



The LIGO HET Response (LIGHETR) Project to Discover and Spectroscopically Follow Optical Transients Associated with Neutron Star Mergers*

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Abstract

The LIGO HET Response (LIGHETR) project is an enterprise to follow up optical transients (OTs) discovered as gravitational-wave merger sources by the LIGO/Virgo collaboration (LVC). Early spectroscopy has the potential to constrain crucial parameters such as the aspect angle. The LIGHETR collaboration also includes the capacity to model the spectroscopic evolution of mergers to facilitate a real-time direct comparison of models with our data. The principal facility is the Hobby–Eberly Telescope. LIGHETR uses the massively replicated VIRUS array of spectrographs to search for associated OTs and obtain early blue spectra, and in a complementary role, the low-resolution LRS2 spectrograph is used to obtain spectra of viable candidates as well as a densely sampled series of spectra of true counterparts. Once an OT is identified, the anticipated cadence of spectra would match or considerably exceed anything achieved for GW170817 = AT2017gfo for which there were no spectra in the first 12 hr and thereafter only roughly once daily. We describe special HET-specific software written to facilitate the program and attempts to determine the flux limits to undetected sources. We also describe our campaign to follow up OT candidates during the third observational campaign of the LIGO and Virgo Scientific Collaborations. We obtained VIRUS spectroscopy of candidate galaxy hosts for five LVC gravitational-wave events and LRS2 spectra of one candidate for the OT associated with S190901ap. We identified that candidate, ZTF19abvionh = AT2019pip, as a possible Wolf–Rayet star in an otherwise unrecognized nearby dwarf galaxy.

Unified Astronomy Thesaurus concepts: [Gravitational wave astronomy \(675\)](#)

1. Introduction

The first gravitational-wave event discovered by LIGO, the merger of two black holes in a binary system (BBH), opened an exciting new vista in multimessenger astronomy (Abbott et al. 2016). Utilizing this new window into the universe is one of the most exciting prospects in astrophysics. Optical astronomers

and observers in other electromagnetic bands were invited to this arena with the advent of GW170817, the result of the inspiral and merger of two neutrons stars (BNS), and its gamma-ray (GW170817A) and optical (AT2017gfo) counterparts (Abbott et al. 2017a, 2017b; Coulter et al. 2017; Evans et al. 2017; Haggard et al. 2017; Hallinan et al. 2017; Kasen et al. 2017; Kasliwal et al. 2017; Kilpatrick et al. 2017; Margutti et al. 2017; Murguia-Berthier et al. 2017; Nicholl et al. 2017; Tanvir et al. 2017; Troja et al. 2017; Valenti et al. 2017; Villar et al. 2017; Lyman et al. 2018; Margutti et al. 2018; Mooley et al. 2018; Pooley et al. 2018; Hajela et al. 2019; Lamb et al. 2019; Piro et al. 2019; Troja et al. 2019; Gillanders et al. 2022; Hajela et al. 2022; Kilpatrick et al. 2022). AT2017gfo corresponded to a kilonova, a transient about 1000 times brighter than a classical nova (Li & Paczyński 1998; Metzger et al. 2010).

* Based on observations obtained with the Hobby–Eberly Telescope (HET), which is a joint project of the University of Texas at Austin, the Pennsylvania State University, Ludwig-Maximilians-Universität München, and Georg-August Universität Göttingen. The HET is named in honor of its principal benefactors, William P. Hobby and Robert E. Eberly.



Many groups have reported their techniques for searching for optical transients (OTs) associated with gravitational-wave events in the third LIGO/Virgo collaboration (LVC) run that began on 2019 April 1 (O3; Gomez et al. 2019; Hosseinzadeh et al. 2019; Lundquist et al. 2019; Yang et al. 2019; Ackley et al. 2020; Gompertz et al. 2020; Kasliwal et al. 2020; Becerra et al. 2021; Chang et al. 2021; Kim et al. 2021; Oates et al. 2021; Paterson et al. 2021; de Jaeger et al. 2022; Rastinejad et al. 2022). Here we report on the LIGO HET Response (LIGHETR) project that employed spectrographs on the Hobby–Eberly Telescope (HET) at McDonald Observatory to conduct a spectroscopic search and follow up during O3.

The first spectra were taken 0.88 and 1.84 hr after the detection of the OT of AT2017gfo (Shappee et al. 2017), fully 12 hr after the original LVC GW signal. The next few spectra were obtained at ~ 1.18 and 1.5 days (McCully et al. 2017; Nicholl et al. 2017; Pian et al. 2017; Buckley et al. 2018) and then a few at daily intervals after that, including VLT/X-shooter data extending into the infrared (Pian et al. 2017). The target varied rapidly over that timescale and was unobservable spectroscopically after about 10 days. Pian et al. (2017) argue for the detection of shallow, broad features at about 8100 and 12300 Å in a spectrum obtained at 1.5 days and several subsequent spectra. Watson et al. (2019) suggest that the former may be evidence of the neutron-capture element strontium. Perego et al. (2022) and Tarumi et al. (2023) explore the possibility that the feature is He I $\lambda 10830$ and stress the possible role of non-LTE effects in differentiating strontium from helium.

Although it was fortuitously nearby, practical factors prevented a higher cadence of spectroscopy of AT2017gfo. The southern location at a decl. of $-17^{\circ}51'$ meant that it was primarily visible in the south and only for about 2 hr per night. A delay in the release of the refined location meant that it was too late to observe from Africa and too early for Chile. Once the OT was accessible in Chile, discovery came quickly and spectroscopy shortly thereafter. The result was an inadequate cadence for such a rapidly changing event. We were left to wonder about the nature of earlier spectra.

LIGHETR is a program designed to complement the global effort to obtain spectroscopy of merger OT components. HET, with the massively replicated wide-field VIRUS IFU spectrographs (Hill et al. 2018a, 2018b; Gebhardt et al. 2021; Hill et al. 2021) and the Marcario Low-Resolution Spectrograph (LRS2; Hill et al. 1998; Chonis et al. 2016; G. J. Hill et al. 2023, in preparation) can play a significant role in multi-messenger astronomy since (1) the queue-scheduled HET is designed to respond quickly to events, (2) the blue sensitivity of VIRUS that extends to 3500 Å can be an important discriminant for various models of the mergers, (3) VIRUS provides the largest sky area coverage of any spectrograph by a factor of 70, (4) the location of GW170817 at the edge of a nearby galaxy highlights the advantage of using an array like VIRUS, and (5) LRS2 can provide frequent wide wavelength spectroscopic coverage of these rapidly evolving events.

With the HET, we can, in principle, discover the OT, get the first spectrum, and then a dense time series of spectra. Ideally, we could obtain several spectra per night, depending on the decl., and maintain this schedule until the OT fades below detectability, in roughly a week. In this time, under ideal circumstances, we could obtain of order 20 spectra and capture in detail the structure and evolution of these incredible events.

O3 was a resounding success for the detection of gravitational-wave signals from compact star mergers, but disappointing in the lack of observable OTs. Here we describe our attempts to detect six merger events spectroscopically. We failed to see any OT associated with a gravitational-wave event, but did reveal an interesting object, ZTF19abvionh, which is of intrinsic interest for its own sake.

Section 2 gives an outline of the LIGHETR observational program. Section 3 presents some results of the program in the first portion of O3 that spanned 2019 April 1 to September 31. Section 3.5 presents our follow up of ZTF candidate ZTF19abvionh that proved to be interesting, but not a merger counterpart. Section 4 explores attempts to quantify our ability to assess the upper limits of undetected sources. Section 5 presents models that illustrate the potential power of early, frequent, spectral observations. Section 6 summarizes the results to date and our future plans. Appendix A gives a table of compositions used in kilonovae light curve and spectral simulations, Appendix B gives some details of LIGHETR operating procedures, and Appendix C presents a notional timetable for responding to an LVC event and initiating a spectroscopic search and follow up.

2. Observational Program

LIGHETR was designed to observe about one LVC event per trimester during O3 using the VIRUS IFU array spectroscopically to search for and perhaps discover the OT and to use the LRS2 low-resolution spectrograph to do intensive, high-cadence spectroscopic follow up. LIGHETR was prepared rapidly to generate the appropriate target information from a trigger and work with the HET staff in order to implement the search and subsequent spectroscopic monitoring as rapidly as possible. Our observational strategy involved digesting public alerts with our custom software DIAGNOSIS and assessing if the candidate is visible to HET. After further human vetting of the candidate, the VIRUS array was utilized to take spectra over a wide field of view (FoV) centered on the candidate. We have the capacity to carry out further intensive spectroscopic follow up with the LRS2 instrument. LIGHETR has the capability to be on target 20 minutes after an alert, achieving dense spectral coverage of the target. LIGHETR can, in principle, observe in bright time since the spectrographs are always mounted.

The LVC alerts give approximate locations, distances, and estimates of the nature of the merger: BNS, BBH, or neutron star–black hole (NSBH). The LIGHETR program was triggered if an appropriate event, BNS or NSBH, fell within the R.A. and decl. attainable with HET at that epoch and promised a reasonably bright OT.

Details of how our program was implemented are given below. A proposed timetable for the LIGHETR response in O3 is given in Appendix C

2.1. The Instruments

HET¹⁷ is equipped with an 11 m spherical primary mirror made of hexagonal segments fixed at a zenith angle of 35° (Hill et al. 2021). It can be moved in azimuth, in such way that over a day, it can cover 70% of the sky visible at McDonald Observatory. The pupil, which is 10.0 m in diameter at the

¹⁷ HET is a collaboration of the University of Texas at Austin, Pennsylvania State University, Stanford University, Georg-August-Universität, Göttingen, and Ludwig-Maximilians-Universität, Munich.

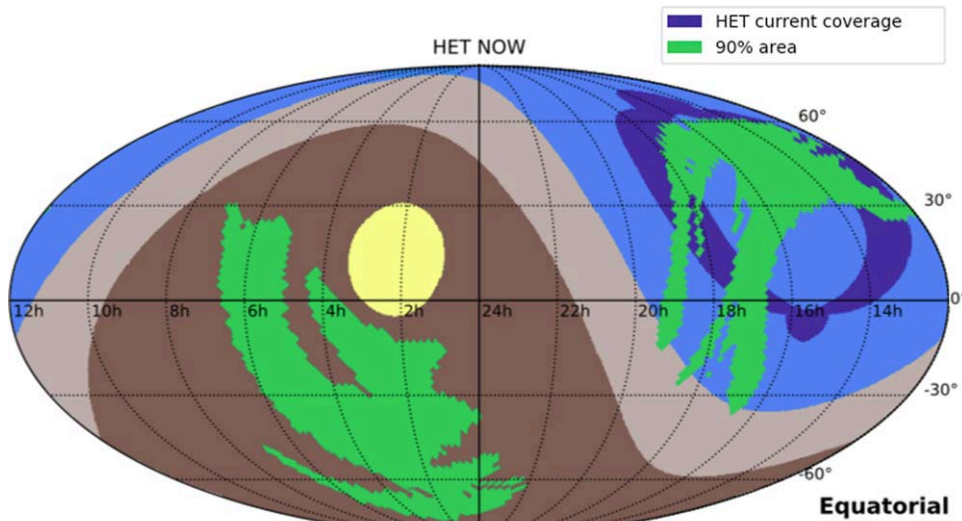


Figure 1. The sky map produced by using LVC localization data for the merger event S190425z and the HET track corresponding to the epoch of the alert. The LVC location of this alert is given by the green bands. The region accessible to the HET above decl. = -10° and below $+71^\circ$ is given by the dark blue band on the right. The lighter blue represents air mass less than 2.5. The location within 18° of the Sun is given by the large yellow patch. The lighter brown area is for air mass greater than 2.5 and the dark brown region on the left represents the portion of the sky below the horizon. The HET track overlapped with the large LVC sky map at the time of this alert. This figure provides a visual aid for observers and illustrates the technique by which we decided whether to trigger on a given LVC alert and to construct a prioritized list of galaxies to be searched for a new OT.

center of the track, can be moved to follow an object for different lengths of time depending on the decl. This tracking varies from 40 minutes at the lowest point (at $\delta = -10^\circ.3$) to 2.8 hr at the highest (at $\delta = +71^\circ.6$) with a maximum of 5 hr at 65° .

