



# Over 500 Days in the Life of the Photosphere of the Type Iax Supernova SN 2014dt

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## Abstract

Type Iax supernovae (SNe Iax) are the largest known class of peculiar white dwarf SNe, distinct from normal Type Ia supernovae (SNe Ia). The unique properties of SNe Iax, especially their strong photospheric lines out to extremely late times, allow us to model their optical spectra and derive the physical parameters of the long-lasting photosphere. We present an extensive spectral timeseries, including 21 new spectra, of SN Iax 2014dt from +11 to +562 days after maximum light. We are able to reproduce the entire timeseries with a self-consistent, nearly unaltered deflagration explosion model from Fink et al. using TARDIS, an open source radiative-transfer code. We find that the photospheric velocity of SN 2014dt slows its evolution between +64 and +148 days, which closely overlaps the phase when we see SN 2014dt diverge from the normal spectral evolution of SNe Ia (+90 to +150 days). The photospheric velocity at these epochs,  $\sim 400\text{--}1000\text{ km s}^{-1}$ , may demarcate a boundary within the ejecta below which the physics of SNe Iax and normal SNe Ia differ. Our results suggest that SN 2014dt is consistent with a weak deflagration explosion model that leaves behind a bound remnant and drives an optically thick, quasi-steady-state wind creating the photospheric lines at late times. The data also suggest that this wind may weaken at epochs past +450 days, perhaps indicating a radioactive power source that has decayed away.

*Unified Astronomy Thesaurus concepts:* Supernovae (1668); Type Ia supernovae (1728); Radiative transfer (1335); Radiative transfer simulations (1967); Supernova dynamics (1664)

*Supporting material:* data behind figure, machine-readable tables

## 1. Introduction

Type Ia supernovae (SNe Ia) are important standard candles in cosmology and a major source of chemical enrichment in the universe. Despite their importance, a complete understanding of the progenitor system and explosion mechanics of SNe Ia remains elusive. The use of SNe Ia as standard candles implies that they

emerge from a homogeneous origin, but an increasing number of observations shows there is diversity within the group. The largest class of peculiar versions of SNe Ia—with over 60 known members—are Type Iax supernovae (SNe Iax; Foley et al. 2013; Jha 2017). With an occurrence of approximately 15%–30% of the SNe Ia rate (Foley et al. 2013; Miller et al. 2017; Srivastav et al. 2022), it is imperative to improve our understanding of this growing SN Iax population.

SNe Iax, also called SN 2002cx-like SNe (Li et al. 2003; Filippenko 2003; Jha et al. 2006; Phillips et al. 2007), start off resembling normal SNe Ia; both have early-time optical spectra dominated by iron-group elements (IGEs) like Cr, Fe, Co, and Ni, as well as intermediate-mass elements (IMEs) like Si, Na, and Ca. The biggest differences at early times are that SNe Iax have

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strongly mixed, low-velocity ejecta and lower luminosities relative to normal SNe Ia. Nonetheless, SNe Iax are not a homogeneous class. They cover a wide range of peak absolute magnitudes, ranging from nearly as bright as normal SNe Ia to objects 100 times fainter (Jha 2017). SNe Iax also cover a broad range of photospheric velocities,  $v_{\text{phot}} \sim 2000\text{--}8000 \text{ km s}^{-1}$  (Foley et al. 2013) at maximum light, lower than typical SNe Ia velocities,  $\sim 9000\text{--}15,000 \text{ km s}^{-1}$  (Silverman et al. 2015).

Despite some early-time similarities, SNe Iax spectra diverge significantly from those of normal SNe Ia at late times (traditionally defined as  $\gtrsim 100$  days past maximum light; Bowers et al. 1997; Branch et al. 2008; Silverman et al. 2013; Friesen et al. 2014; Black et al. 2016). Normal SNe Ia are nebular at late times, dominated by forbidden emission lines of Fe, Co, and Ni. SNe Iax have not been observed to go fully nebular; although they do exhibit some forbidden lines, their late-time spectra continue to be dominated by permitted features of Fe, Ca, and Na (Jha et al. 2006; Sahu et al. 2008; Foley et al. 2010, 2016). Not only is this distinct from normal SNe Ia, it is unlike any other class of SNe.

It is difficult to observe a SN during the full intermediate phase between early and late times ( $\sim +80$  to  $+200$  days) since a SN discovered in the night sky will generally have moved to a region of unobservable daytime sky during this time. Because of sparse observations, the precise epochs when SNe Iax spectra transition away from those of normal SNe Ia is not well studied. In this paper we analyze SN Iax 2014dt, which was fortuitously found in a very easterly location in the sky and has extensive spectral coverage throughout the intermediate phase.

SN 2014dt was discovered at 13.6 mag on 2014 October 29.8 (UT dates are used throughout this paper) in the nearby galaxy M61 ( $d \approx 14.6$  Mpc; see Section 2; Nakano et al. 2014). It was classified as a SN Iax by Ochner et al. (2014) using the Asiago Copernico Telescope on 2014 October 31. Although SN 2014dt was discovered past peak brightness, several attempts have been made to estimate its date of maximum. In this paper, we adopt the  $B_{\text{max}}$  date from Kawabata et al. (2018), 2014 October 20.4 (JD =  $2,456,950.9 \pm 4.0$  days), which is consistent with other measurements (Foley et al. 2016; Fox et al. 2016; Singh et al. 2018).

SN 2014dt is a moderately luminous SN Iax ( $M_B$  between approximately  $-17$  and  $-18$  at maximum light, depending on the assumed distance), though still dimmer than normal SNe Ia with similar decline rates (Kawabata et al. 2018; Singh et al. 2018). Its spectra resemble those of other SNe Iax at both early and late times (Kawabata et al. 2018; Singh et al. 2018), placing it solidly within the SNe Iax class.

Deep preexplosion images from the Hubble Space Telescope (HST) at the location of SN 2014dt did not detect a progenitor system, which excludes a high-mass main-sequence star ( $M_{\text{init}} \gtrsim 16 M_{\odot}$ ) or a high-mass evolved star ( $M_{\text{init}} \gtrsim 8 M_{\odot}$ ) as the companion (Foley et al. 2015). We note that another SN Iax, SN 2012Z, is the only thermonuclear SN with a direct observation of a progenitor system in a preexplosion image, which was interpreted as a luminous, blue, helium-star companion (McCully et al. 2014). This direct detection implies that SNe Iax can arise from the single-degenerate progenitor scenario.

Pure deflagration models—explosions propagated with only a subsonic flame—can explain the lower luminosities of SNe Iax as well as their unique late-time spectra (Branch et al. 2004; Phillips et al. 2007; Jha 2017). In particular, a weak deflagration model in which a Chandrasekhar-mass carbon–

oxygen white dwarf (WD) explodes but does not completely unbind the star has been especially promising (Jordan & Perets 2012; Kromer et al. 2013; Fink et al. 2014). In this model, a gravitationally bound remnant—a most peculiar clump of surviving matter—is left over, in contrast to normal SNe Ia that are thought to completely disrupt their progenitor WD. Fox et al. (2016) discovered a mid-infrared (MIR) excess in SN 2014dt at late times, which they interpreted as dust emission. Foley et al. (2016) find a lack of evidence for dust and suggest that the MIR excess may instead come from a bound remnant driving an optically thick wind.

Other possible explosion models for SNe Iax include the core-collapse (CC) model (Valenti et al. 2009), the fallback CC model (Moriya et al. 2010), and the merger model of a carbon–oxygen WD and an oxygen–neon WD (Kashyap et al. 2018; Karambelkar et al. 2021).

We model the optical spectra of SN 2014dt, from  $+11$  to  $+562$  days after  $B_{\text{max}}$ , using a nearly unaltered deflagration explosion model from Fink et al. (2014). To our knowledge, this is the latest-epoch spectrum modeled for any type of SN using only permitted lines. We describe our observations in Section 2, and explore the spectral divergence of SN 2014dt compared to normal SNe Ia in Section 3. In Section 4 we describe our spectral modeling using the radiative-transfer code TARDIS (Kerzendorf & Sim 2014; Kerzendorf et al. 2023), including an analysis of different explosion models and the evolution of the photospheric velocity. We summarize our results and possible interpretations in Section 5.

## 2. Spectral Observations and Reduction

We present 21 new spectra of SN 2014dt, ranging from  $+32$  to  $+456$  days past  $B_{\text{max}}$ .

Eleven new spectra of SN 2014dt were taken using the Robert Stobie Spectrograph (RSS; Smith et al. 2006) on the 10 m Southern African Large Telescope (SALT). We used a  $1''.5$  slit (behind an atmospheric dispersion compensator (ADC)) and the PG0900 grating, typically tilted to four positions with appropriate order-blocking filters, to cover the entire optical wavelength range without detector chip gaps. The SALT/RSS data were reduced using a custom pipeline (RUSALT) based on PySALT (Crawford et al. 2010), including wavelength calibration, 1D extraction, relative flux calibration with spectrophotometric standard stars, combination of the grating tilt positions, telluric absorption removal, and heliocentric correction.

Four new spectra of SN 2014dt were taken with the DEEP Imaging Multi-Object Spectrograph (DEIMOS; Faber et al. 2003) on the 10 m Keck II telescope at the W.M. Keck Observatory. The first of these spectra was taken on 2014 November 21 using a  $0''.8$  slit and the 1200G grating (+GG455 blocking filter) centered at  $6170 \text{ \AA}$ . These data were reduced with the open source software PyeIt (Prochaska et al. 2020a, 2020b). Two DEIMOS spectra (2014 December 20 and 2015 March 18) were taken with the 600ZD grating (+GG455 blocking filter) centered at  $7200 \text{ \AA}$ , with a  $1''.2$  and  $1''.0$  slit, respectively. One additional DEIMOS spectrum was taken on 2015 December 16 with a  $1''.0$  slit and the 600ZD grating (+GG455 blocking filter) centered at  $7000 \text{ \AA}$ . These last three spectra were processed with the DEEP2 DEIMOS data reduction pipeline (Cooper et al. 2012; Newman et al. 2013).

Four new spectra of SN 2014dt were taken with FLOYDS, a pair of nearly identical, low-resolution spectrographs installed on the 2 m Faulkes Telescope North (FTN) at Haleakalā Observatory and Faulkes Telescope South (FTS) at Siding Spring Observatory; both telescopes are part of the Las Cumbres Observatory Global Telescope Network. The three observations taken with FTN (2015 April 8, 2015 April 19, and 2015 May 20) used a 1''/6 slit and one observation with FTS (2015 April 24) used a 2''/0 slit. All four spectra were observed at the parallactic angle (Filippenko 1982), and reduced with the FLOYDS pipeline<sup>23</sup> (Valenti et al. 2014).

One new spectrum taken with the Gemini Multi-Object Spectrograph (GMOS; Hook et al. 2004) on the 8.1 m Gemini-North telescope is a combination of eight exposures from two consecutive nights. All exposures were taken at the parallactic angle, and used the R400 grating with the G5305 filter; half at a central wavelength of 6900 Å and half at 7100 Å. After standard data reduction with IRAF, we applied our own IDL routines to flux calibrate the data and remove telluric lines using the well-exposed continua of spectrophotometric standard stars (Wade & Horne 1988; Foley et al. 2003). Details of the reduction can be found in Silverman et al. (2012a).

One new spectrum was taken with the Goodman High Throughput Spectrograph (GHTS; Clemens et al. 2004) on the 4 m Southern Astrophysical Research Telescope (SOAR) on 2015 April 10. A 1''/07 slit was used with the SYZY 400 grating, and the spectrum was reduced with the same techniques as other GHTS spectra in Foley et al. (2016), following Silverman et al. (2012b).

All new spectra have minimal differential refraction losses due to being observed at the parallactic angle, using an ADC, and/or being taken at low airmass ( $<1.45$ ).

We also collect published optical spectra of SN 2014dt, ranging from +11 to +562 days past  $B_{\max}$ , from Ochner et al. (2014), Foley et al. (2015, 2016), Lyman et al. (2017),<sup>24</sup> Kawabata et al. (2018), Singh et al. (2018),<sup>25</sup> and Stahl et al. (2020). These spectra were obtained, in part, through the Open Supernova Catalog (Guillochon et al. 2017), WISEREP<sup>26</sup> (Yaron & Gal-Yam 2012), and the UC Berkeley Filippenko Group's Supernova Database<sup>27</sup> (Silverman et al. 2012a).

After reduction or collection, all spectra went through the following series of steps: (1) manual cleaning, (2) scaling/mangling, (3) Galactic extinction correction, and (4) host-galaxy redshift correction. First, we manually identified and removed any remaining cosmic rays or telluric features, and masked narrow interstellar emission lines from the host galaxy. Next, for absolute flux calibration, we “mangled” the spectra to match the broadband optical *BVRI* photometry of SN 2014dt from Singh et al. (2018), Kawabata et al. (2018), and J. Johansson et al. (2023, in preparation). The photometry is fit via a Gaussian process (GP) model that produces smooth, continuous *BVRI* light curves with uncertainties that account for data uncertainties and scatter. In the mangling procedure,

we multiply the spectra by a smooth function in wavelength so that the passband-integrated flux matches the value of the corresponding GP light curve at that time. Some of the spectra had enough wavelength coverage to match multiple optical filters; for these spectra we verified that the mangling did not significantly change the relative line strengths—generally, the relative flux adjustments were no more than  $\sim 20\%$ . For a few spectra where the mangling procedure unacceptably altered the spectrum (most likely owing to uncertainties in the photometry or background contamination in the spectrum), we only apply a uniform rescaling to one filter. Spectra that only had enough wavelength coverage to match one optical filter were rescaled to that single photometric band. The filters used to scale/mangle each spectrum are listed in Table 1.

