tacconi, Genzel & Sternberg 2020; Walter et al. 2020). Some theoretical studies have attempted to separate gas accretion into cold-mode and hot-mode accretion. The former refers to the accretion of cold gas from filaments in the intergalactic medium and is expected to dominate at redshift $z \gtrsim 3$ (Kereš et al. 2005; Nelson et al. 2013). Simulations predict that the accretion produces flows that co-rotate with the galaxy disc out to ~ 100 kpc (Stewart et al. 2011; van de Voort et al. 2011; Stewart et al. 2013) and can be traced by dense H I absorbers (van de Voort et al. 2012; Theuns 2021). In contrast, the cooling of hot halo gas is also expected to grow gas reservoirs and dominate at lower redshifts and/or higher masses (Dekel et al. 2009a; Dekel, Sari & Ceverino 2009b). The accretion of hot gas transitions from a spherical geometry to a disc when it cools because of angular momentum conservation (Mo, Mao & White 1998; Hafen et al. 2022). While gas accretion at $z \lesssim 1$ has been probed using Mg II absorbers (Ho et al. 2017; Martin et al. 2019; Zabl et al. 2019), studies of accretion using Ly- α absorbers remain limited.

In contrast to accretion, neutral gas in outflows has been ubiquitously observed in galaxies when the SFR surface density is sufficiently high (e.g. Heckman & Thompson 2017; Hayes 2023) and in galaxies containing an active galactic nucleus (e.g. Ciccone et al. 2014, 2015). While there remains uncertainty regarding how the cool gas survives in the hot and turbulent wind environment, recent idealized simulations of the cold phase in a turbulent medium find that cold gas clouds can survive and even grow depending on their size (Gronke et al. 2022). Traditionally, outflowing gas has been observed using down-the-barrel observations by identifying blueshifted absorption against the galaxy's stellar continuum (Martin 2005; Veilleux et al. 2005; Rubin et al. 2014; Heckman & Thompson 2017). Identifying outflows using absorbers towards background

sources is a complementary method where the location of the absorbing gas is precisely measured (Bouché et al. 2012; Kacprzak et al. 2014; Schroetter et al. 2019). However, the velocity degeneracy between inflows and outflows increases the difficulty in identifying gas flows for these transverse absorption-line studies. Observations and simulations of gas outflows that emerge from the galaxy centre suggest they form an expanding biconical shape perpendicular to the galaxy disc, as this is the path of least resistance (Bordoloi et al. 2011; Lan, Ménard & Zhu 2014; Nelson et al. 2019b). Hence, quasar absorbers have been interpreted to be outflowing when they are aligned with the projected minor axis of galaxies (Rahmani et al. 2018b; Schroetter et al. 2019). Transverse absorption-line studies of outflowing gas at $z > 0.2$ typically use the Mg II ion, but there have been limited studies on Ly- α absorbers that directly trace the neutral hydrogen required for star formation.

On the other hand, Ly- α absorbers at $z \lesssim 1$ have a bimodal distribution in the gas-phase metallicity (Lehner et al. 2013, 2022b). Here, metal-poor systems are expected to trace the accretion of cold gas, while metal-rich absorbers trace outflows, recycled or tidally stripped gas (Péroux et al. 2020; Lehner et al. 2022b). Indeed, studies of high-velocity clouds (HVCs) in the Milky Way supplement velocity measurements of the HVC with metallicity to distinguish between inflows and outflows (Fox et al. 2019; Ramesh, Nelson & Pillepich 2023). For extragalactic sources, the imprint of inflows and outflows in the CGM might also be seen in the bimodal distribution of azimuthal angles (Φ) between a galaxy's major axis and the absorber sightline (where values close to 90° indicate the gas is located near the projected galaxy minor axis; Bouché et al. 2012; Kacprzak, Churchill & Nielsen 2012; Bordoloi et al. 2014; Schroetter et al. 2019). Naturally, it follows that there should be a metallicity dependence on azimuthal angle, where gas located towards the minor axis is metal-enriched as outflows eject metals produced in stars (Nelson et al. 2019b; Péroux et al. 2020; Truong et al. 2021; van de Voort et al. 2021). However, in observations, this phenomenon remains unclear, as studies have found no such trend (Péroux et al. 2016; Kacprzak et al. 2019; Pointon et al. 2019), perhaps due to uncertainties in the ionization correction and dust modelling of absorbers or poor metal-mixing in the CGM (Zahedy et al. 2019, 2021; Bordoloi et al. 2022; Sameer et al. 2022). Simulations of Mg II absorption around galaxies find that the equivalent width (EW) of the line decreases as the azimuthal angle of the sightline increases for a fixed impact parameter (DeFelippis et al. 2021). Interestingly, Wendt et al. (2021) finds that gas located near the minor axis is more dust-depleted and perhaps metal-enriched than absorbers near the major axis. More carefully selected observational samples are required to reach a consensus on the metallicity–azimuthal angle dependence, especially because the relationship is dependent on the galaxy's stellar mass, impact parameter, redshift, and H I column density (Péroux et al. 2020).

The MUSE-ALMA Haloes survey targets 32 Ly- α absorbers with column densities $\log[N(\text{H I})/\text{cm}^{-2}] > 18.0$ at redshift $0.2 \lesssim z \lesssim 1.4$ (see Péroux et al. 2022, for an overview). In total, 79 galaxies are detected within $\pm 500 \text{ km s}^{-1}$ of the absorbers at impact parameters ranging from 5 to 250 kpc (Weng et al. 2023). The stellar masses of these associated galaxies have been measured from broad-band imaging in Augustin et al. (in preparation) and their morphologies studied in Karki et al. (in preparation). Gas flows for a subsample of galaxies have already been studied in earlier works (Péroux et al. 2017; Klitsch et al. 2018; Rahmani et al. 2018a, b; Hamanowicz et al. 2020; Szakacs et al. 2021), and our work is a continuation of these studies for the full MUSE-ALMA Haloes survey. With the complete sample of absorbers and their galaxy counterparts, we examine the

Table 1. Summary of kinematic properties for the modelled galaxies. The kinematic and stellar properties of galaxies within $\pm 500 \text{ km s}^{-1}$ of H I absorbers (highlighted in bold). For each modelled galaxy, the identification and emission line used for modelling is given. Velocity differences (Δv) are calculated with respect to the absorber redshift. Star-formation rates (SFRs) and limits are based on Kennicutt (1998), and we use the H α empirical relation when available. At $z \gtrsim 0.4$, we use the [O II] $\lambda\lambda 3726, 3729$ luminosity to estimate the SFR. Dust corrections are performed for galaxies with measured H β and H α emission-line fluxes. Stellar mass estimates are derived from SED fitting with LEPHARE using *HST* broad-band imaging (Arnouts & Ilbert 2011, Augustin et al., in preparation). The maximum rotational velocity (V_{max}), velocity dispersion (σ), and inclination (i) measurements are derived from GALPAK. We adopt the photometric position angle measured using GALFIT (Peng et al. 2002) on the *HST* photometry (Karki et al., in preparation.) when the kinematic PA is not available from GALPAK. The azimuthal angle, Φ , is preferentially calculated using PA_{kin} when available. A value of 90° means that the absorber is orientated along the minor axis of the galaxy. Whether the absorber is tracing gas that is part of the disc (Disc), inflowing (In) or outflowing (Out), part of the intragroup medium (IGrM) or ambiguous (A) is given in the last column. This label is assigned to only one galaxy in an absorber-galaxy system unless the gas is in the intragroup medium or ambiguous. A machine-readable table of associated galaxies and their properties is available online. We note that the online table contains both kinematic and photometric inclination and position angle measurements. Entries where no measurements are available are filled as –999.

| ID | Line | b (kpc) | Δv (km s^{-1}) | SFR ($M_\odot \text{ yr}^{-1}$) | $\log M_*/M_\odot$ | V_{max} (km s^{-1}) | σ (km s^{-1}) | i ($^\circ$) | PA ($^\circ$) | Φ ($^\circ$) | Flow |
|--|------------|--------------|--------------------------------------|--------------------------------------|--------------------|--|------------------------------------|---------------------|--------------------|------------------------|------|
| Q0138–0005, $z_{\text{quasar}} = 1.34$, $z_{\text{abs}} = 0.7821$, $\log[N(\text{H I})/\text{cm}^{-2}] = 19.81 \pm 0.08$ | | | | | | | | | | | |
| Q0138m0005_14 | [O II] | 82 | –3.4 | 6.9 ± 2.4 | 9.8 ± 0.2 | 100 ± 22 | 141 ± 6 | 67 ± 9 | 100 ± 6 | 80 ± 6 | A |
| Q0152–2001, $z_{\text{quasar}} = 2.06$, $z_{\text{abs}} = 0.383$, $\log[N(\text{H I})/\text{cm}^{-2}] < 18.78$ | | | | | | | | | | | |
| Q0152m2001_5 | H α | 60 | 84 | 0.73 ± 0.4 | 11.25 ± 0.14 | 116 ± 1 | 14 ± 2 | 79 ± 1 | 142.1 ± 0.3 | 9.3 ± 0.3 | In |
| Q0152m2001_7 | H α | 150 | 350 | 0.18 ± 0.4 | 11.00 ± 0.13 | 162 ± 3 | 4 ± 4 | 79 ± 2 | 129 ± 2 | 58 ± 2 | |
| Q0152m2001_13 | H α | 84 | 330 | 0.10 ± 0.4 | 10.5 ± 0.2 | 75 ± 6 | 5 ± 4 | 54 ± 6 | 153 ± 4 | 6 ± 4 | |
| ... | ... | ... | ... | ... | ... | ... | ... | ... | ... | ... | ... |

azimuthal dependence of metallicity in the circumgalactic medium and identify the gas flows being probed by QSO sightlines. We adopt the following Λ CDM cosmology: $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$.

2 IONIZED GAS MAPS

To determine the origin of the gas, that is, whether it traces inflows, outflows or other phenomena, we need to compare the ionized gas kinematics with the neutral-phase absorber kinematics. We create continuum-subtracted cubes centred on the [O II] $\lambda\lambda 3727, 3729$, H β $\lambda 4861$, [O III] $\lambda 5007$, and H α $\lambda 6563$ emission lines for each galaxy when available within the MUSE wavelength range (4700 to 9300 Å). We use the GALPAK algorithm v1.32 to extract intrinsic galaxy parameters such as the inclination (i) and kinematic position angle (PA_{kin}) and velocity maps from data cubes. This forward modelling approach assumes a disc model, and we refer readers to Bouché et al. (2015) for a complete description of the method. For the code to provide stable results, the signal-to-noise ratio (S/N) of the nebular line at the brightest spaxel needs to meet the requirement $\text{S/N} > 3$. Additionally, the algorithm is robust provided the galaxy is not too compact; the prerequisite $R_e/R_{\text{PSF}} > 1.50$ should be met, where R_{PSF} is half the full width at half maximum (FWHM) of the PSF and R_e is the effective radius.

From the initial sample of 79 associated galaxies, 48 satisfy the signal-to-noise ratio and projected galaxy size criteria. We run the algorithm for all the nebular lines ([O II], H β , [O III], and H α). For galaxies where multiple emission lines are available, we find the measured parameters agree within 10 per cent for any given galaxy. We preferentially adopt the values returned by the H β , [O III], and H α lines to avoid complications associated with fitting the [O II] doublet. The maximum rotational velocity, inclination, and kinematic position angle are tabulated in Table 1. We also include photometric position angles (PA_{phot}) from GALFIT fits of the galaxy using the *HST* broad-band imaging in the online version of the table (Karki et al., in preparation).

The smoothed rotational velocity maps for three example galaxies are shown in Fig. 1 and the remainder are found in Appendix A (Figs A1 to A9). Each map is accompanied by dashed lines marking projected angles of $\pm 30^\circ$ away from the minor axis. The arrow points in the direction of the QSO sightline and consists of two halves coloured by the impact parameter and line of sight velocity difference between galaxy and absorber. All galaxies located at the smallest impact parameter to the absorber have their velocity maps emboldened. We provide the emission line used to fit each galaxy, identification and absorber redshift below the velocity map. For each fitted galaxy, we also display the observed flux and flux residual maps from the GALPAK fit. Regions of the flux map that have a white background are pixels that were masked during fitting or are outside the MUSE field of view. Pixels were masked to remove the flux contribution from the nearby QSO. The arrow below the residual maps demarcates the physical size of the cube used to fit the emission.

Ten galaxy fits have significant flux residuals and can be attributed to several causes. The residuals for four galaxies in the Q1130–1449 field associated with the absorber at $z_{\text{abs}} = 0.3127$ arise because of an extended ionized gas nebula permeating the large galaxy group (Kacprzak et al. 2010; Chen et al. 2019; Péroux et al. 2019). Galaxies Q1130–1449_6 and Q1130–1449_8 additionally have significant residuals due to possible mergers and outflows, respectively. There are also significant flux residuals for galaxy Q0454–220_4 due to the flux saturation of IFUs in the MUSE data from a nearby bright star. Another two galaxies (Q1211+1030_7 and Q1229–021_6) likely host active galactic nuclei based on their positions on the [O III] $\lambda 5007/\text{H } \beta$ versus [O II] $\lambda\lambda 3727, 29/\text{H } \beta$ classification diagram (Lamareille 2010; Weng et al. 2023). As GALPAK compares a disc model to the data using forward modelling, the flux excess is likely caused by the AGN. Finally, the significant residuals at the centre of source Q0454–220_69 are a result of flux contamination from the QSO less than two arcseconds away. We exclude these objects with unreliable fits from further analysis but include and note them in the figures for completeness.

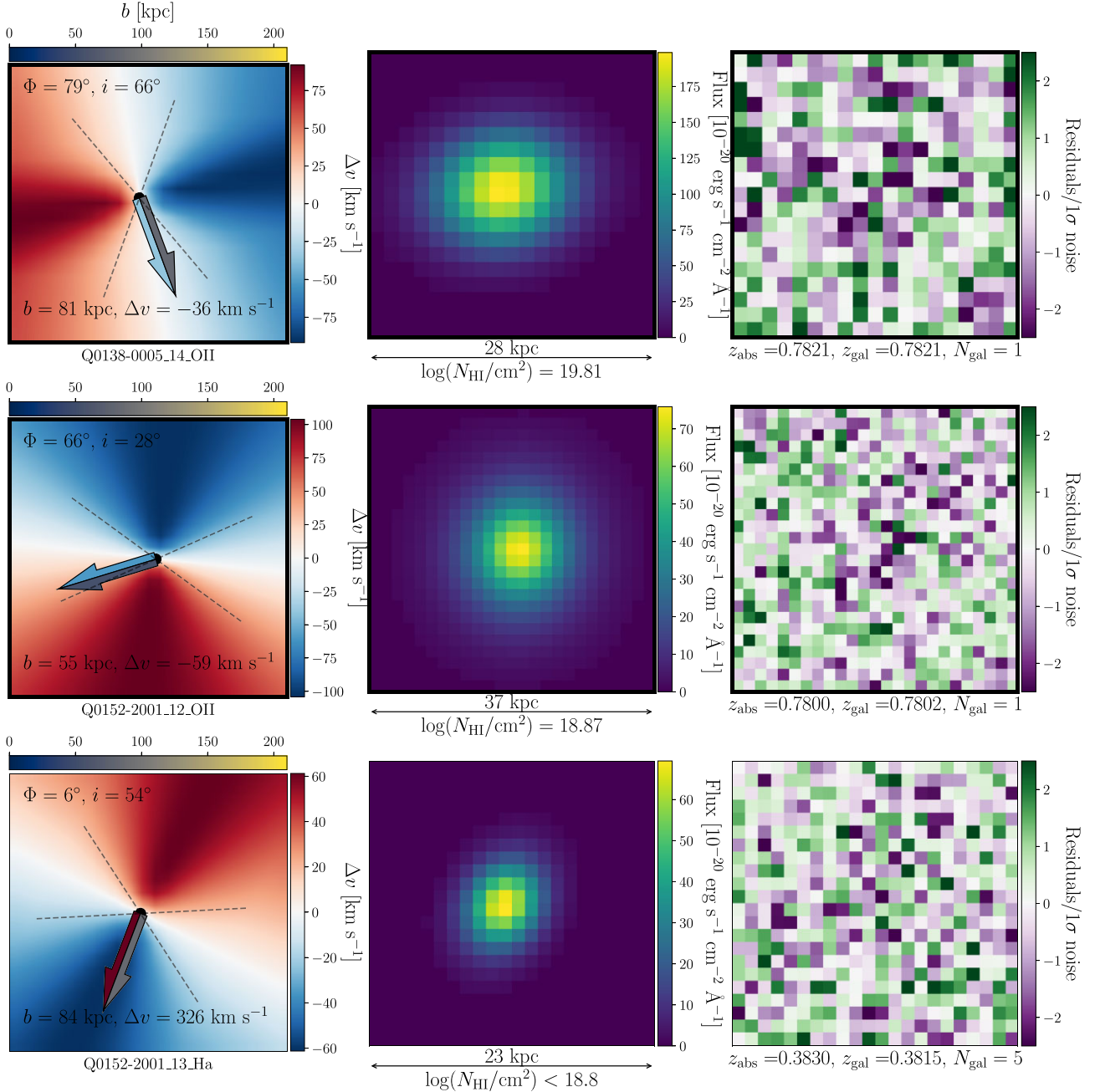


Figure 1. The modelled velocity, observed flux, and residual maps of three randomly chosen galaxies associated with QSO absorbers that were fitted using GALPAK. Each galaxy is represented by three maps. The modelled velocity maps in column one have been smoothed using repeated linear interpolation. A black dot indicates the modelled galaxy centre and two dashed lines represent 2D projected angles of $\pm 30^\circ$ from the minor axis. The arrow points in the direction of the H I absorber, and consists of two coloured halves. The halves are coloured by the impact parameter and line of sight velocity difference between galaxy and absorber. When pointing up, the left half corresponds to the impact parameter, and the reverse is true when pointing down. We include text with the azimuthal angle (Φ) and inclination (i) along with the impact parameter (b) and line of sight velocity difference (Δv) for each velocity map. Below each velocity map is the galaxy identification and emission line fitted. The observed flux maps are shown in column two. White pixels (not shown here) correspond to regions that are outside the MUSE field-of-view or were masked during fitting. The extent of the spatial co-ordinates of the cube used to model the emission and column density of the absorber are given below each residual map. Residual flux maps are found in the third column, where each pixel is coloured by the residual (data-model) normalized by the noise. The absorber and galaxy redshift along with the number of associated galaxies within $\pm 500 \text{ km s}^{-1}$ of the absorber, are written below the residual map. Associated galaxies that have the smallest impact parameter on the absorber have a bold border around the plots.