Mounted on the HET are two low-resolution integral field unit (IFU) instruments: VIRUS and LRS2. VIRUS is made up of 156 spectrographs, each fed by a fiber-integral unit, having an overall total fill factor of $1/4.6$. Each unit consists of 248 fibers with an individual $1/3$ fill factor, filled out by a three-point dither pattern, moving approximately $1''.5$ between each position. While this dither fills in the gaps between fibers in each IFU, it does not fill in the gaps between IFUs. Each fiber has a diameter of $1''.5$ and each spectrograph covers an area of $50'' \times 50''$. The VIRUS wavelength coverage goes deeply in the blue, ranging from $350 < \lambda(\text{nm}) < 500$ at a resolving power of $R = \lambda/\delta\lambda \approx 700$. LRS2, on the other hand, has a much smaller FoV of $12'' \times 6''$, with a higher resolution and broader wavelength coverage, spread between two independent fiber-fed dual-channel spectrographs. The blue spectrograph, LRS2-B, with $R = \lambda/\delta\lambda$ of 1900 and 1100 (for the respective channels) covering $370 < \lambda(\text{nm}) < 700$ and the red spectrograph, LRS2-R, with $R = \lambda/\delta\lambda$ of 1800 covering $650 < \lambda(\text{nm}) < 1050$. The response functions of LRS2 will be presented in G. J. Hill et al. (2023, in preparation).

Typical observing conditions at HET have a median seeing of $1''.7$. Under such conditions, a baseline 20 minute observation with three dithered exposures of 360 s has a line sensitivity at 5σ of $\sim 6 \times 10^{-17} \text{ erg cm}^{-2} \text{ s}^{-1}$. The fill factor of VIRUS of $1/4.6$ limits its capacity to get several galaxies in one dither pattern. This is why we do targeted galaxy follow up, placing individual galaxies at the center of an individual IFU.

2.2. Alert

In O3, if automated vetting detected a potential merger event, LVC then sent an automated public GCN Preliminary

Alert within about 10 s to 1 minute of the detection. Each alert was subject to human vetting and either an Initial Alert or a retraction was sent on a timescale of tens of minutes to an hour. The GCN Preliminary Alert pointed to the LVC sky map that gave the 3D probability distribution of the source.

Both due to the rapidly varying nature of kilonovae and the fact that HET is a fixed zenith telescope, which gives it limited tracking capability, quick knowledge of when the region of highest probability for a given event would be visible required a custom alert system. With this in mind, we created DIAGNOSIS,¹⁸ a low-latency alert system and observation planner for HET. DIAGNOSIS takes into account the geometry of the HET pupil and its position at any given point in time to submit an observation request to the HET resident astronomers and to alert the LIGHETR team.

DIAGNOSIS continuously listened to the Gamma-ray Coordinates Network/Transient Astronomy Network (GCN/TAN)¹⁹ for alerts on gravitational-wave events. When triggered, if the event was likely to have an electromagnetic counterpart,²⁰ DIAGNOSIS downloaded the associated sky map, identified if and when the 90% confidence region fell within the HET pupil, and if so, informed the observers. As an example, Figure 1 shows the sky map produced by DIAGNOSIS for the merger event S190425z, which turned out to be the second confident BNS detection by the LVC (Coughlin et al. 2019; Antier et al. 2020). DIAGNOSIS also queried a galaxy catalog for galaxies within the observable 90% probability region, organized them by probability, gave their local sidereal times (LSTs) to start observations and made a submission file for HET observations. An example of a graphic representation of the notice sent to the observers is in Figure B1 in Appendix B.

¹⁸ <https://github.com/Majoburo/Diagnosis>

¹⁹ <https://gcn.gsfc.nasa.gov/>

²⁰ If it was likely to be a BNS, NSBH, or in the mass gap (one or more of the binary objects is in the range of 3 to 5 M_\odot). We did not pursue any BBH candidates in O3.

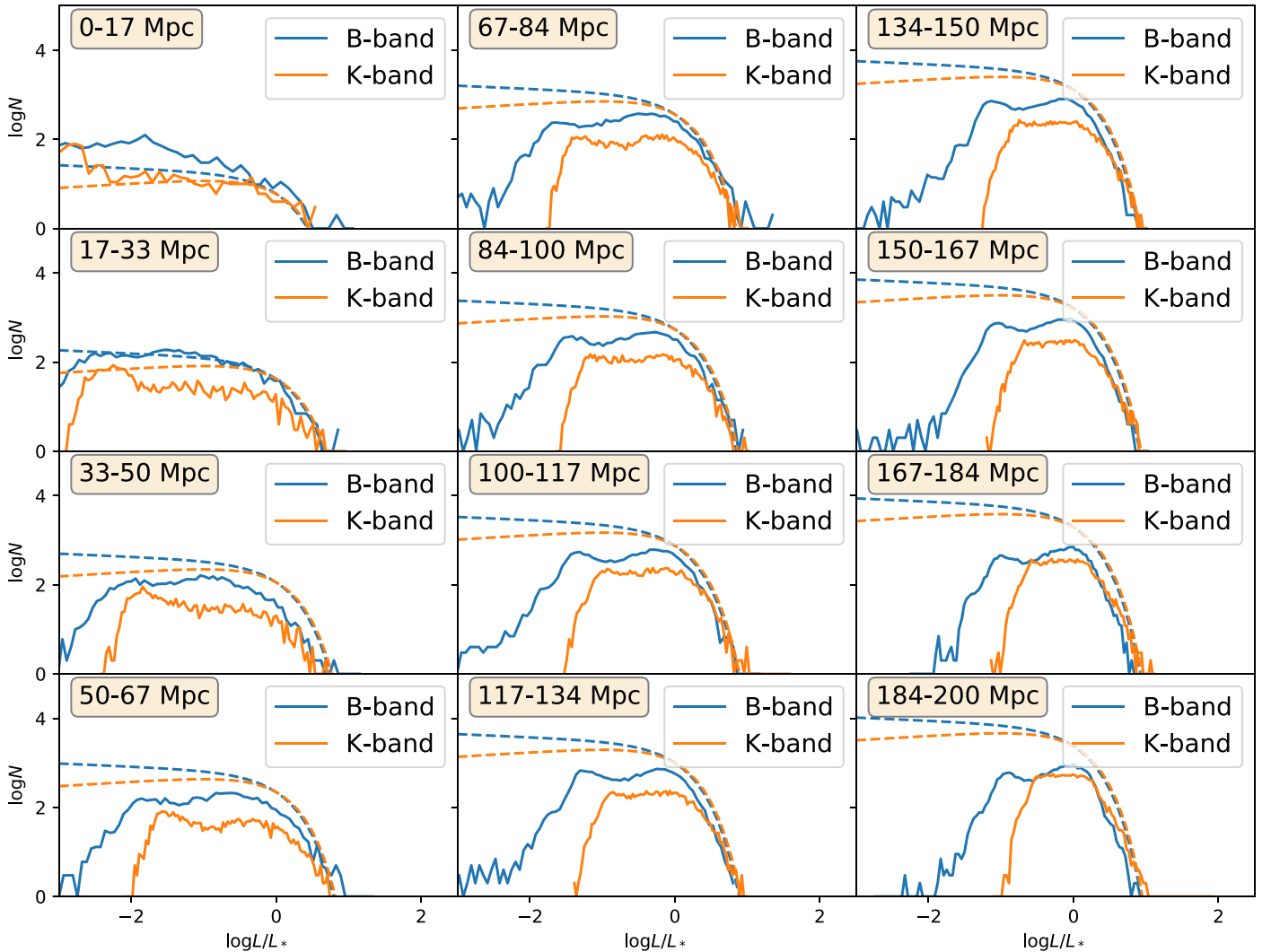


Figure 2. Completeness of the GLADE catalog in the B and K bands for different redshift ranges. The solid lines are luminosity histograms of GLADE galaxies within different distance shells in terms of their measured B -band and K -band luminosities. Dotted lines are our expectations for complete catalogs based on Schechter function measurements for each band. The GLADE catalog is designed to be more complete in the B band, but we use the K band since we aim to use the luminosity as an indicator of mass.

2.3. Galaxy Catalog

We used version 2.3 of the GLADE galaxy catalog (Dálya et al. 2018). The GLADE catalog contains around 3.26 million objects (149 globular clusters, 297,014 quasars, and 2,965,717 galaxies). It was constructed combining data from five astronomical catalogs: GWGC (White et al. 2011), 2MPZ (Bilicki et al. 2013), Two Micron All Sky Survey (2MASS) XSC (Skrutskie et al. 2006), HyperLEDA (Makarov et al. 2014), and the Sloan Digital Sky Survey (SDSS)-DR12Q (Pâris et al. 2017). The catalog is complete up to $d_L = 37^{+3}_{-4}$ Mpc in terms of cumulative B -band luminosity out to such distance, decreasing to 50% completeness at a d_L of 91 Mpc. Figure 2 shows the catalog's completeness in both the B and K bands as a function of luminosity for different redshift bins. The expected detection depth for LVC merger events in O3 was 100–140 Mpc. This substantially exceeds the completeness depth of the Glade catalog, but the catalog contains some galaxies extending to the LVC limit.

We constrained the selection of galaxies in the catalog to those visible by HET and those for which we could estimate their mass, that is, galaxies within declinations of $71^\circ 6$ to

$-10^\circ 3$ and with luminosity distances and K -band magnitude measurements. This cut our list of objects to 19% of the original list, leaving only galaxies. Figure 3 shows the number density in the sky of our selection of GLADE galaxies.

2.4. Galaxy Sorting

Gehrels et al. (2016) estimated that for bright galaxies contributing $\sim 50\%$ of the light in a given area of the sky, there will be ~ 20 galaxies inside a typical LVC error box that are also consistent with the LVC distance estimates. In order to rank the most probable galaxies, we weighted the localization probability given by the LVC (Singer et al. 2016) by the galaxy mass, which has been shown to be linearly proportional to the K -band luminosity (Kauffmann & Charlot 1998). Previous work took B -band luminosity as a weight factor on the galaxy-selection process (Arcavi et al. 2017; Yang et al. 2019). Even though the GLADE catalog we used is more complete in terms of galaxies with B -band luminosities, these luminosities have the problem of being affected by star formation history and dust extinction, which makes them an unreliable indicator for mass. Using the B band without such considerations would

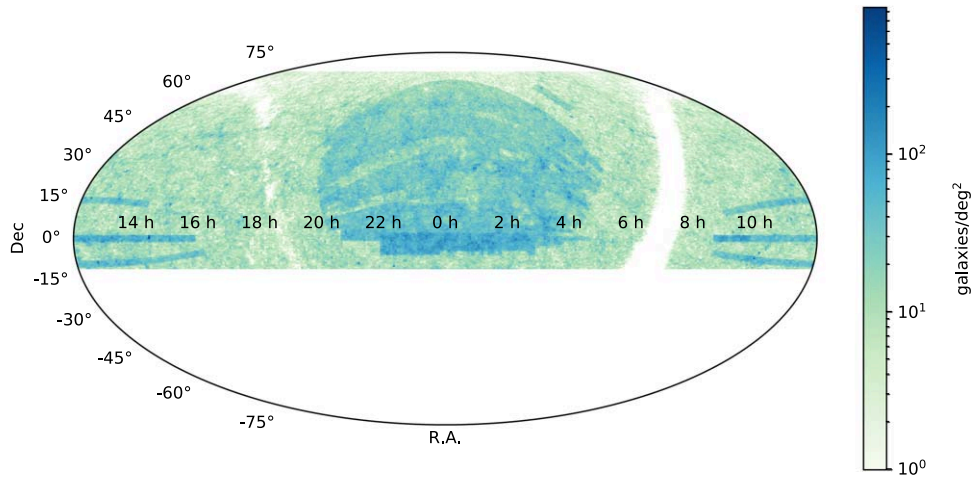


Figure 3. Galaxies contained within the GLADE catalog that are visible from HET.

effectively establish a preference toward star-forming galaxies. Both simulations and measurements have shown that short soft gamma-ray bursts (SGRBs) are produced by BNS mergers (Wiggins et al. 2018). GW170817 confirmed this expectation. A recent kilonova candidate was associated with a smaller galaxy (Troja et al. 2022; Yang et al. 2022) and SGRBs are found in a variety of galaxy types (O’Connor et al. 2022). Using a galaxy ranking based on K -band luminosity might aid in the identification of events like GW170817 but might discriminate other candidates in other host galaxy types. Despite significantly reducing the completeness of our galaxy sample by constraining it to those with K -band luminosities as shown in Figure 2, we chose galaxy mass and hence K -band luminosity as our weight factor in the localization probability.