For all spectra, we next correct for Milky Way extinction of  $E(B - V) = 0.02$  mag with  $R_V = 3.1$  (Schlafly & Finkbeiner 2011). We did not further correct for host-galaxy extinction, thought to be negligible for SN 2014dt (Foley et al. 2015; Singh et al. 2018). Lastly, we transform the spectra to the host-galaxy rest frame, correcting for the redshift of M61,  $z = 0.0052$  (Allison et al. 2014 via the NASA/IPAC Extragalactic Database (2019), hereafter NED).

In total, we present 69 new and previously published SN 2014dt spectra, making SN 2014dt one of the best spectroscopically observed SNe. Details of the observations, including phase, date, telescope+instrument, exposure time, spectral resolution, and which filter(s) were used for scaling/mangling can be found in Table 1. The full spectral evolution of SN 2014dt—after cleaning, scaling/mangling, Galactic extinction correction, and host-galaxy redshift correction—is shown in Figure 1 and is available as data behind the figure. New spectra, denoted with  $\star$  in Figure 1, will be available on WISEREP.<sup>28</sup>

Our spectral modeling in Section 4 requires the additional step of converting our spectra from flux to luminosity units by using the distance to M61. Distances to M61 from NED range from 7.6 Mpc (Bottinelli et al. 1984) to 35.5 Mpc (Sparks 1994) with a mean distance of 14.6 Mpc. Recent papers on SN 2014dt choose varying distances around this value. In particular, Foley et al. (2015, 2016) adopt 12.3 Mpc, Fox et al. (2016) use 19.3 Mpc, Singh et al. (2018) adopt 21.4 Mpc, and Kawabata et al. (2018) use 14.5 Mpc. This range of distances would produce a difference of  $\sim 1$  mag for the absolute maximum magnitude in  $B$ ,  $M_{\max}^B$ , and correspondingly change the luminosity we obtain from our spectra. In this paper we adopt a distance of 14.6 Mpc, but we also explore the effects of a range of distances from 12.3 to 21.4 Mpc. As we describe in Section 4.2, distances at the lower end of this range have a minimal impact on our spectral modeling results, while higher assumed distances produce larger changes in our modeling that are less compatible with the observations. Our analysis favors a distance to SN 2014dt less than 18 Mpc.

### 3. Spectral Divergence Phase

As discussed in Section 1, SNe Iax are spectroscopically similar to normal SNe Ia at early times, but by late times their spectra are much different (Jha et al. 2006; Sahu et al. 2008; Foley et al. 2010, 2016). The extensive spectral coverage of SN 2014dt between +80 and +200 days allows us to pinpoint the phase when it transitions away from normal SNe Ia spectra, which we call the spectral divergence phase of SNe Iax.

<sup>23</sup> <https://github.com/LCOGT/floydspipeline>

<sup>24</sup> J. Lyman (private communication) provided a host-galaxy-subtracted VLT/MUSE spectrum (2016 January 20) using a background annulus around the source. The latest Keck spectrum (2016 May 7) also had host contamination that we modeled and subtracted to isolate the SN flux, better matching other late-time spectra.

<sup>25</sup> For two HCT spectra, 2015 February 1 and 2015 February 5, we manually rescaled the red and blue channels to eliminate a flux discontinuity and to match the other spectra better at similar epochs.

<sup>26</sup> <https://wiserep.weizmann.ac.il>

<sup>27</sup> <http://heracles.astro.berkeley.edu/sndb>

<sup>28</sup> <https://wiserep.weizmann.ac.il>



**Table 1**  
Spectral Observations of SN 2014dt

Phase <sup>a</sup> (days)	UT Date	MJD	Telescope /Instrument	Exposure (s)	Resolution ( $R = \lambda/\Delta\lambda$ )	Scaled/Mangled To Filter(s)	Reference <sup>b</sup>
+11	2014-10-31	56961.2	Copernico/AFOSC	300	450	<i>V</i>	3
+21	2014-11-10	56971.84	Kanata/HOWPol	900–1200	400	<i>V, R, I</i>	5
+22	2014-11-11	56972.98	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+28	2014-11-18	56979.0	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+29	2014-11-18	56979.56	Lick/Kast	600	580–950	<i>V</i>	2
+32	2014-11-21	56982.66	Keck/DEIMOS	360	1400	<i>V</i>	This paper
+34	2014-11-23	56984.97	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+36	2014-11-25	56986.56	Lick/Kast	600	580–950	<i>V</i>	4
+40	2014-11-29	56990.88	Kanata/HOWPol	900–1200	400	<i>R, I</i>	5
+41	2014-12-1	56992.0	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+45	2014-12-4	56995.93	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+48	2014-12-8	56999.0	HCT/HFOSC	600–1800	1330–2190	<i>I</i>	6
+51	2014-12-10	57001.94	HCT/HFOSC	600–1800	1330	<i>V</i>	6
+60	2014-12-19	57010.91	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+61	2014-12-20	57011.66	Keck/DEIMOS	1350	1400	<i>R</i>	This paper
+64	2014-12-23	57014.5	Lick/Kast	600	580–950	<i>B, V, R, I</i>	4
+64	2014-12-23	57014.86	Kanata/HOWPol	900–1200	400	<i>V, R, I</i>	5
+66	2014-12-26	57017.01	HCT/HFOSC	600–1800	1330	<i>B</i>	6
+76	2015-1-5	57027.03	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+78	2015-1-6	57028.84	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+82	2015-1-10	57032.84	Kanata/HOWPol	900–1200	400	<i>R</i>	5
+90	2015-1-18	57040.93	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+92	2015-1-20	57042.51	Lick/Kast	1800	570–980	<i>V, R, I</i>	4
+95	2015-1-24	57046.08	SALT/RSS	2200	1000	<i>V</i>	This paper
+97	2015-1-25	57047.81	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+100	2015-1-28	57050.46	Lick/Kast	1200	580–950	<i>R</i>	4
+104	2015-2-1	57054.77	Kanata/HOWPol	900–1200	400	<i>R</i>	5
+104	2015-2-1	57054.86	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+108	2015-2-5	57058.93	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+117	2015-2-14	57067.53	Lick/Kast	1800	580–950	<i>V</i>	4
+119	2015-2-17	57070.03	SALT/RSS	1550	1000	<i>R</i>	This paper
+124	2015-2-21	57074.54	Lick/Kast	1800	580–950	<i>B, V, R, I</i>	4
+124	2015-2-22	57075.55	Subaru/FOCAS	1200	650	<i>R</i>	5
+126	2015-2-23	57076.86	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+133	2015-3-2	57083.67	Kanata/HOWPol	900–1200	400	<i>R</i>	5
+133	2015-3-3	57084.01	SALT/RSS	1800	1000	<i>V</i>	This paper
+138	2015-3-8	57089.03	SALT/RSS	1650	1000	<i>V</i>	This paper
+140	2015-3-10	57091.0	SALT/RSS	1600	1000	<i>V</i>	This paper
+143	2015-3-12	57093.79	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+148	2015-3-18	57099.42	Lick/Kast	1800	580–950	<i>V</i>	4
+148	2015-3-18	57099.44	Keck/DEIMOS	1800	1400	<i>R</i>	This paper
+151	2015-3-21	57102.55	Kanata/HOWPol	900–1200	400	<i>R</i>	5
+156	2015-3-25	57106.73	HCT/HFOSC	600–1800	1330–2190	<i>V</i>	6
+156	2015-3-26	57107.47	Lick/Kast	1800	570–980	<i>V</i>	4
+159	2015-3-28	57109.9	SALT/RSS	1810	1000	<i>V</i>	This paper
+160	2015-3-29	57110.92	SALT/RSS	1850	1000	<i>V</i>	This paper
+169	2015-4-8	57120.48	FTN/FLOYDS	3600	300–600	<i>B, V, R, I</i>	This paper
+171	2015-4-10	57122.14	SOAR/GTHS	1200	800	<i>V</i>	This paper
+175	2015-4-13	57125.93	SALT/RSS	1800	1000	<i>B, V</i>	This paper
+177	2015-4-16	57128.41	Lick/Kast	1800	610–910	<i>R</i>	1, 4
+180	2015-4-18	57130.9	SALT/RSS	1600	1000	<i>V, R, I</i>	This paper
+180	2015-4-19	57131.49	FTN/FLOYDS	3600	300–600	<i>V</i>	This paper
+184	2015-4-22	57134.89	SALT/RSS	1372	1000	<i>V, R, I</i>	This paper
+185	2015-4-24	57136.6	FTS/FLOYDS	3600	300–600	<i>V</i>	This paper
+187	2015-4-26	57138.85	SALT/RSS	1525	1000	<i>V</i>	This paper
+208	2015-5-17	57159.83	SALT/RSS	1700	1000	<i>B, V, R, I</i>	1
+211	2015-5-20	57162.39	FTN/FLOYDS	3600	300–600	<i>V</i>	This paper
+217	2015-5-26	57168.29	Lick/Kast	1800	570–980	<i>B, V, R, I</i>	1, 4
+233	2015-6-11	57184.72	SALT/RSS	1700	1000	<i>B, V, R, I</i>	1
+237	2015-6-16	57189.07	SOAR/GTHS	3600	800	<i>V</i>	1

**Table 1**  
(Continued)

Phase <sup>a</sup> (days)	UT Date	MJD	Telescope /Instrument	Exposure (s)	Resolution ( $R = \lambda/\Delta\lambda$ )	Scaled/Mangled To Filter(s)	Reference <sup>b</sup>
+238	2015-6-16	57189.33	Keck/LRIS	600	1000–1700	<i>B, V, R, I</i>	1, 4
+243	2015-6-21	57194.26	Subaru/FOCAS	1800	650	<i>V, R, I</i>	5
+275	2015-7-24	57227.04	SOAR/GTHS	3600	800	<i>V</i>	1
+410	2015-12-6	57362.59	Subaru/FOCAS	1800	650	<i>R</i>	5
+415	2015-12-11	57367.64	Keck/LRIS	2400	1000–1700	<i>V</i>	1
+420	2015-12-16	57372.59	Keck/DEIMOS	1200	1400	<i>R, I</i>	This paper
+454	2016-1-20	57407.0	VLT/MUSE	2080	2000–4000	<i>R, I</i>	7
+456	2016-1-22	57409.0	Gemini/GMOS-N	7200	640	<i>R</i>	This paper
+562	2016-5-7	57515.0	Keck/LRIS	1200	1100–1700	<i>R</i>	4

**Notes.**<sup>a</sup> Days since *B* maximum, UT 2014 October 20.4 (MJD 56950.4), corrected for time dilation.<sup>b</sup> 1: Foley et al. (2016), 2: Foley et al. (2015), 3: Ochner et al. (2014), 4: Stahl et al. (2020) 5: Kawabata et al. (2018), 6: Singh et al. (2018), and 7: Lyman et al. (2017).

(This table is available in machine-readable form.)

In Figure 2, the spectral evolution of SN Iax 2014dt (black/gray) is compared with the spectral evolution of other normal SNe Ia (colors) from  $\sim +30$  days to  $\sim +200$  days. In order to account for the lower velocity of SN 2014dt, we convolve its spectra with a Gaussian with a FWHM intensity of  $1200 \text{ km s}^{-1}$  and blueshift by  $4000 \text{ km s}^{-1}$  before comparing it to the higher velocity, normal SNe Ia. The early epoch (+29 days) spectrum of SN 2014dt shows remarkable similarity with normal SNe Ia spectra for all features at optical wavelengths. By  $\sim +90$  days, the prominent double-peaked Fe II feature at  $\sim 5000 \text{ \AA}$  starts to disappear in the normal SNe Ia but remains present in SN 2014dt until very late epochs (disappearing sometime between +275 and +410 days; see Figure 1). The near-IR Ca II triplet  $\lambda\lambda 8498, 8542, 8662$  (hereafter, the Ca II IR triplet) in normal SNe Ia is also nearly gone by  $\sim +90$  days and completely gone by  $\sim +150$  days, but continues to be a strong feature in SN 2014dt. The broad Na I D  $\lambda 5890 \text{ \AA}$  absorption is strong in SN 2014dt at all epochs, but starts to disappear in the normal SNe Ia by  $\sim +150$  days. In general, by +150 days the spectra of normal SNe Ia lose most of their photospheric (permitted) features and are dominated by broad, nebular (forbidden) emissions of [Fe II] and [Fe III], while the spectra of SN 2014dt continue to be dominated by photospheric features and remain largely unchanged after +90 days. For these reasons, we conclude that the spectral divergence phase of SN 2014dt lasts  $\sim 60$  days between +90 days and +150 days past  $B_{\text{max}}$ .