3 METALLICITY DEPENDENCE ON GALAXY ORIENTATION

Out of the initial 79 galaxies associated with 32 Ly- α absorbers, 48 have kinematic position angle measurements from GALPAK. These

measurements are complemented by an additional 19 photometric PA measurements using GALFIT on the *HST* broad-band imaging (Karki et al., in preparation). For galaxies with both measurements, the position angles agree within their 1σ uncertainties. To calculate

Table 2. A sample summary of modelled galaxies. We consider only the sample of modelled galaxies with inclinations $i > 30^\circ$ to limit potential errors in the position angle measurement.

| Criterion | Tool | Number | Total |
|--------------------------|--------|--------|-----------|
| Kinematic modelling | GALPAK | 48 | 79 |
| Photometric modelling | GALFIT | 19 | – |
| Not modelled | – | 12 | – |
| Inclination $> 30^\circ$ | – | 56 | 67 |
| Inclination $< 30^\circ$ | – | 11 | – |

the azimuthal angle (Φ), we preferentially use the kinematic PA when available. We then set a restriction on the inclination $i > 30^\circ$ to ensure face-on galaxies with significant errors on the PA are removed from our sample. This leaves a total of 56 galaxies with robustly measured azimuthal angles associated with 22 Ly- α absorbers. The remaining 23 galaxies were not considered in this study because they were too faint to be modelled (12) or failed to meet the inclination requirement (11). A summary of the sample is shown in Table 2.

3.1 Distribution of azimuthal angles

We show the distribution of H I absorbers around galaxies in Fig. 2. In the left histogram, galaxies with measured kinematic position angles are shown in purple, while the blue represents photometric PAs. Galaxies found at the closest impact parameter to the absorber are given an opaque bar. We impose a restriction on the inclination ($i > 30^\circ$) for the sample to remove face-on galaxies with large errors in their position angle and, by extension, the azimuthal angle. This leaves a total of 56 galaxies. We test for bimodality by performing Hartigan’s dip test (Hartigan & Hartigan 1985). To account for errors in the azimuthal angle, we repeat the dip test on 5000 realizations with randomly sampled errors. The combined frequency distribution of the 5000 iterations is depicted in the middle panel of Fig. 2, and we find that the initial peak in the number count near $\Phi = 75^\circ$ diminishes in significance after accounting for errors in the measurement of Φ .

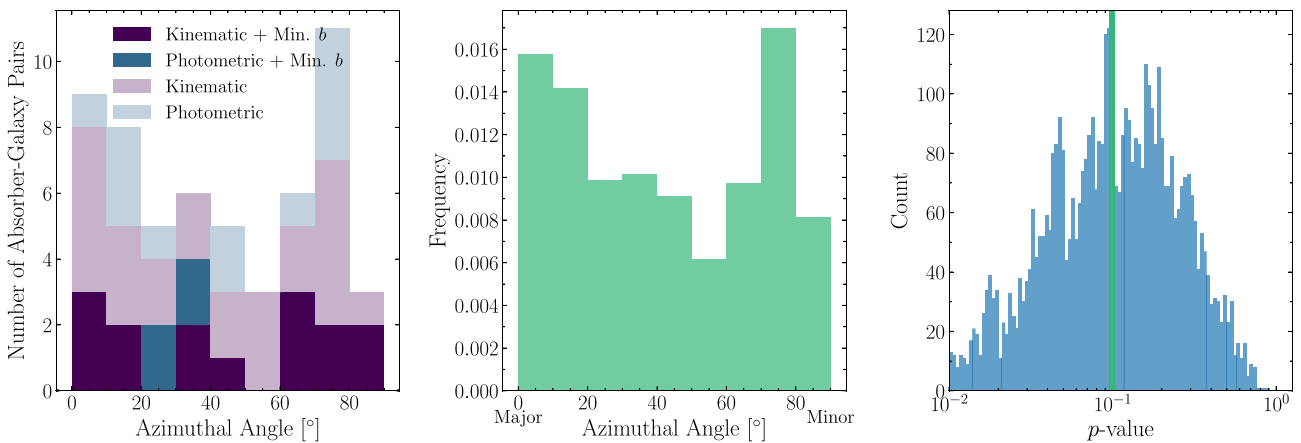


Figure 2. The distribution of azimuthal angles between absorbers and galaxies with inclination $i > 30^\circ$. We set this restriction on the inclination to remove face-on galaxies with large errors in their modelled position angle. In the left histogram, the purple and blue colours, respectively, represent galaxies with measured kinematic and photometric PAs, derived from modelling the MUSE cubes and *HST* imaging, respectively. Opaque bars represent galaxies found at the closest impact parameter to the QSO sightline, while translucent bars represent PA measurements of the other galaxy counterparts further from the absorber. To test whether the distribution of azimuthal angles is bimodal, we generate 5000 iterations of the initial distribution by randomly sampling the uncertainties in Φ and perform the Hartigan’s dip test for each iteration (Hartigan & Hartigan 1985). The final frequency distribution from the 5000 samples is shown in the middle histogram, and the corresponding distribution of p -values is shown on the right. The median p -value is ≈ 0.1 (indicated by the vertical green bar), which suggests there is marginal evidence for a bimodal distribution in azimuthal angles.

The corresponding distribution of p -values is displayed in the right-hand panel, where the vertical green bar shows the median p -value of the iterations is $\lesssim 0.1$. This suggests there is marginal evidence for a bimodality in the azimuthal angle distribution of H I absorbers around galaxies. A larger sample of galaxy-absorber pairs is required to confirm this signal with more certainty.

We show polar plots in Fig. 3 to visualize the distribution of H I absorbers more clearly. In the left-hand panel, we see the anticorrelation between H I column density and impact parameter (see Weng et al. 2023, for more analysis of this effect). There is also little correlation between the line of sight velocity difference between absorber and galaxy, and the polar position, as the median $|\Delta v_{\text{LOS}}|$ values in the two bins $0^\circ < \Phi < 30^\circ$ and $60^\circ < \Phi < 90^\circ$ are, respectively, 75 and 82 km s $^{-1}$.

3.2 Metallicity dependence on azimuthal angle

In Fig. 4, we show the absorber metallicity against azimuthal angle using data from previous works (Péroux et al. 2011, 2012; Bouché et al. 2013, 2016) and the new results presented here from the MUSE-ALMA Haloes survey. These dust-free metallicities are determined from zinc abundances of dense neutral gas. We colour each point by the stellar mass using 1 dex bin from $\log(M_*/M_\odot) = 8.0$ to 11.0. Result from the TNG50 simulation (Pillepich et al. 2019; Nelson et al. 2019a) are plotted for an impact parameter of $b = 100$ kpc, redshift $z = 0.5$ and stellar masses $\log(M_*/M_\odot) = 8.5, 9.5, \text{ and } 10.5$ (Péroux et al. 2020) and the shaded bands represent deviations of 1σ . The stellar mass significantly affects the normalization of the trend, but the metallicity difference between gas near the major and minor axes is approximately 0.3 dex for all stellar mass bins. We note that Péroux et al. (2020) find that the relationship diminishes for lower impact parameters ($b < 50$ kpc) and higher redshifts ($z > 1.5$). Additionally, limiting the column density to $\log[N(\text{H I})/\text{cm}^{-2}] > 17.0$ washes out the signal in the simulations because of the limited statistics. Hence, it is unsurprising that we find little evidence for a correlation between absorber metallicity and azimuthal angle given we have limited the sample to strong H I ($\log[N(\text{H I})/\text{cm}^{-2}] > 19.0$) absorbers with

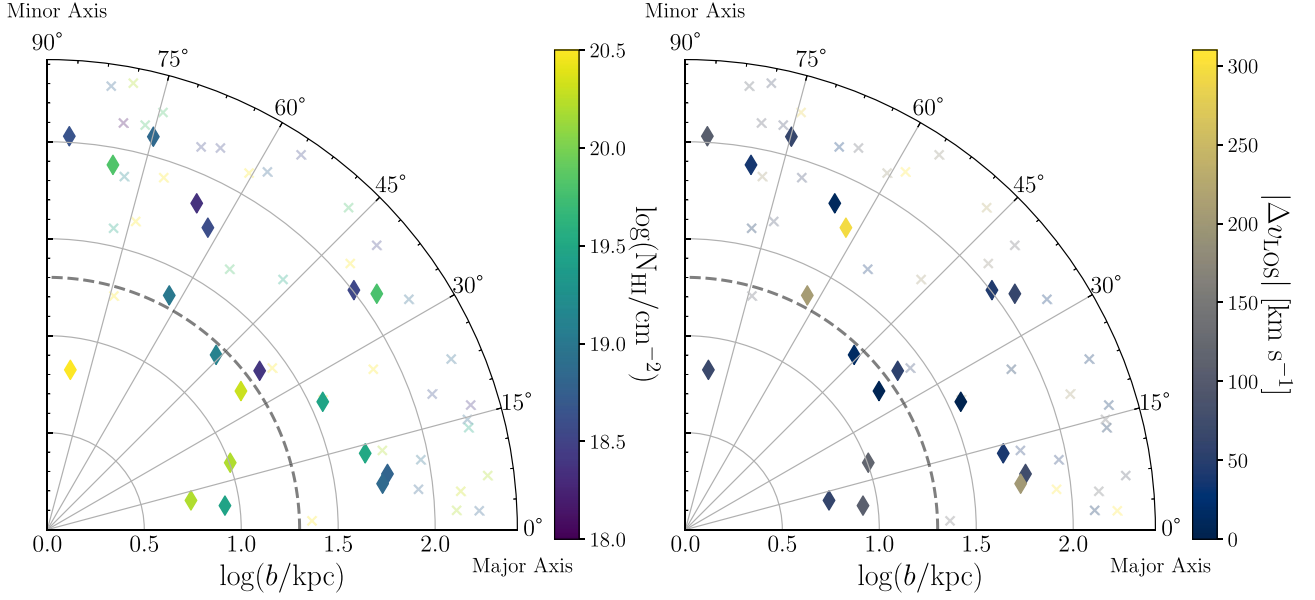


Figure 3. Polar plots illustrating the distribution of the H I absorbers as a function of azimuthal angle and impact parameter. Diamonds represent galaxies that are found at the closest impact parameter to a given absorber, and each point is coloured by the H I column density (left) or the absolute value of the line of sight velocity difference between galaxy and absorber (right). Crosses show the location of associated galaxies at larger impact parameters when multiple galaxies are associated with a single absorber. Points located near 90° are found near the galaxy’s minor axis. The black dashed line marks an impact parameter of $b = 20$ kpc.

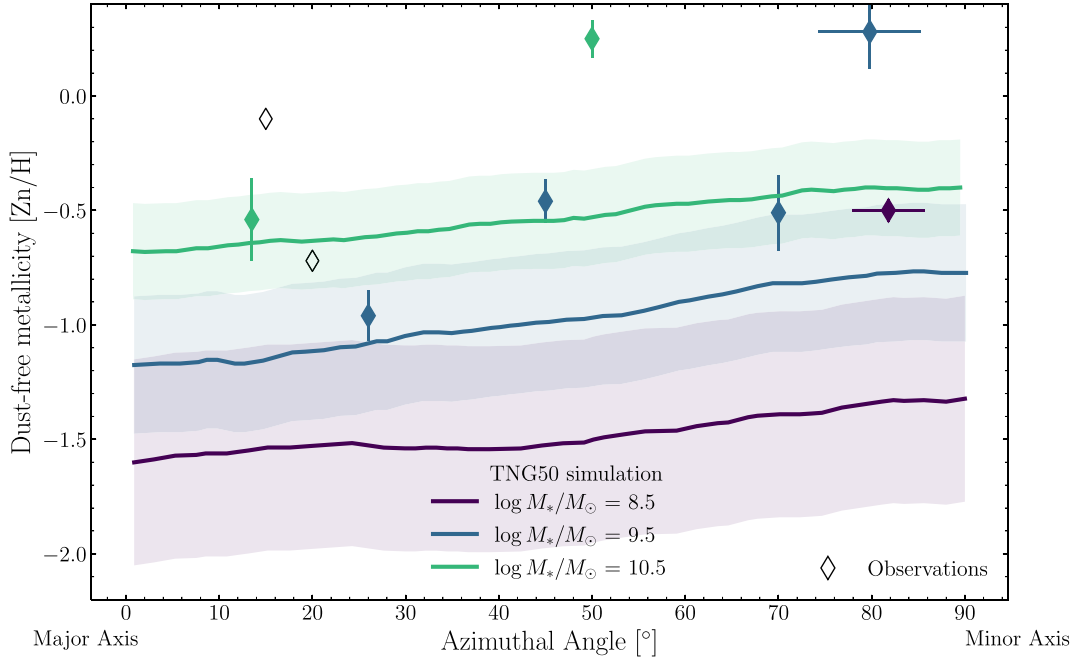


Figure 4. The absorber’s metallicity as a function of azimuthal angle. Diamonds represent results from this survey and previous works (Péroux et al. 2011, 2012; Bouché et al. 2013, 2016) that measured metal abundances directly. The coloured lines are the trends predicted by the TNG50 simulation (Péroux et al. 2020) for galaxies of different stellar masses at $z = 0.5$ and $b = 100$ kpc. The shaded regions represent a 1σ deviation. Each galaxy is coloured by their stellar mass corresponding to the coloured lines, while the unfilled black diamonds are galaxies without M_* measurements. At present, there are too few H I absorbers with robust metallicities and galaxy counterparts to confirm that gas near the minor axis is more metal-enriched than gas near $\Phi = 0^\circ$.

dust-free metallicity measurements. We suggest that a sample of $\sim 100 \log(N(\text{H I})/\text{cm}^{-2}) > 19.0$ absorbers (where the fraction of ionized gas is expected to be small) with reliable metallicities is required to confirm that gas near the minor axis is more metal-enriched than gas near the major axis.

4 CHARACTERIZING GAS FLOW ORIGINS

Identifying broader relationships between absorber properties and galaxy orientation is useful to characterize the global properties of the circumgalactic medium, but identifying gas flows for individual systems enables us to connect galaxy properties with the galactic

baryon cycle. In this section, we examine the rotation maps of galaxies associated with individual absorbers to infer whether they originate from the galaxy disc, are inflowing or outflowing, or arise from other phenomena.

Previous works have already performed a kinematic analysis of several galaxy-absorber systems in the MUSE-ALMA Haloes survey, and we summarize it here. The two absorbers at $z_{\text{abs}} = 0.3830$ and 0.7802 towards quasar Q0152–0005 were found to, respectively, trace inflowing and outflowing gas (Rahmani et al. 2018a, b). In a study of the neutral, ionized, and molecular gas phases, the $z_{\text{abs}} = 0.633$ absorber towards Q0420–0127 is found to uncover signatures of outflows or intragroup gas (Klitsch et al. 2018). A rich galaxy group is found at $z_{\text{abs}} = 0.313$ towards Q1130–1449 and the Ly- α absorber appears to trace intragroup gas (Péroux et al. 2019). Finally, the absorber towards Q2131–1207 is consistent with co-rotating gas accreting onto the galaxy at the lowest impact parameter to the absorber (Péroux et al. 2017; Szakacs et al. 2021). The remaining 22 absorbers with galaxy counterparts are analysed in this work, and the final designation for each absorber is noted in Table 1. An individual discussion for each absorber is contained in Appendix B.