The LVC provided a localization probability, ρ_i , as a 2D HEALpix sky map. For each pixel, i , within the sky map, the LVC also included Gaussian parameters (σ_i, μ_i) for the distance probability. Thus, we took the galaxies within our catalog falling in the 90% confidence region in terms of the 2D sky map, calculated their distance probability in terms of their cataloged luminosity distances, r_g , and multiplied this by their K -band luminosity, L_K^g , and 2D localization probability, ρ_i , to get their final ranking,

$$R(r_g, L_K^g) \propto L_K^g \rho_i \frac{1}{\sigma_i} \exp \left[-\frac{(r_g - \mu_i)^2}{2\sigma_i^2} \right] r_g^2. \quad (1)$$

This ranking was normalized to add to unity when summing over the most probable 100 galaxies.

2.5. Search for the OT

As described above, the probability distributions of the LVC sky maps and HET track maps illustrated in Figure 1 were automatically employed by DIAGNOSIS to query the GLADE catalog to produce a prioritized list of galaxies to search for the OT. Such galaxy lists were the basis for the VIRUS search. DIAGNOSIS also generated a Phase II target submission list (TSL) for the prioritized list of galaxies. The result was a prioritized list of galaxies in the HET overlap region and within the distance window provided in the LVC alert. We then systematically observed that sample of galaxies with the VIRUS array. Examples of the LIGO alert as processed by DIAGNOSIS for local redistribution, the galaxy priority list, and the selection of galaxy priorities are given in Appendix B.

Within O3, our observations all fell within the ~ 10 most likely galaxies according to our prioritized list.

Once we start the search, a key capacity is to reduce the data rapidly to decide whether there is an OT in a given galaxy or to continue to the next galaxy on the priority list. VIRUS observations always produce a zeroth-moment image as one of the first steps. The VIRUS IFU units do not physically touch one another. The wavelength region of 3500–5500 Å in a selected aperture is used to make a collapsed image, as illustrated in Figure 4. This initial collapsed image was compared to archival imaging data (specifically from Pan-STARRS), a step which could (but did not in O3) promptly reveal the OT.

During O3, VIRUS spectrograph units were steadily being added to the array. At the time, the array observed about 80% of its field that is 21' on a side (Figure 4). Each IFU unit covers an area of 50'' on a side. Normal VIRUS operation is to do a three-point dither to fill the fiber spaces. We employed this operation since it enhanced our ability to calibrate the spectrum.

For rapid data reduction, we used REMEDY v0.1, a stripped-down version of the LRS2 reduction code, PANACEA (Zeimann & Hagen 2021).²¹ The initial success of REMEDY encouraged an evolved version of the code to become the HET’s default pipeline for VIRUS observations. A full description of REMEDY can be found in G. Zeimann et al. (2023, in preparation). REMEDY has the capability of reducing the full frame or just the data from a dithered IFU exposure of the host galaxy and its vicinity. We then produced a collapsed image from the IFU and a reduced spectrum, all in 30 s.

We pointed a single VIRUS IFU at a target host galaxy while getting spectra of all the other galaxies in the VIRUS field simultaneously. Depending on the target density, circumstances might have allowed us to multiplex, acquiring multiple target galaxies in a single VIRUS field. We note that the merger event GW170817 was about 10'' away from its host galaxy. With the 50'' square VIRUS IFU FoV, had we pointed at the host, we would easily have picked up the OT and acquired its spectrum (Figure 4).

Each VIRUS set of three dithered exposures of a target galaxy required a total of ~ 15 minutes, including shuffling

²¹ Panacea v0.1 documentation can be found at https://github.com/grzeimann/Panacea/blob/master/README_v0.1.md.

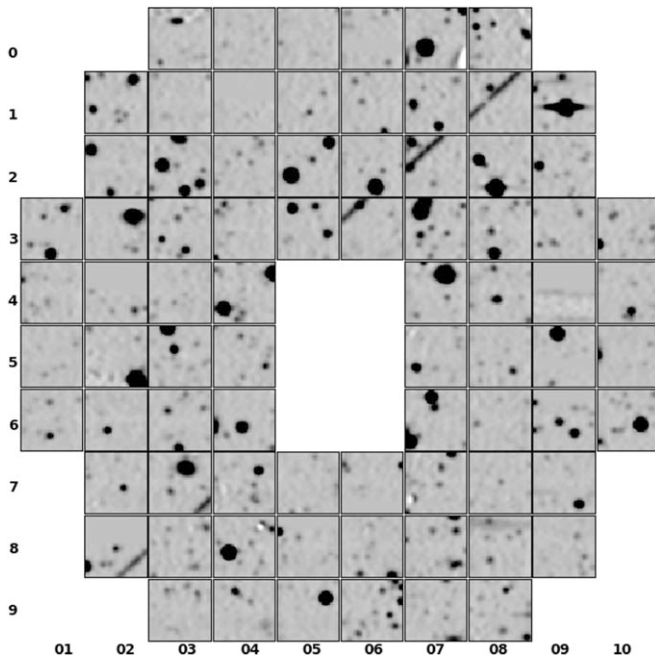


Figure 4. A sample VIRUS (Gebhardt et al. 2021; Hill et al. 2021) image that would be used in our proposed search for the OT associated with a gravitational-wave event. Each IFU unit covers an area of $50''$ on a side. The VIRUS array automatically produces a zeroth-moment image using the wavelength region $3500\text{--}5500\text{ \AA}$ and an adopted aperture to make a collapsed image. The whole array is $21'$ on a side. This image butts the IFUs together, preserving their relative position. Each of the 20,000 fibers is positioned to an accuracy of $0.1''$. The faintest objects in the image are $g \sim 22\text{--}23$. The streak in the upper right is a satellite or asteroid artifact.

between galaxies. The total time to observe 20 galaxies was thus ~ 5 hr. We were not able to reach this limit in O3 because our targets only became accessible to HET a few hours before sunrise.

2.6. Dense Spectral Follow Up

Once an OT candidate was identified by our search or others, we planned to implement follow-up spectroscopy with LRS2. The spectra of a kilonova are expected at first to change on timescales of hours and later on timescales of days. We could, in principle, take spectra for about a week's span with as rapid a cadence as the HET would allow using all the viewable time.

Table 1 presents a short summary of track lengths and number of spectra expected as a function of the brightness and decl. of the target. We assume that the total setup time is 10 minutes for a single track following a single object. The actual setup time is often less than 10 minutes, and there is very little setup time, perhaps one or two minutes, to switch from LRS2-B and LRS2-R if both observe the same object. For an OT as bright as AT2017gfo, ~ 17.5 mag in the i band at discovery, we assume 20 minute exposures for each early observation and hence that a single spectrum would require $10 + 20 = 30$ minutes and that a full LRS2-B and LRS2-R spectrum requires $10 + 20 + 20 = 50$ minutes. For long tracks allowing multiple exposures of the same object, the total exposure would be 10 minutes plus the number of 20 minute intervals that sum to less than the track length. If we elect to get only LRS2-B in the early phases, then we would get roughly twice as many spectra per track than obtaining LRS2-B and LRS2-R spectra, with the caveat that we could only get an even number of B+R spectra utilizing a single track.

Thus, in an ideal northern decl. case of $\sim 70^\circ$ with a track length of ~ 156 minutes and target of brightness ~ 17.5 mag, we can expect no more than overheads plus seven consecutive 20 minute shots per track. We could get three full LRS2-B+LRS2-R spectra per track. If the object is at a lower decl., $\sim 20^\circ$, then we can still get three single spectra or two full spectra per track.

Events discovered in O3 were on the average more distant and fainter than AT2017gfo, requiring longer exposures. AT2017gfo was at only ~ 40 Mpc. LIGO–Virgo–KAGRA (LVK) will reach to $160\text{--}190$ Mpc in O4, implying a factor of $16\text{--}23$ in luminosity and $3\text{--}3.4$ mag in brightness. A target in this distance range would be about $r = 20.5\text{--}21$, requiring an exposure of about an hour on the HET. A single spectrum including overheads would require 70 minutes. To get both LRS2-B and LRS2-R would require 130 minutes. At this brightness at a decl. of $\sim 70^\circ$ with a track length of ~ 156 min and a target of brightness ~ 21 mag, we might acquire two single spectra per track and one complete LRS2-B+LRS2-R spectrum per track. At decl. $\sim 20^\circ$ and $\sim 45^\circ$, we could expect only a single spectrum per track.

If we assume the target is visible throughout the whole night, once on an east track and once on a west track, a somewhat unlikely circumstance in practice, then we could get about twice as many spectra per night. For a decl. of $\sim 70^\circ$ at ~ 17.5 mag, we could get 14 single spectra per night or six complete LRS2-B+LRS2-R spectra. At lower declinations of $\sim 20^\circ$ and $\sim 45^\circ$ and ~ 21 mag, we would need this special circumstance to acquire LRS2-B on one track and LRS2-R on the other to get even a single complete LRS2-B+LRS2-R spectrum per night.

While somewhat optimistic, this cadence of spectra would match or considerably exceed anything achieved for AT2017gfo. At later epochs, the OT will be dimmer, also calling for longer exposures, but the cadence can be more relaxed. A timetable for alert, search, dense spectral sampling, and analysis is given in Appendix B.

3. Results

Details on the number of exposures on all the events we followed in O3 are summarized in Table 2.

With all the elements in place for a successful program, we recognized that this is a complex process that had never been employed on the HET. For this reason, we did a dry run during the day with a practice alert trigger.

No event in O3 revealed an OT to any facility, so we did not employ LRS2 (but see Section 3.4).

3.1. Dress Rehearsal: GW190412m

To complement the dry run, we also planned a full dress rehearsal initial run. For this, we proposed to trigger on the first LVC alert of O3 that was accessible to the HET. The probability was that this first alert would be a BBH merger event and hence most likely not accompanied by an electromagnetic signal. There was a small possibility that the first alert of O3 might correspond to a BNS or NSBH merger and an associated OT. Theoretical modeling and observations hinted that a BBH merger within the gas-dense environment of a supermassive black hole accretion disk might generate an OT (Cheng & Wang 1999; Bartos et al. 2017; Graham et al. 2020; Perna et al. 2022). In any case, responding to this alert would give us critical feedback that all our interlocking components,

Table 1
Estimated Number of Single Track Spectra Per Night as a Function of Decl. and OT Magnitude

Decl. (degrees)	Track Length (minutes)	Magnitude (mag)	Exposure Time (minutes)	Number of Single Spectra	Number of B+R Spectra
20	72	17.5	20	3	1
45	96	17.5	20	4	2
70	156	17.5	20	7	3
20	72	21.0	60	1	0
45	96	21.0	60	1	0
70	156	21.0	60	2	1

hardware, software, and human, would function as needed for a merger with an OT.