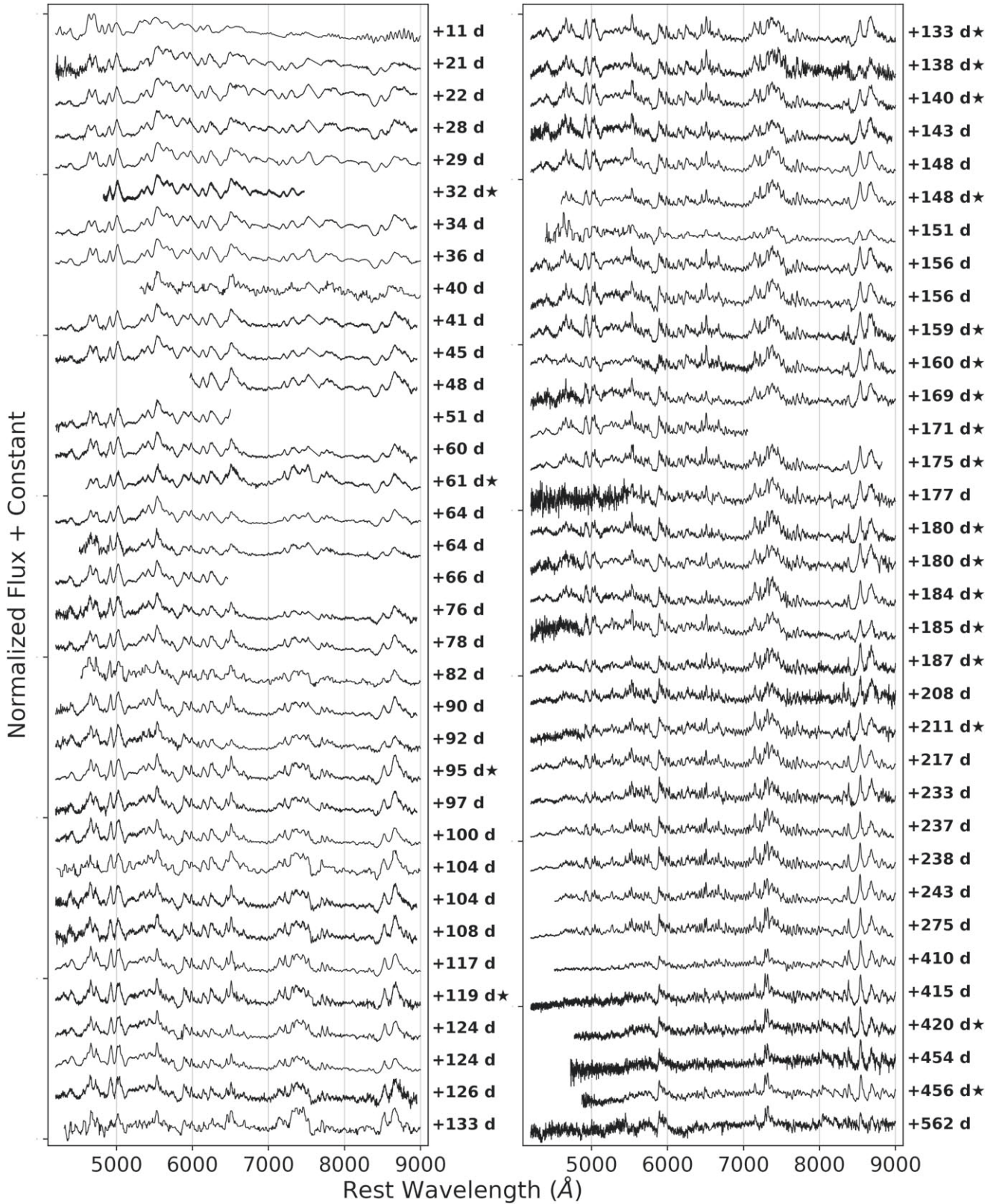
**4. Spectral Models and Analysis**

Attempts at modeling the early-time ( $< +41$  days) spectra of SN 2014dt have been carried out by Singh et al. (2018) and Kawabata et al. (2018). These early-time models were used primarily for line identification and show prominent, often blended, lines of IMEs and IGEs; most notably Fe II, Co II, and Ca II. We present spectral modeling that spans from +11 days, 10 days earlier than the earliest spectra previously modeled, to +562 days, 521 days later than the latest spectra previously modeled. Our early-time models generally agree with the line identifications found in these previous studies and offer additional insights into the evolving contributions from different elements and ions.

We use TARDIS (Kerzendorf & Sim 2014; Kerzendorf et al. 2023), a 1D open source Monte Carlo (MC) radiative-

transfer code to model the spectra of SN 2014dt. It generates model spectra based on simple assumptions: spherical symmetry, homologous expansion ( $r = vt$ ), and a sharp photosphere that emits a blackbody spectrum (i.e., it cannot generate forbidden line emission). Indivisible MC packets representing bundles of photons with the same frequency are followed as they go from the photosphere through the ejecta, until they escape to create the model spectrum. TARDIS has been used to model many different kinds of SNe, including other SNe Iax: SN 2002cx (Barna et al. 2018), SN 2005hk (Magee et al. 2017, 2019; Barna et al. 2018), SN 2007J (Magee et al. 2019), SN 2010ae (Magee et al. 2019), SN 2011ay (Barna et al. 2017, 2018), SN 2012Z (Barna et al. 2018; Magee et al. 2019, 2022), PS1-12bwh (Magee et al. 2017), SN 2015H (Magee et al. 2016, 2019; Barna et al. 2018), SN 2019muj (Barna et al. 2021), SN 2019gsc (Srivastav et al. 2020), and SN 2020sck (Dutta et al. 2022).

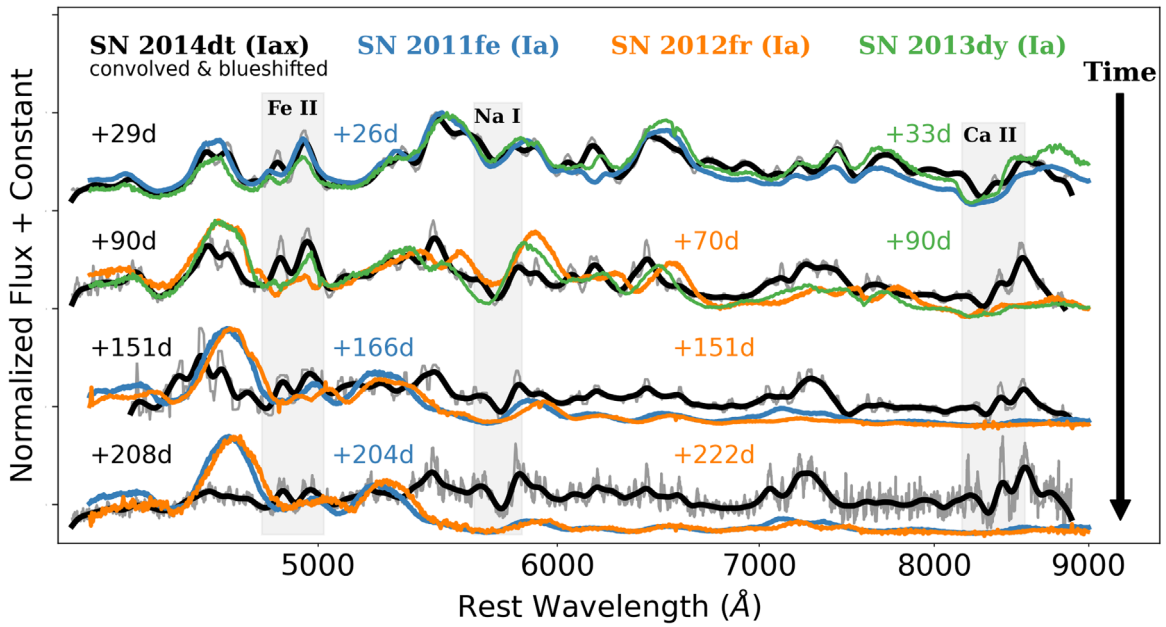
TARDIS is typically only able to model relatively early-time spectra (less than a few weeks after peak magnitude) because the assumption of a sharp photosphere starts to break down at this time for most SNe. However, since SN Iax spectra never become dominated by forbidden line emission, we are able to model the optical spectra of SN 2014dt out to extremely late times. We construct TARDIS models for all SN 2014dt spectra from +11 days to +562 days past  $B_{\text{max}}$  with generally good to excellent fits for all epochs. In Figure 3, a few early, mid, and late epochs are chosen to demonstrate how the model spectra (generated using the formal integral method from TARDIS to reduce the MC noise) and observed spectra of SN 2014dt compare. The model spectra correctly recreate the continua, spectral lines, and relative strengths of most of the photospheric features in the observed spectra. When deciding on the best-fit model, we tried to match the full wavelength range of each spectrum while emphasizing a good match to the blue end of the spectrum in the range  $\sim 3500\text{--}5500 \text{ \AA}$ , a region with many permitted lines of IGEs. In the latest two epochs in Figure 3, +217 and +456 days past  $B_{\text{max}}$ , the gray region includes nebular emission at  $\sim 7200 \text{ \AA}$ ; these nebular regions were excluded during our fitting since TARDIS is not able to model these forbidden Fe, Ca, and Ni lines (Foley et al. 2016).



**Figure 1.** Optical spectral evolution of SN 2014dt. Stars (\*) denote new spectra presented in this paper. Phases are calculated relative to the time of  $B_{\max} = 2014$  October 20.4 (Kawabata et al. 2018) and corrected for time dilation to the host-galaxy rest frame. The spectra are available as data behind the figure.

(The data used to create this figure are available.)





**Figure 2.** The spectrum of Type Iax SN 2014dt (black/gray) diverges from the spectrum of normal Type Ia SNe (blue, orange, green) over time. The black SN 2014dt spectrum has been convolved with a velocity of  $1200 \text{ km s}^{-1}$  and blueshifted by  $4000 \text{ km s}^{-1}$  in order to compare its spectral features with the higher velocities of normal SNe Ia. The unconvolved, but still blueshifted, SN 2014dt spectrum is shown in gray. All spectra of the normal SNe Ia were collected from The Open Supernova Catalog (Guillochon et al. 2017) and came from the following sources: SN 2011fe (Silverman et al. 2012a; Yaron & Gal-Yam 2012; Mazzali et al. 2014, 2015; Stahl et al. 2020), SN 2012fr (Yaron & Gal-Yam 2012; Zhang et al. 2014; Childress et al. 2015), and SN 2013dy (Silverman et al. 2012a; Stahl et al. 2020).

The input parameters for each model spectrum are discussed subsequently in Section 4.1. A detailed analysis of the resulting synthetic spectra follows in Section 4.2, including a closer look at the individual contributions from different ions and elements.

#### 4.1. Synthetic Spectra: Input Parameters

TARDIS needs a number of input parameters in order to generate each synthetic spectrum. Some inputs are used to set up the computational framework (e.g., number of iterations, number of MC packets, etc.) for each simulation. Of particular importance, the modes used to calculate the state and behavior of plasma for all of our simulations are

1. ionization: nebular
2. excitation: dilute-lte
3. radiative\_rates\_type: dilute-blackbody
4. line\_interaction\_type: macroatom

Other input parameters pertaining to our specific SN include the requested luminosity, time since explosion, chemical abundance profile, density profile, and inner ( $v_{\text{phot}}$ ), and outer ( $v_{\text{outer}}$ ) boundaries of the line-forming region; details can be found in Table 3. All TARDIS input files and output synthetic spectra are available on GitHub<sup>29</sup> and 10.5281/zenodo.7838690.

The luminosity,  $L_{\text{requested}}$ , comes directly from integrating each observed spectrum. The  $L$  integration range,  $\lambda_{\text{min}}-\lambda_{\text{max}}$ , is based on the observation limits, sometimes restricted to exclude noisy regions at the ends of a spectrum.

We set the time since explosion to be equal to rest-frame days past  $B_{\text{max}}$ , plus 11 days to account for the rise time from explosion to peak. Unfortunately, SN 2014dt was found after peak brightness and does not have data for a precise rise-time

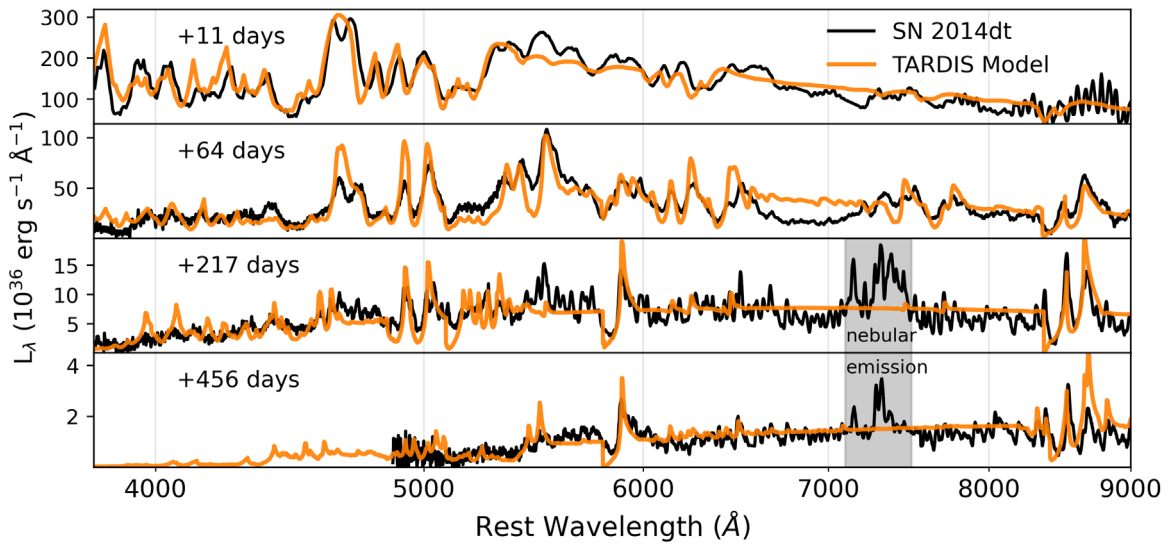
measurement. Kawabata et al. (2018) estimate a wide range of possible rise times, 11.8–21.0 days, for the bolometric light curve of SN 2014dt (which may differ slightly from the rise time to  $B_{\text{max}}$  that we desire). Our choice of an 11 day rise time is at the lower end, but roughly consistent with their results. It is also within the  $\sim 10$ –20 day rise time typical for SNe Iax (Jha 2017), and it is in the range of rise times found in the relevant Fink et al. (2014) deflagration models that we use. For completeness, we explore a range of rise times up to 20 days, and find that our spectral-fitting results are not very sensitive to the choice of rise time, particularly for epochs after about +40 days. Our choice of rise time gives the best fits to the earlier epochs, but different choices are not definitively worse.