4.1 Galaxy discs

We find one case where the absorber evidently intersects the galaxy disc. Such absorbers are characterized by a strong H I absorber ($\log[N(\text{H I})/\text{cm}^{-2}] > 20.0$) found within ~ 20 kpc of a galaxy (Wolfe et al. 1986; Péroux et al. 2005; Zwaan et al. 2005). The clearest example of disc gas being probed is the absorber towards Q1110+0048 at $z_{\text{abs}} = 0.5604$, where the host galaxy is found at an impact parameter of 6 kpc. We find the $\log[N(\text{H I})/\text{cm}^{-2}] = 20.2$ absorber at the expected velocity sign and magnitude as the modelled rotational map (see top left-hand panel of Fig. 5). For the remaining three cases, no kinematic modelling is possible for the galaxies found at impact parameters $b < 20$ kpc. Nevertheless, we also attribute these absorbers to arise from the galaxy disc given their column densities ($\log[N(\text{H I})/\text{cm}^{-2}] > 20.0$) and low impact parameters. The $|\Delta v_{\text{LOS}}|$ values for these absorbers range from 5 to 150 km s $^{-1}$ and are consistent with typical rotation speeds. We calculate the azimuthal angle between the QSO sightline and galaxy using the *HST* photometry to range from 20° to 40°. Thus, we do not find the absorbers near the minor axes, where outflowing gas may be expected. While we lack direct evidence the absorber is rotating with the galaxy disc without velocity maps, damped Lyman- α absorbers have long been associated with the rotating disc of galaxies (Wolfe et al. 1986; Wolfe et al. 2005). More recent simulations predict a covering fraction of 50 per cent for DLAs within $0.1 R_{\text{vir}}$ (15 kpc) of galaxies with halo mass $10^{12} M_{\odot}$ at $z \sim 0$ (Stern et al. 2021). These findings support the idea that the strong H I absorbers at $z < 1$ in this study arise from galaxy discs, whereas DLAs at $z \gtrsim 2$ increasingly probe the inner circumgalactic medium.

4.2 Accretion

Similar to absorbers that probe the galaxy disc, gas that is co-rotating with a galaxy can be detected by comparing the velocity of the absorber with the ionized gas velocity field. The difference is that gas accreting onto the galaxy is lower column density and can be found at larger impact parameters. We find that the $z_{\text{abs}} = 0.6057$ absorber towards Q1345–0035, where the host galaxy is found at an impact parameter of 56 kpc, is consistent with infalling gas. The $\log[N(\text{H I})/\text{cm}^{-2}] = 18.85$ absorber is found with the same velocity

sign and magnitude to the modelled rotational map (bottom right-hand panel of Fig. 5) and is likely tracing gas in the circumgalactic medium co-rotating with the galaxy disc.

Less certain cases include the $\log[N(\text{H I})/\text{cm}^{-2}] < 19.1$ absorber towards Q1130–1449 at $z_{\text{abs}} = 0.1906$. The impact parameter of 18 kpc suggests the gas originates from the galaxy disc, but the upper limit on the H I column density signifies the gas is perhaps accreting. We also find dense gas consistent with co-rotation for the absorber towards Q0152+0023 at $z_{\text{abs}} = 0.4818$ (top right-hand panel of Fig. 5). However, it is unclear whether such cool dense gas ($\log[N(\text{H I})/\text{cm}^{-2}] = 19.78$) can be found co-rotating at impact parameters > 100 kpc. Importantly, the stellar mass of the galaxy Q0152+0023_20 associated with the absorber is $\log(M_{*}/M_{\odot}) = 8.1 \pm 0.1$ and the estimated virial radius, R_{vir} , is ~ 70 kpc (Rodríguez-Puebla et al. 2017). Despite the alignment in velocity, the absorber is beyond the halo of the associated galaxy and it is possible that a passive, low-mass galaxy ($\text{SFR} < 0.19 M_{\odot} \text{ yr}^{-1}$, $\log(M_{*}/M_{\odot}) < 8.5$) below the MUSE detection threshold nearer the QSO sightline hosts the dense absorber.

Within the full MUSE-ALMA Haloes Survey sample of H I absorbers, we find only three cases that are consistent with gas accretion. Cold streams in the circumgalactic medium are expected to be observed in absorption because of their high densities and temperatures $\sim 10^4$ K (Fumagalli et al. 2011b; Faucher-Giguère et al. 2015; Hafen et al. 2017). Simulations predict that dense H I absorbers at $z \gtrsim 3$ trace cold-mode accretion (van de Voort et al. 2012) and detections of low-metallicity Lyman-limit systems (LLSs) have been suggested to trace these inflows (Ribaud et al. 2011; Fumagalli, O’Meara & Prochaska 2011a; Lehner et al. 2013). However, at lower redshifts, direct detections of gas accretion are sparse for both down-the-barrel and transverse absorption-line studies (Kacprzak et al. 2010; Martin et al. 2012; Rubin et al. 2012; Ho et al. 2017; Zabl et al. 2019). Our results echo the findings of studies without pre-selection of targets and find only ~ 10 per cent of H I absorbers consistent with accretion.

4.3 Outflows

In total, we find seven absorbers that are within 30° of the projected minor axis of nearby galaxies. However, only three out of these seven absorbers can be confidently attributed to gas that is outflowing. One clear case of outflowing gas is seen in the $z_{\text{abs}} = 0.7869$ absorber towards Q1554–203. The absorber is found at an impact parameter of 23 kpc and an azimuthal angle of 62° (see row 2 in Fig. A9). The gas is unlikely to be inflowing or associated with the galaxy disc as the velocity of the absorber is opposite in sign to the ionized gas velocity field. A line-of-sight velocity difference between the galaxy and absorber of ~ 200 km s $^{-1}$ is also typical of outflows. The absorber towards Q0454+039 at $z_{\text{abs}} = 1.1532$ with column density $\log[N(\text{H I})/\text{cm}^{-2}] = 18.59$ aligns with the minor axis ($\Phi = 62^{\circ}$) of the nearest galaxy (see second row of Fig. A3). At an impact parameter of 60 kpc and a line of sight velocity difference of -290 km s $^{-1}$, the gas is likely tracing neutral gas entrained in an outflow. While the absorber and ionized gas are both blueshifted with respect to the galaxy’s systemic redshift, there is a > 250 km s $^{-1}$ discrepancy between the maximum rotational velocity of the galaxy and the absorber velocity. Finally, we find the DLA at $z_{\text{abs}} = 0.3950$ towards Q1229–021 is consistent with cold gas entrained in an outflow. At an impact parameter of 6 kpc and an azimuthal angle of 81°, the neutral gas velocity is inconsistent with the ionized gas disc. We measure a line of sight velocity difference

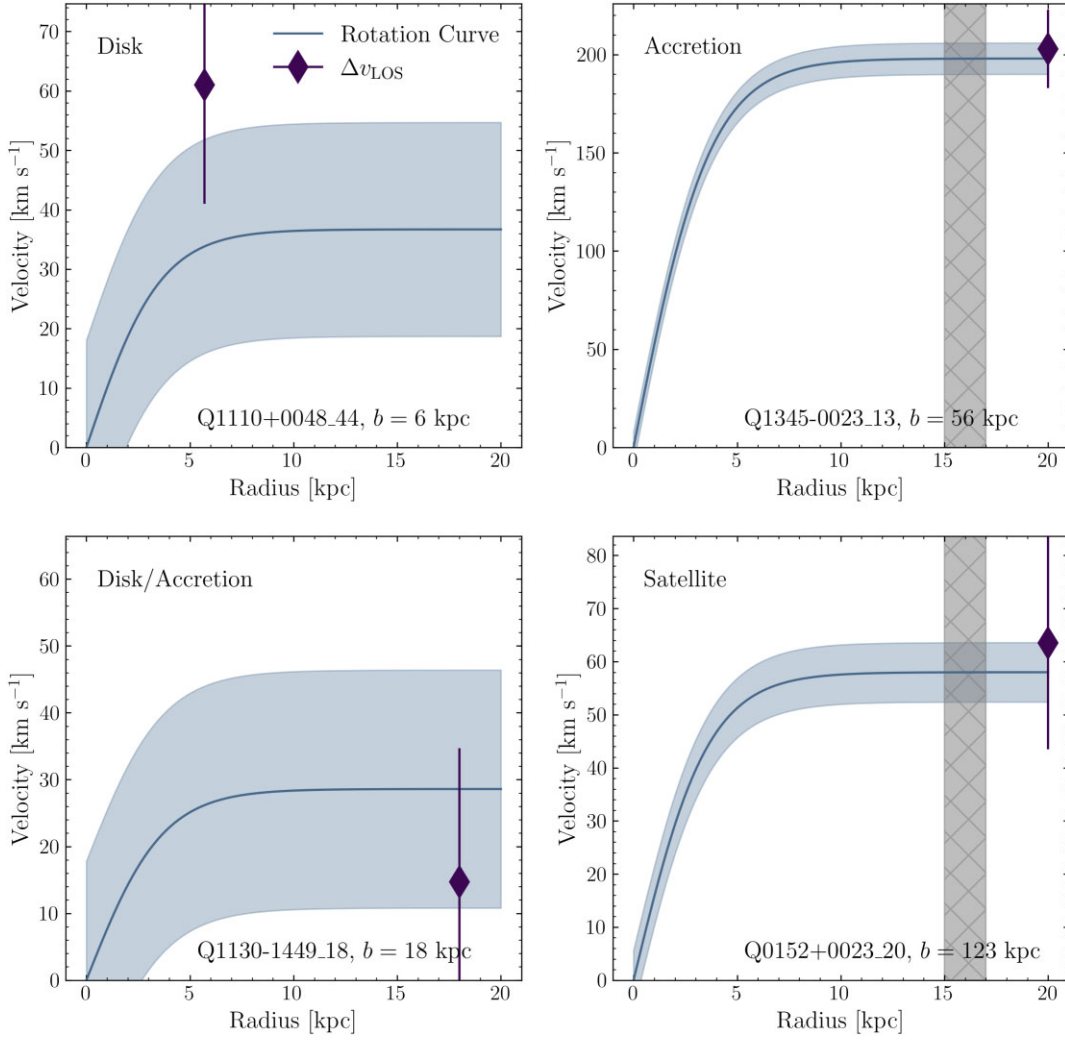


Figure 5. Comparisons between the galaxy rotation curve and absorber line of sight velocities for cases, where the disc or gas accretion is being traced. The blue line represents the galaxy rotation curve modelled by a hyperbolic tangent, and the shaded region represents the 1σ error. We plot the velocity of the absorber relative to the systemic redshift of the galaxy as a purple diamond. A vertically hatched region is used to illustrate that the impact parameter of the absorber is beyond the limits of the x -axis. We provide the gas flow origin in text at the top-left, while the galaxy ID and absorber impact parameter are found at the bottom of each plot. In three cases, we find the absorber velocity is consistent with the rotating disc or co-rotating halo of the galaxy at the nearest impact parameter. In the bottom-right case, the absorber is beyond the virial radius of the nearest galaxy and likely arises from a faint, quiescent galaxy.

of 70 km s^{-1} , which suggests the gas is still accelerating at this small distance from the galaxy.

The remaining four absorbers that align with the minor axis possess unclear gas flow origins. The uncertainty arises from two issues: the absorber velocity is inconsistent with outflowing gas or there are multiple galaxies with similar impact parameters to the absorber. Addressing the former concern first, we find that the absorber towards Q0138–0005 at $z_{\text{abs}} = 0.7821$ has only one galaxy counterpart at an impact parameter of 80 kpc (first row of Fig. 1). The absorber is aligned with the minor axis of the associated galaxy and has a zinc abundance of $[\text{Zn}/\text{H}] = 0.28 \pm 0.16$. While both of these properties are consistent with metal-enriched gas expelled by the single galaxy counterpart, the absorber velocity is $<40 \text{ km s}^{-1}$ from the galaxy systemic redshift. Another similar example is the absorber towards Q1229–021 at redshift $z_{\text{abs}} = 0.7572$ (last row in Fig. A6), which is separated by only 15 km s^{-1} from the redshift of its galaxy counterpart. A possible explanation for these low velocities is that there is a large velocity component orthogonal to the sightline, as

these galaxies have inclinations $i > 65^\circ$. The line of sight velocity is not necessarily well-correlated with the radial velocity of the gas. Alternatively, the neutral gas may be static and not co-rotating with the halo.

For the two absorbers at $z_{\text{abs}} = 0.7691$ and 0.8311 towards Q1229–021, there are at least three galaxies found at the absorber redshift and at similar distances from the QSO sightline. It becomes difficult to determine the gas flow origin of the absorber as gravitational interactions between the galaxies can cause gas to be stripped. The individual cases are discussed in Appendix B, and a more detailed exploration of intragroup gas is explored in the following subsection. We choose to label these four unclear cases as possible outflows.

4.3.1 Does the gas escape?

We can determine whether the neutral gas escapes the galaxy halo by comparing the absorber velocity with the escape velocity (V_{esc}). The

Table 3. Comparison of line-of-sight absorber velocity with the escape velocity for galaxy-absorber pairs consistent with outflows. The line of sight velocity of the absorber relative to the galaxy is roughly equal to the escape velocity, assuming a singular isothermal sphere. Given that Δv_{LOS} does not account for velocities orthogonal to the line of sight, it is possible that the H I absorbers trace neutral gas escaping the potential of the galaxy. While $|\Delta v_{\text{LOS}}| < V_{\text{esc}}$, the radial velocity of the gas may be larger than the escape velocity.

| QSO | z_{abs} | $\log N(\text{H I})$ $\log(\text{cm}^{-2})$ | b kpc | $ \Delta v_{\text{LOS}} $ km s^{-1} | V_{esc} km s^{-1} |
|-----------|------------------|--|------------|---|--|
| Q0454+039 | 1.1532 | 18.59 | 60 | 290 | 300 |
| Q1229-021 | 0.3950 | 20.75 | 6 | 70 | 80 |
| Q1554-203 | 0.7869 | <19.0 | 23 | 210 | 270 |

escape velocity at a given radius, r , is calculated assuming a singular isothermal sphere (Veilleux et al. 2005):

$$V_{\text{esc}} = V_{\text{vir}} \times \sqrt{2 \left(1 + \ln \frac{R_{\text{vir}}}{r} \right)}, \quad (1)$$

where V_{vir} and R_{vir} are the virial velocity and radius, respectively. Here, we assume the radius to be the impact parameter ($r \approx b$). We estimate V_{vir} using the prescription in Schroetter et al. (2019), where $V_{\text{vir}} \approx 1.2 \times S_{0.5}$. Here, $S_{0.5} = \sqrt{0.5 \times V_{\text{max}}^2 + \sigma^2}$ is the kinematic estimator and is a function of the rotational velocity and velocity dispersion, σ (Weiner et al. 2006). Using V_{vir} , we can then approximate the virial radius to be $R_{\text{vir}} \approx V_{\text{vir}}/10H(z)$, where $H(z)$ is the Hubble constant at redshift z . The estimated escape velocity values are tabulated in Table 3.

While the absorber velocity relative to the galaxy is less than the escape velocity for all three likely cases of outflows, we do not take into account the velocity component orthogonal to the line of sight. Hence, it is possible that the radial velocity of the outflowing neutral gas exceeds the escape velocity and will be ejected from the galaxy halo.

4.4 Alternative phenomena

Beyond the absorber origins discussed earlier in this section, there are other phenomena that may produce H I absorption around galaxies. In particular, recent studies using MUSE reveal that roughly 50 per cent of Ly- α absorbers have multiple galaxies within a velocity window of $\pm 500 \text{ km s}^{-1}$ (Hamanowicz et al. 2020; Chen et al. 2020a; Weng et al. 2023; Berg et al. 2023). This suggests the absorption may arise from the intragroup medium between galaxies (Gauthier 2013; Nielsen et al. 2018; Dutta et al. 2023). It is difficult to identify intragroup gas because we require kinematic modelling for all galaxies in the overdensity to exclude gas flows. The two galaxies associated with the $z_{\text{abs}} = 0.3283$ absorber towards Q1130-1449 are roughly equidistant in projected distance from the absorber (75 and 90 kpc). There is no hint of inflows or outflows from the relative velocity and geometry of the absorber. Instead, we find that the velocity of the absorber is between the two galaxy systemic redshifts, which suggests we may be tracing intragroup gas with $\log[N(\text{H I})/\text{cm}^{-2}] < 18.9$. There are no other clear examples of absorbers tracing gas between galaxies, but there are several cases where multiple explanations are viable, such as the two absorbers towards Q1229-021 discussed in the previous section.