For our dress rehearsal, we responded to BBH candidate GW5190412m, only the second event of O3 and the first in range of the HET. The LVC alert for GW190412m was ingested by our DIAGNOSIS pipeline, which created a prioritized list of target galaxies, a corresponding TSL, and alerted our team. We were awakened at 4 am local Texas time and were on the sky and taking data within about 20 minutes. We observed seven of the top priority list of 13 target galaxies in the overlapped LVC/HET field. We drafted and sent a GCN notice about our observations in real time and submitted it as the Sun rose on the observatory (Rosell et al. 2019a). In hindsight, the LVC alert had been issued about two hours earlier than our local alert. This did not impede our observations in practice, and the problem leading to the local delay was addressed and eliminated. We also found that a vestigial instruction in the TSL requested shorter exposures than ideal. This allowed us to make more observations, but with smaller signal-to-noise ratios (S/Ns) than preferable. We corrected this in subsequent observations. Despite the expected absence of an OT for S5190412m, the dress rehearsal served its purpose.

3.2. GW190425z

The next event with the potential to produce an optical signal was GW190425z with a probability greater than 99% of being a BNS merger. This was the sort of event for which our program was designed. The local alert was released at 4:01 am local time. Our target list of high priority galaxies was loaded into the HET queue at 4:07 am. Unfortunately, the LVC sky map was very large since only LVC Livingston and Virgo were operating. The distance was estimated by LVC to be 155 ± 44 Mpc, much more distant than GW170817 at 40 ± 8 Mpc. Pan-STARRS images of our target galaxies were shared at 4:48 am. We sampled a prioritized list of four galaxies from the GLADE catalog that overlapped with the LIGO probability map and the observable pupil of the HET. The effective limiting magnitude in the *B* band was ~ 22 mag. A GCN was submitted at 6:51 am (Rosell et al. 2019b). We did not detect an OT, and neither did any other team. Some transient candidates were reported, but they all proved to be more prosaic, mostly supernovae. We elected not to continue our search of the LVC sky map on a second night. LANL team members (C.F., O.K., and R.W.) worked through this period making models corresponding to the LVC distance (see Section 5) and even of some of the transient candidates. LIGHETR group member J.V. and collaborators observed the five galaxies targeted by HET at Konkoly Observatory. CCD frames were taken with the 0.6/0.9 m Schmidt telescope (FoV

70×70 arcmin², unfiltered, limiting mag 21.5) and the 0.8 m RC telescope (FoV 18×18 arcmin², *g* and *r* band, limiting *r* magnitude ~ 20.4).

3.3. GW190426c

The LVC announced an alert on GW190426c at 10:47 am local time the next morning. This event had a 49% chance of being a BNS but a 13% chance of being a NSBH merger and a 14% chance of being noise. The estimated distance was very large, 375 ± 108 Mpc. The LANL group produced new light-curve models based on the estimated distance. Half the HET track was lost to solar occlusion. Target galaxies in the usable track became observable at about 4:30 am. After an extensive discussion of the distance, the expected S/N ratio, the Moon location, the expected rarity of targets, the competition from other groups, and the likelihood of tentative (but ultimately irrelevant) OT candidates, we elected to do our VIRUS search on this event. We sampled a prioritized list of five galaxies from the GLADE catalog and submitted a GCN at 6:31 am (Rosell et al. 2019c).

3.4. S190510g

The next potentially interesting LIGO event, S190510g, had a probability of 42% of being a BNS but 48% of being noise. The distance was also estimated to be rather large, 280 Mpc. After some debate, we passed on this event.

3.5. GW190901ap

There were no other alerts that required a response from us until GW190901ap, which had an 86% chance of being a BNS merger. We observed one target galaxy with VIRUS and sent a GCN (Rosell et al. 2019d). The next morning, the Zwicky Transient Facility (ZTF) announced four possible OT candidates consistent with the redshift estimated by LIGO (Kool et al. 2019). Three were ruled out by other groups.

We elected to observe the fourth ZTF candidate, ZTF19abvionh, with VIRUS on 2019 September 2. We detected a nearly featureless continuum with narrow emission lines corresponding to [O II] $\lambda 3727$ and $\lambda 5370$ and H β at a redshift of 0.1 from a galaxy near to ZTF19abvionh on the sky. That galaxy, GALEXASC J165500.03+140301.3, was about 2.5σ more distant than the estimated distance of the merger candidate. We submitted a GCN (Rosell et al. 2019e) at 12:05 am and another (Rosell et al. 2019f) at 12:39 am after estimating a blackbody temperature of about 10,500 K.

The next night, 2019 September 3, we observed ZTF19abvionh with both VIRUS and LRS2-B. These data showed broadened emission features at about 4686 and 6560 Å as shown in Figure 5 that were compatible with He II $\lambda 4686$ and

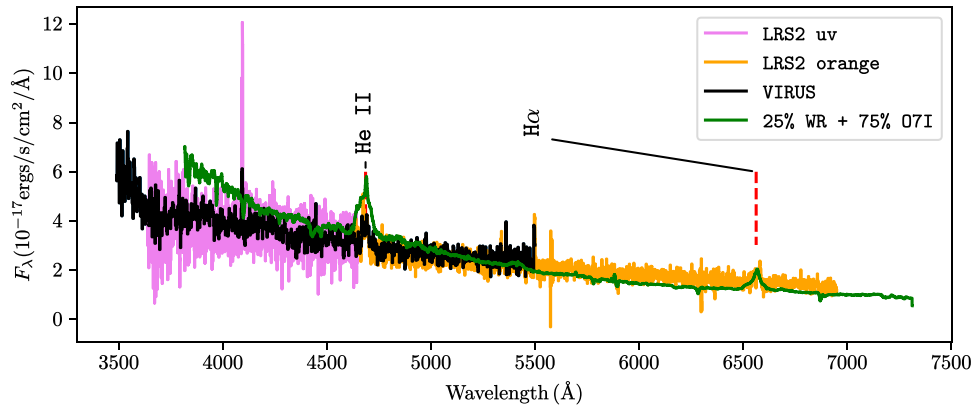


Figure 5. Spectra of ZTF19abvionh. The narrow [O II] $\lambda 3727$ is from a galaxy at redshift ~ 0.1 that is near to ZTF19abvionh on the sky, but unrelated. The figure shows all the data from VIRUS and LRS2-B overlaid with a spectrum of a WN type Wolf-Rayet star (25%) plus an O7I star (75%) at zero redshift. The WR components corresponding to [He II] $\lambda 4686$ and $H\alpha$ align with features in the spectra of ZTF19abvionh.

Table 2
Summary of LIGHETR Observations.

Date (GMT)	Event	Exposures	Instruments	Searched	Followed up
19/04/12	S190412m	21	VIRUS	7	0
19/04/25	S190425z	18	VIRUS	6	0
19/04/27	S190426c	15	VIRUS	5	0
19/08/22	S190822c ^a	9	VIRUS	3	0
19/08/29	S190829u ^a	6	VIRUS	2	0
19/09/03	S190901ap	3	VIRUS	1	0
19/09/03	S190901ap	3	VIRUS	0	1
19/09/06	S190901ap	5	VIRUS, LRS2	0	1

Notes. Summary of all events we followed. Retracted events are shown with an “a”. (1) Date searched; (2) GCN/TAN event name; (3) number of exposures; (4) number of searched galaxies; and (5) number of external EM triggers followed.

^a Retracted event.

$H\alpha$ (or maybe He II $\lambda 6560$) at a negligible redshift. This spectrum is compatible with a Wolf-Rayet star of type WN at ~ 1 Mpc. Since we saw no host galaxy, the host would have to be a previously unknown dwarf galaxy (e.g., McQuinn et al. 2020) accidentally along the line of sight, but unrelated to GALEXASC J165500.03+140301.3.

We checked an alternative hypothesis that ZTF19abvionh was an ultraluminous X-ray source (ULX) that can have similar spectra, but could be at much greater distances, ~ 10 Mpc, by requesting a Swift observation. Swift/XRT observed the field on 2019 September 20 with an exposure of 3921 s and on 2019 September 23 with an exposure of 435 s. The source was not detected in either observation. We combined the two exposures and used a source region of $45''$ radius centered on 16:55:00.212, +14:03:04.67. In this region, there are six counts in the 0.2–10 keV. In a nearby source-free region of $5'$ radius, there are 251 counts in 0.2–10 keV, for an average of 0.000888 background counts per square arcsecond, or 5.65 in our source region. Using the Bayesian method of Kraft et al. (1991), we calculated a 90% confidence upper limit of 5.67 counts, for a count rate of 0.0013 c s^{-1} . We assumed an absorbed power-law spectrum with photon index of two and column density of $n_H = 4.81 \times 10^{20} \text{ cm}^{-2}$ (HI4PI Collaboration et al. 2016). We used WebPIMMS to calculate a 90% upper limit of $5.6 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$ in the 0.2–10 keV band.

The lack of detection could be because the source is a WN star at 1 Mpc, but could also be because ULXs are variable, and

the source had evolved to an X-ray minimum. ZTF19abvionh then underwent solar occlusion. We submitted a third GCN specifically concerning ZTF19abvionh (Rosell et al. 2019g).

4. Quantitative Assessment of Null Results

While a careful visual inspection of all the data cubes obtained by the LIGHETR program has been performed and no obvious electromagnetic counterpart has been found, the need for a quantitative error on our assessment is still necessary. For this purpose, we performed two different subsequent sets of analyses on our data, one using the standard continuum grid-search and point-source extraction routines provided by HETDEX (described in full in Gebhardt et al. 2021) and one using an adaptation of SCARLET (Melchior et al. 2018), a package that performs source separation on multiband images designed for the Legacy Survey of Space and Time (LSST) Science Pipeline. Our preliminary investigation of both methods highlights the challenges of making quantitative estimates to our sensitivity to OTs. We now discuss these analyses and provide a prescription for a path forward.

4.1. Point-source Detection and Extraction via HETDEX’s Standard Pipeline

Most of LIGHETR’s target fields are centered, by the nature of the search, on extended sources (galaxies), with the occasional centering on an externally detected transient. Although HETDEX’s tools are not designed to deal with