Remarkably, we are able to use a self-consistent description of the ejecta (fixed chemical abundance and density profiles as a function of velocity) to generate the model spectra at the wide range of epochs covered by the data. As mentioned in Section 1, deflagration models are generally favored for SNe Iax. We model the spectra of SN 2014dt using numerous deflagration models from Fink et al. (2014). These models are 3D deflagration explosions of a Chandrasekhar-mass carbon-oxygen WD. In order to be used by the 1D TARDIS code, we obtain the spherically averaged density and isotopic abundance profiles from the Heidelberg Supernova Model Archive<sup>30</sup> (HESMA; Kromer et al. 2017).

We make three small adjustments to the Fink et al. (2014) model isotope abundances. First, we remove all elements past Zn, the atomic data limit of TARDIS. The total abundance (mass fraction) of removed elements is small,  $\sim 10^{-6}$ , and is not expected to have a noticeable effect on the model spectrum. Second, we use a uniform, averaged isotope abundance value for all shells. We find that this has little to no impact on the resulting spectrum since the model ejecta are already well mixed (i.e., the abundances are

<sup>29</sup> <https://github.com/ycamacho/SN-2014dt-TARDIS>

<sup>30</sup> <https://hesma.h-its.org/doku.php>



**Figure 3.** Model spectra (orange) from TARDIS for SN 2014dt (black) at +24, +59, +212, and +415 days past  $B_{\max}$ . Most features and line widths agree with the observations. Nebular emission at later epochs cannot be modeled by TARDIS (gray region).

nearly constant as a function of velocity; see Section 4.3 for details of modeling with velocity-dependent abundances). Third, we increase the abundance of  $^{23}\text{Na}$  from  $\sim 0.01\%$  to  $\sim 1.7\%$  (discussed further in Section 4.4).

The only modification made to the Fink et al. (2014) model density profiles was to extend them below the original tabulation limit of  $\sim 200 \text{ km s}^{-1}$  via a cubic spline extrapolation down to zero velocity. This extrapolated velocity range is only relevant for modeling the latest epochs.

The line-forming region is specified by the inner velocity boundary,  $v_{\text{phot}}$ , and the outer velocity boundary,  $v_{\text{outer}}$ . In general, the synthetic spectrum is much more sensitive to the choice of  $v_{\text{phot}}$  than to the choice of  $v_{\text{outer}}$ . Our choice of  $v_{\text{outer}}$  is based on fitting the Ca II IR triplet, the main feature affected by  $v_{\text{outer}}$ , and is discussed further in Section 4.4. This leaves  $v_{\text{phot}}$  as our primary fitting parameter for each spectral epoch.

#### 4.2. Synthetic Spectra: Results and Analysis

We find that using the n1def deflagration model density and abundance from Fink et al. (2014) produces the best-fit synthetic spectra for SN 2014dt. The results for a wide range of epochs are presented in Figure 3. The level of agreement between the observed spectra of SN 2014dt and the synthetic spectra generated by a nearly unaltered deflagration model (as described in Section 4.1) is quite remarkable. The continua, spectral lines, and relative strengths of most of the photospheric features in the observed spectra are well reproduced by our model for all epochs, covering over 500 days of evolution.

Figure 4 shows the observed spectra in gray compared to the synthetic spectra, which are broken down and color coded by contributions from different elements and ions. Contributions come from each photon packet’s last interaction, with emission shown as a positive luminosity and absorption as a negative luminosity. Some wavelength regions in each spectrum are dominated by the underlying blackbody photosphere (labeled “No Interaction”), while other regions have significant contributions from one or many different elements and ions. In particular, the blue end ( $< 5500 \text{ Å}$ ) of the spectra up to +217 days exhibits complex line contributions. Using both Figures 3

and 4, we now discuss the details of how well the synthetic spectra reproduce the observed spectra of SN 2014dt.

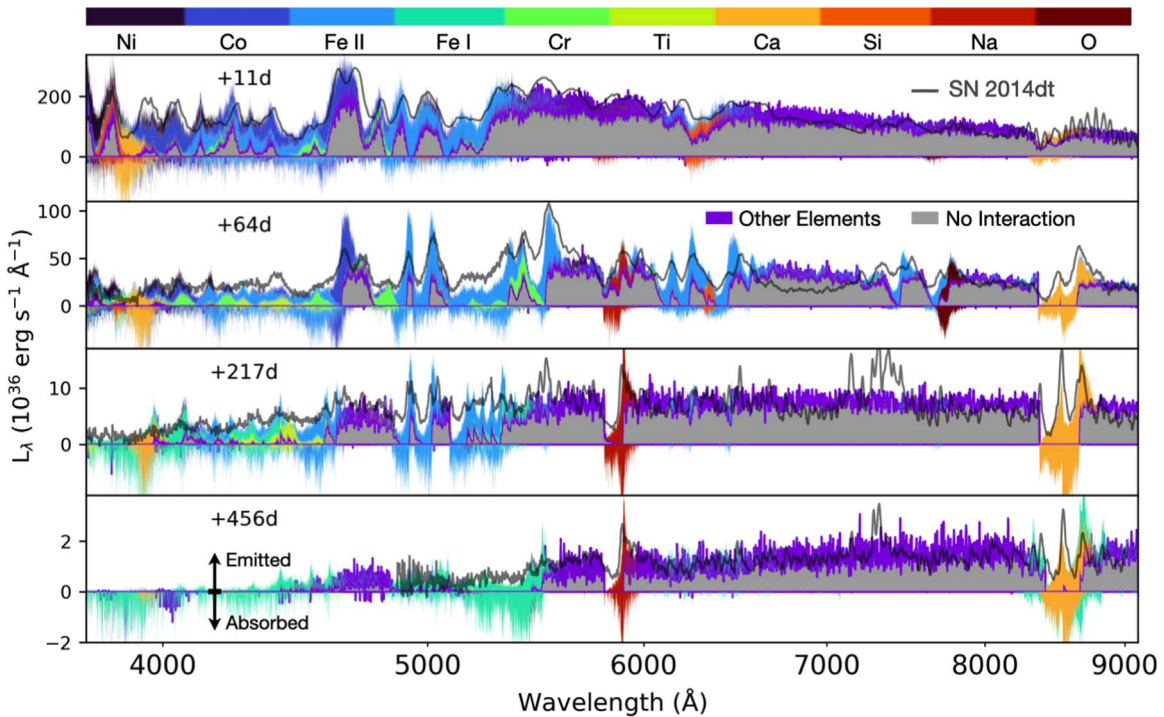
In the earliest epoch, +11 days, the region between  $\sim 3900$  and  $\sim 5500 \text{ Å}$  is particularly well matched with almost all the spectral features present. This is a blended region of IGEs, dominated by permitted lines of Co and Fe with some contributions from Ni and Cr. The redder end ( $> 5500 \text{ Å}$ ) of the +11 day spectrum is also mostly reproduced by our model, with contributions from Ca II, O I, and Si II.

The +64 day synthetic spectrum reproduces most of the photospheric features, and matches the Ca II IR triplet, as well as the Fe II features at  $\sim 5000 \text{ Å}$  and  $\sim 6200 \text{ Å}$  particularly well. By +217 days, while the relative strengths of the photospheric lines have decreased as the nebular emission at  $\sim 7200 \text{ Å}$  increases, many photospheric lines continue to be reproduced well by our model, especially Fe II at  $\sim 5000 \text{ Å}$  and the Ca II IR triplet. More than 200 days later, the +456 day spectrum continues to match the now Fe I features at  $\sim 5500 \text{ Å}$  and the Ca II IR triplet.

Although the model allows us to explore the contribution of each ionization state for each element, for most elements, one ionization state dominates all epochs, e.g., O I, Na I, Si II, or Ca II. For IGEs, multiple ionization states are relevant. We group together these contributions for Cr, Co, and Ni for clarity in Figure 4, but show separately the contributions of Fe II and Fe I to illustrate how the ionization evolves with time. In the earliest epochs, +11 and +64 days, Fe II dominates the hot ejecta. As the ejecta cool and expand with time, we see the contribution from Fe I increase and by +217 days it makes up most of the Fe at  $< 4500 \text{ Å}$  while Fe II still dominates at  $\sim 5000 \text{ Å}$ . By +456 days, all the Fe lines in the optical regime come from Fe I.

For the +11 day epoch, while the relative line strengths and location of prominent features are well matched, there are some features that appear slightly too blueshifted when compared with the observed spectrum. In particular, the model line profiles around Si II  $\lambda 6355 \text{ Å}$  shows a blueshift  $\sim 1000 \text{ km s}^{-1}$  more than expected from observations. This might be due to too much Si in the faster, outer layers of the ejecta of our model or some other asymmetry in the ejecta structure or explosion propagation. We have assumed uniform abundance profiles for all elements, supported by Barna et al. (2017) and Magee et al.





**Figure 4.** Element decomposition of the model spectra at +11, +64, +217, and +456 days past  $B_{\max}$ . SN 2014dt (dark-gray line) is shown for comparison. Elements and ions with the greatest contributions are shown on the color map. Contributions from all other elements are in purple and packets that did not interact with the ejecta (blackbody emission from the assumed photosphere) are in light gray. Emission is shown as positive (above zero luminosity) and absorption is negative.

(2022), but their work focuses on the outer layers around maximum light, and it is nonetheless possible the ejecta of SNe Iax require some degree of stratification.

One notable missing feature in the +11 to +64 day spectra is the absorption at  $\sim 4670$  Å. This exact absorption feature was also not recovered in the very similar +10.4 day SN Iax 2015H spectrum that was modeled by Barna et al. (2018). Both Kawabata et al. (2018) and Singh et al. (2018) are able to reproduce the appearance, but not the correct strength, of this absorption feature in their model SN 2014dt spectra, but the former credits the feature to Fe II and the latter thinks it might be [Fe III]. Although we are unable to produce this feature with the n1def model, we were able to recreate it with the n20def model for only the earliest epoch, +11 days. This model showed that the absorption at  $\sim 4670$  Å is very blended with contributions from many IGEs and IMEs—the most notable difference being more C II interaction and less contribution from the photospheric blackbody. However, besides the  $\sim 4670$  Å absorption, the n20def model does not match the earliest spectra of SN 2014dt nearly as well as the n1def model, and is unable to reproduce this absorption feature for any later epoch. The identity of the prominent absorption at  $\sim 4670$  Å therefore remains unclear and is likely a complex blended line.

