



First M87 Event Horizon Telescope Results. IX. Detection of Near-horizon Circular Polarization

The Event Horizon Telescope Collaboration
(See the end matter for the full list of authors.)

Received 2023 June 19; revised 2023 September 5; accepted 2023 September 12; published 2023 November 8

Abstract

Event Horizon Telescope (EHT) observations have revealed a bright ring of emission around the supermassive black hole at the center of the M87 galaxy. EHT images in linear polarization have further identified a coherent spiral pattern around the black hole, produced from ordered magnetic fields threading the emitting plasma. Here we present the first analysis of circular polarization using EHT data, acquired in 2017, which can potentially provide additional insights into the magnetic fields and plasma composition near the black hole. Interferometric closure quantities provide convincing evidence for the presence of circularly polarized emission on event-horizon scales. We produce images of the circular polarization using both traditional and newly developed methods. All methods find a moderate level of resolved circular polarization across the image ($\langle |v| \rangle < 3.7\%$), consistent with the low image-integrated circular polarization fraction measured by the Atacama Large Millimeter/submillimeter Array ($|v_{\text{int}}| < 1\%$). Despite this broad agreement, the methods show substantial variation in the morphology of the circularly polarized emission, indicating that our conclusions are strongly dependent on the imaging assumptions because of the limited baseline coverage, uncertain telescope gain calibration, and weakly polarized signal. We include this upper limit in an updated comparison to general relativistic magnetohydrodynamic simulation models. This analysis reinforces the previously reported preference for magnetically arrested accretion flow models. We find that most simulations naturally produce a low level of circular polarization consistent with our upper limit and that Faraday conversion is likely the dominant production mechanism for circular polarization at 230 GHz in M87*.

Unified Astronomy Thesaurus concepts: Polarimetry (1278); Galaxies (573); Supermassive black holes (1663); Very long baseline interferometry (1769); Low-luminosity active galactic nuclei (2033)

1. Introduction

The Event Horizon Telescope (EHT) has produced the first images of the event-horizon-scale millimeter emission around the supermassive black hole (BH) in the core of the massive elliptical galaxy M87 at the center of the Virgo Cluster. Using very long baseline interferometry (VLBI) at 230 GHz, these initial EHT observations from 2017 recovered a ring-like structure with a diameter similar to the predicted “BH shadow” of an $M_{\text{BH}} \approx 6.5 \times 10^9 M_{\odot}$ BH at the distance of M87* ($D \approx 16.8$ Mpc). The resolved total intensity images of the ring were consistent with models of synchrotron emission from ultrahot magnetized plasma near the event horizon (Event Horizon Telescope Collaboration et al. 2019a, 2019b, 2019c, 2019d, 2019e, 2019f, hereafter Papers I–VI).

The EHT observes in full polarization, recording simultaneous data from orthogonally polarized feeds at each antenna. Images of the near-horizon linearly polarized radiation were published and analyzed in Event Horizon Telescope Collaboration et al. (2021a, hereafter Paper VII). These linear polarimetric images provide essential new information about the magnetic field structure near the event horizon of M87’s supermassive BH, indicating that the near-horizon magnetic

fields are ordered and dynamically important (Event Horizon Telescope Collaboration et al. 2021b, hereafter Paper VIII).

In this Letter we report on the search for resolved *circularly* polarized radiation (CP) on event horizon scales in M87* from EHT observations in 2017. The circularly polarized signal from synchrotron radiation near the BH should contain unique information about the magnetic field and the nature of the radiating particles that cannot be inferred from total intensity or linear polarization alone. These include the possibility of directly measuring the strength of the magnetic field and determining whether the observed radiation is mainly from an electron–positron or an electron–ion plasma (e.g., Wardle et al. 1998). However, the circularly polarized signal is expected to be weaker than the linear polarization (Jones & O’Dell 1977; Jones 1988), requiring high sensitivity and accurate calibration of each antenna to be detected.

Previous radio and millimeter-wavelength observations of CP in M87* are quite limited, whether at low or high angular resolution. Homan & Lister (2006), using the Very Long Baseline Array (VLBA) at 15 GHz, measured fractional circular polarization in the core of $-0.49\% \pm 0.10\%$. Interestingly, they found no linear polarization ($LP < 0.07\%$), reversing the Stokes parameter hierarchy described above. M87 is the only source in their list to show such behavior. On the other hand, using the Very Large Array (VLA) at 8.4 GHz, Bower et al. (2002) detected weak fractional linear polarization ($LP = 1.74\% \pm 0.06\%$) but no circular polarization ($|CP| < 0.1\%$). Single-dish measurements of M87 with the Effelsberg 100 m radio telescope at the same range of

¹⁵³ NASA Hubble Fellowship Program, Einstein Fellow.