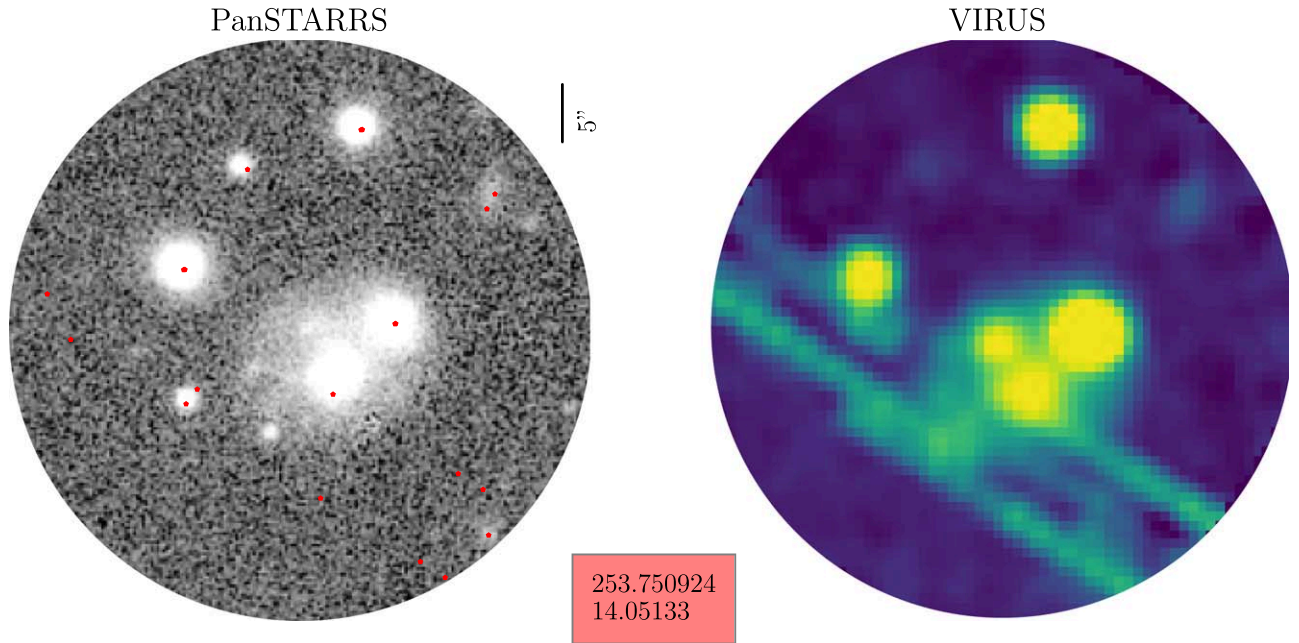


Figure 6. Target galaxy GALEXASC J165500.03+140301.3. On the left is Pan-STARRS imaging of the field overlaid with Pan-STARRS cataloged sources (blue) and VIRUS point sources (red). On the right is the collapsed spectral image of a VIRUS shot of the same field cropped around a $25''$ radius. The streaks across the IFU field correspond to CCD artifacts at the time of the observations and which affected point-source identification. These artifacts are primarily hot pixels and charge traps that have been subsequently identified and removed from the latest reductions.

extended sources and deblending, they do routinely perform point-source detection and extraction. Under optimal conditions, they can detect point sources up to magnitude 23 in the g band (Gebhardt et al. 2021). While the one data point at hand, AT2017gfo, did appear as a point source in the outskirts of its host galaxy, uncertainties in the intensity of supernova kicks and the orbital energy dissipated in the explosion debris leave unclear whether kilonovae are to be preferentially found in the outskirts of galaxies. Even if that were to be the case, it is likely that many will be placed within the line of sight of the continuum of the galaxy. With this caveat in mind, we briefly present HETDEX’s standard continuum grid-search and point-source extraction techniques and the results of applying them to our data set to assess the existence or absence of a transient within our fibers.

For the continuum search, spectra from a VIRUS data cube are flagged as possible continuum sources if they have at least $0.5 \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$ either in the blue (from 3700 to 3900 \AA) or in the red (from 5100 to 5300 \AA). Each possible continuum source then gets searched around a $1''.5 \times 1''.5$ grid of $0''.1 \times 0''.1$ resolution, and the spatial location with the maximum flux is selected as the location of the source. Around each location, a point-spread function (PSF) extraction of $3''$ radius in aperture is subsequently performed. The particular aperture and specifics of the weights that go into the PSF are determined both by typical seeing conditions and known systematics in the design of the instrument (particularly, regarding lack of an atmospheric dispersion corrector and the fiber dithering pattern by which we fill in our collected area). With a set of possible continuum sources in hand for the fields around all of our targets, we compare them with archival sources from the Pan-STARRS catalog.

Figure 6 shows, for reference, one of our target galaxies, GALEXASC J165500.03+140301.3. As can be seen in the image, the CCDs had artifacts across a large fraction of the

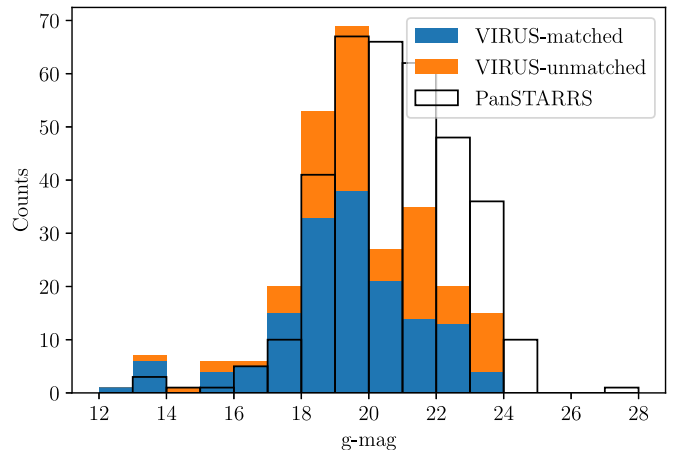


Figure 7. Point-source magnitudes for all of our target fields. Transparent bins correspond to Pan-STARRS cataloged sources. Blue and orange are stacked and correspond, respectively, to VIRUS sources matched with Pan-STARRS or only detected by VIRUS. The point-source identification algorithm fails with the brightest targets, where saturation leads it to identify more than one source per target. For lower brightnesses, the mismatch can be explained by spurious continuum point-source identification both from Pan-STARRS and VIRUS.

IFU’s footprint. These artifacts made for spurious point-source detections and hid some cataloged sources from detection. The problem has since been fixed with improved reductions. Extended sources like galaxies or effectively extended sources like bright stars were also tagged as more than one source, due to oversaturation and to the point-source algorithm failing to identify them as single sources. Figure 7 shows a binned histogram of our point-source detections for all of our targets. Point sources detected by VIRUS that are less than $2''$ from a Pan-STARRS source, with a g -band magnitude difference of less than 1 magnitude, are considered coincident. The histogram shows overidentification of bright sources and underidentification of low-brightness ones. Visual inspection

of our targets shows this is mostly due to saturation and spurious point-source identification in extended sources both from Pan-STARRS and VIRUS. While we expected the identification of point sources within the continuum of the galaxy to have significant difficulties, the large number of misidentifications outside of that continuum both for bright and dim sources makes it impossible for us to place limiting magnitudes on point sources for the fields observed. We thus turn to an altogether different detection algorithm for our search.

4.2. Nonparametric Extended Source Identification with SCARLET

Many of the difficulties of our search with VIRUS data cubes and the standard pipeline are due to the intrinsic spatial unevenness of a dithering IFU. The fibers fall in different places of the CCD and have inherent flux variations that can be as large as a factor of ~ 2 . While REMEDY does a superb job on the weighting of this variation, software can only do so much and will never be comparable with direct imaging in terms of spatial evenness. The power of spectral data cubes is in that extra spectral dimension, and a tool that fully uses this advantage should be employed.

It is in the context of these difficulties that SCARLET, a software tool designed for multiband source separation for the LSST Science Pipeline, appears to be a promising alternative (Melchior et al. 2018). SCARLET makes use of a constrained matrix factorization in which each source is modeled with a spectral energy distribution (a spectrum) and a nonparametric morphology (the equivalent of a PSF). The code also allows the imposition of priors on the shape of each source, allowing further constraints. SCARLET seems particularly suitable for our problem. The algorithm can make use of the spectral information to identify neighboring fibers contributing to the same flux, so systematics in the PSF extraction could be significantly mitigated. In addition, artifacts could also be removed from the CCD, since again the bleeding into adjacent pixels should show similar spectral features.

Figure 8 shows an individual source from a VIRUS observation extracted by SCARLET. The package treats point sources, on which a symmetry prior is enforced, or multiple overlapping sources with nonparametric morphologies. Constraints on the shape of the spectra can also be enforced. The ability to use analytic priors to fit against, say, a blackbody spectrum, were implemented recently in SCARLET due to a suggestion from M.J.R.-B. One of the major limitations on the immediate use of SCARLET for our purposes is that it requires approximate positions for the sources to be given beforehand. As shown in Figure 6, in its current state the standard continuum grid search when applied to our data suffers from too high a degree of mismatch to be usable as a source of positions.

Given these difficulties, our final approach was to simulate a set of point sources with a blackbody spectrum and place them randomly within a VIRUS data cube to assess under what conditions could they be recovered. We performed a maximum likelihood analysis fit on the data cube, with which we matched a PSF-weighted 3D source (2D for space, 1D for spectrum) to all positions in the data cube that hosted the injected sources. We observed that we could only detect sources down to a few magnitudes brighter than the nominal value quoted for HETDEX’s pipeline, but we suspect flux calibration issues

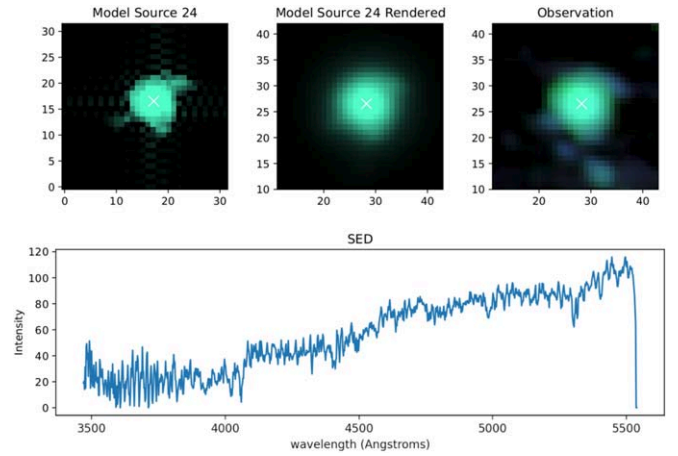


Figure 8. Preliminary results of an adaptation of SCARLET to work on VIRUS spectral cubes. Images are on the same scale but with inherently different resolutions and displayed in pixel space. Upper left shows one of the extended sources modeled by SCARLET. Upper middle shows the source, rendered to the given PSF. Upper right shows the actual observation compressed along the spectrum. The colors have been mapped to the corresponding wavelengths of maximum intensity. In the bottom is a picture of the overall collected spectrum. The results are promising but need to be combined with a customized continuum grid-search routine to be usable.

were at the heart of this difference. Proper knowledge of the PSF weights accessible to HETDEX’s point-source extraction algorithms will be of extreme value for this method to be usable, and further investigations are ongoing. A combination of the continuum grid search and this method might also be of great value, since the overidentification of continuum sources could be mended if the nonparametric PSF of SCARLET engulfs them all.

5. Models Guiding Observations and Motivation for Spectral Follow Up

Observations of GW170817/AT2017gfo closely confirmed expectations of theoretical models. The merging NSs create a tidal tail of material ripped from their surfaces and ejected in the orbital plane. This material is expected to be rich in lanthanides, with very large opacity, strong absorption in the blue, and hence characterized by a red continuum (Kasen et al. 2017; Tanaka et al. 2017). Orthogonal to the orbital plane, ejecta are expected to be expelled at about $0.3c$ in a jet or cocoon of somewhat lower-opacity material radiating in the blue. AT2017gfo first showed a blue continuum evolving over the course of a few days into a red continuum (McCully et al. 2017; Nicholl et al. 2017; Pian et al. 2017; Shappee et al. 2017; Buckley et al. 2018). This was interpreted as first seeing the lower-opacity material ejected along the orbital axis and then the high-opacity, lanthanide-dominated matter from the orbital plane. Even if other events are very similar, the relative proportions of these components may be different and the orientation will surely be different, so spectral observations promise a rich new haul of insight. An NSBH merger event will be quantitatively and perhaps even qualitatively different (Bhattacharya et al. 2019; Desai et al. 2019; Darbha et al. 2021; Zhu et al. 2021).

Since we expect any lines to be broadened by the fraction of the speed of light velocities at which the ejecta expand, there could be shallow, broad features in the early spectra as well as fast, transient emission lines. There is a slim chance that one

portion of the ejecta, a jet, could irradiate another portion of the ejecta, a tidal stream, and yield a photoionized region that would briefly produce emission lines. A fortuitous orientation might mean that broadening is minimized, but even the transverse Doppler shift might be an issue in broadening and obscuring lines. Models of kilonovae show that blue emission is sensitive to the radii of the NSs, and that particular aspect angles can show broad yet distinct spectral features as well as a very strong dependence of the flux on angle in the optical band. Early spectral observations in the blue, like the ones LIGHETR is particularly designed to perform, promise strong constraints on the nature of the ejecta.