Another important consideration is that low-mass, quiescent galaxies or satellites hosting the absorber may not be detected in our observations. Indeed, there is a possible example of such a

case with the four galaxies at impact parameters of 122 to 190 kpc associated with the absorber towards Q0152+0023. We previously discussed in this section whether the gas may arise from accretion due to an alignment between the rotational velocity and absorber velocity (top right-hand panel of Fig. 5), but we noted the absorber is beyond the virial radius of the nearest galaxy. In fact, the $\log[N(\text{H I})/\text{cm}^{-2}] = 19.78$ absorber is beyond the virial radii of all four galaxies. This points to the hypothesis that there is a galaxy with stellar mass $\log(M_*/M_{\odot}) < 8.6$ and $\text{SFR} < 0.24 M_{\odot} \text{ yr}^{-1}$ below the detection threshold in the MUSE data near the strong H I absorber (Weng et al. 2023). While recent works suggest that the CGM extends beyond the virial radius (Wilde et al. 2021, 2023), it is unlikely to find such dense absorbers with column densities $\log[N(\text{H I})/\text{cm}^{-2}] \approx 20$ at such large distances from galaxies.

A final point is that the line of sight velocity is not necessarily a good predictor of the physical line of sight distance between an absorber-galaxy pair. In fact, simulations have shown that absorbers can be found at $> 1 \text{ pMpc}$ when applying a velocity cut-off $|\Delta v_{\text{LOS}}| < 500 \text{ km s}^{-1}$ (Rahmati et al. 2015; Ho, Martin & Schaye 2020, 2021). The velocity difference between absorber and galaxy is influenced by peculiar motions of the gas and the Hubble flow at larger separations. This suggests that there is always the possibility of chance associations between galaxies and absorbers. Quantifying this probability is the goal of an upcoming work using the TNG50 simulations.

5 DISCUSSION

In this section, we discuss the incidence rate of the various gas flows in the MUSE-ALMA Haloes survey. We also examine the relationships between measured galaxy properties and the origin of the gas. Additionally, we discuss the limitations of using simple geometric arguments to distinguish outflows and inflows and future improvements to this analysis by studying individual gas components.

5.1 The origins of gas in the CGM

The H I Ly- α absorbers in the MUSE-ALMA Haloes survey appear to trace various phenomena. In summary, out of the 32 absorbers in the survey, 27 are found to have at least one galaxy within $\pm 500 \text{ km s}^{-1}$. The 27 absorbers are comprised of four absorbers tracing the galaxy disc, three tracing accretion, four tracing outflows, two tracing gas in the intragroup medium, and one likely tracing an undetected, low-mass galaxy. Six absorbers may arise from multiple phenomena, while the remaining five (two) have associated galaxies without kinematic modelling (with inclinations $i < 30^\circ$). While this sample of absorbers only provides limited statistics, we do find that the accretion of gas onto galaxies is difficult to trace. Whether this is caused by the criteria used to identify accretion or an intrinsic property of the accretion itself (Faucher-Giguère & Kereš 2011) is unclear, but the percentage of accreting absorbers in this work is similar to down-the-barrel studies (Martin et al. 2012; Rubin et al. 2012) and transverse absorption-line works using Mg II (Zabl et al. 2019). We find four convincing cases of H I outflows, which is proportionally far less than other surveys (e.g. Schroetter et al. 2019). However, we note that there are six other cases where the absorber kinematics and geometry are consistent with outflows, where the interpretation is unclear because of possible interactions with other galaxies in the field or the modelled inclination not meeting our threshold ($i < 30^\circ$). Ultimately, we have found that strong H I Ly- α absorbers probe gas in a variety of environments and of many

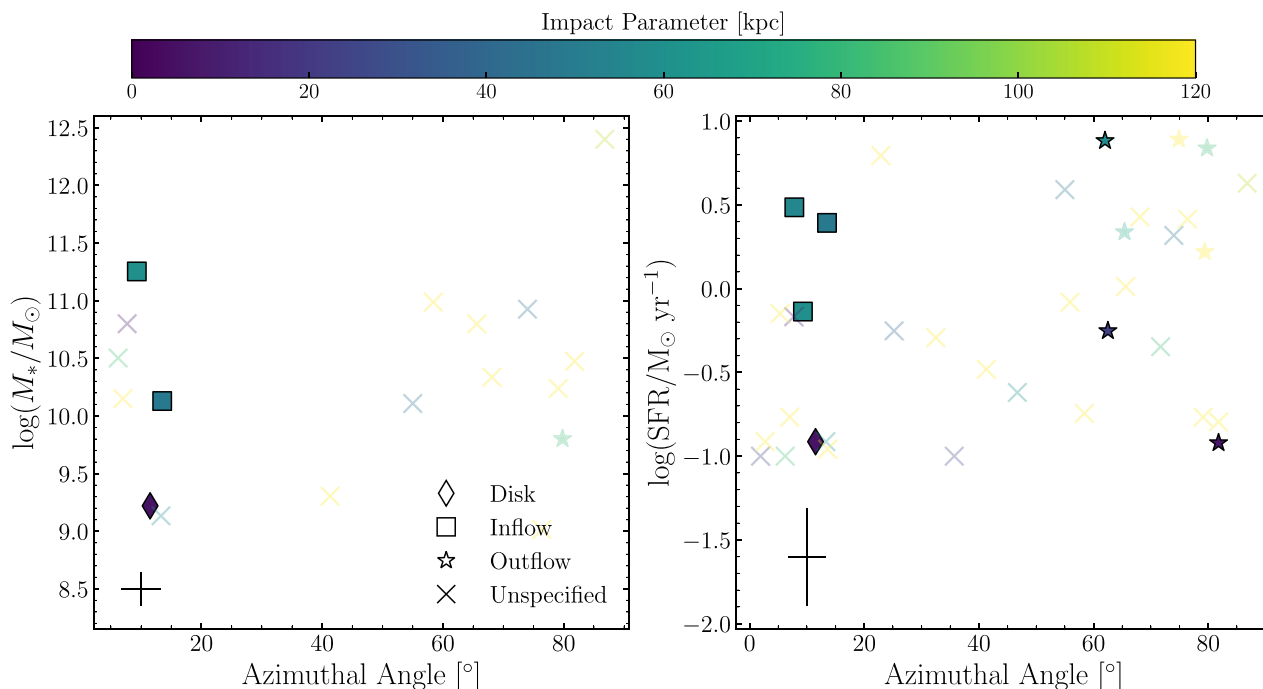


Figure 6. The stellar properties of the galaxies associated with an absorber. The left plot shows the stellar mass of galaxies associated with absorbers at a given azimuthal angle. Stellar masses are derived from spectral energy distribution (SED) fitting of the *HST* broad-band imaging (Augustin et al., in preparation). Different symbols represent galaxy-absorber pairs that are found to trace inflowing, outflowing or gas in the disc, while faint crosses represent pairs with ambiguous origins. Symbols with a black border are cases where the gas origin has been confidently identified, while those with lower transparency are possible cases. Each galaxy is coloured by the impact parameter of the absorber. We find absorbers associated with galaxies that span four dex in stellar mass. On the right, we show the dust-uncorrected star-formation rate of galaxies measured using the $H\alpha$ or $[O\ II]$ emission lines (Weng et al. 2023). We find that the galaxies associated with inflows and outflows do not differ significantly in their SFRs. The median errors in the plotted properties are shown as a cross at the bottom left of both plots.

origins in the complex circumgalactic medium. In future work using the TNG50 simulation, we will address whether the various gas origins discussed here can be distinguished from each other using the available observables of impact parameter, metallicity, line of sight velocity difference, and azimuthal angle.

5.2 Gas flows and galaxy properties

With our physical interpretations of the absorber, we examine the properties of the galaxies associated with these absorbers. In Fig. 6, we show the stellar masses and star-formation rates of galaxies associated with the absorbers of various origin. Diamonds represent absorbers that likely intersect the galaxy disc, stars, and squares represent outflows and inflows, respectively, while crosses signify there are multiple origins for the gas. For both the stellar mass and SFR, a ~ 3 dex range of values is observed.

Inflows are found to be associated with galaxies with large stellar masses 10^{10} to $10^{11} M_{\odot}$, but more photometry is required to estimate M_* for the other galaxies. While the star-formation rate of galaxies associated with the various gas flows span two dex, a more important indicator is the star-formation rate per unit area (Σ_{SFR}). Two of the three galaxies where the absorber likely traces outflows have a $\Sigma_{\text{SFR}} > 0.1 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$ (Heckman 2002, 2003). The third galaxy does not meet the threshold in Σ_{SFR} , but we note that the SFR is not corrected for dust and the starburst responsible for the galactic wind might be more localized. Currently, a larger sample is required to connect the gas flows traced by absorbers to the galaxies that host the neutral gas.

A recent study of galaxies associated with Mg II absorbers finds that galaxies with large inflow rates are located above the $\text{SFR}-M_*$ main sequence (Langan et al. 2023). Using the three galaxies associated with inflows that have both SFR and M_* measurements, we find no such signal and most of the associated galaxies lie on the $\text{SFR}-M_*$ main sequence (Karki et al., in preparation).

5.3 The azimuthal distribution of H I and metals in the CGM

Outflows of gas driven by supernovae or AGN are expected to be preferentially aligned with the minor axis and collimate into a biconical shape (Veilleux et al. 2005). Both hot- and cold-mode accretion is expected to channel cool gas onto galaxies via cool flows that align with the disc (e.g. Kereš et al. 2005; Hafen et al. 2022). In absorption-line studies of ions such as Mg II and O VI in the CGM, the inflow-outflow dichotomy manifests in the bimodal distribution of azimuthal angles between the quasar line of sight and nearby associated galaxies (e.g. Bouché et al. 2012; Kacprzak et al. 2015). In this work, we also find marginal evidence for a bimodality in the distribution of H I absorbers. In principle, this bimodality in Φ should extend to metallicity measurements, as outflowing gas is typically more metal-enriched than gas accreting onto a galaxy. However, the dependence of metallicity on azimuthal angle is far less clear from observations (Péroux et al. 2016; Kacprzak et al. 2019; Pionton et al. 2019; Wendt et al. 2021). Here, we find very little evidence for a correlation between absorber metallicity and azimuthal angle for three different stellar mass bins, and a larger and more homogeneous sample is required to determine whether a relationship exists.

The reasons for the absence of a distinct signal are manifold. The trend of gas-phase metallicity versus azimuthal angle depends on properties such as the impact parameter, H I column density and most significantly, the stellar mass (Péroux et al. 2020). These parameters not only affect the normalization of the signal, but also the magnitude of the discrepancy between low and high azimuthal angles. A larger sample of absorbers reaching column densities $\log[N(\text{H I})/\text{cm}^{-2}] = 13.0$ that are found at impact parameters larger than 50 kpc from their host galaxies that have measured stellar masses is required to fully test the azimuthal dependence of metallicity in the CGM. Such a sample will be difficult to construct because of the challenge in measuring gas-phase metallicity due to the uncertainties surrounding photoionization and dust modelling. Furthermore, there is the intrinsic inhomogeneity of gas properties at small spatial scales in the CGM that will require large samples to take into account. Estimates of line of sight cloud sizes range from subparsec to ~ 100 parsec from ionization modelling (e.g. Churchill et al. 2003; Werk et al. 2014), and studies find significant ranges in metallicity for different components along a single line of sight (e.g. Zahedy et al. 2019; Nielsen et al. 2022), suggesting metals may be poorly mixed in the CGM (Péroux et al. 2018; Tejos et al. 2021). This is captured by the TNG50 predictions depicted in Fig. 4, where the 1σ errors range ~ 1 dex in metallicity. Moreover, recent work from Berg et al. (2023) suggests there is a population of low-metallicity absorbers residing in overdense regions away from galaxy haloes. Indeed, the baryon cycle is more complex than a linear combination of gas being expelled out via the minor axis and accreting along the major axis; these processes interact to form the complex, multiphase circumgalactic medium.

5.4 The fidelity of geometric and kinematic arguments

In this work, we adopt an opening angle of 60° (corresponding to $\pm 30^\circ$ from the minor axis) to identify gas that is being expelled (Chen et al. 2010; Lan et al. 2014), but a diverse range of opening angles have been observed in the literature (e.g. Veilleux, Shopbell & Miller 2001), and these angles differ depending on the observed wavelength (e.g. see the opening angles for the Circinus galaxy ranging from 15° to 100° Harnett et al. 1990; Elmoutie et al. 1995; Veilleux & Bland-Hawthorn 1997; Curran et al. 1999). It is unclear whether several of the absorbers with azimuthal angles just below the threshold should be considered outflowing. Similarly, while inflowing gas from hot- and cold-mode accretion is expected to align in angular momentum with the disc, there has been growing evidence for the condensation of ambient gas in the halo caused by interactions between gas ejected by stellar winds and hot coronal gas (Marinacci et al. 2010; Fraternali 2017). Known as the galactic fountain, the phenomenon is expected to cause neutral gas to ‘rain’ down onto the galaxy disc rather than align with the major axis. There is mounting evidence for galactic fountains in the Milky Way, with the kinematics of high- and intermediate velocity clouds consistent with a mixture of outflowing and diffuse inflowing gas (Marasco et al. 2022; Lehner et al. 2022a). Beyond the local Universe, signatures of fountain flows have been found to persist at impact parameters $b > 5$ kpc (Rubin et al. 2022). Hence, the absorbers in this sample may also trace this process at $z \sim 0.5$.

Another consideration is that the strong H I Ly- α absorbers in this sample are comprised of multiple components at varying velocities relative to the DLA redshift. Recently, photoionization modelling of individual components within H I systems has been performed to determine cloud properties such as temperature and size (Cooper et al. 2021; Zahedy et al. 2021; Nielsen et al. 2022). The individual clouds and their properties have then been related to outflows,

inflows, or intragroup gas. We leave the modelling of the various metal-line components and estimations of the typical cloud properties embedded in outflowing and inflowing gas to future works.

6 SUMMARY AND CONCLUSION

The MUSE-ALMA Haloes survey combines multiwavelength observations of galaxies associated with $32 \log[N(\text{H I})/\text{cm}^{-2}] > 18.0$ absorbers. In this work, we have modelled the ionized gas kinematics of 48 galaxies associated these absorbers using the forward modelling algorithm GALPAK to extract properties such as the rotational velocity, velocity dispersion, inclination, and position angle. By determining the position and geometry of the absorber with respect to the modelled galaxies, we seek to determine the distribution of gas in the circumgalactic medium and identify the possible origins of these strong H I absorbers. To summarize, we find:

(i) An excess of absorption sightlines passing near the major and minor axes of galaxies. There is marginal evidence for a bimodal distribution in azimuthal angles between galaxy and absorber after performing the Hartigan’s dip test on 5000 iterations of the data by randomly sampling the errors ($p \approx 0.1$). This is similar to previous studies of the Mg II and O VI ions, suggesting inflows and outflows of gas in the CGM can also be traced by neutral hydrogen.

(ii) That there is little evidence for a dependency of the metallicity on the azimuthal angle for the absorbers in the MUSE-ALMA Haloes survey as predicted by simulations. This suggests that gas in the circumgalactic medium is not merely a linear combination of metal-poor inflows and metal-enriched outflows and that other phenomena such as gas recycling and poor metal mixing are significant. The results from simulations also show that the scatter in metallicities at any given azimuthal angle is comparable to the actual metallicity discrepancy at the minor and major axes. At this stage, we argue that a larger sample of ~ 100 strong H I absorbers with dust-free metallicity measurements is still required to recover any signal predicted by simulations.

(iii) That H I absorbers have a variety of origins in the CGM. Absorbers with column densities $\log[N(\text{H I})/\text{cm}^{-2}] > 20.0$ at impact parameters of $b < 20$ kpc from the nearest galaxy are considered associated with the galaxy disc. Only 15 per cent of absorbers are found to trace the disc, suggesting other processes must account for the remaining absorbers. We find that roughly 10 per cent of absorbers are co-rotating with the halo out to distances up to 60 kpc and these are suspected to trace gas accretion. The rarity of such cases is in line with previous works. Up to ≈ 30 per cent of absorbers are found within $\pm 30^\circ$ of the minor axis of galaxies and are consistent with outflows, but we only identify four clear cases in the sample. The remaining absorbers trace gas in the intragroup medium, low-mass galaxies below the detection limit of the MUSE data or do not have sufficient data to uncover a physical origin.

In the future, larger surveys such as the ByCycle survey (PI: Péroux) using the 4MOST instrument will enable us to compare the kinematics of galaxies with absorbers on a much larger scale (de Jong et al. 2012, 2019). The combination of high-resolution ($R \sim 20000$) background QSO spectroscopy with deep and complete foreground galaxy surveys (Driver et al. 2019) will enable better constraints on how the absorber metallicity varies with azimuthal angle (Szakacs et al., in preparation). While we have identified the various gas flows traced by dense H I absorbers, future modelling of individual gas components will provide information on the gas properties of inflows and outflows (e.g. temperature, density, and cloud size).

The proliferation of absorber follow-up surveys that use integral field spectroscopy has led to a re-characterization of the physical processes that absorbers trace in the circumgalactic medium and how gas is distributed with respect to galaxies. In this work, we find and emphasize that even strong Ly- α absorbers trace gas flows in the circumgalactic medium. The fact that dense, neutral gas required for star formation is seen accreting and being expelled highlights that these processes have significant impacts on galaxy evolution and that the CGM plays an important role in regulating the baryon cycle.