Original content from this work may be used under the terms of the [Creative Commons Attribution 4.0 licence](https://creativecommons.org/licenses/by/4.0/). Any further distribution of this work must maintain attribution to the author(s) and the title of the work, journal citation and DOI.

frequencies showed a similar trend with $LP \sim 1.5\%$ and $|CP| < 0.2\%$ (Myserlis et al. 2018), while at 86 GHz the POLAMI monitoring program observed a fairly stable $LP \sim 5\%$ and $CP \sim -1.5\%$ over a period of 12 yr with the IRAM 30 m telescope (Thum et al. 2018). Simultaneous observations during the 2017 EHT campaign with the Atacama Large Millimeter/submillimeter Array (ALMA) did not result in a significant detection of the unresolved fractional CP at 221 GHz ($\approx -0.3\% \pm 0.6\%$; Goddi et al. 2021).

Beyond M87*, accurate measurements of circular polarization are generically difficult to obtain in VLBI. Homan & Lister (2006) detected CP at the level of 3σ or better in 17 sources out of their sample of 133 active galactic nuclei (AGNs) using a “gain transfer” technique (Homan & Wardle 1999), in which all sources are used for polarization calibration. These detections at 15 GHz all had fractional polarizations between 0.25% and 0.70%. Gabuzda et al. (2008) also detected circular polarization in eight AGN jets with 15 GHz VLBI measurements and associated the observed CP signs with CP production by Faraday conversion in helical jet magnetic fields. At higher frequencies up to 43 GHz, Vitriishchak et al. (2008) found circular polarization fractions up to $\sim 1\%$ in a sample of AGN cores resolved with VLBI.

This paper presents the details of the 2017 EHT observations and data calibration for circular polarization, procedures and results for circular polarimetric imaging, and their theoretical interpretation in constraining parameters in a library of simulation models. In Section 2, we summarize the EHT 2017 observations, describe evidence for the detection of circular polarization in M87*, and describe our a priori calibration procedure. In Section 3, we describe our methods of circular polarimetric image reconstruction and gain calibration from EHT data. In Section 4, we present image reconstructions from all methods across all EHT 2017 data sets. We derive an upper limit on the average, resolved circular polarization fraction in M87* at 230 GHz.

In Section 5, we examine circular polarization in a library of general relativistic magnetohydrodynamic (GRMHD) simulations of M87*. We add upper limits to the circular polarization on event horizon scales from our EHT observations to the list of constraints applied to theoretical models in Paper VIII. We discuss the effect of these new constraints on our preferred models for M87*'s accretion flow, and we investigate the physical origin of circular polarization in passing GRMHD models. We summarize the work in Section 6.

2. EHT Observations of M87*

2.1. Conventions

A radio interferometer, such as the EHT, samples the Fourier plane (visibility domain) of the brightness distribution (image) on the sky (e.g., Thompson et al. 2017). Each image domain Stokes parameter (\mathcal{I} , \mathcal{Q} , \mathcal{U} , \mathcal{V}) has a corresponding visibility, which we denote with a tilde (e.g., $\tilde{\mathcal{V}}$). At most stations, the EHT data are recorded in two orthogonal circular polarizations, right and left (R , L). Interferometric visibilities are computed through the complex correlation between each pair of sites (j , k) and polarization (i.e., $R_j R_k^*$, $R_j L_k^*$, $L_j R_k^*$, and $L_j L_k^*$). These measurements can be transformed to the Stokes representation through linear algebraic relations. In particular, the visibility

domain circular polarization $\tilde{\mathcal{V}}$ on the $j - k$ baseline is given by

$$\tilde{\mathcal{V}}_{jk} = 0.5(R_j R_k^* - L_j L_k^*). \quad (1)$$

In practice, polarimetric measurements are corrupted through a multitude of systematic effects, which can be conveniently represented in a matrix formalism (e.g., Hamaker et al. 1996; Smirnov 2011):