The LIGHETR program sought to integrate modeling constraints by coordinating with the LANL group simulating merger models, especially their spectra. The observational properties of kilonovae are subject to many variables, including the morphology of winds, disks, and jets, the composition on a given line of sight, and, crucially, the aspect angle. To illustrate how early and densely sampled spectra can serve to constrain these variables, we leverage a large, and continually growing, LANL database of kilonova spectra and light-curve calculations. To demonstrate the importance of the spectra, we present the results from two separate studies: a morphology study (Korobkin et al. 2021) and a new study designed for this paper varying the composition based on yields from detailed postmerger disk calculations (Miller et al. 2020). The ultimate goal is to be able to query a database of models in real time as spectra of candidate OTs are obtained.

Expected kilonova emission is subject to a broad range of uncertainties, both in the ejecta properties (that depend on aspects of both merger and disk-wind calculations) and in the physics and its numerical implementation. Examples of these uncertainties span all facets of kilonova emission modeling efforts: e.g., the amount of matter that is ejected in the dynamical phase of the explosion and distributed in a disk still depends sensitively on the numerics and physics (Henkel et al. 2022) and on the composition; the ejecta angular and velocity distributions are sensitive to physics modeling such as neutrino transport (Miller et al. 2020); and the emission depends on the nuclear network simulations and the model for energy deposition (Barnes et al. 2021). The model database is built on a large suite of models using the the SuperNu code (Wollaeger et al. 2013; van Rossum et al. 2016), adapting initial conditions to study a range of uncertainties in the ejecta composition. This work includes a suite of models varying the ejecta mass (Wollaeger et al. 2021), composition (Wollaeger et al. 2018; Even et al. 2020), morphology (Wollaeger et al. 2018; Korobkin et al. 2021), energy sources (Wollaeger et al. 2019), and atomic (Fontes et al. 2020) and nuclear (Zhu et al. 2018) physics.²²

The significant modeling uncertainties currently make it difficult to identify quickly and unambiguously which transients in the LVC localization maps to follow up with LIGHETR observations. We hope to use our growing suite of models to help identify the transients to follow up with LIGHETR and then use LIGHETR observations to constrain further the properties of the kilonova ejecta. Figure 9 shows *g*- and *r*-band light curves from models produced by Korobkin et al. (2021). This work varied the morphology, mass, and velocity of a “wind” (modest electron fraction) and

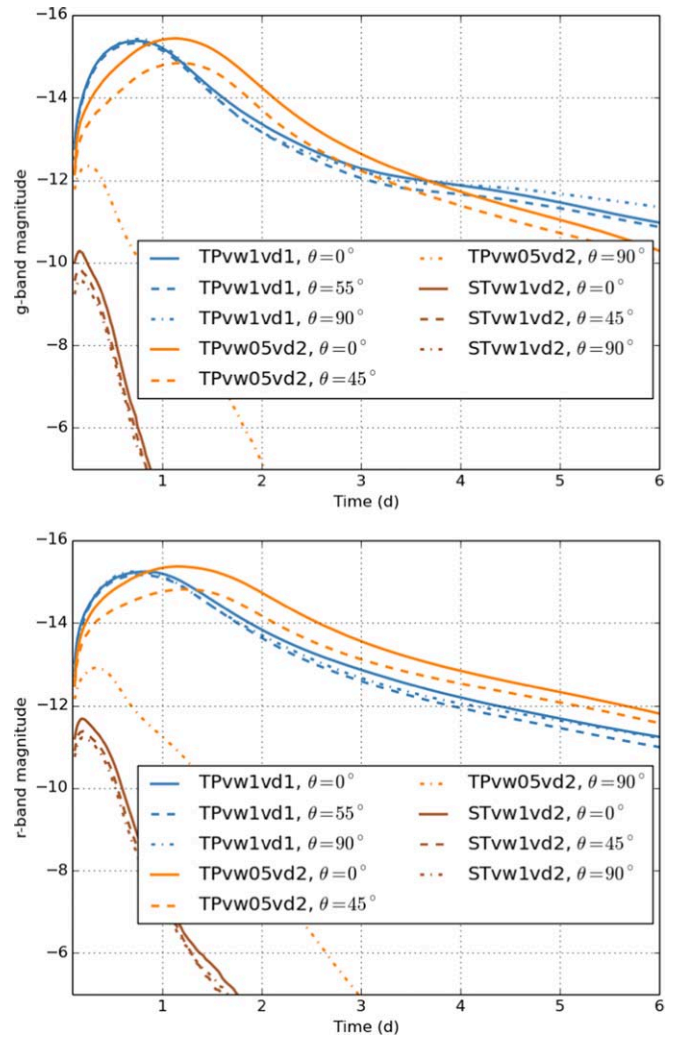


Figure 9. *g*-band (top) and *r*-band (bottom) magnitudes of three different two-component models with varying morphologies and average ejecta velocities (see Korobkin et al. 2021 for details). Despite the fact that the ejecta masses are the same, the different morphologies and velocities of the components produce different light curves. Because of the complex morphologies, the light curves can also vary with viewing angle.

“dynamical” (neutron rich) component for the ejecta. The morphologies follow the nomenclature described in the paper including a toroidal (“T”), peanut-shaped (“P”), and spherical (“S”) morphology. In this figure, we limit ourselves to models with $0.01 M_{\odot}$ in each component, but we vary the morphology and average ejecta velocity of each component. All of the models are characterized by a fast rise (fraction of a day) and decay time (from a fraction of a day to 2 days) for the light curves.

With 2–3 days of observations, fast evolution in the optical would be a strong indicator of a kilonova observation. To guide follow-up observations, we need observational indicators already on day one. In an attempt to differentiate these models quickly, we are investigating a range of discriminating light-curve features. For example, in Figure 10 we calculate the change in magnitude with the time of the models presented in Figure 9. Most of our light curves are characterized by an initial rapid rise followed by a slower decay. The rate of this decay depends on the model. In most cases, the rapid evolution all occurs within the first day. After the first 6 hr, some models are already decaying. Others continue to rise for 12–24 hr. In all

²² Many of these models are available at https://ccsweb.lanl.gov/astro/transient/transients_astro.html.

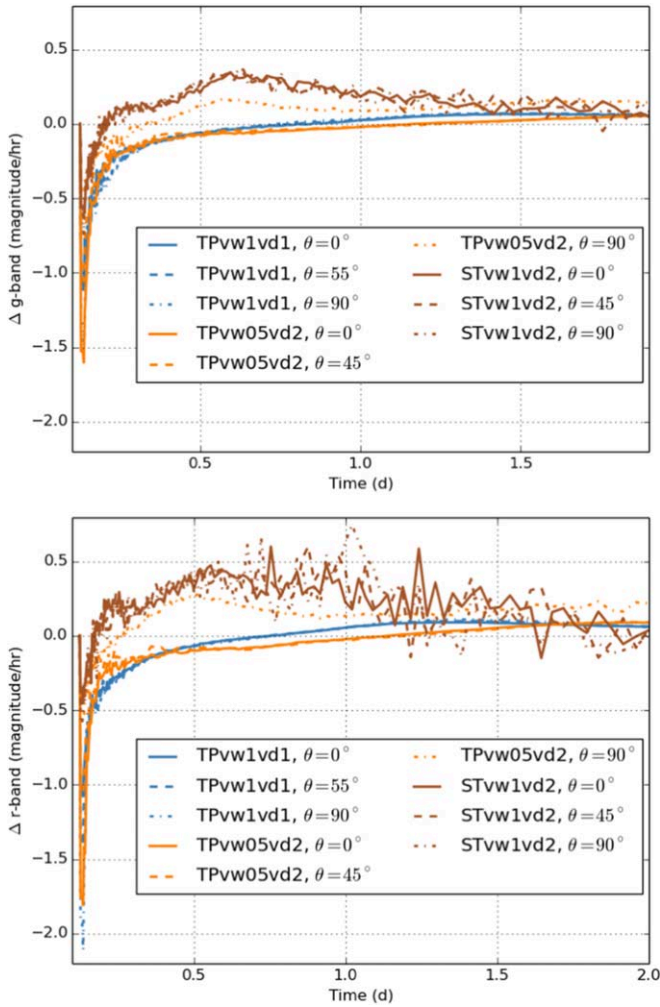


Figure 10. The rage of change (magnitudes/h) in the g and r bands for the models in Figure 9. All of the models are characterized by a sharp rise in the emission. But if we miss this initial rise, the evolution can vary dramatically depending on the model. Within the first 12 hr, some models decline, others continue to rise, but at a slower pace. All are characterized by rapid variability and, although other transients may exhibit this variability, this rapid variability is a strong indicator of a potential kilonova.

cases, the variability is high in kilonova models and variation of a few tenths of a magnitude in a few hours is an indication of a potential kilonova.

Once we identify potential kilonovae and obtain observations, we plan to use these same simulations to help interpret the observations. A variety of ejecta yields were predicted from different analyses of AT2017gfo (Côté et al. 2018). The results of studies of potential kilonovae associated with GRBs have also shown a variety of interpretations for a single data set (Fong et al. 2021; O’Connor et al. 2021). Constraining this range of interpretations of a single data set requires a broad set of observations and the HET can play a critical role in differentiating model interpretations. Differentiating the models can be done with light-curve bands, but spectral observations also provide insight into the properties of the ejecta. For example, although the fast ejecta velocities and dense forest of lines produce continuum-like spectra, both observations of AT2017gfo and a series of models suggest that line features might be observed across a broad spectral range (Watson et al. 2019; Korobkin et al. 2021; Domoto et al. 2022).

Many of these spectral features are difficult to detect, requiring high S/N spectra, and oftentimes these features are line blends from multiple elements. Despite these difficulties, broad line features and spectral slopes can help us constrain the ejecta composition. Broad spectral features can also probe the morphology and viewing angle of the merger. Figure 11 shows the spectra at four epochs in a series of as-yet-unpublished spherically symmetric wind models using ejecta compositions spanning a range of distributions. The compositions in Figure 11 are listed in Table A1. Compositions C_1 and C_2 have high mass fractions of heavier elements and display relatively few features beyond some broad features above 8000 Å at 2 days. As we decrease the fraction of these heavy elements (C_4 and C_5 have the lowest mass fractions), a number of spectral features from elements up to the first r -process peak are visible. The wide variation in this spherically symmetric wind model, particularly in the optical bands, demonstrates just how sensitive the HET spectra could be to the ejecta characteristics. We can also use these broad spectral features to help us confirm an electromagnetic counterpart, allowing us to notify the community of a true kilonova detection.

Composition is just one of the properties of model kilonovae that we have studied. Figure 12 shows the spectra in the 3000–7000 Å wavelength range focusing on one of the models from the morphology study (Korobkin et al. 2021) with a spherical wind of steady velocity $0.5c$ and a toroidal disk component with a velocity of $0.2c$ (TSvw0.5vd2), comparing spectra at different times and different viewing angles using the standard “wind 1” composition from Korobkin et al. (2021). The variation with respect to viewing angle for this particular morphology is dramatic. Without constraining the viewing angle, we have yet to identify distinct model spectral features in the early blue kilonova spectra. We do expect any line features to be broad and, as the LANL team improves its opacities, we will continue to look for specific spectral features. These features are essential to help us distinguish between the morphology, composition, and viewing angle effects. Even so, extensive observations are required to disentangle these effects successfully, and HET spectra can be a critical part of this effort.

Our model database continues to grow as we incorporate new physics and initial conditions into our models and import the simulations of other groups into our studies. The kilonova community is conducting comparison projects to constrain better issues in physics and numerics. These studies also will ultimately allow us to do better uncertainty quantification as well.