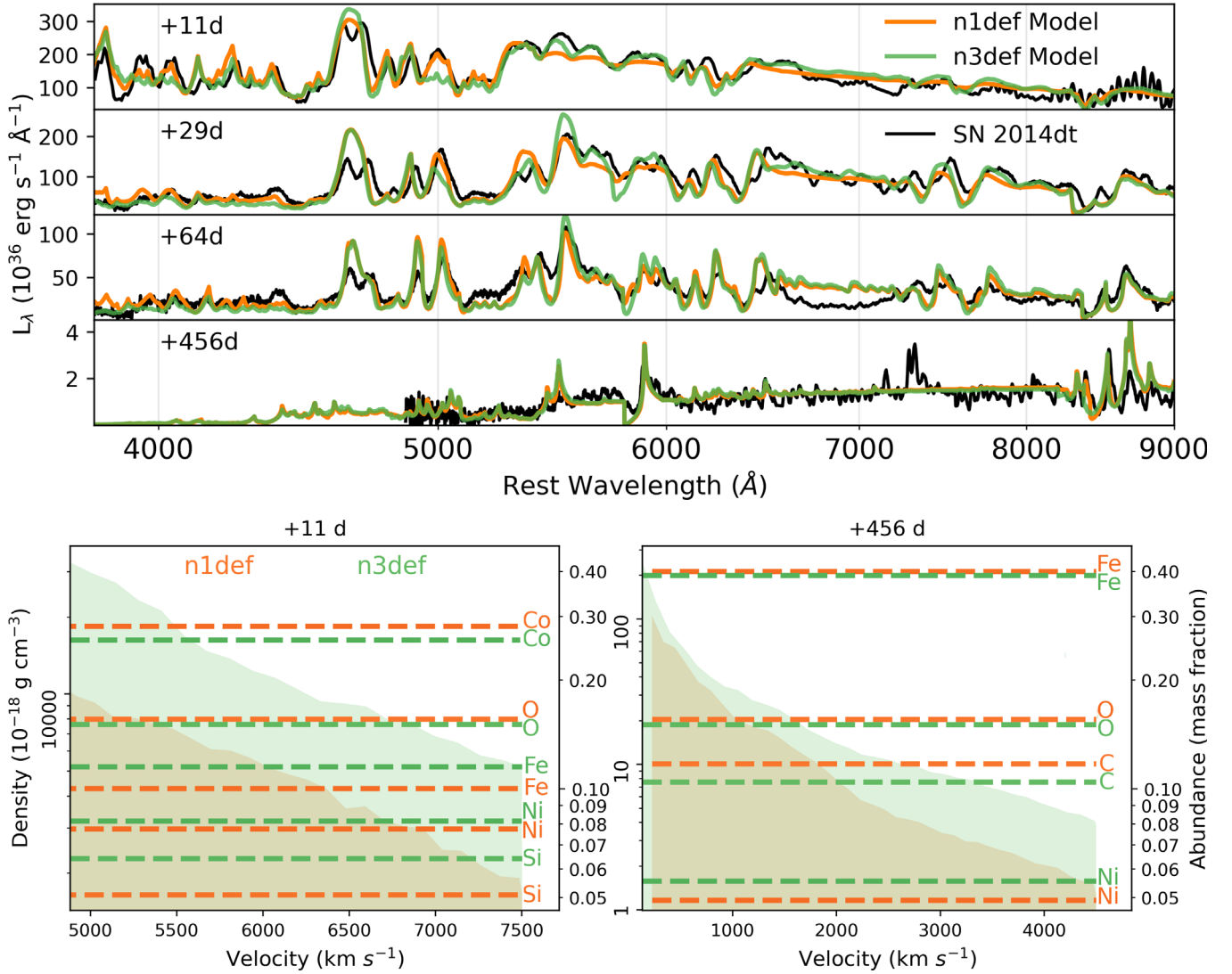
As described at the end of Section 2, we adopted a distance of 14.6 Mpc for SN 2014dt to calculate the luminosity input parameter,  $L_{\text{requested}}$ . Because of the large uncertainty in the distance to M61 we also explored a wide range of distances in the range 12.3–21.4 Mpc. We find that the late-time spectra, after about +64 days, are insignificantly affected by changes in the assumed distance, so we focus our varying-distance analysis on earlier epochs. We find that using distances below 14.6 Mpc has a small adverse effect on the overall fit of the spectrum produced that can be easily remedied by using a slightly lower model  $v_{\text{phot}}$  value. However, using distances at

the higher end of the plausible range produced poor fits to the observed spectra, even when we tested a wide range of  $v_{\text{phot}}$  values. In those cases, the best  $v_{\text{phot}}$  value to match the spectral features produced a continuum that was too hot and blue, and the best  $v_{\text{phot}}$  value that matched the continuum ( $\sim 2500$  km s $^{-1}$  higher than our original  $v_{\text{phot}}$ ) produced a featureless spectrum beyond  $\sim 5200$  Å and an overall poor fit to the blue end under  $\sim 5200$  Å. Based on the results of this analysis, along with the remarkable agreement of the n1def model to the spectra of SN 2014dt when using shorter distances, we suggest that the correct distance to M61 is likely less than about 18 Mpc.

Maeda & Kawabata (2022) recently analyze the Type Iax SN 2019muj with spectral modeling similar to that presented here. They find evidence for a dense inner core of a few hundredths of a solar mass, in contrast with the density profiles of the deflagration models we use. We find that we are able to satisfactorily model SN 2014dt without a dense inner core. We tested the Maeda & Kawabata (2022) inner core density profile for SN 2014dt and find that it produces reasonable fits to the early-time spectra,  $\lesssim +100$  days, but the n1def density profile without a central density enhancement still produces better fits. Fitting the late-time spectra with the dense inner core model requires a higher  $v_{\text{phot}}$  than in our n1def model and does not appear to improve the overall fit to the late-time spectra. Thus, while we cannot rule out an inner core in SN 2014dt, we do not require one. Despite these different inner density profiles, we find remarkable agreement in the photospheric velocities of SN 2014dt and those measured for SN 2019muj by Maeda & Kawabata (2022), even at the latest epochs at  $\sim +500$  days.

#### 4.3. Comparing Deflagration Models

Although the n1def model produced synthetic spectra that most closely resembled SN 2014dt, other deflagration models from



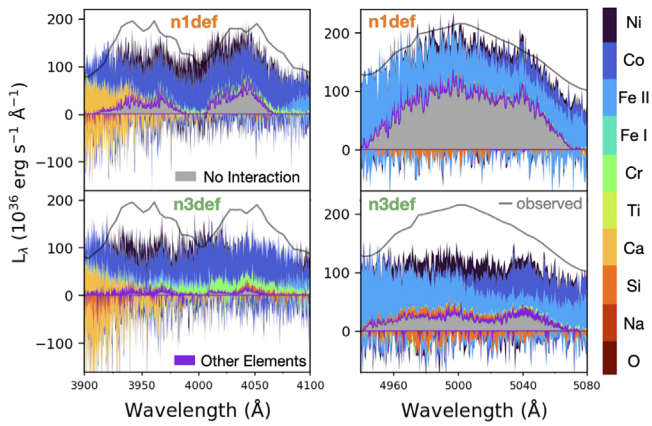
**Figure 5.** Top panel: early-time and late-time TARDIS spectra using the n1def (orange) and n3def (green) deflagration models from Fink et al. (2014). Both deflagration models fit the observed spectrum of SN 2014dt well; early epochs show the biggest differences, favoring the n1def model, and both models become extremely similar after +64 days. Bottom panels: density and select element abundances of the n1def and n3def models at early and late times.

Fink et al. (2014) with few ignition points (n3def in particular) also do a reasonable job. The density profile and abundances of n3def were generated in the same way as n1def, as described in Section 4.1. We keep all other TARDIS input parameters the same for each epoch, with only a very small change in the  $v_{\text{phot}}$  parameter for the latest epochs,  $>+410$  days.

In Figure 5, we compare the n1def and n3def deflagration models at various epochs. The top panel shows the resulting synthetic spectra while the bottom panel illustrates the density and abundance structure of each model. Though both models produce extremely similar spectra, as expected, the earliest epoch, +11 days, exhibits the greatest difference in spectral features. There are three main regions in the optical spectra that differ: 3900–4100  $\text{\AA}$ , 5000  $\text{\AA}$ , and 5300–6000  $\text{\AA}$ . The best-fit model, n1def, was chosen based on fitting SN 2014dt better in the 3900–4100  $\text{\AA}$  and 5000  $\text{\AA}$  regions. In particular, the peak at 5000  $\text{\AA}$  is fit extremely well by n1def while n3def is completely unable to produce this prominent feature. By +29 days, the difference between n1def and n3def at 5000  $\text{\AA}$  has declined but the differences at 3900–4100  $\text{\AA}$  and 5300–6000  $\text{\AA}$  remain. By +64 days, both deflagration models produce very similar

spectra with only small differences, mostly  $<4100 \text{ \AA}$ . This similarity continues for all subsequent epochs, as shown by the +456 days spectrum.

The bottom panel in Figure 5 compares the density profiles and element abundances of n1def and n3def at early and late times. The velocity range for each epoch is set from  $v_{\text{phot}}$  to  $v_{\text{outer}}$ ; note that at late times we probe much deeper (lower-velocity) layers in the ejecta. Both n1def and n3def have similarly declining slopes in their density profiles but n1def always has a lower density at a given velocity. The abundance profiles of both models are very similar and we select a few elements with high abundances to show in Figure 5. N1def has slightly more C and O, while n3def has more Si. Since the C, O, and Si abundances are stable over time, they are not all shown for both epochs. The abundances of elements with radioactive isotopes do change with time—this radioactive decay is done internally by TARDIS. At early times, +11 days, n1def has more Co, while n3def has more Fe. By +456 days, almost all of the radioactive Ni and Co has decayed to Fe. The abundance of Co becomes so small that it falls below the lower limit of our plot. The remaining Ni shown at +456 days is



**Figure 6.** Element decomposition of n1def and n3def model spectra at wavelengths where they differ most, compared to the observed SN 2014dt spectrum at +11 days.

stable  $^{58}\text{Ni}$ . The abundance of Fe at this late time is high for both models, with n1def having slightly more Fe than n3def.

A closer look at the 3900–4100 Å and 5000 Å regions in Figure 6 shows that more emission from the underlying blackbody (labeled “No Interaction”) in n1def is one of the reasons why this model matches SN 2014dt better at early epochs; we note that the blackbody radiation we see in our models is limited by TARDIS originating photons solely from the photosphere. The lower density of n1def leads, in part, to a lower optical depth at these wavelengths and allows more of the blackbody light through. N1def and n3def show similar, but not identical, contributions from elements and isotopes at these wavelengths; the main differences include n3def having more Cr emission at 3900–4100 Å, and more Co emission and Si absorption around 5000 Å.

The differences seen in the early-time spectra of n1def and n3def are mostly due to their different density profiles rather than abundances, such that a model with an n1def (n3def) density and n3def (n1def) abundance produces a spectrum that looks like a pure n1def (n3def) model. This sensitivity to the density profile, coupled with both models producing very similar spectra after +64 days, leads us to conclude that the density profile between  $\sim 1000$  and  $7500 \text{ km s}^{-1}$  is a key fitting parameter when reproducing the spectra of SN 2014dt.

In Table 2 we compare the observed properties of SN 2014dt to those generated by the n1def and n3def models. Owing to the large difference in the distances used by Kawabata et al. (2018) and Singh et al. (2018), it is difficult to compare their absolute brightness,  $M_{\text{max}}^B$ , so we focus on  $B_{\text{max}}$  instead. Kawabata et al. (2018) estimate a  $B_{\text{max}}$  of 13.76–13.88 mag while Singh et al. (2018) find a similar, but not overlapping,  $13.59 \pm 0.04$  mag. The observed  $B_{\text{max}}$  is a close match to the n3def model for our adopted extinction and distance of 14.6 Mpc. Even though the n1def model better matches our spectra, it predicts a  $B_{\text{max}}$  that is slightly too dim, even at a closer distance of 13.2 Mpc, the lower-limit distance in Kawabata et al. (2018).

As previously discussed in Section 4.1, SN 2014dt was discovered past peak brightness and Kawabata et al. (2018) estimated the bolometric light-curve rise time in the range 11.8–21.0 days using other SNe Iax. This range of rise times for SN 2014dt is slightly longer than what is expected for either n1def or n3def, though it is largely consistent with other

deflagration models in Fink et al. (2014). The decline rate  $\Delta m_{15}^B$  estimated by Kawabata et al. (2018) is 1.43–1.57 mag, while Singh et al. (2018) obtain  $1.35 \pm 0.06$  mag; both of these decline rates are slower than what the n1def and n3def models predict. Overall, the light-curve properties of SN 2014dt suggest a slightly brighter and more slowly evolving object than the n1def model, and to a lesser extent the n3def model. Thus, there is a slight discrepancy between the results of our spectral modeling (favoring n1def) and the photometric properties of SN 2014dt; we encourage further research into whether updated deflagration models can reconcile this issue.

The total ejecta masses,  $M_{\text{ejecta}}$ , predicted for n1def and n3def are  $0.0843 M_{\odot}$  and  $0.195 M_{\odot}$ , respectively. This is within the 0.08–0.35  $M_{\odot}$  range found by Kawabata et al. (2018) but far below the  $0.95 M_{\odot}$  predicted by Singh et al. (2018). This difference is due to the different assumptions and methods used in each paper, but is primarily driven by the greater distance and thus higher luminosity adopted by Singh et al. (2018). Both papers use a similar analysis of the late-time bolometric light curve to calculate the amount of  $^{56}\text{Ni}$  ( $M_{\text{Ni56}}$ ) synthesized in SN 2014dt, but come up with different results. Singh et al. (2018) suggest at least  $0.14 M_{\odot}$  of  $^{56}\text{Ni}$  was synthesized while Kawabata et al. (2018) estimate a smaller  $M_{\text{Ni56}}$  of 0.035–0.105  $M_{\odot}$ , which is a close match to both n1def and n3def. The important explosion parameters,  $M_{\text{ejecta}}$  and  $M_{\text{Ni56}}$ , calculated by Kawabata et al. (2018) for SN 2014dt suggest it is consistent with both the n1def and n3def models while the higher  $M_{\text{ejecta}}$  and  $M_{\text{Ni56}}$  from Singh et al. (2018) are more consistent with the higher densities of the n20def model, though we find that this model is not a good fit to the early-time spectra (see the discussion in Section 4.2). This again leads us to favor a lower distance and luminosity for SN 2014dt.

We explored fitting SN 2014dt with deflagration models having more ignition points (n10def, n20def, n150def, and n1600def) that span the full breadth of model densities and abundances from Fink et al. (2014). The flattening of abundances in velocity space is a reasonable simplification for models up to and including n10def; n20def and higher models were left with their velocity-dependent abundances (note that the adjustments of removing elements past Zn, and the increase of  $^{23}\text{Na}$ , still remain). This means that the densities of models with fewer ignition points (n10def and lower) were able to be extrapolated down to zero velocity with a cubic spline, while n20def and above were not extrapolated. This limits our modeling of n20def and higher models to  $\sim +420$  days, instead of more than +500 days, since that is when  $v_{\text{phot}}$  dips below the velocity limits of those models. By this very late epoch, all deflagration models look remarkably similar to one another and we do not expect n20def and higher models to deviate much from lower models past  $\sim +420$  days. All other TARDIS input parameters were the same as n1def for each epoch.