ACKNOWLEDGEMENTS

This research is supported by an Australian Government Research Training Program (RTP) Scholarship. EMS, GK, and SW acknowledge the financial support of the Australian Research Council through grant CE170100013 (ASTRO3D). VPK and AK acknowledge partial support for GO program 15939 (PI: Peroux) provided through a grant from the STScI under NASA contract NAS5-26555 and NASA grant 80NSSC20K0887 (PI: Kulkarni). VPK also gratefully acknowledges additional support from the National Science Foundation grants AST/2007538 and AST/2009811 (PI: Kulkarni). DN acknowledges funding from the Deutsche Forschungsgemeinschaft (DFG) through an Emmy Noether Research Group (grant number NE 2441/1-1). This research was supported by the International Space Science Institute (ISSI) in Bern, through ISSI International Team project #564 (The Cosmic Baryon Cycle from Space).

This research also made use of several PYTHON packages: ASTROPY (Astropy Collaboration 2013, 2018), MATPLOTLIB (Hunter 2007), and NUMPY (Harris et al. 2020).

DATA AVAILABILITY

Data directly related to this publication and its figures are available upon request. The catalogues for the MUSE-ALMA Haloes survey have been made public with the publication of Péroux et al. (2022). The raw data can be downloaded from the public archives with the respective project codes.

REFERENCES

- Arnouts S., Ilbert O., 2011, *Astrophysics Source Code Library*, record ascl:1108.009
- Astropy Collaboration, 2013, *A&A*, 558, A33
- Astropy Collaboration, 2018, *AJ*, 156, 123
- Bacon R. et al., 2023, *A&A*, 670, A4
- Berg M. A. et al., 2023, *ApJ*, 944, 101
- Bordoloi R. et al., 2011, *ApJ*, 743, 10
- Bordoloi R., Lilly S. J., Kacprzak G. G., Churchill C. W., 2014, *ApJ*, 784, 108
- Bordoloi R. et al., 2022, *Nature*, 606, 59
- Bouché N., Hohensee W., Vargas R., Kacprzak G. G., Martin C. L., Cooke J., Churchill C. W., 2012, *MNRAS*, 426, 801
- Bouché N., Murphy M. T., Kacprzak G. G., Péroux C., Contini T., Martin C. L., Dessauges-Zavadsky M., 2013, *Science*, 341, 50
- Bouché N., Carfanten H., Schroetter I., Michel-Dansac L., Contini T., 2015, *AJ*, 150, 92
- Bouché N. et al., 2016, *ApJ*, 820, 121
- Burchett J. N., Rubin K. H. R., Prochaska J. X., Coil A. L., Vaught R. R., Hennawi J. F., 2021, *ApJ*, 909, 151
- Cameron A. J. et al., 2021, *ApJ*, 918, L16
- Chen Y.-M., Tremonti C. A., Heckman T. M., Kauffmann G., Weiner B. J., Brinchmann J., Wang J., 2010, *AJ*, 140, 445
- Chen H.-W., Boettcher E., Johnson S. D., Zahedy F. S., Rudie G. C., Cooksey K. L., Rauch M., Mulchaey J. S., 2019, *ApJ*, 878, L33
- Chen H.-W. et al., 2020a, *MNRAS*, 497, 498
- Chen Y. et al., 2020b, *MNRAS*, 499, 1721
- Churchill C. W., Mellon R. R., Charlton J. C., Vogt S. S., 2003, *ApJ*, 593, 203
- Cicone C. et al., 2014, *A&A*, 562, A21
- Cicone C. et al., 2015, *A&A*, 574, A14
- Cooper T. J. et al., 2021, *MNRAS*, 508, 4359
- Curran S. J., Rydbeck G., Johansson L. E. B., Booth R. S., 1999, *A&A*, 344, 767
- Daddi E. et al., 2010, *ApJ*, 713, 686
- de Jong R. S. et al., 2012, in McLean I. S., Ramsay S. K., Takami H. eds, *Proc. SPIE Conf. Ser. Vol. 8446, Ground-based and Airborne Instrumentation for Astronomy IV*. SPIE, Bellingham, p. 84460T
- de Jong R. S. et al., 2019, *The Messenger*, 175, 3
- DeFelippis D., Bouché N. F., Genel S., Bryan G. L., Nelson D., Marinacci F., Hernquist L., 2021, *ApJ*, 923, 56
- Dekel A. et al., 2009a, *Nature*, 457, 451
- Dekel A., Sari R., Ceverino D., 2009b, *ApJ*, 703, 785
- Driver S. P. et al., 2019, *The Messenger*, 175, 46
- Dutta R. et al., 2023, *MNRAS*, 522, 535
- Elmouttie E., Haynes R. F., Jones K. L., Ehle M., Beck R., Wielebinski R., 1995, *MNRAS*, 275, L53
- Epinat B. et al., 2018, *A&A*, 609, A40
- Faucher-Giguère C.-A., Kereš D., 2011, *MNRAS*, 412, L118
- Faucher-Giguère C.-A., Oh S. P., 2023, preprint ([arXiv:2301.10253](https://arxiv.org/abs/2301.10253))
- Faucher-Giguère C.-A., Hopkins P. F., Kereš D., Muratov A. L., Quataert E., Murray N., 2015, *MNRAS*, 449, 987
- Fox A. J., Richter P., Ashley T., Heckman T. M., Lehner N., Werk J. K., Bordoloi R., Peebles M. S., 2019, *ApJ*, 884, 53
- Fraternali F., 2017, in Fox A., Davé R. eds, *Astrophysics and Space Science Library*, Vol. 396, *Gas Accretion onto Galaxies*. Springer-Verlag, Berlin, p. 323
- Fraternali F., Binney J. J., 2008, *MNRAS*, 386, 935
- Fumagalli M., O'Meara J. M., Prochaska J. X., 2011a, *Science*, 334, 1245
- Fumagalli M., Prochaska J. X., Kasen D., Dekel A., Ceverino D., Primack J. R., 2011b, *MNRAS*, 418, 1796
- Gauthier J.-R., 2013, *MNRAS*, 432, 1444
- Genzel R. et al., 2010, *MNRAS*, 407, 2091
- Gronke M., Oh S. P., Ji S., Norman C., 2022, *MNRAS*, 511, 859
- Hafen Z. et al., 2017, *MNRAS*, 469, 2292
- Hafen Z. et al., 2019, *MNRAS*, 488, 1248
- Hafen Z. et al., 2022, *MNRAS*, 514, 5056
- Hamanowicz A. et al., 2020, *MNRAS*, 492, 2347
- Harnett J. I., Whiteoak J. B., Reynolds J. E., Gardnere F. F., Tzioumis A., 1990, *MNRAS*, 244, 130
- Harris C. R. et al., 2020, *Nature*, 585, 357
- Hartigan J. A., Hartigan P. M., 1985, *Ann. Stat.*, 13, 70
- Hayes M. J., 2023, *MNRAS*, 519, L26
- Hayes M., Melinder J., Östlin G., Scarlata C., Lehnert M. D., Mannerström-Jansson G., 2016, *ApJ*, 828, 49
- Heckman T. M., 2002, in Mulchaey J. S., Stocke J. T. eds, *ASP Conf. Ser. Vol. 254, Extragalactic Gas at Low Redshift*. Astron. Soc. Pac., San Francisco, p. 292
- Heckman T. M., 2003, in Avila-Reese V., Firmani C., Frenk C. S., Allen C. eds, *Revista Mexicana de Astronomía y Astrofísica Conference Series Vol. 17, Revista Mexicana de Astronomía y Astrofísica Conference Series*, Instituto de Astronomía, México 47
- Heckman T. M., Thompson T. A., 2017, in Alsabti A. W., Murdin P., eds, *Handbook of Supernovae*. Springer, p. 2431
- Helton J. M., Johnson S. D., Greene J. E., Chen H.-W., 2021, *MNRAS*, 505, 5497
- Ho S. H., Martin C. L., Kacprzak G. G., Churchill C. W., 2017, *ApJ*, 835, 267
- Ho S. H., Martin C. L., Schaye J., 2020, *ApJ*, 904, 76
- Ho S. H., Martin C. L., Schaye J., 2021, *ApJ*, 923, 137
- Hunter J. D., 2007, *Comput. Sci. Eng.*, 9, 90
- Johnson S. D. et al., 2018, *ApJ*, 869, L1

- Kacprzak G. G., Churchill C. W., Ceverino D., Steidel C. C., Klypin A., Murphy M. T., 2010, *ApJ*, 711, 533
- Kacprzak G. G., Churchill C. W., Nielsen N. M., 2012, *ApJ*, 760, L7
- Kacprzak G. G. et al., 2014, *ApJ*, 792, L12
- Kacprzak G. G., Muzahid S., Churchill C. W., Nielsen N. M., Charlton J. C., 2015, *ApJ*, 815, 22
- Kacprzak G. G., Pointon S. K., Nielsen N. M., Churchill C. W., Muzahid S., Charlton J. C., 2019, *ApJ*, 886, 91
- Kennicutt Robert C. J., 1998, *ApJ*, 498, 541
- Kereš D., Katz N., Weinberg D. H., Davé R., 2005, *MNRAS*, 363, 2
- Klitsch A., Péroux C., Zwaan M. A., Smail I., Oteo I., Biggs A. D., Popping G., Swinbank A. M., 2018, *MNRAS*, 475, 492
- Lamareille F., 2010, *A&A*, 509, A53
- Lan T.-W., Ménard B., Zhu G., 2014, *ApJ*, 795, 31
- Langan I. et al., 2023, *MNRAS*, 521, 546
- Leclercq F. et al., 2022, *A&A*, 663, A11
- Lehner N. et al., 2013, *ApJ*, 770, 138
- Lehner N., Howk J. C., Marasco A., Fraternali F., 2022a, *MNRAS*, 513, 3228
- Lehner N. et al., 2022b, *ApJ*, 936, 156
- Lofthouse E. K. et al., 2020, *MNRAS*, 491, 2057
- Lundgren B. F. et al., 2021, *ApJ*, 913, 50
- Marasco A., Fraternali F., Lehner N., Howk J. C., 2022, *MNRAS*, 515, 4176
- Marinacci F., Binney J., Fraternali F., Nipoti C., Ciotti L., Londrillo P., 2010, *MNRAS*, 404, 1464
- Martin C. L., 2005, *ApJ*, 621, 227
- Martin C. L., Shapley A. E., Coil A. L., Kornei K. A., Bundy K., Weiner B. J., Noeske K. G., Schiminovich D., 2012, *ApJ*, 760, 127
- Martin C. L., Ho S. H., Kacprzak G. G., Churchill C. W., 2019, *ApJ*, 878, 84
- Mo H. J., Mao S., White S. D. M., 1998, *MNRAS*, 295, 319
- Momose R. et al., 2014, *MNRAS*, 442, 110
- Muzahid S. et al., 2020, *MNRAS*, 496, 1013
- Nelson D., Vogelsberger M., Genel S., Sijacki D., Kereš D., Springel V., Hernquist L., 2013, *MNRAS*, 429, 3353
- Nelson D. et al., 2019a, *Comput. Astrophys. Cosmol.*, 6, 2
- Nelson D. et al., 2019b, *MNRAS*, 490, 3234
- Nielsen N. M., Kacprzak G. G., Pointon S. K., Churchill C. W., Murphy M. T., 2018, *ApJ*, 869, 153
- Nielsen N. M., Kacprzak G. G., Pointon S. K., Murphy M. T., Churchill C. W., Davé R., 2020, *ApJ*, 904, 164
- Nielsen N. M., Kacprzak G. G., Sameer, Murphy M. T., Nateghi H., Charlton J. C., Churchill C. W., 2022, *MNRAS*, 514, 6074
- Norris J. M., Muzahid S., Charlton J. C., Kacprzak G. G., Wakker B. P., Churchill C. W., 2021, *MNRAS*, 506, 5640
- Peng C. Y., Ho L. C., Impy C. D., Rix H.-W., 2002, *AJ*, 124, 266
- Péroux C., Howk J. C., 2020, *ARA&A*, 58, 363
- Péroux C., Dessauges-Zavadsky M., D'Odorico S., Sun Kim T., McMahon R. G., 2005, *MNRAS*, 363, 479
- Péroux C., Bouché N., Kulkarni V. P., York D. G., Vladilo G., 2011, *MNRAS*, 410, 2237
- Péroux C., Bouché N., Kulkarni V. P., York D. G., Vladilo G., 2012, *MNRAS*, 419, 3060
- Péroux C. et al., 2016, *MNRAS*, 457, 903
- Péroux C. et al., 2017, *MNRAS*, 464, 2053
- Péroux C., Rahmani H., Arrigoni Battaia F., Augustin R., 2018, *MNRAS*, 479, L50
- Péroux C. et al., 2019, *MNRAS*, 485, 1595
- Péroux C., Nelson D., van de Voort F., Pillepich A., Marinacci F., Vogelsberger M., Hernquist L., 2020, *MNRAS*, 499, 2462
- Péroux C. et al., 2022, *MNRAS*, 516, 5618
- Pillepich A. et al., 2019, *MNRAS*, 490, 3196
- Pointon S. K., Kacprzak G. G., Nielsen N. M., Muzahid S., Murphy M. T., Churchill C. W., Charlton J. C., 2019, *ApJ*, 883, 78
- Rahmani H. et al., 2018a, *MNRAS*, 474, 254
- Rahmani H. et al., 2018b, *MNRAS*, 480, 5046
- Rahmati A., Schaye J., Bower R. G., Crain R. A., Furlong M., Schaller M., Theuns T., 2015, *MNRAS*, 452, 2034
- Ramesh R., Nelson D., Pillepich A., 2023, *MNRAS*, 522, 1535
- Ribaud J., Lehner N., Howk J. C., Werk J. K., Tripp T. M., Prochaska J. X., Meiring J. D., Tumlinson J., 2011, *ApJ*, 743, 207
- Rodríguez-Puebla A., Primack J. R., Avila-Reese V., Faber S. M., 2017, *MNRAS*, 470, 651
- Rubin K. H. R., Prochaska J. X., Koo D. C., Phillips A. C., 2012, *ApJ*, 747, L26
- Rubin K. H. R., Prochaska J. X., Koo D. C., Phillips A. C., Martin C. L., Winstrom L. O., 2014, *ApJ*, 794, 156
- Rubin K. H. R. et al., 2022, *ApJ*, 936, 171
- Sameer et al., 2022, *MNRAS*, 510, 5796
- Schroetter I. et al., 2019, *MNRAS*, 490, 4368
- Steidel C. C., Bogosavljević M., Shapley A. E., Kollmeier J. A., Reddy N. A., Erb D. K., Pettini M., 2011, *ApJ*, 736, 160
- Stern J. et al., 2021, *MNRAS*, 507, 2869
- Stewart K. R., Kaufmann T., Bullock J. S., Barton E. J., Maller A. H., Diemand J., Wadsley J., 2011, *ApJ*, 738, 39
- Stewart K. R., Brooks A. M., Bullock J. S., Maller A. H., Diemand J., Wadsley J., Moustakas L. A., 2013, *ApJ*, 769, 74
- Szakacs R. et al., 2021, *MNRAS*, 505, 4746
- Tacconi L. J., Genzel R., Sternberg A., 2020, *ARA&A*, 58, 157
- Tejos N. et al., 2021, *MNRAS*, 507, 663
- Theuns T., 2021, *MNRAS*, 500, 2741
- Truong N., Pillepich A., Nelson D., Werner N., Hernquist L., 2021, *MNRAS*, 508, 1563
- Tumlinson J., Peebles M. S., Werk J. K., 2017, *ARA&A*, 55, 389
- van de Voort F., Schaye J., Booth C. M., Haas M. R., Dalla Vecchia C., 2011, *MNRAS*, 414, 2458
- van de Voort F., Schaye J., Altay G., Theuns T., 2012, *MNRAS*, 421, 2809
- van de Voort F., Bieri R., Pakmor R., Gómez F. A., Grand R. J. J., Marinacci F., 2021, *MNRAS*, 501, 4888
- Veilleux S., Bland-Hawthorn J., 1997, *ApJ*, 479, L105
- Veilleux S., Cecil G., Bland-Hawthorn J., 2005, *ARA&A*, 43, 769
- Veilleux S., Shopbell P. L., Miller S. T., 2001, *AJ*, 121, 198
- Walter F. et al., 2020, *ApJ*, 902, 111
- Wang B., 1993, *ApJ*, 415, 174
- Weiner B. J. et al., 2006, *ApJ*, 653, 1027
- Wendt M., Bouché N. F., Zabl J., Schroetter I., Muzahid S., 2021, *MNRAS*, 502, 3733
- Weng S. et al., 2023, *MNRAS*, 519, 931
- Werk J. K. et al., 2014, *ApJ*, 792, 8
- Wilde M. C. et al., 2021, *ApJ*, 912, 9
- Wilde M. C. et al., 2023, preprint (arXiv:2301.02718)
- Wolfe A. M., Gawiser E., Prochaska J. X., 2005, *ARA&A*, 43, 861
- Wolfe A. M., Turnshek D. A., Smith H. E., Cohen R. D., 1986, *ApJS*, 61, 249
- Zabl J. et al., 2019, *MNRAS*, 485, 1961
- Zahedy F. S., Chen H.-W., Johnson S. D., Pierce R. M., Rauch M., Huang Y.-H., Weiner B. J., Gauthier J.-R., 2019, *MNRAS*, 484, 2257
- Zabl J. et al., 2021, *MNRAS*, 507, 4294
- Zahedy F. S. et al., 2021, *MNRAS*, 506, 877
- Zwaan M. A., van der Hulst J. M., Briggs F. H., Verheijen M. A. W., Ryan-Weber E. V., 2005, *MNRAS*, 364, 1467

SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

t.paper.csv

Please note: Oxford University Press is not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

APPENDIX A: ADDITIONAL VELOCITY MAPS

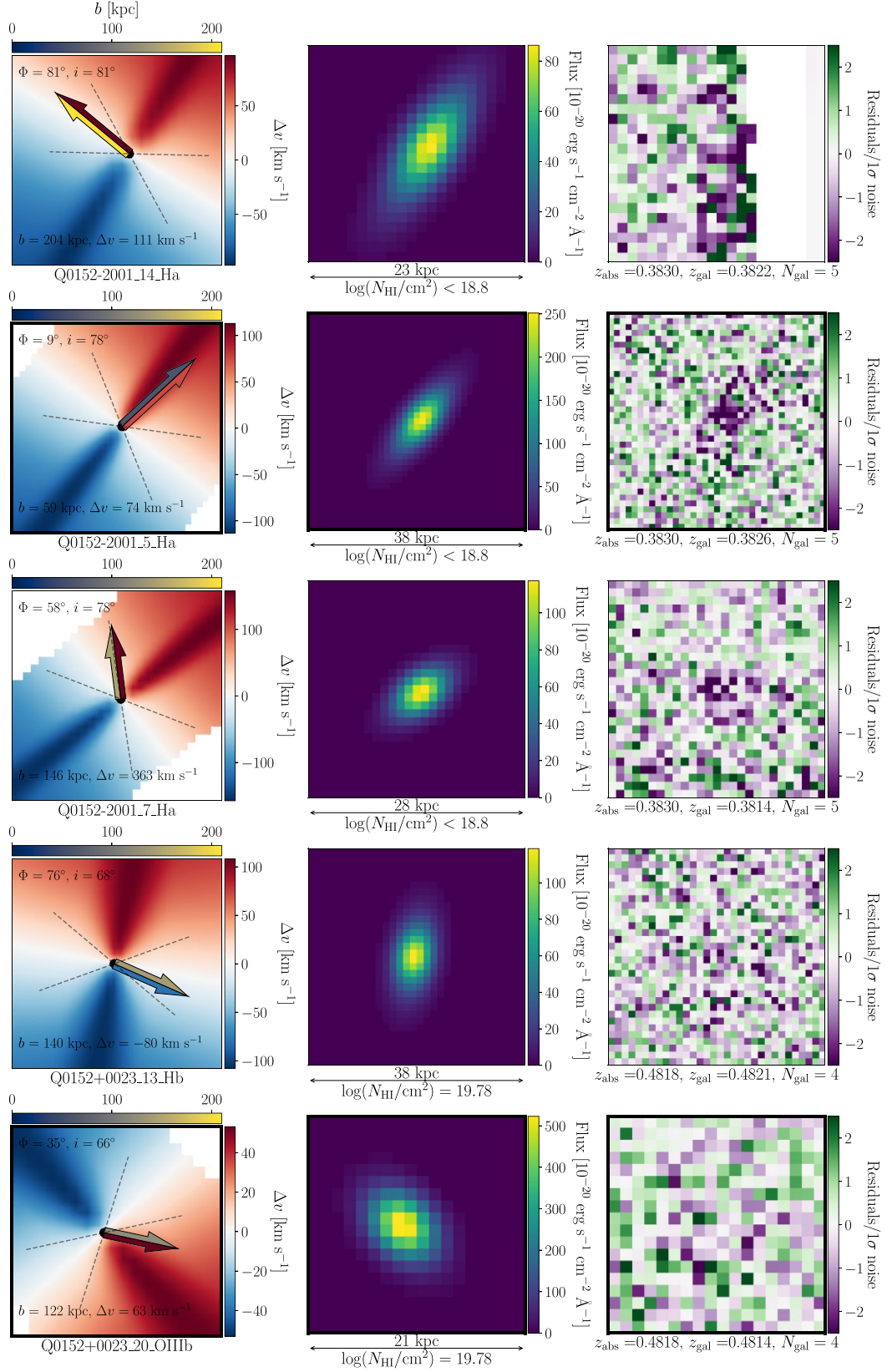


Figure A1. Continuation of Fig. 1, refer to its caption.

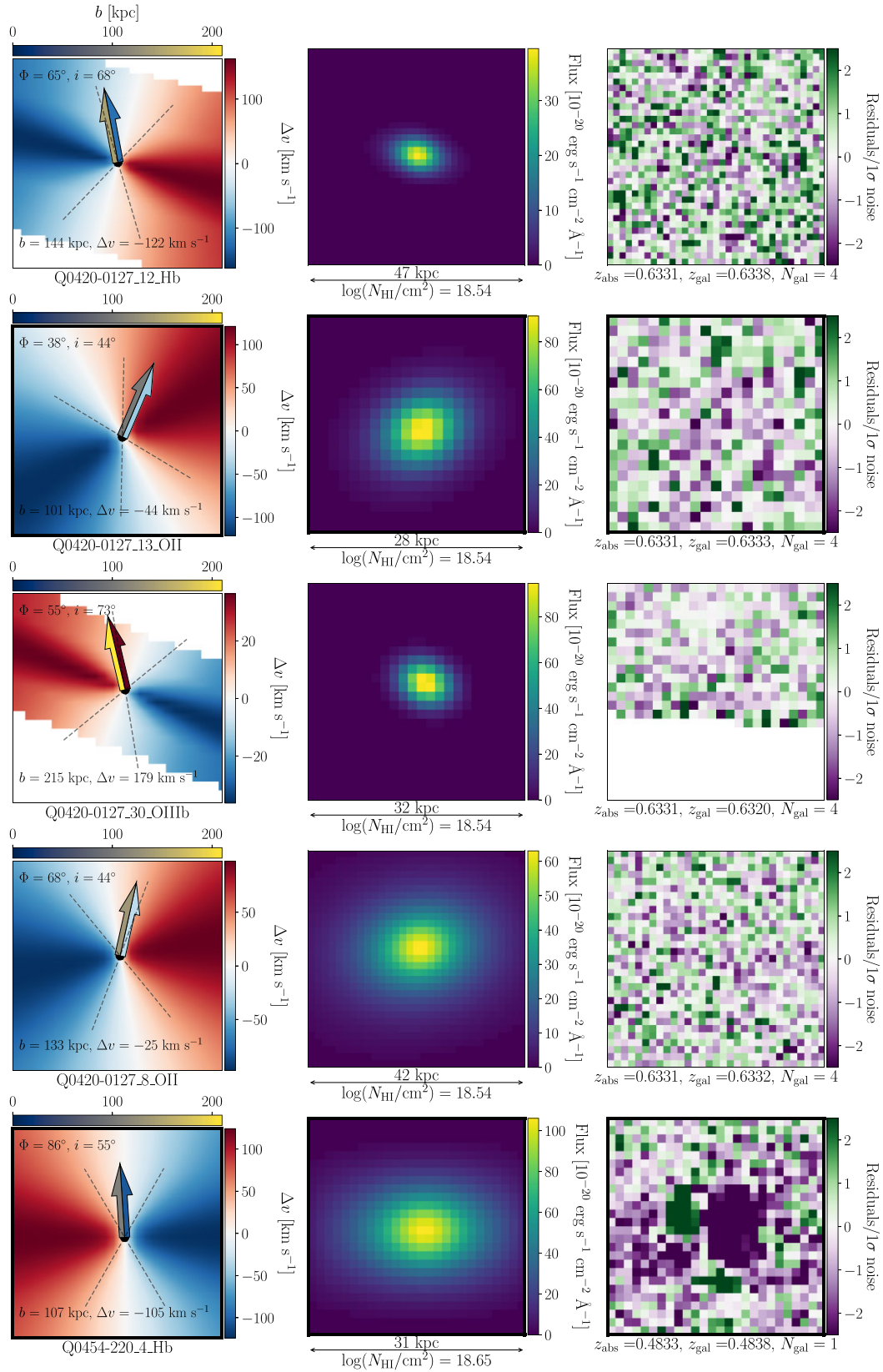


Figure A2. Continuation of Fig. 1 refer to its caption. The significant flux residuals for object Q0454–220.4 shown in the last row are caused by the saturation of a nearby bright star in the MUSE field.

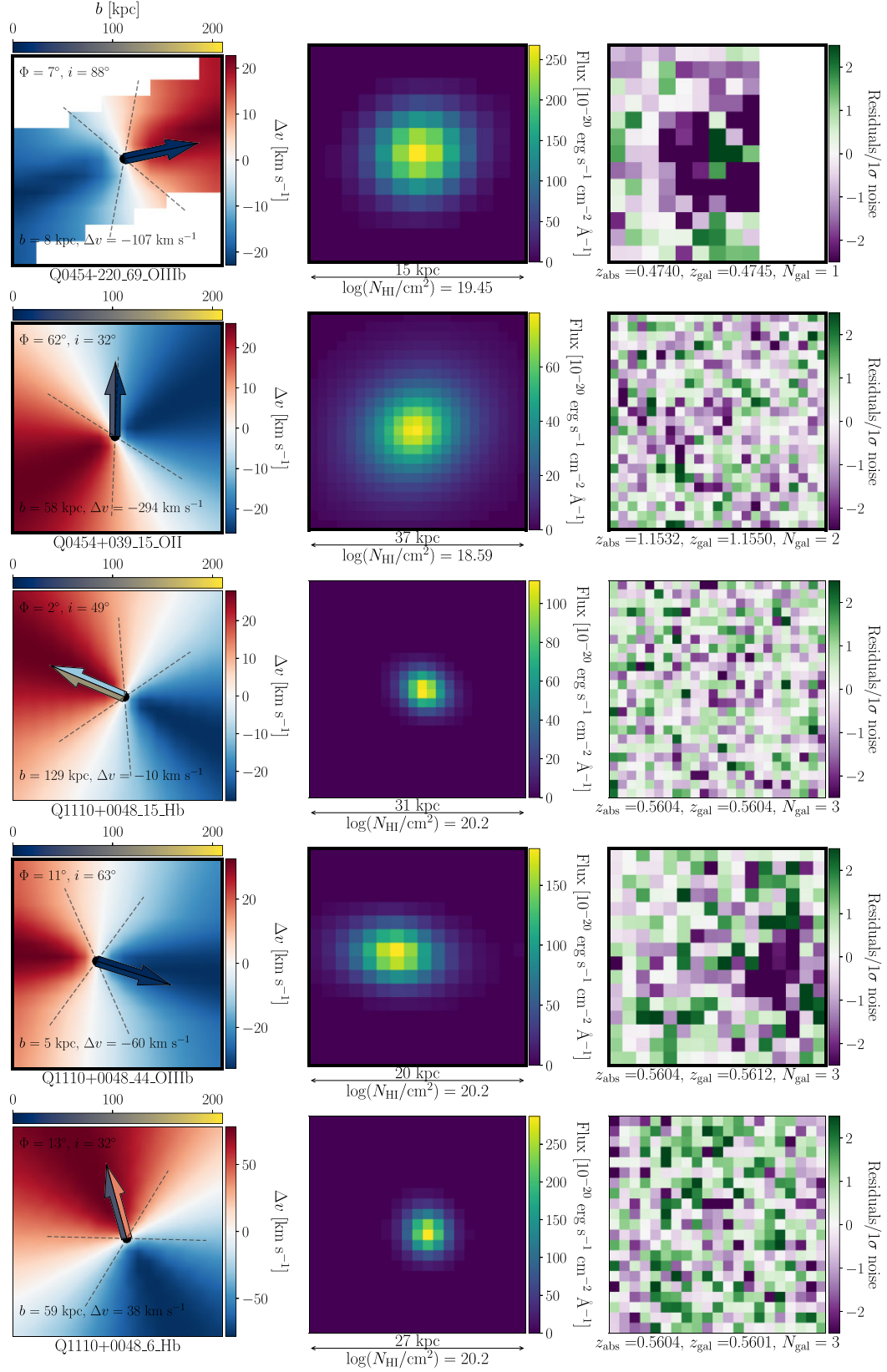


Figure A3. Continuation of Fig. 1 refer to its caption. Object Q0454–220.69 (top row) is located less than two arcseconds from the background QSO. The residuals arise from the QSO flux contaminating the galaxy flux.

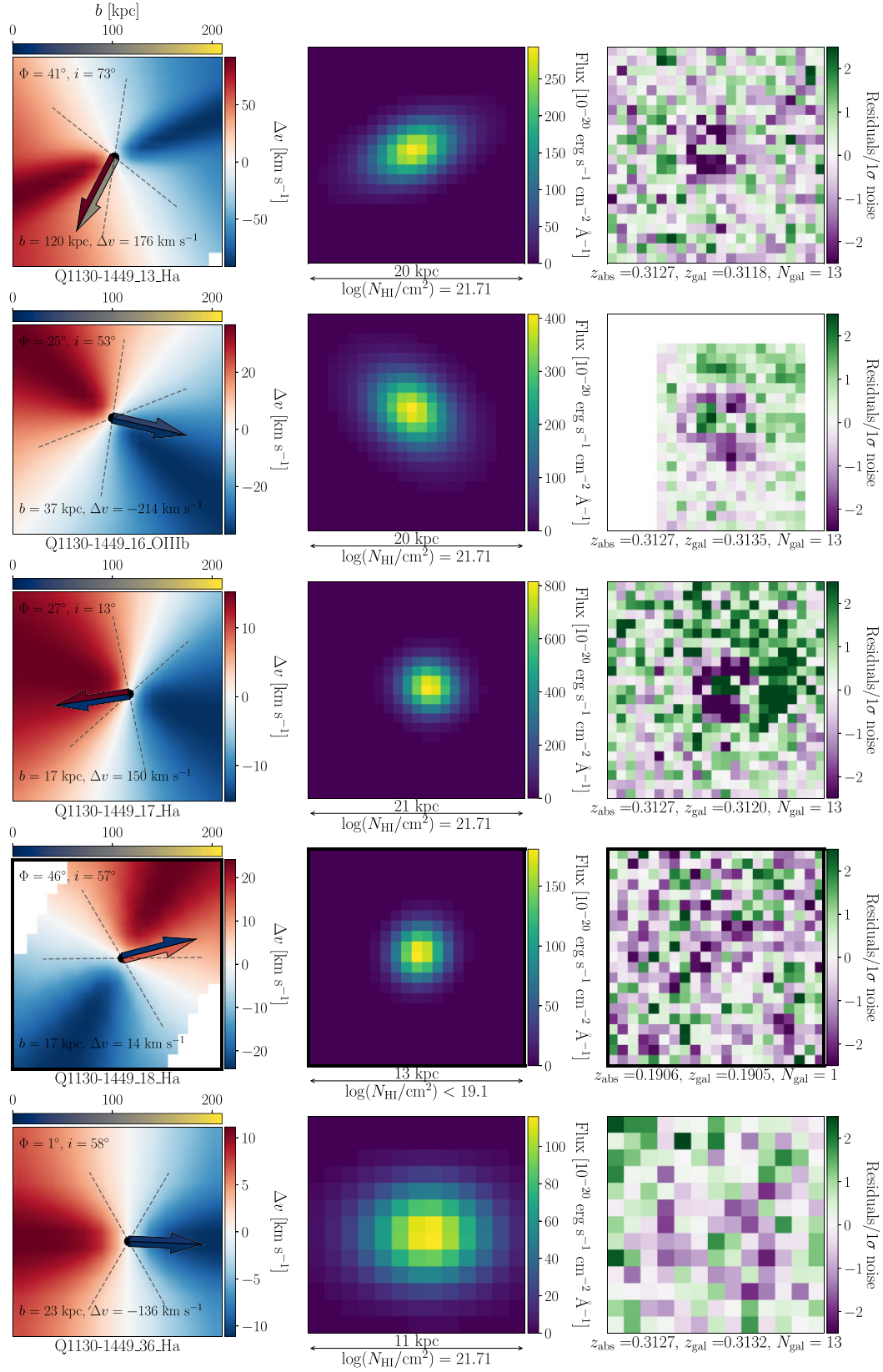


Figure A4. Continuation of Fig. 1 refer to its caption. The galaxies in the top three rows have significant flux residuals due to the ionizing nebula associated with a galaxy group at the absorber redshift.