$$\rho'_{jk} = \mathbf{J}_j \rho_{jk} \mathbf{J}_k^\dagger. \quad (2)$$

In this expression, ρ_{jk} is a 2×2 matrix of the true visibilities on the $j - k$ baseline, ρ'_{jk} gives the measured visibilities, and \mathbf{J} is a complex time- and site-dependent 2×2 matrix that describes the aggregate systematic effects. The latter can be further decomposed into a product of three terms, $\mathbf{J} = \mathbf{G} \mathbf{D} \mathbf{\Phi}$, where \mathbf{G} is a diagonal matrix that describes the time-dependent “gains” of the two feeds, \mathbf{D} is a matrix with diagonal entries that are unity and off-diagonal entries that describe “leakage” between the feeds, and the matrix $\mathbf{\Phi}$ describes the overall rotation of the feeds. For studies of circular polarization with circularly polarized feeds, the gain matrix is the most important source of contamination, predominantly through the gain ratio $G_{R/L}$,

$$\mathbf{G} \equiv \begin{pmatrix} G_R & 0 \\ 0 & G_L \end{pmatrix} = G_L \begin{pmatrix} G_{R/L} & 0 \\ 0 & 1 \end{pmatrix}. \quad (3)$$

In particular, while the gains G_L and G_R have rapid variations, especially in their phase, the gain ratio $G_{R/L}$ can be stable over timescales of multiple days. For additional details on polarimetric relations, these representations, and our conventions, see Paper VII.

2.2. Observations and Data Reduction

The EHT observed M87* on 2017 April 5, 6, 10, and 11 with a VLBI array of seven telescopes located at five geographical sites: ALMA and the Atacama Pathfinder Experiment (APEX) in the Atacama Desert in Chile, the Large Millimeter Telescope Alfonso Serrano (LMT) on the Volcán Sierra Negra in Mexico, the IRAM 30 m telescope (PV) on Pico Veleta in Spain, the Submillimeter Telescope (SMT) on Mt. Graham in Arizona (USA), and the Submillimeter Array (SMA) and the James Clerk Maxwell Telescope (JCMT) on Maunakea in Hawai'i. The South Pole Telescope (SPT) also participated in the EHT observations but cannot see M87*.

Each telescope recorded two frequency bands, each 2 GHz wide, centered at 229.1 GHz (high band, HI) and at 227.1 GHz (low band, LO). Most sites recorded right and left circular polarization simultaneously, except for the JCMT, which recorded a single circular polarization each night, and ALMA, which recorded orthogonal linear polarizations that were subsequently converted to a circular basis (Martí-Vidal et al. 2016; Matthews et al. 2018; Goddi et al. 2019). Paper II provides a detailed description of the EHT array and observations.

The EHT data sets were first correlated and then calibrated using two independent pipelines, EHT-HOPS (Blackburn et al. 2019) and CASA rPICARD (Janssen et al. 2019), for the stabilization of the source signal in time and frequency (e.g., Janssen et al. 2022). The data presented in this paper correspond to the EHT-HOPS pipeline, following a verification of the interpipeline consistency (Paper III), although in

Appendix A we also give a brief summary of the gain calibration in the CASA rPICARD pipeline, and we compare imaging results from EHT-HOPS and rPICARD pipeline data in Section 4. The amplitude flux density calibration was performed with custom-built EHT-HOPS postprocessing scripts, based on the metadata provided by the participating telescopes. A more extensive description of other aspects of the EHT M87* data set calibration was presented in Paper III. The data sets analyzed in this paper are identical to the ones used in Papers VII and VIII, following the same polarimetric calibration procedures. The full-Stokes calibrated VLBI data from the 2017 EHT observations of M87 are publicly available through the EHT data portal¹⁵⁴ under the code 2023-D01-01 (Event Horizon Telescope Collaboration 2023).

2.3. R/L Gain Calibration

We calibrated the gain ratio $G_{R/L}$ for each site using multisource and multiday polynomial fits to the differences between RR^* and LL^* visibilities. Since ALMA observes in a linear polarization basis, it provides the absolute electric vector position angle (EVPA) information and is used as the reference station; its $G_{R/L}$ is fixed to unity (Martí-Vidal et al. 2016; Goddi et al. 2019).¹⁵⁵

For the other sites, the R/L gains were modeled separately for each band using visibilities on baselines to ALMA, using custom-built Python scripts. The use of multiple sources (10 sources were observed with ALMA during the EHT 2017 campaign) in the fitting procedure helps to distinguish between instrumental effects, which are largely source independent, and contributions from circular polarization, which are source dependent (Steel et al. 2019).