As before, the biggest differences between these deflagration models with more ignition points is seen at early times. At the earliest epoch, +11 days, n10def and higher models resemble n3def with an overall slightly worse fit. N20def and higher models were extremely similar, even at our earliest epoch of +11 days, with a small but noticeable difference in the strength of absorption at  $\sim 5500$  Å (n20def has the strongest absorption with decreasing strength toward n1600def). We find that past +64 days, all deflagration models produce a very similar spectrum—with n1def and n3def fitting the blue end of the



**Table 2**  
Comparison of SN 2014dt Properties Inferred in Published Works with Our Best-fit Spectral Models

	Distance (Mpc)	$M_{\max}^B$ (mag)	$B_{\max}$ (mag)	$t_{\max}^B$ (days)	$\Delta m_{15}^B$ (mag)	$M_{\text{ejecta}}$ ( $M_{\odot}$ )	$M_{\text{Ni56}}$ ( $M_{\odot}$ )
Singh et al. (2018)	$21.44 \pm 0.03$	$-18.13 \pm 0.04$	$13.59 \pm 0.04$	...	$1.35 \pm 0.06$	0.95	$>0.14$
Kawabata et al. (2018)	$13.24\text{--}15.92$	$-(16.82\text{--}17.34)$	$13.76\text{--}13.88$	$11.8\text{--}21.0^a$	$1.43\text{--}1.57$	$0.08\text{--}0.35$	$0.035\text{--}0.105$
Fink et al. (2014) n1def model	$14.6^b$	$-16.55$	$14.3^b$	7.6	2.15	0.0843	0.0345
Fink et al. (2014) n3def model	$14.6^b$	$-17.16$	$13.7^b$	9.6	1.91	0.195	0.0730

**Notes.** The different assumed distances affect the derived absolute magnitude and mass estimates. Details are discussed in Section 4.3.

<sup>a</sup> Estimated rise time of bolometric light curves from other SNe Iax.

<sup>b</sup> Using our adopted distance, 14.6 Mpc, and Milky Way extinction,  $A_V = 0.062$  mag.

spectrum slightly better than higher models—and are all able to do a reasonable job at fitting the optical spectrum of SN 2014dt. This reinforces our earlier conclusion that the choice of densities in the range  $\sim 1000\text{--}7500$  km s<sup>−1</sup> is a key input parameter for modeling the spectra of SN 2014dt. Though our models do not seem to be as sensitive to the density profile below  $\sim 1000$  km s<sup>−1</sup>, this could be partly due to the overall very low densities for epochs past +64 days.

#### 4.4. Fitting Strong Resonant Lines

We increase <sup>23</sup>Na from the original n1def abundance of  $\sim 0.01\%$  to  $\sim 1.7\%$  in order to fit the strong Na I  $\lambda\lambda 5890, 5896$  doublet absorption in SN 2014dt. A smaller <sup>23</sup>Na abundance of  $\sim 1\%$  was sufficient for reproducing the absorption feature for very late epochs,  $>+400$  days, but insufficient for reproducing the feature at +60–275 days. A <sup>23</sup>Na abundance of  $\sim 1.7\%$  is high, but reasonable because the model isotopes reflect only the nucleosynthetic yield of a pure carbon–oxygen WD and do not include any preexisting <sup>23</sup>Na in the system. We can speculate that a lower <sup>23</sup>Na abundance is possible if its distribution is not spherically symmetric or uniform across all ejecta layers. In particular, the inner ejecta layers that correspond to the epochs that require more <sup>23</sup>Na ( $300\text{--}2000$  km s<sup>−1</sup>) may have a higher <sup>23</sup>Na abundance than the outer layers. This stratified abundance structure is possible to model with TARDIS but is beyond the scope of this work.

Using  $v_{\text{outer}}$  from the Fink et al. (2014) models produced some spectral features that were too broad and blueshifted compared to the observed spectrum. Our choice of  $v_{\text{outer}}$ , always less than the Fink et al. (2014) model  $v_{\text{outer}}$ , is based on fitting the feature that was most sensitive to the choice of  $v_{\text{outer}}$ : the Ca II IR triplet. In fact, this was the only noticeable feature affected after +200 days. In earlier epochs, the Fe II features around  $5000$  Å were also slightly weak and blueshifted. This strong Ca sensitivity to the choice of  $v_{\text{outer}}$  may point to Ca being stratified—requiring less Ca in the outer layers—or not being spherically symmetric. Another possibility is that the Ca II IR triplet is sensitive to the line treatment and might require non-LTE effects, which are not included in our radiative-transfer model, as is suggested by Kasen (2006). Further analysis of the effects these options might have on the Ca distribution, especially beyond  $v_{\text{outer}}$ , could be the subject of future work. Our lower  $v_{\text{outer}}$  values also support the possibility that SNe Iax may need steeper density profiles, relative to pure deflagration models, as has been suggested by Barna et al. (2018) and Magee et al. (2022).

#### 4.5. Evolution of the Photosphere: Signs of a Quasi-steady-state Wind

With our spectral modeling we are able to obtain a robust evolution of the photospheric velocity,  $v_{\text{phot}}$ , that takes into account both the underlying thermal continuum as well as the complexity of fitting individual spectral lines. Our  $v_{\text{phot}}$  results from our best-fit model are listed in Table 3 and plotted in Figure 7.

For SN 2014dt, we find that the evolution of  $v_{\text{phot}}$  is not simply fit by a pure exponential, power law, or linear function; though the fast decline of  $v_{\text{phot}}$  up to +150–200 days can be roughly fit by an exponential or power law, the subsequent slower evolution is better fit linearly in time. The  $v_{\text{phot}}$  of SN 2014dt decreases very rapidly from  $4800$  km s<sup>−1</sup> to  $\sim 1000$  km s<sup>−1</sup> from +11 to +64 days, and continues to decrease more slowly until it is  $\sim 400$  km s<sup>−1</sup> by +148 days.  $v_{\text{phot}}$  then plateaus at  $\sim 400$  km s<sup>−1</sup> for the next 50 days (we fit 16 spectra during this time) before slowly decreasing to  $\sim 130$  km s<sup>−1</sup> by +456 days. Since we have two spectra (+454 and +456 days) from two different telescopes at roughly the same epoch, we are confident in both our reduction of the observed spectra and our model’s ability to fit the spectra at this epoch. For +562 days, we obtain a  $v_{\text{phot}}$  of  $\sim 40$  km s<sup>−1</sup>. While we were able to model the continuum and some photospheric features in this last spectrum, the fit was overall worse; in particular, the strength of the Ca IR triplet was higher in our model than what was observed. Owing to the faintness of SN 2014dt by this time, and without any other spectrum at a similar phase, it may be that this last spectrum has some contamination from its host galaxy that could alter the continuum and the  $v_{\text{phot}}$  of our best-fit spectral model by tens of kilometers per second. We also note that the low velocities we find in our late-time models are now more sensitive to any peculiar motion of the SN relative to its host, for which we do not account.

With the assumption of homologous expansion, we also show the evolution of the photospheric radius  $R_{\text{phot}}$  in Figure 8. In order to convert the original  $R_{\text{phot}}$  results from Kawabata et al. (2018) and J. Johansson et al. (2023, in preparation) to  $v_{\text{phot}}$ , we add our adopted rise time, 11 days, to the phase in order to get the time since explosion,  $t_{\text{exp}}$ . Assuming homologous expansion,  $v_{\text{phot}} = R_{\text{phot}} \times t_{\text{exp}}^{-1}$ .

We compare our  $v_{\text{phot}}$  results from spectral fitting with other works that derive  $v_{\text{phot}}$  from the photometry of SN 2014dt. Our results are in excellent agreement with Kawabata et al. (2018) at +230 days; they derive the bolometric luminosity and temperature using optical and near-IR photometry, and find the photosphere at  $R_{\text{phot}} \approx 6.2 \times 10^{14}$  cm ( $\sim 300$  km s<sup>−1</sup>). Using the additional assumption that the bolometric luminosity follows the <sup>56</sup>Co decay, they obtain  $R_{\text{phot}} \approx 2.8 \times 10^{14}$  cm ( $\sim 80$  km s<sup>−1</sup>) at +410 days. This second  $v_{\text{phot}}$ ,  $\sim 80$  km s<sup>−1</sup>, is much lower than our  $v_{\text{phot}}$  of

**Table 3**  
TARDIS Parameters

Phase <sup>a</sup> (days)	$t_{\text{exp}}$ (days)	$L_{\text{requested}}$ ( $L_{\odot}$ )	$L: \lambda_{\text{min}} - \lambda_{\text{max}}$ (Å)	$v_{\text{outer}}$ ( $\text{km s}^{-1}$ )	$v_{\text{phot}}$ ( $\text{km s}^{-1}$ )
+11	21.7	$1.65 \times 10^8$	4064–8757	7502	4800
+21	32.3	$1.35 \times 10^8$	4479–8951	7507	3637
+22	33.5	$1.32 \times 10^8$	4500–8452	7507	3513
+28	39.5	$1.05 \times 10^8$	4280–8452	7510	2855
+29	40.0	$1.27 \times 10^8$	3982–9446	7510	2800
+32	43.1	$4.80 \times 10^7$	5328–6970	7511	2580
+34	45.4	$9.11 \times 10^7$	4280–8452	7512	2417
+36	47.0	$1.03 \times 10^8$	3982–9446	7513	2304
+40	51.3	$4.96 \times 10^7$	5808–8951	7284	2103
+41	52.4	$7.37 \times 10^7$	4300–8452	7225	2051
+45	56.3	$6.37 \times 10^7$	4363–8452	7016	1869
+48	59.4	$2.22 \times 10^7$	6469–8453	6852	1727
+51	62.3	$2.33 \times 10^7$	4500–6000	6696	1591
+60	71.2	$4.36 \times 10^7$	4300–8452	6218	1175
+61	71.9	$3.54 \times 10^7$	5098–9296	6179	1140
+64	74.8	$4.58 \times 10^7$	3982–9446	4901	1015
+64	75.1	$3.46 \times 10^7$	5009–8951	4795	1000
+66	77.3	$1.46 \times 10^7$	4145–5965	4782	984
+76	87.2	$2.84 \times 10^7$	4465–8452	4721	913
+78	89.0	$3.17 \times 10^7$	4280–8452	4710	900
+82	93.0	$2.13 \times 10^7$	5032–8951	4685	871
+90	101.1	$2.60 \times 10^7$	4170–8452	4636	813
+92	102.6	$2.70 \times 10^7$	4181–9446	4626	795
+95	106.2	$2.84 \times 10^7$	3973–8861	4605	755
+97	107.9	$2.25 \times 10^7$	4377–8452	4594	735
+100	110.5	$2.13 \times 10^7$	4185–9446	4578	706
+104	114.8	$1.59 \times 10^7$	4714–8951	4552	679
+104	114.9	$2.07 \times 10^7$	4480–8451	4551	678
+108	119.0	$2.04 \times 10^7$	4381–8451	4526	653
+117	127.5	$2.23 \times 10^7$	3982–9446	4478	600
+119	130.0	$1.50 \times 10^7$	4246–8837	4480	600
+124	134.5	$1.76 \times 10^7$	4016–9446	4483	586
+124	135.5	$1.61 \times 10^7$	4280–9448	4484	583
+126	136.8	$1.51 \times 10^7$	4386–8452	4485	569
+133	143.6	$1.13 \times 10^7$	4814–8951	4490	497
+133	143.9	$1.83 \times 10^7$	3977–8795	4490	495
+138	148.9	$1.80 \times 10^7$	3978–8844	4494	464
+140	150.9	$1.61 \times 10^7$	3979–8844	4496	452
+143	153.7	$1.16 \times 10^7$	4677–8452	4498	435
+148	159.2	$1.60 \times 10^7$	3982–9446	4502	400
+148	159.3	$1.09 \times 10^7$	5103–9299	4502	400
+151	162.4	$9.78 \times 10^6$	4893–8951	4502	400
+156	166.5	$1.22 \times 10^7$	4343–8452	4502	400
+156	167.3	$1.49 \times 10^7$	3932–9402	4502	400
+159	169.7	$1.36 \times 10^7$	3982–8800	4502	400
+160	170.7	$1.49 \times 10^7$	3978–8844	4502	400
+169	180.2	$1.27 \times 10^7$	3782–9448	4502	400
+171	181.8	$5.82 \times 10^6$	4480–6554	4502	400
+175	185.6	$1.04 \times 10^7$	3982–8321	4502	400
+177	188.1	$9.96 \times 10^6$	4537–9446	4502	400
+180	190.6	$1.09 \times 10^7$	4181–8842	4502	400
+180	191.2	$1.36 \times 10^7$	3902–9447	4502	400
+184	194.5	$1.10 \times 10^7$	4120–8795	4502	400
+185	196.2	$1.04 \times 10^7$	3952–9447	4502	400
+187	198.5	$1.19 \times 10^7$	3977–8842	4502	400
+208	219.3	$9.23 \times 10^6$	3980–8843	4502	395
+211	221.9	$1.03 \times 10^7$	3684–9447	4502	370
+217	227.8	$9.64 \times 10^6$	3932–9446	4502	310
+233	244.1	$8.17 \times 10^6$	3979–8843	4501	302
+237	248.4	$7.10 \times 10^6$	4479–8472	4501	298
+238	248.7	$8.73 \times 10^6$	4051–9448	4501	298
+243	253.6	$8.08 \times 10^6$	5017–9447	4501	294
+275	286.2	$5.43 \times 10^6$	4483–8469	4501	270