6. Future Observations

The LVC third observing run brought two new factors into play. The greater sensitivity increased the search volume, thus increasing the likelihood of detecting a BNS event. This also meant, however, that the average merger was more distant and the OT more difficult to detect. It is also true that the factors influencing the intrinsic luminosity and color of a BNS kilonova are many, especially including the aspect angle, so it may not be reliable to scale expectations with the observations of AT2017gfo. The bottom line is that no group anywhere detected an OT in O3, despite a massive deployment of resources. That said, we are confident that our original analysis of the capabilities of the HET stands, and that the

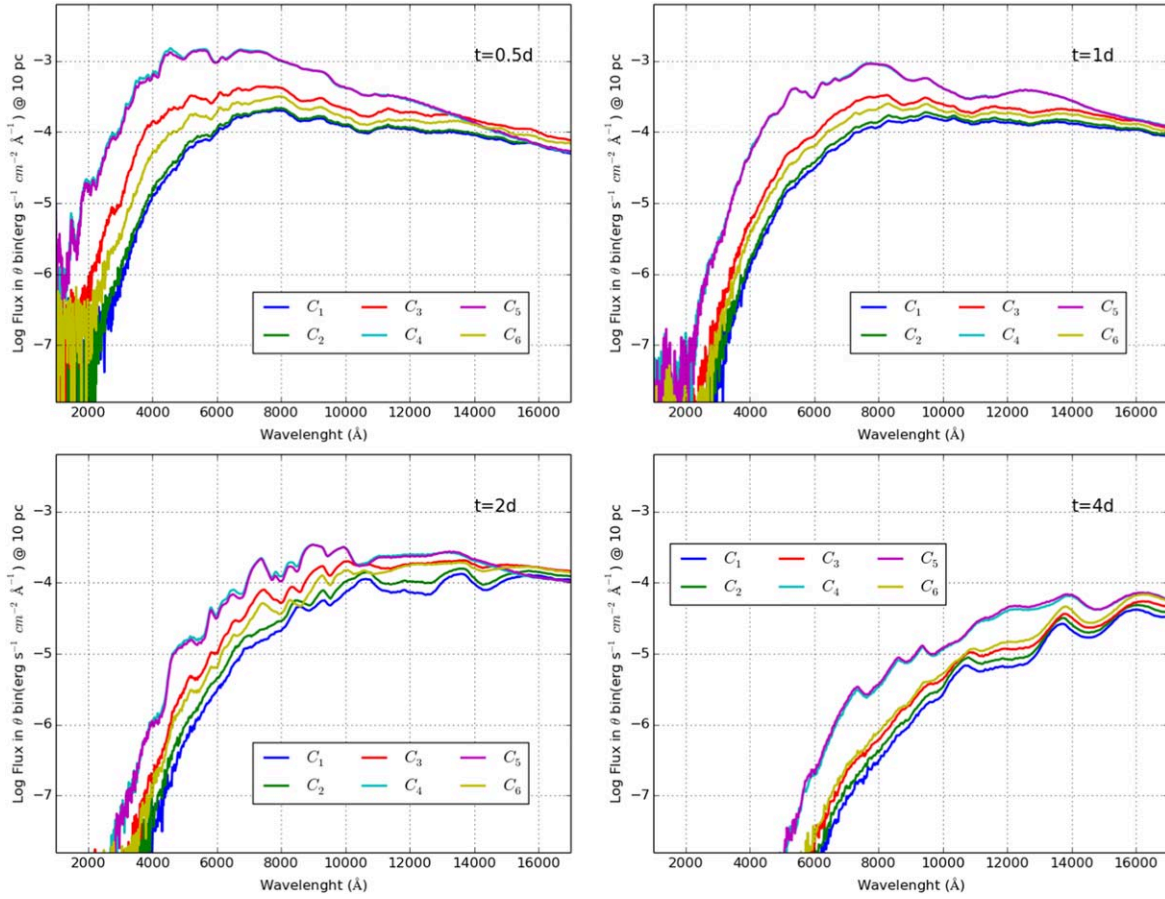


Figure 11. Spectra from a set of compositions produced by a high-fidelity spherical wind-ejecta model at a variety of times for a range of angles using the wind disk model from Miller et al. (2020) at 0.05, 1, 2, and 4 days. Even though these results are focused on a single wind profile, varying the abundances produce a range of spectral fluxes, especially in the optical bands that vary by over an order of magnitude. The compositions studied in this plot are listed in Table A1.

program merits continuing with the same commitment of HET resources.

For the O4 observing run, DIAGNOSIS will listen to alerts sent by LVK on Kafka Notices via SCIMMA,²³ which is the method by which LVK alerts will be sent in O4. In addition, galaxies within the region observable by HET will be weighted more accurately by mass using the updated GLADE catalog.

LIGHETR collected a considerable amount of ancillary data during O3. We will make these data available upon request. We will make our O3 pointings available to the Treasure Map²⁴ record of O3 data and intend to implement deposition of our real time pointing record in O4.

The O3 science run drove home the difficulty in identifying the kilonova resulting from a compact object merger among the extensive lists compiled by optical astronomers of contemporary transients that fell within the gravitational-wave localization errors. To help us focus on the most likely candidates to follow up with our telescope time, we plan to leverage a large, and continually growing database of model kilonova spectra and light-curve calculations from our LANL collaborators. These simulations can be used both to guide our observations by helping us determine which transients to follow up and to interpret our results better once observations are made.

More work remains to be done to strengthen the connection between models and the observational data. Kilonovae

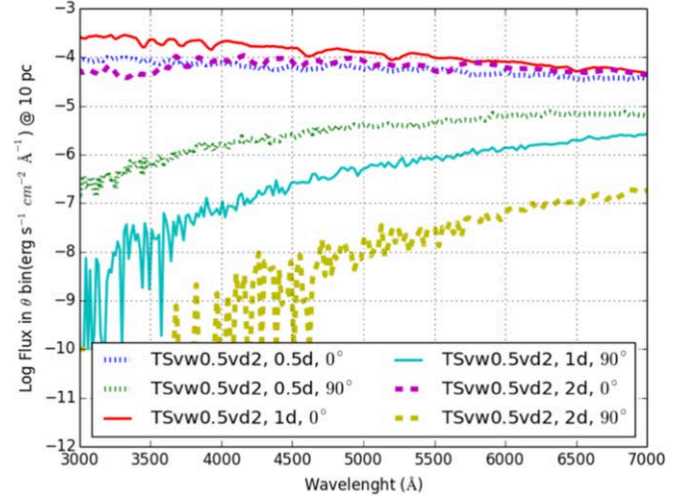


Figure 12. Model spectra at 0.5, 1.0, and 2.0 days for the model TSvw0.5vd2 from Korobkin et al. (2021) along two different viewing angles (along the angular momentum axis and perpendicular to this axis).

templates make use of a broad suite of models from the LANL simulation effort, with physics uncertainty studies covering opacity, composition, power source, and morphology. Template matching to these models could improve our sensitivity to low-brightness kilonovae and also help in the effort of distinguishing kilonovae from other transients in real time,

²³ <https://rtd.igwn.org/projects/userguide/en/stable/quickstart.html>

²⁴ <http://treasuremap.space/>

which could significantly reduce the amount of telescope time the community devotes to LVC alerts. When and if a source were to be detected by LIGHETR, we could also wield the templates to make inferences on the kilonova properties and thus provide the community with a more detailed picture of the physics behind the transient.

The pipelines we are developing for LIGHETR are designed to incorporate the growing database of LANL models. We have recently incorporated a neural-network-based classifier that scans each pixel in an IFU data cube and can identify kilonova candidates by comparison to theoretical models. This classifier was trained on the LANL suite of models and will be updated with any new theoretical developments. During an active search, the software can spectroscopically identify a kilonova candidate and then rapidly compare to the LANL model grid to constrain model parameters such as the critical aspect angle. We will thus have a substantial theoretical and simulation effort to complement our observational program.

The challenges faced by both the standard HETDEX pipeline and SCARLET in extracting candidate OTs from the VIRUS data illustrate the value of repeated observations of our target galaxies. With repeated observations we could perform data cube differentiation and any transients, be they systematic or real, would be identified and characterized. This would also help assess the performance of our methods, since we could compare our debleded extraction of some of the transients we followed up with the differentiated data.

While it is clear much needs to be done to obtain a quantitative figure of merit for nondetections, we should also highlight the great success of the LIGHETR program. The LIGHETR collaboration achieved the production of an effective alert system for possible kilonova transients, DIAGNOSIS, customized to the specifics of our instrument. The creation of an extremely quick and powerful reduction pipeline, REMEDY, has also been successfully achieved for the purposes of this endeavor. These tools allow members of the collaboration to have, within minutes of a gravitational-wave alert, access to spectral data cubes of galaxies that have a finite probability of hosting the transient. During O3, these data cubes were obtained, reduced, and immediately inspected for possible continuum sources but no evident kilonova transients were found.

The data and tools collected during this campaign are of great scientific value regardless of the identification of a transient. Kinematic maps for several galaxies have been collected and already REMEDY has become a standard reduction tool for the HETDEX consortium. Some of the transients we followed up motivated new observing proposals. The Wolf-Rayet star we suspect to be hosted by a low-brightness dwarf galaxy remains a puzzle we hope to solve.

For O4, the VIRUS array is fully implemented with 78 VIRUS units with a total of 34,944 fibers spanning an area $21'$ on a side (Figure 4). Previous VIRUS operations have shown that the typical setup time is about 6 minutes. An exposure of 9 minutes will get down to about 19th magnitude. Although in this project the average time for setup was only 2.3 minutes, we still plan on a single exposure on a target galaxy to require a total time of 15 minutes. There will be a great deal of competition from wide-field robotic photometric facilities to discover the OT, but if we find the OT first, we will automatically acquire the first spectrum. Even if we do not find the OT first, by concentrating on a high-confidence region

of the candidate sky map HET will be roughly positioned quickly to get early spectra. If we see an early blue component, as widely expected, that will already be interesting. If we do not see an early blue component, that will also be interesting. In any case, the blue response of the HET VIRUS array is well designed to get this early data in the blue.

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The University of Texas at Austin and McDonald Observatory sit on indigenous land. The Tonkawa lived in central Texas and the Comanche and Apache moved through this area. The Davis Mountains that host McDonald Observatory were originally husbanded by Lipan Apache, Warm Springs Apache, Mescalero Apache, Comanche and various tribes of the Jumanos. We acknowledge and pay our respects to all the Indigenous Peoples and communities who are or have been a part of these lands and territories in Texas.

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Software: astropy (Astropy Collaboration et al. 2013, 2018), emcee (Foreman-Mackey et al. 2013), numpy (Harris et al. 2020), scipy (Virtanen et al. 2020), matplotlib (Hunter 2007), and pandas (Wes 2010; pandas development team T 2020). Swift data were reduced with XRTDAS (v0.13.5), CALDB (v20190910), XRTPipeline, and XSELECT. Chandra data were processed with SPECEXTRACT.

Appendix A Model Light-curve Compositions

Table A1 gives the compositions employed in Figure 11.