For most sites, $G_{R/L}$ could be successfully modeled using a constant value across the full observing campaign, separately for each frequency band. The sites that required a time-dependent $G_{R/L}$ were APEX, which showed a linear phase drift between RR^* and LL^* visibilities, and the SMA, which showed irregular phase variation that we modeled using third-order polynomials specific to each day and frequency band. We calibrated the amplitudes $|G_{R/L}|$ assuming a constant value for each site.

In Appendix A, we provide more details on this strategy for relative gain calibration, as well as an example plot detailing how gains were estimated for the LMT (Figure 14). We also compare the above strategy for relative gain calibration used for the main results in this paper with the alternative strategy employed by the CASA-based rPICARD pipeline for EHT data reduction, which calibrates relative gains assuming the intrinsic $\mathcal{V} = 0$.

2.4. Evidence for Circular Polarization

Under a perfect calibration, we could directly interpret the signal-to-noise ratio (S/N) of $\tilde{\mathcal{V}}$, calculated from observables with Equation (1), to identify robust detections of CP. In Figure 1 we show the S/N of $\tilde{\mathcal{V}}$ as a function of a projected baseline length. Unfortunately, as circular polarization is encoded in the difference between RR^* and LL^* visibilities,

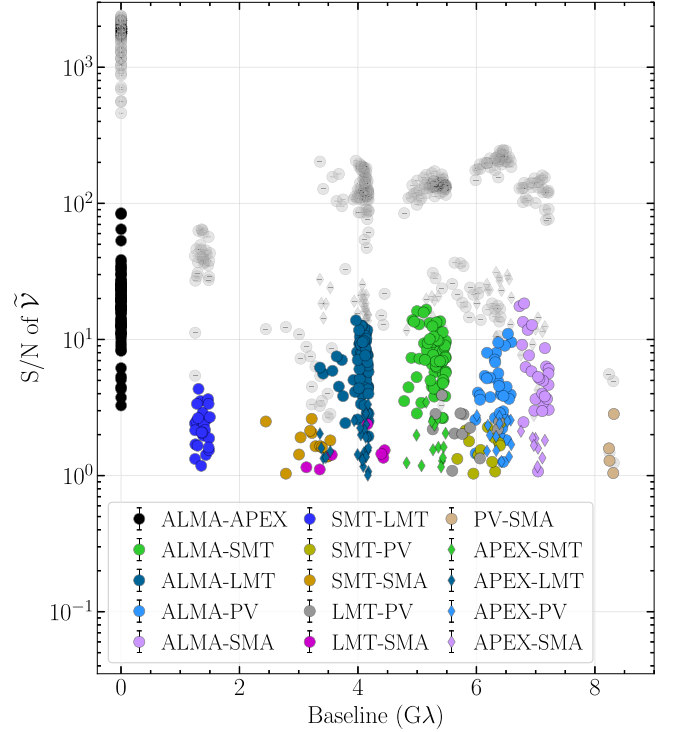


Figure 1. Stokes $\tilde{\mathcal{V}}$ visibility S/N in scan-averaged EHT 2017 data, as a function of the projected baseline length. Data from both frequency bands and all observing days are shown. No systematic uncertainties, like the imperfect calibration of the gain ratios $G_{R/L}$, were accounted for. Hence, the plotted S/N represents upper limits on the $\tilde{\mathcal{V}}$ detections. Gray points in the background indicate the S/N of Stokes $\tilde{\mathcal{I}}$ detections on the corresponding baselines.

residual errors in the calibration of $G_{R/L}$ will create spurious $\tilde{\mathcal{V}}$ signatures. As a consequence, the S/N will be inflated, and hence values shown in Figure 1 can only be treated as upper limits. This is further illustrated in Appendix A, where we discuss a more aggressive calibration strategy for $G_{R/L}$, which reduces the S/N of $\tilde{\mathcal{V}}$ significantly.

There are, however, quantities that are robust against the effects of complex antenna gains. An example of such quantities are the closure phases (e.g., Thompson et al. 2017; Blackburn et al. 2020), defined for a triangle of antennas (A , B , and C) as

$$\psi = \phi_{AB} + \phi_{BC} + \phi_{CA}, \quad (4)$$

where ϕ_{ij} is the visibility phase measured by baseline $i - j$. One way to minimize the effects from possible antenna miscalibrations and infer the presence of resolved source-intrinsic circular polarization in our observations is to compute the difference of closure phases between the RR^* and LL^* visibilities; by construction, the difference of these closures will not be affected by $G_{R/L}$ terms at any antenna.¹⁵⁶

In Figure 2 (top panels), we show the closure phases, ψ , of the antenna triangle formed by ALMA, SMT, and PV, computed for the whole campaign (left column, open markers for low band; right column, filled markers for high band). Closure phases for the RR^* and LL^* visibilities are shown in red and blue, respectively. It can be seen that there is an offset in

¹⁵⁴ <http://eventhorizontelescope.org/for-astronomers/data>

¹⁵⁵ A clockwise 45° rotation was also applied to all observations across the whole array (i.e., 90° added to the $G_{R/L}$ phases of all stations), in order to account for the orientation of the ALMA Band 6 cartridges with respect to the antenna mounts (Paper VII).