**Table 3**  
(Continued)

Phase <sup>a</sup> (days)	$t_{\text{exp}}$ (days)	$L_{\text{requested}}$ ( $L_{\odot}$ )	$L: \lambda_{\text{min}} - \lambda_{\text{max}}$ (Å)	$v_{\text{outer}}$ ( $\text{km s}^{-1}$ )	$v_{\text{phot}}$ ( $\text{km s}^{-1}$ )
+410	421.1	$2.80 \times 10^6$	5014–9446	4500	170
+415	426.1	$3.16 \times 10^6$	3985–9448	4500	166
+420	431.0	$2.04 \times 10^6$	5274–9087	4500	162
+454	465.2	$1.34 \times 10^6$	5226–8759	4500	135
+456	467.2	$1.07 \times 10^6$	5373–8513	4500	128
+562	572.7	$3.96 \times 10^5$	4616–9204	4499	42

**Note.**

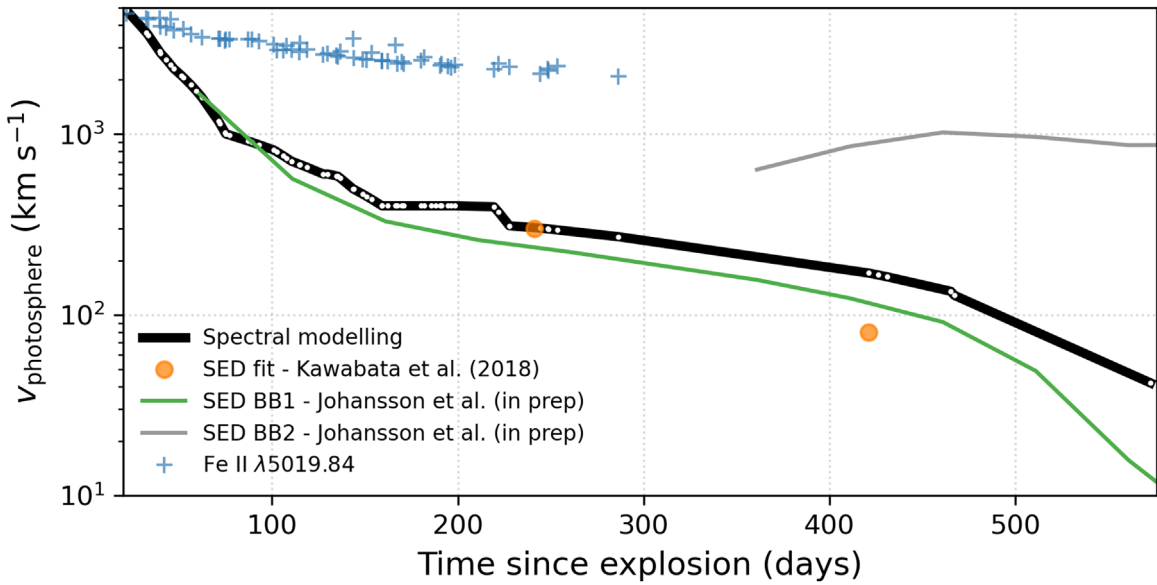
<sup>a</sup> Days since B maximum, 2014 October 20.4 (MJD 56950.4), corrected for time dilation.

(This table is available in machine-readable form.)

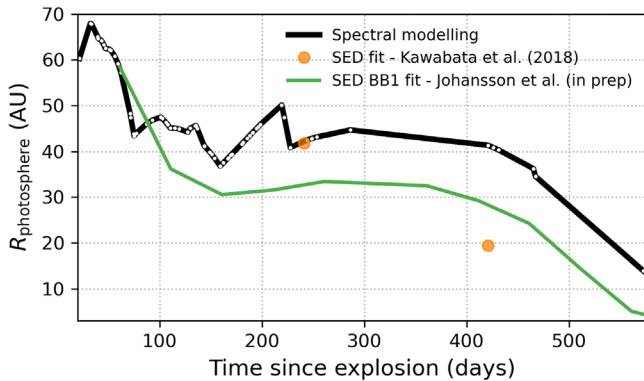
$\sim 170 \text{ km s}^{-1}$  for that same epoch. Our results suggest  $R_{\text{phot}}$  is  $\sim 6.2 \times 10^{14} \text{ cm}$  at +410 days, essentially unchanged from +230 days. If our  $R_{\text{phot}}$  at +410 days is correct, this would suggest that either the bolometric luminosity of SN 2014dt is holding steady between +230 and +410 days (assuming a constant temperature; Kawabata et al. 2018), or that a modest decrease in temperature,  $\sim 1000 \text{ K}$ , is needed to match the decreased bolometric luminosity of SN 2014dt at +410 days from Kawabata et al. (2018). We believe the second scenario to be more likely as the modest decrease in temperature would still produce a blackbody curve that closely matches the +410 day SED from Kawabata et al. (2018).

Our  $v_{\text{phot}}$  results also broadly agree with the photometric analysis of J. Johansson et al. (2023, in preparation). They fit two blackbody components, BB1 and BB2, to the optical and near-IR SED of SN 2014dt. BB1 is used mainly to fit the SED below  $\sim 2 \mu\text{m}$ , and BB2 fits the SED above  $\sim 2 \mu\text{m}$  after +350 days. Our  $v_{\text{phot}}$  results match well with the general evolution of the BB1 component, with our  $v_{\text{phot}}$  being slightly ( $\lesssim 100 \text{ km s}^{-1}$ ) lower than their  $v_{\text{phot}}$  after +100 days. The BB2 component is at a much larger radius than the photosphere derived from our spectral modeling. This component is associated with the MIR excess seen by Fox et al. (2016), and though our data and models cannot directly bear upon the question of whether this long-wavelength flux is from dust emission or optically thick wind interaction (Foley et al. 2016), it is clear that it involves faster-moving material at larger distances than the optical photosphere.

Traditionally,  $v_{\text{phot}}$  is measured by line shifts of individual line profiles. In Figure 4, we show how most line profiles are largely blended and, with further exploration of the individual line interactions, we note that labeling any absorption or emission feature with a single line identification is often misleading. We are able to identify one absorption feature that is largely from a single line for all epochs when it appears, Fe II  $\lambda 5019.84$ , and measure its velocity with a Gaussian fit to the line profile. At our earliest epoch, +11 days, the velocity measured from Fe II  $\lambda 5019.84$  agrees well with our results from spectral modeling. In the next epoch, just 10 days later, the velocity of Fe II is already  $\sim 700 \text{ km s}^{-1}$  higher than our spectral modeling  $v_{\text{phot}}$ . By +275 days, the last epoch where we could reliably identify Fe II  $\lambda 5019.84$ , the line velocity was  $\sim 2000 \text{ km s}^{-1}$  while the  $v_{\text{phot}}$  from spectral modeling was  $\sim 270 \text{ km s}^{-1}$ . Overall, the velocity measured from Fe II  $\lambda 5019.84$  declines much more slowly than our spectral modeling  $v_{\text{phot}}$ , suggesting that the Fe II line is tracing faster, outer material in the ejecta which is distinct from the true



**Figure 7.** The photospheric velocity,  $v_{\text{phot}}$ , of SN 144dt from modeling the optical spectra is shown as a solid black curve, with individual spectral fits as white dots. The line velocities of Fe II  $\lambda 5019.84$  are shown as blue crosses. For comparison, we display results from Kawabata et al. (2018; solid orange dots) who calculated the  $v_{\text{phot}}$  of SN 144dt using the bolometric luminosity and temperature from an analysis of optical and near-IR photometry. We also show the results from J. Johansson et al. (2023, in preparation); they calculate  $v_{\text{phot}}$  from spectral energy distribution (SED) blackbody fits to the optical and near-IR photometry of SN 144dt. Their blackbody curve 1 (BB1; solid green) fits the SED below  $\sim 2 \mu\text{m}$  while their blackbody curve 2 (BB2; solid gray) fits the SED above  $\sim 2 \mu\text{m}$ .



**Figure 8.** Photospheric radius,  $R_{\text{phot}}$ , of SN Iax 144dt from the spectral models. For comparison, the results from the photometric analyses by Kawabata et al. (2018) and J. Johansson et al. (2023, in preparation) are included.

underlying photosphere. Our results also suggest that measuring  $v_{\text{phot}}$  from a single line is no longer appropriate a few weeks past maximum brightness.

The line velocities of Fe II  $\lambda 6149$  and  $\lambda 6247$  in SN 144dt are measured by Kawabata et al. (2018) from +21–120 days and are in rough agreement with our Fe II  $\lambda 5019.84$  velocities. Our extension of the line velocity out to +275 days demonstrates that Fe II continues to stay above  $\sim 2000 \text{ km s}^{-1}$  even at these very late epochs. Our analysis of the Fe II  $\lambda 6149$  and  $\lambda 6247$  lines with our spectral models show that the Fe II  $\lambda 6149$  absorption line is correctly identified and generally uncontaminated, but the Fe II  $\lambda 6247$  absorption is actually mostly Fe II  $\lambda 6240.12$ . This different identification would only cause a small change in the resulting velocities,  $\sim 300 \text{ km s}^{-1}$ , which is within the  $\sim 400 \text{ km s}^{-1}$  line velocity scatter noted by Kawabata et al. (2018). Though we are able to confirm the line velocities from Kawabata et al. (2018), we find that the Fe II  $\lambda 6149$  and  $\lambda 6247$  lines have weaker absorption strengths and last for less time than Fe II  $\lambda 5019.84$ . We identify Fe II  $\lambda 5019.8$ , a seldom used, but very useful line that can provide

an approximate  $v_{\text{phot}}$  value for very early times around maximum light and can also trace a physically important region for Fe out to later times than more commonly used lines can.

Based on Figure 8, we see that  $R_{\text{phot}}$  holds very steady at  $\sim 40$ – $50 \text{ au}$  from +60 to 420 days and then decreases to  $\sim 14 \text{ au}$  by +562 days. While it is hard to be certain about the significantly lower  $R_{\text{phot}}$  at +562 days—the  $v_{\text{phot}}$  uncertainty at this epoch was discussed at the beginning of this section—the  $R_{\text{phot}}$  falling to  $\sim 35 \text{ au}$  by +454 days may indicate the smaller size at this epoch is real. These results are in general agreement with the evolution of the BB1 component from J. Johansson et al. (2023, in preparation), though our  $R_{\text{phot}}$  estimate is  $\sim 10 \text{ au}$  higher after  $\sim +100$  days. Our results also concur with the estimate from Kawabata et al. (2018) at +230 days; at +410 days, our  $R_{\text{phot}}$  is a factor of 2 larger than their estimate, but we discuss how a reasonable decrease in temperature could bring their estimate in agreement with ours earlier in this section. We do not include the BB2 component from J. Johansson et al. (2023, in preparation) and the Fe II  $\lambda 5019.84$  line in Figure 8 because they are poor tracers of the optical photosphere and produce an  $R_{\text{phot}}$  that is significantly larger than our spectral modeling results.

A photosphere with a constant velocity and density could be evidence for a steady-state wind coming from a bound remnant. Our results support this interpretation since we see a constant  $v_{\text{phot}}$  from +148 to 208 days and then a very slow decrease from +208 to 456 days. The wind model is further supported by the constant density seen at  $\sim +150$ – $450$  days inferred from the steady ratio of forbidden to permitted Ca II in the spectra of SN 144dt (other SNe Iax also show a steady late-time density; McCully et al. 2014). A constant density breaks the TARDIS model assumption of simple homologous expansion where density decreases as  $t^{-3}$ , and suggests more sophisticated models may be required. Even though TARDIS cannot directly model the steady-state wind, we can estimate a wind mass-loss rate starting at  $\sim +150$  days when the photospheric velocity levels off. At this epoch  $v_{\text{phot}} = 400 \text{ km s}^{-1}$ ,  $R_{\text{phot}} = 5.5 \times 10^{14} \text{ cm}$ , and  $\rho$  at  $v_{\text{phot}} = 1.7 \times 10^{-15} \text{ g cm}^{-3}$ ; these correspond to a wind mass-loss rate of  $\sim 4 \times 10^{-3} M_{\odot} \text{ yr}^{-1}$ . If this



mass-loss rate remains steady out to  $\sim +450$  days, the total mass in this wind would be  $\sim 3 \times 10^{-3} M_{\odot}$ , a small fraction of the mass of the bound remnants found in weak deflagration models (Fink et al. 2014). The slow decrease in  $v_{\text{phot}}$  over the range  $+208$ – $456$  days implies that the total mass in the wind may be a bit lower than this estimate. The drop in photospheric velocity at late times further suggests that the power source of the wind might be waning, perhaps consistent with the depletion of radioactive material.

## 5. Summary

We present an extensive spectral timeseries—69 total (21 new) spectra—of the Type Iax SN 2014dt, making it one of the best spectroscopically observed SNe. The spectra show over 500 days of evolution in the long-lasting photospheric phase of SNe Iax. In no other class of SN do photospheric lines persist to such late times. We identify the epochs between  $+90$  and  $+150$  days past maximum brightness as the spectral divergence phase, when SNe Iax transition away from typical SNe Ia spectral evolution.