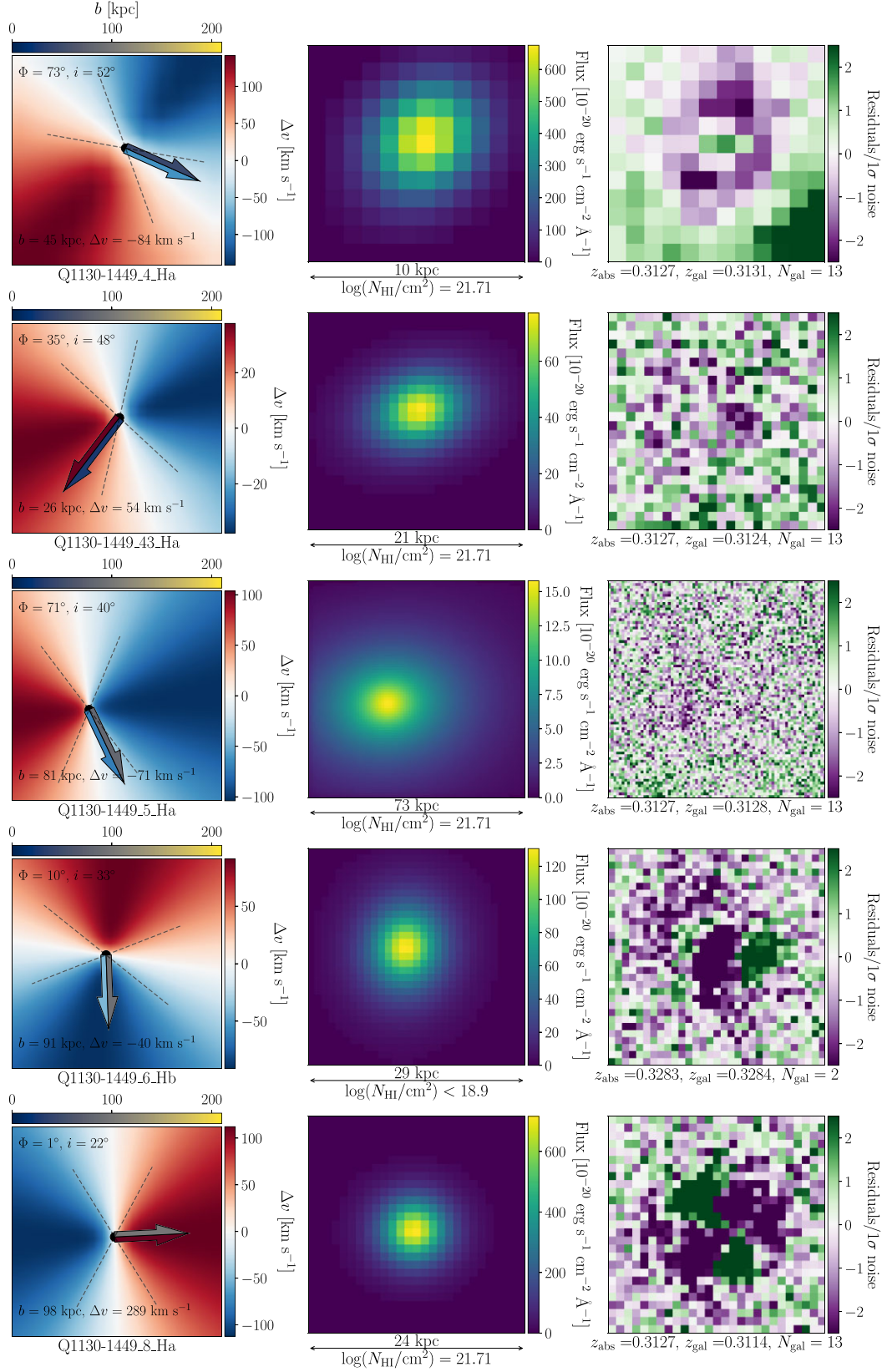


Figure A5. Continuation of Fig. 1 refer to its caption. The galaxy in the top row is similarly affected by the emission-line flux of the ionizing nebula. In addition, galaxies Q1130–1449.6 and Q1130–1449.8 have asymmetric emission-line profiles and the flux residuals are likely caused by a merger and strong outflows, respectively.

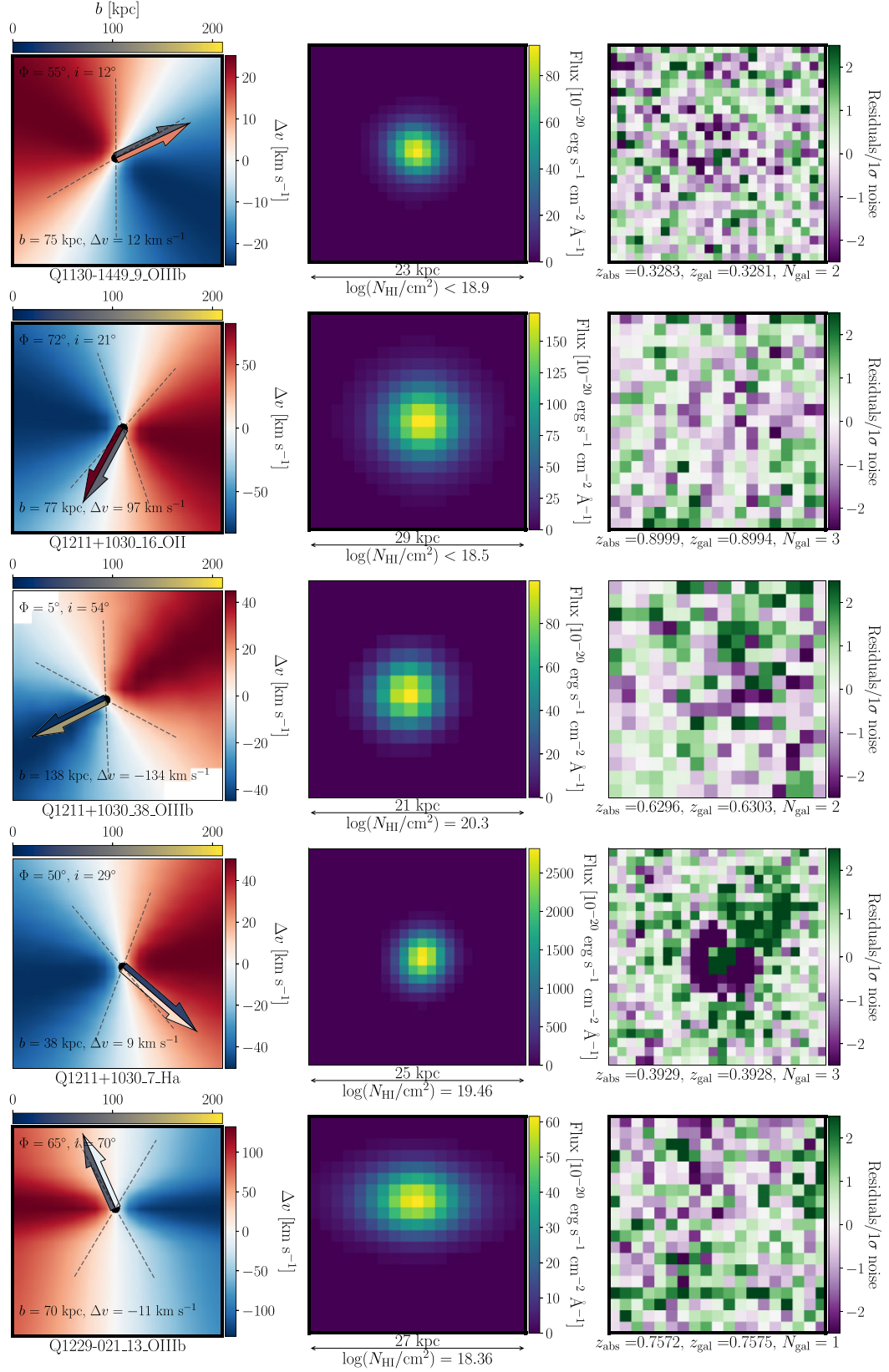


Figure A6. Continuation of Fig. 1 refer to its caption. Galaxy Q1211+1030.7 likely hosts an active galactic nucleus from its position on the [O III]/H β versus [O II]/H β classification diagram (Lamareille 2010; Weng et al. 2023). The flux excess away from the galaxy disc is likely caused by the AGN ionizing gas.

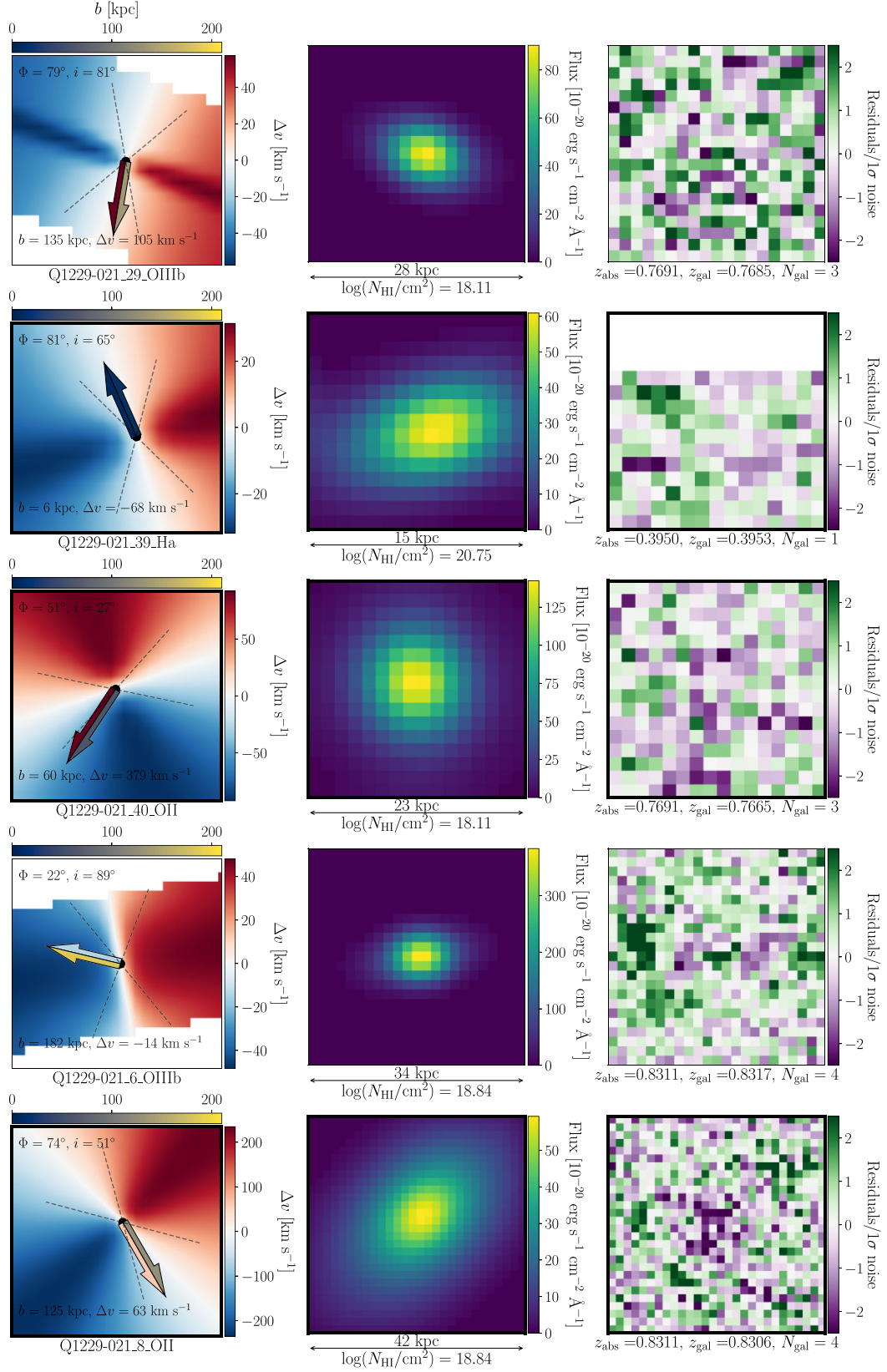
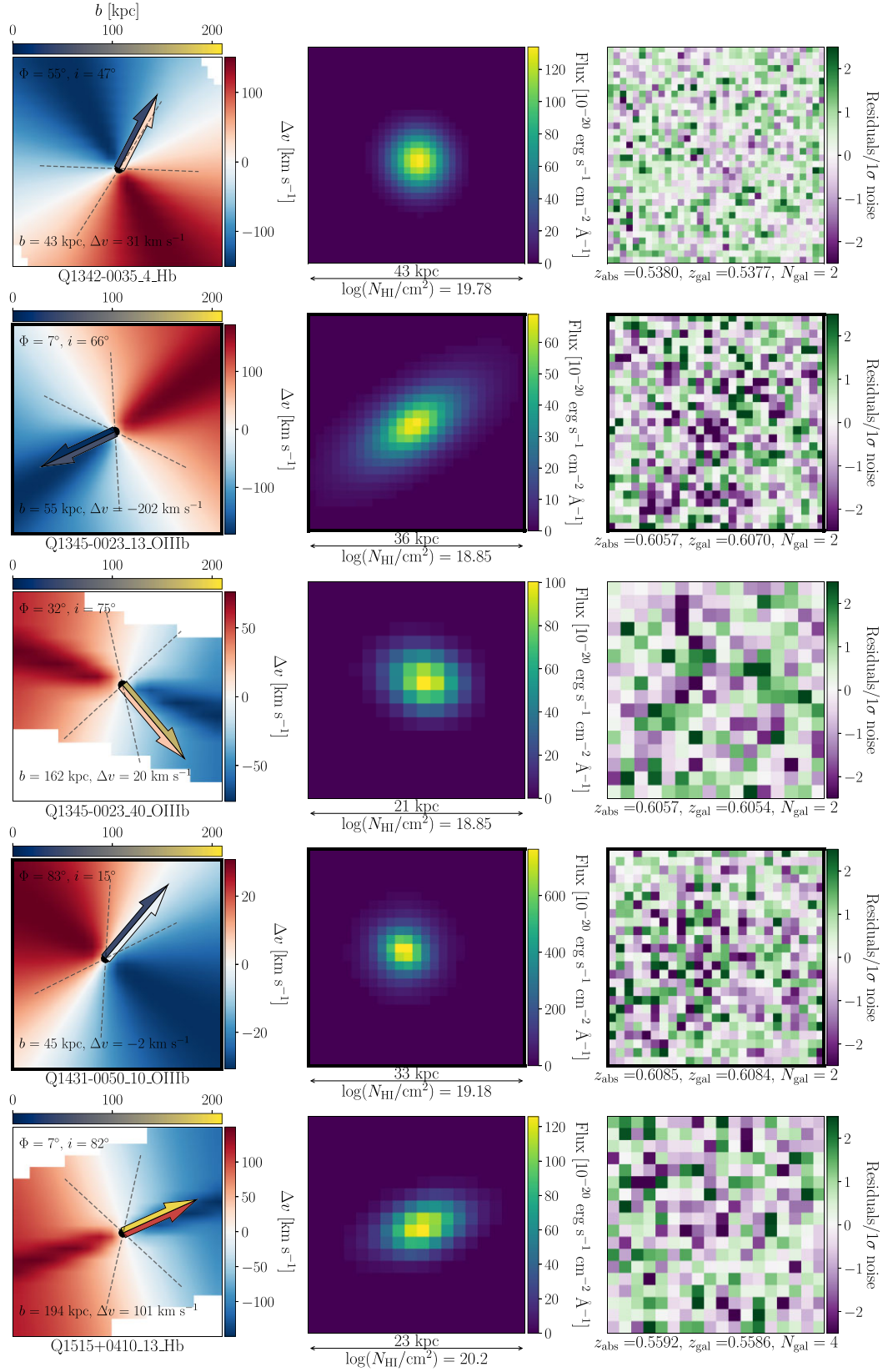


Figure A7. Continuation of Fig. 1 refer to its caption. The flux residuals seen in the fit for Q1229–021.6 are likely also caused by an AGN (Weng et al. 2023).