¹⁵⁶ Note that care must be taken when time-averaging visibilities before computing closures, since the averaging of miscalibrated visibilities can introduce systematics into the post-average closure distributions (e.g., Martí-Vidal & Marcaide 2008).

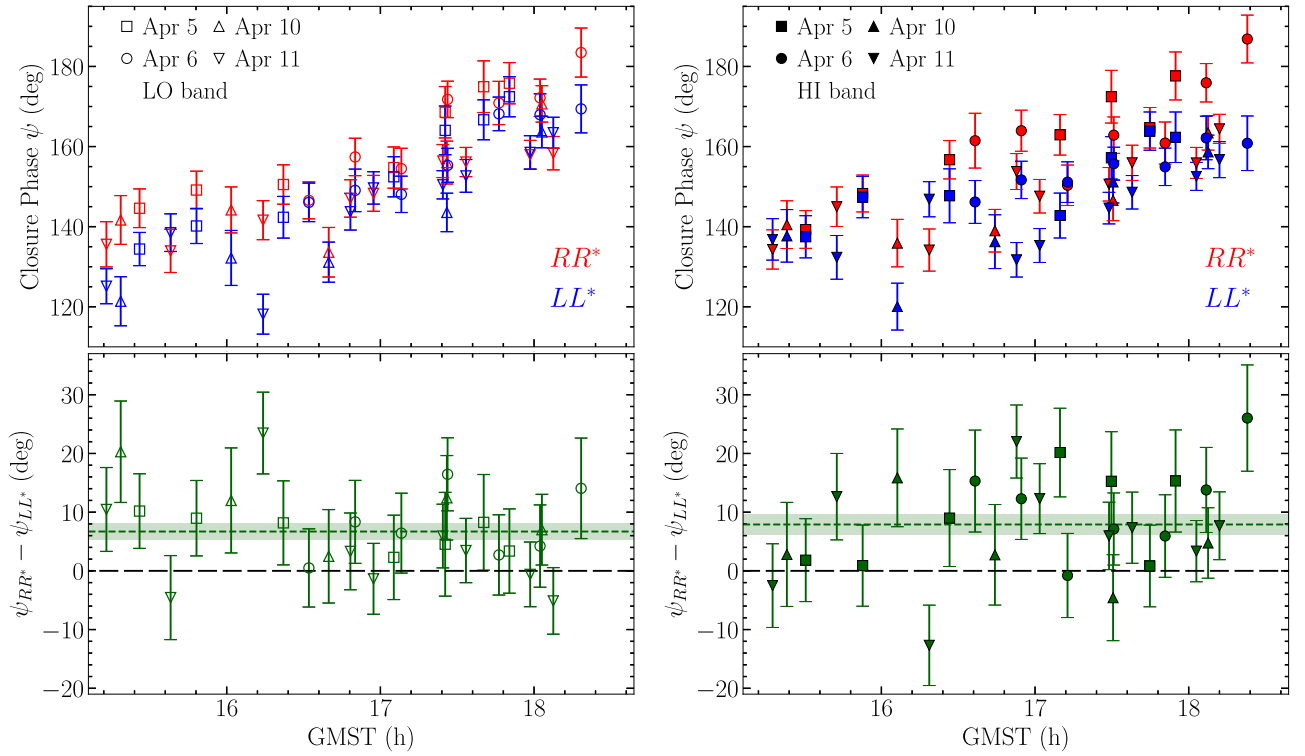


Figure 2. Closure phases observed on the ALMA-SMT-PV triangle during M87* observations on 2017 April 5–11 in low band (left column) and in high band (right column). Top: closure phases of scan-averaged visibilities for all epochs, RR^* in red, LL^* in blue. Bottom: difference of closure phases between RR^* and LL^* . The zero level of the closure difference (i.e., no ∇ detected) is marked with a black dashed line. A light-green band shows the $RR^* - LL^*$ difference, inferred by band, under the constant difference assumption.

the closure phases between the two polarization channels, which generates a nonzero closure difference (indicated in green in the bottom panels). This difference in the RR and LL closure phases is present in all epochs and in both bands.