We are able to reproduce the optical spectra of SN 2014dt with a nearly unaltered, weak deflagration explosion model (n1def) from Fink et al. (2014) using the radiative-transfer code TARDIS. Our models are self-consistent, using the same chemical abundance and density profiles throughout, and are able to model the spectra from  $+11$  to  $+562$  days after  $B_{\text{max}}$ . To our knowledge, this is the latest-epoch SN spectrum TARDIS has ever been able to model successfully, only possible because of the continuing photospheric features of SNe Iax at epochs where normal SNe Ia are dominated by nebular features. It is astonishing that the n1def deflagration model, created well before the appearance of SN 2014dt, can successfully reproduce the spectral evolution of this SN over 500 days in its evolution; such a match between a SN model prediction and observations may be unprecedented.

We explore in detail the evolution of the photospheric velocity—our main fitting parameter,  $v_{\text{phot}}$ . We find that the  $v_{\text{phot}}$  of SN 2014dt decreases very rapidly from  $+11$  to  $+64$  days, then slows down significantly between  $+64$  and  $+148$  days, and continues to have a very slow, sometimes unchanging, decrease until  $+456$  days. The significant slowdown in the evolution of  $v_{\text{phot}}$  between  $+64$  and  $+148$  days closely overlaps the spectral divergence phase we find for SN 2014dt,  $+90$  to  $+150$  days. The transition happens at a photospheric velocity of  $\sim 400$ – $1000 \text{ km s}^{-1}$ , which at these epochs corresponds to a photospheric radius of  $\sim 40$ – $50 \text{ au}$ . The correspondence in time between this slowdown and the spectral divergence phase suggests that this range of photospheric velocities may demarcate a boundary in the ejecta below which SNe Iax most strongly differ from their normal SNe Ia cousins.

The slow, nearly flat, evolution of  $v_{\text{phot}}$  from our spectral models, combined with a roughly constant density implied by the steady ratio of forbidden to permitted Ca II, is consistent with the picture that a leftover bound remnant from SN 2014dt began to drive an optically thick, quasi-steady-state wind. We estimate a wind mass-loss rate of a few times  $10^{-3} M_{\odot} \text{ yr}^{-1}$ , corresponding to a total mass loss of a few times  $10^{-3} M_{\odot}$ , during this phase from approximately  $+150$  to  $+450$  days after maximum light. Intriguingly, after about  $+450$  days, the photospheric velocity begins to decrease more rapidly, leading us to speculate that the wind may be weakening, perhaps as the radioactive power source decays away. Future observations of more SNe Iax will illuminate this issue. Of special interest is whether the IR excess emission seen in SN 2014dt (Fox et al.

2016; J. Johansson et al. 2023, in preparation), starting at about the same time as the wind dissipates (but seemingly originating at a much larger photospheric radius), will be commonly seen in SNe Iax. JWST IR observations of SNe Iax may play an important role in answering these questions.

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## References

- Allison, J. R., Sadler, E. M., & Meekin, A. M. 2014, *MNRAS*, **440**, 696  
 Astropy Collaboration, Price-Whelan, A. M., Sipőcz, B. M., et al. 2018, *AJ*, **156**, 123  
 Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, **558**, A33  
 Barna, B., Szalai, T., Jha, S. W., et al. 2021, *MNRAS*, **501**, 1078  
 Barna, B., Szalai, T., Kerzendorf, W. E., et al. 2018, *MNRAS*, **480**, 3609  
 Barna, B., Szalai, T., Kromer, M., et al. 2017, *MNRAS*, **471**, 4865  
 Black, C. S., Fesen, R. A., & Parrent, J. T. 2016, *MNRAS*, **462**, 649  
 Bottinelli, L., Gouguenheim, L., Paturel, G., & de Vaucouleurs, G. 1984, *A&AS*, **56**, 381  
 Bowers, E. J. C., Meikle, W. P. S., Geballe, T. R., et al. 1997, *MNRAS*, **290**, 663  
 Branch, D., Baron, E., Thomas, R. C., et al. 2004, *PASP*, **116**, 903  
 Branch, D., Jeffery, D. J., Parrent, J., et al. 2008, *PASP*, **120**, 135  
 Childress, M. J., Hillier, D. J., Seitzmann, I., et al. 2015, *MNRAS*, **454**, 3816  
 Clemens, J. C., Crain, J. A., & Anderson, R. 2004, *Proc. SPIE*, **5492**, 331  
 Cooper, M. C., Newman, J. A., Davis, M., Finkbeiner, D. P., & Gerke, B. F. 2012, spec2d: DEEP2 DEIMOS Spectral Pipeline, Astrophysics Source Code Library, ascl:1203.003  
 Crawford, S. M., Still, M., Schellart, P., et al. 2010, *Proc. SPIE*, **7737**, 25  
 Dutta, A., Sahu, D. K., Anupama, G. C., et al. 2022, *ApJ*, **925**, 217  
 Faber, S. M., Phillips, A. C., Kibrick, R. I., et al. 2003, *Proc. SPIE*, **4841**, 1657  
 Filippenko, A. V. 1982, *PASP*, **94**, 715  
 Filippenko, A. V. 2003, in From Twilight to Highlight: The Physics of Supernovae, ed. W. Hillebrandt & B. Leibundgut, 171  
 Fink, M., Kromer, M., Seitzmann, I. R., et al. 2014, *MNRAS*, **438**, 1762  
 Foley, R. J., Brown, P. J., Rest, A., et al. 2010, *ApJL*, **708**, L61  
 Foley, R. J., Challis, P. J., Chornock, R., et al. 2013, *ApJ*, **767**, 57  
 Foley, R. J., Jha, S. W., Pan, Y.-C., et al. 2016, *MNRAS*, **461**, 433  
 Foley, R. J., Papenkova, M. S., Swift, B. J., et al. 2003, *PASP*, **115**, 1220  
 Foley, R. J., Van Dyk, S. D., Jha, S. W., et al. 2015, *ApJL*, **798**, L37  
 Fox, O. D., Johansson, J., Kasliwal, M., et al. 2016, *ApJL*, **816**, L13  
 Friesen, B., Baron, E., Wisniewski, J. P., et al. 2014, *ApJ*, **792**, 120  
 Guillochon, J., Parent, J., Kelley, L. Z., & Margutti, R. 2017, *ApJ*, **835**, 64  
 Harris, C. R., Millman, K. J., van der Walt, S. J., et al. 2020, *Natur*, **585**, 357  
 Hook, I. M., Jørgensen, I., Allington-Smith, J. R., et al. 2004, *PASP*, **116**, 425  
 Hunter, J. D. 2007, *CSE*, **9**, 90  
 Jha, S., Branch, D., Chornock, R., et al. 2006, *AJ*, **132**, 189  
 Jha, S. W. 2017, in Handbook of Supernovae, ed. S. W. Jha & P. Murdin (Cham: Springer International Publishing), 375  
 Jordan, George C., & Perets, I. 2012, *ApJL*, **761**, L23  
 Karambelkar, V. R., Kasliwal, M. M., Maguire, K., et al. 2021, *ApJL*, **921**, L6  
 Kasen, D. 2006, *ApJ*, **649**, 939  
 Kashyap, R., Haque, T., Lorén-Aguilar, P., García-Berro, E., & Fisher, R. 2018, *ApJ*, **869**, 140  
 Kawabata, M., Kawabata, K. S., Maeda, K., et al. 2018, *PASJ*, **70**, 111  
 Kerzendorf, W., Sim, S., Vogl, C., et al. 2023, tardis-sn/tardis: TARDIS v2023.01.11, Zenodo, [10.5281/zenodo.7525913](https://doi.org/10.5281/zenodo.7525913)  
 Kerzendorf, W. E., & Sim, S. A. 2014, *MNRAS*, **440**, 387  
 Kluyver, T., Ragan-Kelley, B., Pérez, F., et al. 2016, in Positioning and Power in Academic Publishing: Players, Agents and Agendas, ed. F. Loizides & B. Schmidt (Amsterdam: IOS Press), 87  
 Kromer, M., Fink, M., Stanish, V., et al. 2013, *MNRAS*, **429**, 2287  
 Kromer, M., Ohlmann, S., & Röpke, F. K. 2017, *Mem. Soc. Astron. Italiana*, **88**, 312  
 Li, W., Filippenko, A. V., Chornock, R., et al. 2003, *PASP*, **115**, 453  
 Lyman, J. D., Taddia, F., Stritzinger, M. D., et al. 2017, *MNRAS*, **473**, 1359  
 Maeda, K., & Kawabata, M. 2022, *ApJ*, **941**, 15  
 Magee, M. R., Gillanders, J. H., Maguire, K., Sim, S. A., & Callan, F. P. 2022, *MNRAS*, **509**, 3580  
 Magee, M. R., Kotak, R., Sim, S. A., et al. 2016, *A&A*, **589**, A89  
 Magee, M. R., Kotak, R., Sim, S. A., et al. 2017, *A&A*, **601**, A62  
 Magee, M. R., Sim, S. A., Kotak, R., Maguire, K., & Boyle, A. 2019, *A&A*, **622**, A102  
 Mazzali, P. A., Sullivan, M., Filippenko, A. V., et al. 2015, *MNRAS*, **450**, 2631  
 Mazzali, P. A., Sullivan, M., Hachinger, S., et al. 2014, *MNRAS*, **439**, 1959  
 McCully, C., Jha, S. W., Foley, R. J., et al. 2014, *Natur*, **512**, 54  
 McCully, C., Jha, S. W., Foley, R. J., et al. 2014, *ApJ*, **786**, 134  
 McKinney, W. 2010, in Proc. IX Python in Science Conf., ed. S. van der Walt & J. Millman (Austin, TX: SciPy), 56  
 Miller, A. A., Kasliwal, M. M., Cao, Y., et al. 2017, *ApJ*, **848**, 59  
 Moriya, T., Tominaga, N., Tanaka, M., et al. 2010, *ApJ*, **719**, 1445  
 Nakano, S., Itagaki, K., Guido, E., et al. 2014, CBET, **4011**  
 NASA/IPAC Extragalactic Database (NED) 2019, NASA/IPAC Extragalactic Database (NED), IPAC, doi:[10.26132/NED1](https://doi.org/10.26132/NED1)  
 Newman, J. A., Cooper, M. C., Davis, M., et al. 2013, *ApJS*, **208**, 5  
 Ochner, P., Tomasella, L., Benetti, S., et al. 2014, *ATel*, **6648**  
 Phillips, M. M., Li, W., Frieman, J. A., et al. 2007, *PASP*, **119**, 360  
 Prochaska, J., Hennawi, J., Westfall, K., et al. 2020a, *JOSS*, **5**, 2308  
 Prochaska, J. X., Hennawi, J., Cooke, R., et al. 2020b, pypeit/PypeIt: Release v1.0.0, Zenodo, doi:[10.5281/zenodo.3743493](https://doi.org/10.5281/zenodo.3743493)  
 Sahu, D. K., Tanaka, M., Anupama, G. C., et al. 2008, *ApJ*, **680**, 580  
 Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, **737**, 103  
 Science Software Branch at STScI 2012, PyRAF: Python alternative for IRAF, Astrophysics Source Code Library, record, ascl:1207.011  
 Silverman, J. M., Foley, R. J., Filippenko, A. V., et al. 2012a, *MNRAS*, **425**, 1789  
 Silverman, J. M., Ganeshalingam, M., Cenko, S. B., et al. 2012b, *ApJL*, **756**, L7  
 Silverman, J. M., Nugent, P. E., Gal-Yam, A., et al. 2013, *ApJS*, **207**, 3  
 Silverman, J. M., Vinkó, J., Marion, G. H., et al. 2015, *MNRAS*, **451**, 1973

- Smith, M. P., Nordsieck, K. H., Burgh, E. B., et al. 2006, *Proc. SPIE*, **6269**, 62692A
- Singh, M., Misra, K., Sahu, D. K., et al. 2018, *MNRAS*, **474**, 2551
- Sparks, W. B. 1994, *ApJ*, **433**, 19
- Srivastav, S., Smartt, S. J., Huber, M. E., et al. 2022, *MNRAS*, **511**, 2708
- Srivastav, S., Smartt, S. J., Leloudas, G., et al. 2020, *ApJL*, **892**, L24
- Stahl, B. E., Zheng, W., de Jaeger, T., et al. 2020, *MNRAS*, **492**, 4325
- Tody, D. 1986, in SPIE Conf. Ser. 627, *Instrumentation in Astronomy VI*, ed. D. L. Crawford (Bellingham, WA: SPIE), 733
- Tody, D. 1993, in ASP Conf. Ser. 52, *Astronomical Data Analysis Software and Systems II*, ed. R. J. Hanisch, R. J. V. Brissenden, & J. Barnes (San Francisco, CA: ASP), 173
- Valenti, S., Pastorello, A., Cappellaro, E., et al. 2009, *Natur*, **459**, 674
- Valenti, S., Sand, D., Pastorello, A., et al. 2014, *MNRAS*, **438**, L101
- Virtanen, P., Gommers, R., Oliphant, T. E., et al. 2020, *NatMe*, **17**, 261
- Wade, R. A., & Horne, K. 1988, *ApJ*, **324**, 411
- Yaron, O., & Gal-Yam, A. 2012, *PASP*, **124**, 668
- Zhang, J.-J., Wang, X.-F., Bai, J.-M., et al. 2014, *AJ*, **148**, 1