**Figure A8.** Continuation of Fig. 1 refer to its caption.

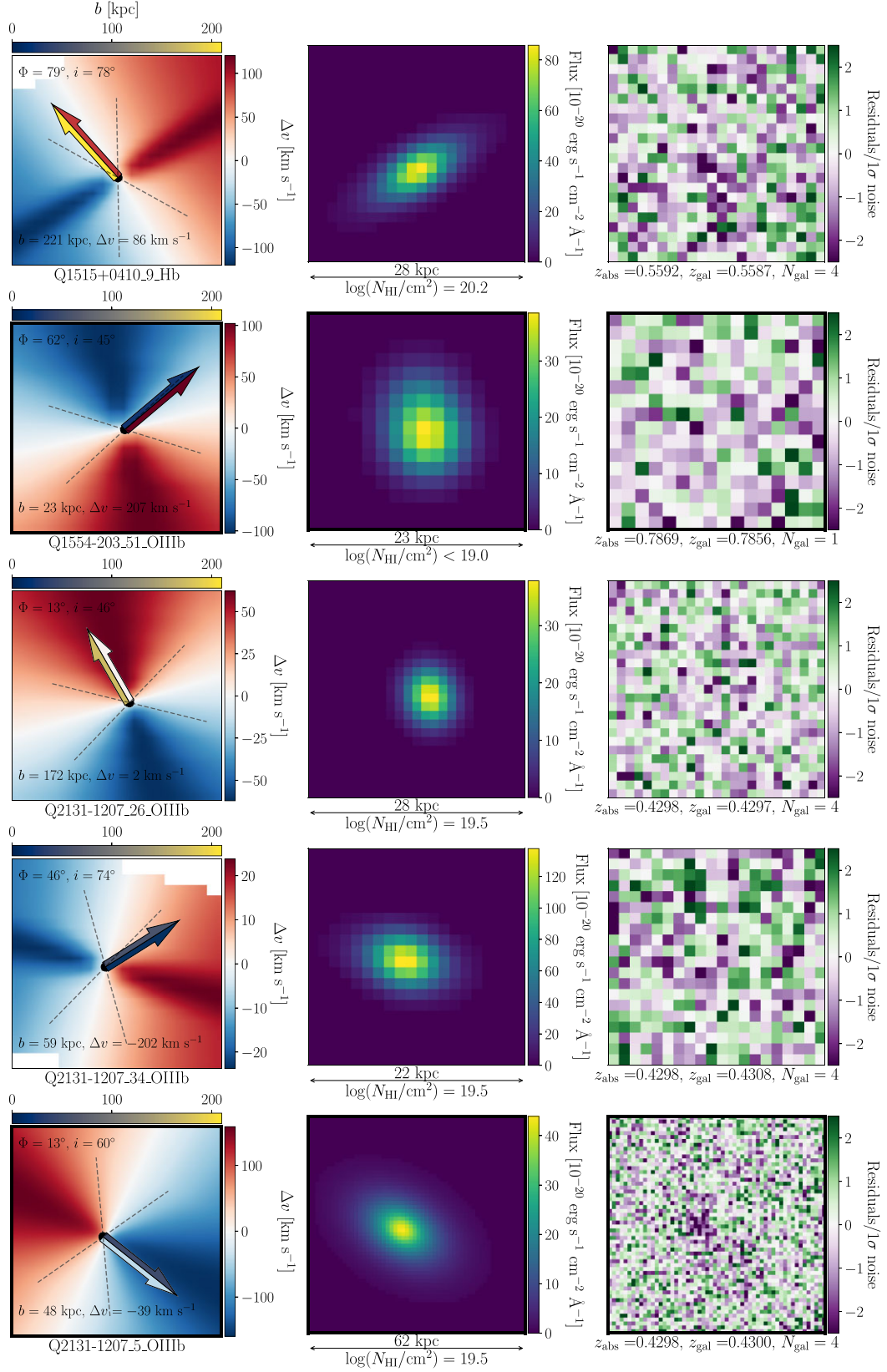


Figure A9. Continuation of Fig. 1 refer to its caption.

APPENDIX B: FLOW CLASSIFICATIONS

In this section, we present the criteria used to classify absorbers as belonging to the galaxy disc, inflowing, or outflowing. Additionally, we explore other origins for the gas, such as the intragroup medium or a low-mass, quiescent satellite at low impact parameters from the absorber that is not detected. We also discuss the origins of the absorber in more ambiguous cases where the data is not sufficient to draw a conclusion, multiple phenomena are possible, or the absorber might trace tidally stripped gas. Of the 27 absorbers with associated galaxies in the MUSE-ALMA Haloes survey, five have their associated galaxies modelled in previous works (Péroux et al. 2017, 2019; Klitsch et al. 2018; Rahmani et al. 2018a,b; Szakacs et al. 2021). We also note that in Hamanowicz et al. (2020), the stellar properties of galaxies associated with a further nine absorbers (towards Q1130–1449, Q1211+1030, and Q1229–021) were studied. In this section, we present the kinematic analysis of 22 absorbers that have not yet analysed in detail.

B1 Absorbing gas associated with the galaxy disc

B1.1 Q0454+039 $z_{\text{abs}} = 0.8596$; $\log[N(\text{H I})/\text{cm}^{-2}] = 20.69$

This damped Ly- α absorber is associated with a single galaxy (Q0454+039_57) at an impact parameter of 18 kpc. The galaxy is detected in emission but is too faint to model with GALPAK. Given the high column density of the absorber and low impact parameter of $b < 20$ kpc, it is likely that the absorber traces gas in the galaxy disc. However, no kinematic nor photometric modelling is available to confirm this hypothesis.

B1.2 Q1110+0048 $z_{\text{abs}} = 0.5604$; $\log[N(\text{H I})/\text{cm}^{-2}] = 20.20$

While there are three galaxies associated with this absorber, the intervening gas is likely hosted by the galaxy at an impact parameter of ~ 5 kpc. The absorber velocity aligns in sign and magnitude with the rotational velocity of the ionized gas disc (row 4 in Fig. A3).

B1.3 Q1211+1030 $z_{\text{abs}} = 0.6296$; $\log[N(\text{H I})/\text{cm}^{-2}] = 20.30$

Q1211+1030_57 is located at $b < 20$ kpc from the QSO and likely hosts the damped Ly- α absorber. However, there is no photometric or kinematic modelling available for this galaxy. Another object (Q1211+1030_38) is found at $b = 140$ kpc and the absorber aligns with this galaxy's ionized gas velocity field (see row 3 in Fig. A6). Given the characteristic radius of DLAs are ~ 20 kpc, it is far more likely the absorber is associated with the galaxy disc of object 57.

B1.4 Q1515+0410 $z_{\text{abs}} = 0.5592$; $\log[N(\text{H I})/\text{cm}^{-2}] = 20.20$

Four galaxies are found at the redshift of this absorber, but object 4 is at an impact parameter of 10 kpc whereas the rest are > 150 kpc away. No kinematic modelling is available for this galaxy, but the photometric azimuthal angle is measured to be 20° . It is most likely that the absorber is probing dense gas in the galaxy disc.

B2 Absorbing gas associated with accretion

B2.1 Q1345–0035 $z_{\text{abs}} = 0.6057$; $\log[N(\text{H I})/\text{cm}^{-2}] = 18.85$

This absorber is consistent with gas that is co-rotating with the galaxy at the lowest impact parameter ($b = 56$ kpc; see the second row in

Fig. A8). With an azimuthal angle of $\Phi = 7^\circ$, the alignment with the major axis suggests the gas is inflowing onto the galaxy.

B3 Absorbing gas associated with outflows

We explore the three most probable cases of outflows first.

B3.1 Q0454+039 $z_{\text{abs}} = 1.1532$; $\log[N(\text{H I})/\text{cm}^{-2}] = 18.59$

The galaxy at lowest impact parameter to the absorber (Q0454+039_15, $b = 60$ kpc) has a projected minor axis aligned with the QSO sightline ($\Phi = 62^\circ$; see the second row in Fig. A3). Q0454+039_65 is another associated galaxy that is located 130 kpc from the absorber, for which we do not have kinematic modelling. We find that the absorber is likely tracing the expelled gas from Q0454+039_15. While the velocity of the absorber is the same sign as the galaxy velocity field, the $\Delta v_{\text{LOS}} \approx -290 \text{ km s}^{-1}$ is multiple times the measured maximum rotational velocity.

B3.2 Q1229–021 $z_{\text{abs}} = 0.3950$; $\log[N(\text{H I})/\text{cm}^{-2}] = 20.75$

There is a single galaxy associated with this absorber at an impact parameter of $b < 7$ kpc. The azimuthal angle of $\Phi = 81^\circ$ suggests the absorber traces outflowing gas with $|v_{\text{LOS}}| = 70 \text{ km s}^{-1}$ (see the second row in Fig. A7). Given the proximity of the sightline to the galaxy centre, the strong absorber could also trace gas in the galaxy disc.

B3.3 Q1554–203 $z_{\text{abs}} = 0.7869$; $\log[N(\text{H I})/\text{cm}^{-2}] < 19.00$

A single galaxy is found at the redshift of this absorber at an impact parameter of $b = 23$ kpc. The kinematic modelling reveals the absorber is found at an azimuthal angle of $\phi = 62^\circ$ and relative velocity of $\sim 200 \text{ km s}^{-1}$ from the systemic redshift of the galaxy. The absorber being traced appears to be outflowing.

The remaining absorbers in this section may arise from several origins.

B3.4 Q0138–0005; $z_{\text{abs}} = 0.7821$; $\log[N(\text{H I})/\text{cm}^{-2}] = 19.81$

A single galaxy was detected in the Q0138–0005 field within $\pm 500 \text{ km s}^{-1}$ of the absorber. We find that the QSO line of sight aligns closely with the major axis of the associated galaxy ($\Phi = 79^\circ$; see the first row in Fig. 1). The zinc abundance is measured to be $[\text{Zn}/\text{H}] = 0.28 \pm 0.16$, perhaps indicating a metal-enriched outflow. However, the line of sight velocity difference between absorber and galaxy is only $\sim 40 \text{ km s}^{-1}$. The true velocity of the absorbing gas may be much larger as the disc has an inclination of $i = 66^\circ$, but the seeing conditions for this observation ($R_{\text{PSF}} = 1.11 \text{ arcsec}$) mean we cannot rule out a faint quiescent galaxy nearer the strong H I absorber.

B3.5 Q1229–021 $z_{\text{abs}} = 0.7572$; $\log[N(\text{H I})/\text{cm}^{-2}] = 18.36$

A single galaxy is found at the absorber redshift in the MUSE data. The absorber appears aligned with the minor axis of object 13 in this field ($\Phi = 65^\circ$ from the last row in Fig. A6). However, the absorber velocity is within 15 km s^{-1} of the systemic redshift of its galaxy counterpart and may be static gas in the halo. It is also possible that the velocity component orthogonal to the sightline is large because the galaxy is nearly viewed edge-on ($i = 70^\circ$).

B3.6 Q1229–021 $z_{\text{abs}} = 0.7691$; $\log[N(\text{H I})/\text{cm}^{-2}] = 18.11$

There are three galaxies associated with this absorber, with impact parameters ranging from 60 to 180 kpc. Curiously, the absorber has the opposite velocity sign to the ionized gas velocity for object Q1229–021_40 at lowest impact parameter. This velocity difference approaches 400 km s^{-1} and there is no clear explanation for how this absorber could arise if it is tied to this galaxy. Such large velocity differences can often equate to physical distances greater than one Mpc (Ho, Martin & Schaye 2021). A possible origin for the absorber is that it traces gas being expelled from galaxy Q1229–021_29 ($\Phi = 79^\circ$, see the middle row in Fig. A7), but we may also be tracing gas in the intragroup medium.

B3.7 Q1229–021 $z_{\text{abs}} = 0.8311$; $\log[N(\text{H I})/\text{cm}^{-2}] = 18.84$

There are four associated galaxies located at impact parameters ranging from 125 to 187 kpc. The orientation and velocity of the absorber is consistent with outflowing gas from the nearest galaxy, Q1229–021_8 (see the last row in Fig. A7). An azimuthal angle of $\Phi = 74^\circ$ and velocity difference of $\Delta v_{\text{LOS}} = 60 \text{ km s}^{-1}$ are measured from the kinematic modelling. However, we also cannot rule out the absorber tracing gas in the intragroup medium.

B4 Absorbing gas of unclear origin

B4.1 Q0152+0023 $z_{\text{abs}} = 0.4818$; $\log[N(\text{H I})/\text{cm}^{-2}] = 19.78$

We find four galaxies associated with this absorber, with impact parameters ranging from 120 to 190 kpc. For the nearest galaxy, Q0152+0023_20, to the QSO sightline, the velocity of the absorber is consistent with co-rotating gas. The absorber also aligns with the ionized gas map of another galaxy (Q0152+0023_13) ~ 10 kpc away from Q0152+0023_20. This can be seen in the last row of Fig. A1. However, despite the consistency in velocities, the high column density absorber is outside the virial radius of both galaxies and, hence, likely traces a quiescent, low-mass galaxy near the sightline.

B4.2 Q0454–220 $z_{\text{abs}} = 0.4833$; $\log[N(\text{H I})/\text{cm}^{-2}] = 18.65$

The GALPAK fits for the single galaxy detected at the absorber redshift are unreliable due to a bright nearby star in the MUSE data saturating the data (see the last row in Fig. A2). This absorber-galaxy pair has already been studied in the literature, and three more associated galaxies are located outside the MUSE field of view (Kacprzak et al. 2010; Norris et al. 2021). The system consists of multiple components with varying metallicities, and Norris et al. (2021) finds that the gas arises from a mix of accretion, outflows, and intergroup material. The kinematic modelling from this work does not shed further light on the origin of the various gas components.

B4.3 Q0454–220 $z_{\text{abs}} = 0.4744$; $\log[N(\text{H I})/\text{cm}^{-2}] = 19.45$

The associated galaxy Q0454–220_69 has an impact parameter of <10 kpc and we expect the absorber to be tracing gas in the galaxy disc. However, the absorber appears to be counter-rotating with a velocity of $\sim 100 \text{ km s}^{-1}$ from the systemic redshift of the galaxy (see, first row in Fig. A3). The modelled fit is unreliable because of the small galaxy size and nearby QSO contaminating the fit. A manual inspection of the [O III] emission line in the MUSE cube reveals that the galaxy emission flux is oriented such that the side with positive velocity is towards the west. Hence, the absorber is

indeed counter-rotating, but the derived kinematic properties from GALPAK are not reliable. There is also no detection in the *HST* broad-band imaging available for this continuum-faint object.

B4.4 Q1130–1449 $z_{\text{abs}} = 0.1906$; $\log[N(\text{H I})/\text{cm}^{-2}] < 19.10$

There is a single galaxy (Q1130–1449_18) detected at the redshift of this absorber. The velocity of the ionized and neutral gas phases are consistent (see the fourth row in Fig. A4). However, the low impact parameter $b = 18$ kpc means we cannot separate gas rotating with the galaxy disc from accreting gas.

B4.5 Q1130–1449 $z_{\text{abs}} = 0.3283$; $\log[N(\text{H I})/\text{cm}^{-2}] < 18.90$

The galaxy Q1130–1449_9 is found nearest the absorber at $b = 75$ kpc. The absorber velocity has the opposite sign to the ionized gas velocity and is found near the projected minor axis ($\Phi = 55^\circ$, first row in Fig. A6). However, the measured inclination of $i = 12^\circ$ suggests the galaxy is almost face-on and it becomes difficult to assign the absorber to any particular process. Instead, we may be tracing gas in the intragroup medium, as there is another galaxy with an impact parameter of 90 kpc.

B4.6 Q1211+1030 $z_{\text{abs}} = 0.3929$; $\log[N(\text{H I})/\text{cm}^{-2}] = 19.46$

There are three galaxies almost equidistant from this absorber on the sky (37 to 39 kpc). The galaxy Q1211+1030_7 with available kinematic modelling likely hosts an AGN, leaving significant residuals in the modelled fit (second last row in Fig. A6). From the photometric modelling, an azimuthal angle of $\Phi = 78^\circ$ is found for object Q1211p1030_9. It is difficult to assess the origin of the absorber, and an alternative hypothesis for all three galaxies is that the gas is intragroup material, especially as the absorber is almost equidistant from each galaxy. Objects 7 and 9 in this field are only several kiloparsecs apart, meaning tidal interactions are taking place and the stripped gas can be probed in absorption.

B4.7 Q1211+1030 $z_{\text{abs}} = 0.8999$; $\log[N(\text{H I})/\text{cm}^{-2}] < 18.50$

There are two galaxies associated with this absorber with impact parameters of 77 and 175 kpc. The absorber aligns with the minor axis of the galaxy at smaller b ($\Phi = 72^\circ$, second row in Fig. A6), but the inclination angle of $i = 21^\circ$ prevents us from making a certain judgement.

B4.8 Q1342–0035 $z_{\text{abs}} = 0.5380$; $\log[N(\text{H I})/\text{cm}^{-2}] = 19.78$

Of the two galaxies associated with this absorber, the nearest galaxy, with an impact parameter of 24.3 kpc has not been modelled with GALPAK. From the GALFIT models of the galaxy using *HST* broad-band imaging, an azimuthal angle of $\Phi_{\text{phot}} = 58^\circ$ is measured. Without the velocity field of the ionized gas, it is not possible to constrain the origin of the absorber. From the velocity map of object 4 at an impact parameter of 44 kpc, the absorber velocity is opposite in sign to the modelled ionized gas velocity.

B4.9 Q1431–0050 $z_{\text{abs}} = 0.6868$; $\log[N(\text{H I})/\text{cm}^{-2}] = 18.40$

There are two galaxies at this absorber's redshift. Both do not have kinematic modelling and from the photometric fitting, it is not possible to make a convincing case for the origin of the gas.

B4.10 Q1431–0050 $z_{abs} = 0.6085$; $\log[N(H\text{I})/\text{cm}^{-2}] = 19.18$

From the second last row of Fig. A4, we see that the absorber aligns with the minor axis of the galaxy at closest impact parameter ($b = 45$ kpc). However, the galaxy appears almost face-on and the absorber

is found at the systemic velocity of the galaxy. Hence, the origin of the gas is unclear.

This paper has been typeset from a \LaTeX file prepared by the author.