For this antenna triangle, the average offset in closure phases between RR^* and LL^* (combining all epochs and assuming a constant residual value) is 6.7 ± 1.3 deg in the low band and 7.9 ± 1.6 deg in the high band, indicated with green bands in the bottom panels. The offset is consistent despite each band being calibrated independently. Moreover, the measured offset is difficult to explain with the conservative systematic nonclosing error budget discussed in Section 8.4 of Paper III, and hence it implies a tentative detection of a weak Stokes ∇ at the level of $S/N \sim 5$.

While this measurement implies the presence of a fractional CP reaching $\sim 3\%$ somewhere in the visibility domain, the measurement cannot be directly translated into a quantitative image domain constraint. The ALMA-SMT-PV triangle shown is the one that produces closure phase differences with the most clear deviation from zero. In Appendix B, we show the results for all other triangles including ALMA. None of these other triangles show an unambiguous detection of a nonzero closure phase difference like that seen on the ALMA-SMT-PV triangle, suggesting that SMT-PV is the baseline most sensitive to the CP signatures in M87*.

Even though the RR^* and LL^* closure phases are robust against antenna gains, they may be affected by instrumental polarimetric leakage (antenna D -terms, the D matrix term in Equation (3)). However, the effect of D -term uncertainties in the parallel-hand visibilities is much smaller than in the cross-hands (e.g., Smirnov 2011), which implies that Stokes ∇ is much less affected by instrumental polarization than Stokes Q

and U . To verify a negligible impact of the polarimetric leakage on the closure phase signal, we have compared the closure phase values between the data with and without D -term calibration. For all triangles related to ALMA, the effect of the D -terms on the $RR^* - LL^*$ closure differences is always less than the standard deviation of the thermal noise on the closure phase difference. For the triangle shown in Figure 2, the maximum effect of the D -terms is only 0.42σ .

Hence, we can conclude that the closure phase differences indicating the presence of circular polarization on EHT baselines, presented in Figure 2, are robust against both antenna gains and polarimetric leakage. In Appendix C, we discuss evidence for polarization in “closure trace” products (Broderick & Pesce 2020), quantities that are insensitive to *all* station-based systematic factors, including D -terms.

3. Polarimetric Imaging Methods

As discussed in Section 2, the circular polarization signal in M87* is both weak and sensitive to calibration errors. Consequently, the inferred circular polarization structure can be sensitive to assumptions made about the residual calibration errors and choices made in producing images from the measured visibilities. To assess the potential impact of these effects, we produce fully polarized images of M87* using five methods. We summarize these methods and their assumptions in Table 1. In Appendix D, we present more detail on each method’s assumptions and procedures used for circular polarization image reconstructions.

The five methods we use can be divided into three general categories. The first category uses adaptations of the standard CLEAN imaging algorithm together with iterative “self-

Table 1
Summary of the Imaging and Modeling Methods Used, Each of Which Produces Images of M87* in All Four Stokes Parameters

Method	Leakage Assumptions	Gain Assumptions
CLEAN Imaging:		
polsolve	Priors from Paper VII	Self-calibration (gains and D -terms). ALMA $G_{R/L} \equiv 1$
DIFMAP	Paper VII values fixed	Self-calibration, assuming $\mathcal{V} = 0$. ALMA $G_{R/L} \equiv 1$
Bayesian Raster:		
DMC	None (leakage fitted)	ALMA gain phases are 0 (all other gains fitted)
THEMIS	None (leakage fitted)	$G_R \equiv G_L$
Geometric Modeling:		
eht-imaging	Paper VII values fixed	None (only analyzes closure products); v_{net} fixed to ALMA measurement.

Note. The circular polarization signal is weak and strongly depends on calibration assumptions, so we summarize the primary differences in calibration assumptions among the methods. For additional details (e.g., priors adopted for fitted values), see the more detailed descriptions in Appendix D.

calibration” to solve for calibration errors; we use the software DIFMAP (Appendix D.1) and polsolve (Appendix D.2) as two implementations of this approach. The second category uses a Bayesian forward modeling approach, jointly solving for both a polarized image raster and residual calibration errors and providing estimates for the posterior distributions of each; we use the software THEMIS and DMC (Appendix D.3). The third category also uses a Bayesian forward modeling approach, but it uses a simple geometric model for the sky image and only fits VLBI “closure” quantities to constrain the circular polarization structure; we use the software eht-imaging (Appendix D.4).