Table A1
Composition of the Models Presented in Figure 11

<i>Z</i>	<i>C</i> ₁	<i>C</i> ₂	<i>C</i> ₃	<i>C</i> ₄	<i>C</i> ₅	<i>C</i> ₆
24	0.0	7.49370e-04	3.84341e-03	9.12667e-03	1.08047e-02	1.50548e-04
26	1.09998e-04	2.94116e-03	1.60909e-02	2.93123e-02	4.22312e-02	5.53921e-03
34	1.20896e-01	1.99367e-01	2.86043e-01	2.92805e-01	2.35532e-01	1.85846e-01
35	3.69738e-02	5.20294e-02	1.11589e-01	2.27298e-01	3.27268e-01	6.73422e-02
40	1.09691e-02	2.26226e-02	6.74217e-02	1.32521e-01	1.65283e-01	1.69854e-01
46	1.47586e-01	1.39364e-01	1.05772e-01	9.16517e-02	7.82567e-02	1.68748e-01
52	6.04172e-01	5.22899e-01	3.86228e-01	2.14772e-01	1.38436e-01	3.76756e-01
57	1.96323e-03	1.39742e-03	7.12728e-04	1.26050e-04	1.12861e-04	9.23346e-04
58	2.95243e-03	2.23794e-03	1.14957e-03	1.81696e-04	1.91118e-04	1.54047e-03
59	7.70276e-04	5.68369e-04	3.06602e-04	1.19901e-04	0.0	3.27700e-04
60	1.92511e-03	1.32451e-03	5.91964e-04	1.19743e-04	1.26560e-04	1.09535e-03
62	2.38889e-03	1.58250e-03	7.01276e-04	1.19743e-04	1.15225e-04	9.87945e-04
63	5.97838e-04	3.86548e-04	1.78528e-04	0.0	0.0	2.69349e-04
64	5.82899e-03	3.87802e-03	1.72436e-03	1.91844e-04	1.31481e-04	1.05073e-03
65	1.54890e-03	1.02126e-03	4.51920e-04	0.0	0.0	3.63695e-04
66	1.31050e-02	8.49339e-03	3.89174e-03	4.32027e-04	2.87312e-04	2.10961e-03
67	2.90705e-03	1.87583e-03	8.73047e-04	0.0	0.0	4.56880e-04
68	9.80968e-03	6.47393e-03	2.92234e-03	3.28006e-04	2.32171e-04	1.71006e-03
69	1.44669e-03	9.58414e-04	4.32732e-04	0.0	0.0	2.95809e-04
70	1.84161e-02	1.24477e-02	5.52962e-03	6.75113e-04	5.96579e-04	4.89435e-03
92	1.56322e-02	1.73818e-02	3.54577e-03	3.39432e-04	3.94817e-04	9.73867e-03

Appendix B

LIGHETR Alert Components

Here we give some illustrations of how the internal LIGHETR alert system works. Figure B1 gives an example

of an LVC alert as processed by DIAGNOSIS for local redistribution. Figure B2 gives the TSL subsequently sent to the resident astronomers at the HET. Figure B3 presents a graphical representation of the galaxy target priority list used in the original search for an OT.

```
### TEST Preliminary GW ALERT ###
Time until 90% probability region: 3.7 hours

internal = 0
Packet_Type = 150
Pkt_Ser_Num = 1
GraceID = MS181101ab
AlertType = Preliminary
HardwareInj = 0
OpenAlert = 1
EventPage = https://example.org/superevents/MS181101ab/view/
Instruments = H1,L1
FAR = 9.11069936486e-14
Group = CBC
Pipeline = gstlal
Search = MDC
skymap_fits = https://emfollow.docs.ligo.org/userguide/static/bayestar.fits.gz
BNS = 0.95
NSBH = 0.01
BBH = 0.03
MassGap = 0.0
Terrestrial = 0.01
HasNS = 0.95
HasRemnant = 0.91
Distance = 141.1453950128411 ± 39.09548411497191

If you happen to find the location of the source, please submit coordinates to GraceDB using submit_gracedb.py.
#####
Alert created with DIAGNOSIS, for more information on the software and data products, please refer to the wiki:
https://github.com/Majoburo/Diagnosis
```

Figure B1. Trial LIGO alert as processed by DIAGNOSIS for local distribution.

COMMON

```

PROGRAM UT19-2-009
VIFU      047
EXP       120
NUMEXP    3
EQUINOX 2000.0
INSTRUMENT      VIRUS
MAG          22
SKYBRIGHT      18.0
SEEING   3.0
SKYTRANS       S
SKYCAL S Y
FLUX      Y
PRI        0
SETUPMETHOD   DirectGuider
COMMENT "Usual Dither, look for new object in target IFU"

```

TRACK_LIST

OBJECT	RA	DEC	PIPRI		
GW142.909637+30.975740	9:31:38.3129	30:58:32.664	1		
GW151.394000+27.514560	10:05:34.56	27:30:52.416	2		
GW149.644000+28.667410	9:58:34.56	28:40:02.676	3		
GW152.599000+27.117610	10:10:23.76	27:07:03.396	4		
GW151.746000+27.689290	10:06:59.04	27:41:21.444	5		
GW141.380000+31.741220	9:25:31.2	31:44:28.392	6		
GW154.878000+25.682640	10:19:30.72	25:40:57.504	7		
GW155.146469+25.383549	10:20:35.1526	25:23:00.7764	8		
GW144.251000+31.001340	9:37:00.24	31:00:04.824	9		
GW155.092000+25.426910	10:20:22.08	25:25:36.876	10		
GW155.244110+25.519281	10:20:58.5864	25:31:09.4116	11		
GW155.160019+25.504938	10:20:38.4046	25:30:17.7768	12		
GW155.164000+25.559960	10:20:39.36	25:33:35.856	13		
GW149.667114+28.877678	9:58:40.1074	28:52:39.6408	14		
GW154.494949+25.493961	10:17:58.7878	25:29:38.2596	15		
GW155.096000+25.631350	10:20:23.04	25:37:52.86	16		
GW142.017000+31.986220	9:28:04.08	31:59:10.392	17		
GW156.012177+25.105753	10:24:02.9225	25:06:20.7108	18		
GW145.218000+30.938840	9:40:52.32	30:56:19.824	19		
GW143.747000+30.257750	9:34:59.28	30:15:27.9	20		

Figure B2. Phase II TSL generated by DIAGNOSIS giving the prioritized list of target galaxies that will be employed to trigger the search for the OT with the VIRUS IFU array. The galaxies in this list are also presented in Figure B3.

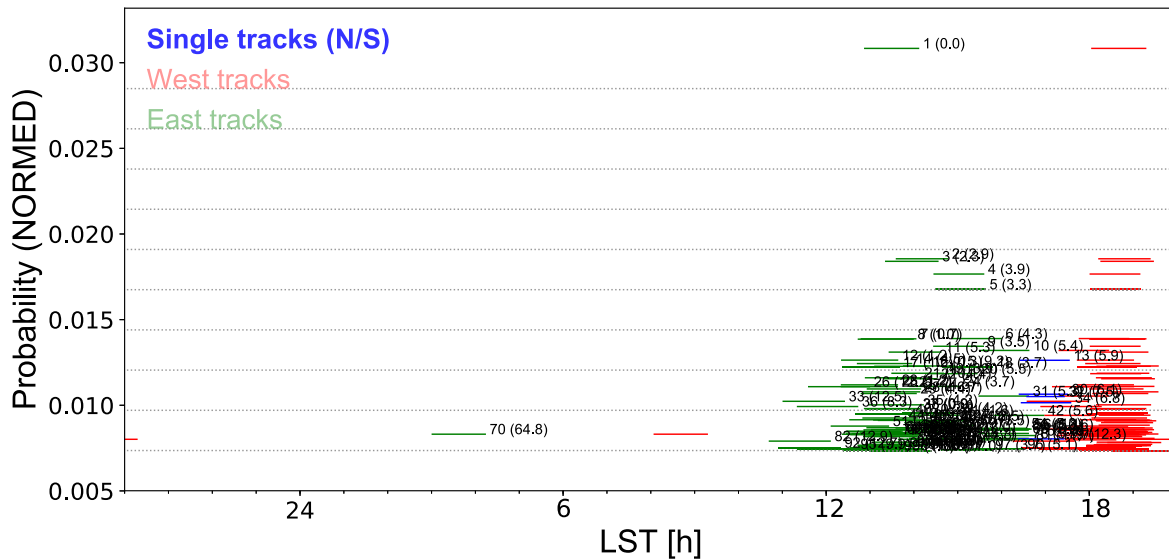


Figure B3. Sample output from combining localization information for LIGO event S190425z as illustrated in Figure 1 with the GLADE galaxy catalog to produce a list of 20 galaxies prioritized by the likelihood of being the host of a LIGO alert burst of gravitational waves. The vertical scale gives the log of the probability, and the horizontal scale gives the position in hour of R.A. Blue (single bands between the east and west extremes) shows galaxies available on a single track, red (right bands) and green (left bands) show the time when a given galaxy can be observed on the west track and the east track, respectively.

Appendix C

Timetable for Alert, Search, Dense Spectral Sampling, and Analysis

VIRUS

00:00+01:00 LIGO alert, precursor or normal, sent to team, resident astronomers.
 01:00+00:30 Automatically run DIAGNOSIS tool in response to the alert.
 01:30+00:30 DIAGNOSIS generates VIRUS Phase II TSL with priority list of galaxy targets.
 02:00+10:00 Scientists in consultation with resident astronomers make decision on whether to trigger telescope.
 12:00+02:00 Resident astronomers ingest galaxy Phase II into HET queue and decide when to implement, depending on current observation.
 14:00+01:00 Select first galaxy target in HET queue
 15:00+05:00 Slew telescope and set up.
 20:00+06:00 Start first 6 minute dither.
 26:00+01:00 Run REMEDY code to produce collapsed image and reduced spectrum. Start second dither.
 27:00+05:00 Inspect REMEDY output. If OT, send alert, submit LRS2 Phase II.
 32:00+06:00 Second dither out, proceed as above with REMEDY, start third dither.
 38:00+05:00 Third dither out, either move to next target or switch to LRS2 if OT found.
 43:00—Proceed as above; start next VIRUS exposure or first LRS2 exposure if OT found.
Total 180 minutes

LRS2

First Night

00:00+05:00 Halt VIRUS observations, submit LRS2 Phase II.
 05:00+08:00 Set up on LRS2-B.
 13:00+20:00 First LRS2-B exposure.
 33:00+02:00 Set up on LRS2-R.

35:00+20:00 First LRS2-R exposure. Reduce and examine first LRS2-B spectrum.
 55:00+02:00 Set up again on LRS2-B.
 57:00+20:00 Second LRS2-B exposure. Reduce and examine first LRS2-R spectrum.
 77:00+02:00 Set up again on LRS2-R.
 79:00—Continue until target no longer visible for the HET.
Total 300 minutes

Second Night

00:00+08:00 Set up on LRS2-B.
 08:00+20:00 First LRS2-B exposure.
 28:00+02:00 Set up on LRS2-R.
 30:00+20:00 First LRS2-R exposure. Reduce and examine first LRS2-B spectrum.
 50:00+02:00 Set up again on LRS2-B.
 52:00+20:00 Second LRS2-B exposure. Reduce and examine first LRS2-R spectrum.
 72:00+02:00 Set up again on LRS2-R.
 74:00+20:00 Second LRS2-R exposure. Reduce and examine second LRS2-B spectrum.
 94:00+02:00 Set up again on LRS2-B.
 96:00+20:00 Third LRS2-B exposure. Reduce and examine second LRS2-R spectrum.
 116:00+02:00 Set up again on LRS2-R.
 118:00+20:00 Third LRS2-R exposure. Reduce and examine third LRS2-B spectrum.
 138:00+12:00 Reduce and examine third LRS2-R spectrum.
Total 150 minutes

Third, Fourth, and Fifth Nights

00:00+08:00 Set up on LRS2-B.
 08:00+30:00 First LRS2-B exposure.
 38:00+02:00 Set up on LRS2-R.
 40:00+30:00 First LRS2-R exposure. Reduce and examine first LRS2-B spectrum.

70:00+02:00 Set up again on LRS2-B.
 72:00+30:00 Second LRS2-B exposure. Reduce and examine first LRS2-R spectrum.
 102:00+02:00 Set up again on LRS2-R.
 104:00+30:00 Second LRS2-R exposure. Reduce and examine second LRS2-B spectrum.
 134:00+06:00 Reduce and examine second LRS2-R spectrum.
Total 140 minutes each night

Day 1—Begin analysis of data, comparison to models.
 Day 5—Finish analysis of data, comparison to models.
 Day 7—Submit paper.

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