In addition to the differences inherent in each approach as a result of the underlying method (CLEAN vs. raster fitting vs. geometric modeling, Bayesian exploration vs. fitting a single image), our methods face additional choices on how they deal with polarimetric leakage (D -terms) and residual gain ratios ($G_{R/L}$). Constraints on the D -terms were derived and discussed extensively in Paper VII; these constraints have been confirmed with analysis of EHT observations of AGNs in Issaoun et al. (2022) and Jorstad et al. (2023). As a result, some methods (DIFMAP, and m -ring modeling) chose to directly apply the Paper VII D -term results to the data and not treat polarimetric leakage further; in contrast, DMC and THEMIS fully explored uncertainties in the D -terms under flat priors as part of their Bayesian posterior exploration. polsolve performed D -term self-calibration as a part of its imaging procedure using priors motivated by the Paper VII results.

Most significantly for circular polarization, each method had freedom to choose how to approach residual uncertainties in the gain ratios $G_{R/L}$, in both amplitude and phase. Three imaging methods—polsolve, DIFMAP, and DMC—solved for separate G_R and G_L terms as part of their self-calibration (polsolve and DIFMAP) or Bayesian forward modeling (DMC) approach. In contrast, THEMIS did not solve for separate right and left circular gains, but only solved for an overall gain term $G = G_R = G_L$, absorbing any uncertainty in the residual gain ratios into the recovered image structure. The geometric m -ring modeling method did not solve for any gain terms, as it directly fit gain-insensitive closure quantities in right and left circular polarizations independently. These choices in the treatment of the gain ratios can have a large impact on the recovered circular polarization structure and its uncertainty; for instance, while the methods are otherwise similar in their Bayesian approach and treatment of the D -terms, the weak priors on the gain ratios in DMC result in

larger error bars in the recovered Stokes \mathcal{V} structure as compared to THEMIS, which locks all gain ratios to unity.

Before applying our imaging methods to EHT observations of M87*, we first tested each method on synthetic data taken from GRMHD simulation images of M87* on EHT 2017 baselines. We present these tests in Appendix E. Our results on synthetic data suggest that our imaging methods are generally not capable of unambiguously determining the horizon-scale structure of circular polarization in M87*, given the low fractional polarization in the source (and GRMHD models), as well as the poor $u - v$ coverage of the EHT in 2017. Indeed, when turning to real data in Section 4, we also find that our methods are unable to arrive at a single consistent image of the circular polarization in M87*. However, we are able to use the images presented in Section 4 to derive an upper limit for the circular polarization fraction on horizon scales.

4. M87* Imaging Results

4.1. Imaging Results on Individual Days and Bands

Each imaging method introduced in Section 3 was used to produce Stokes \mathcal{I} , \mathcal{Q} , \mathcal{U} , and \mathcal{V} images from the eight individual 2017 EHT data sets of M87*, corresponding to the four observation days in both low and high band. Imaging methods were free to use data that were precalibrated for the zero-baseline D -terms derived in Paper VII or not, and to solve for residual $G_{R/L}$ gain errors or assume they are fixed to unity in self-calibration. The choices made by each imaging method are summarized in Table 1. Of the images presented here, the DMC and THEMIS posteriors for the M87* data are identical to those already presented for linear polarization in Paper VII. The polsolve results were generated with a similar procedure to the linear polarization results in Paper VII, but with an additional series of imaging and self-calibration for recovering CP and the $G_{R/L}$ offsets. The DIFMAP and m -ring modeling results are new to this work.

Figure 3 shows the April 11 low-band reconstructions in both total intensity and linear polarization in the top row (in the style of Paper VII) and in an “ellipse-plot” representation of the total linear plus circular polarization in the bottom row. The “ellipse plots” illustrate the degree of linear polarization relative to circular polarization by the eccentricity of small ellipses plotted across the image. As in Paper VII, all imaging methods recover consistent images of total intensity and linear polarization. The $\approx 40 \mu\text{as}$ diameter ring structure is recovered in all methods, as is the $\approx 15\%$ peak linear polarization fraction and predominantly azimuthal EVPA pattern in the southwest

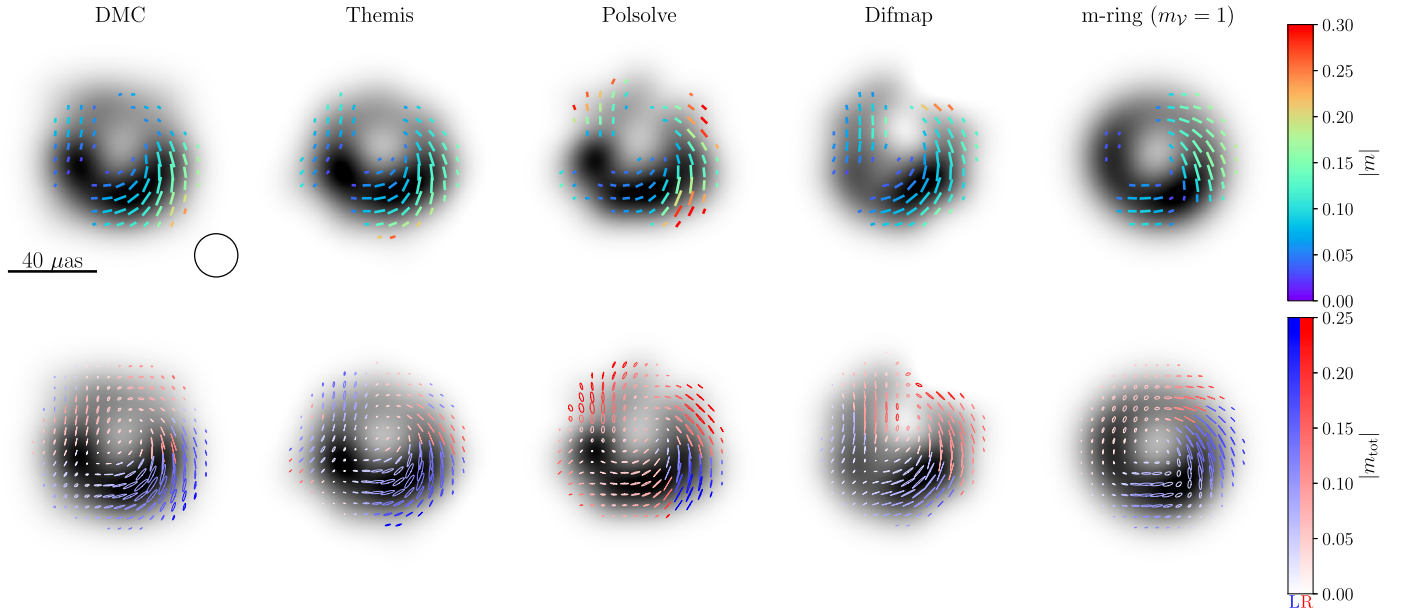


Figure 3. Reconstructions of 2017 EHT M87* data from April 11, low band. The top row shows total intensity images from all reconstruction methods in gray scale and fractional linear polarization in colored ticks as in Paper VII. The second row shows the same grayscale total intensity image overlaid with colored ellipses indicating the *total* polarization fraction $|m_{\text{tot}}| = \sqrt{Q^2 + U^2 + V^2}/I$. The size of each ellipse indicates the total polarized brightness; the orientation of each ellipse indicates the linear EVPA, and axis ratio indicates the relative fraction of circular polarization. The color of each ellipse indicates the sign of circular polarization.

part of the image. Minor differences between images in the fractional linear polarization appear at the edges of the ring and in the m -ring pattern, which is limited by a small number of degrees of freedom in the $m = 3$ mode fit in linear polarization. The ellipse plots in the bottom row show that in all cases the circularly polarized brightness recovered is a small fraction ($\lesssim 20\%$) of the total polarized brightness, which is indicated by the large axis ratio/eccentricity of each ellipse. The colors of each ellipse show the sign of circular polarization recovered at each point. For April 11 low band there is a consistent negative sign of V at the total intensity maximum in the southwest, but across the rest of the image there is little consistency between methods in the sign of V .

In Figures 4 and 5 we focus on the results of our EHT 2017 M87* circular polarization imaging and modeling results by showing circular polarization images within contours indicating the I brightness. In the color map chosen, red corresponds to a positive sign of CP, while blue corresponds to negative CP. As in the synthetic data tests, different methods show consistent structure in total intensity but significantly different structure in circular polarization. The imaging results are most consistent for low-band data, where on most days most methods recover negative circular polarization at and near the total intensity maximum. In general, however, the Stokes V structures across the image are not consistent from method to method. Furthermore, the V images are not consistent between bands when imaged with the same method. An exception is m -ring modeling, which consistently indicates an approximately north–south asymmetry, with more negative circular polarization in the south. These results must be interpreted carefully, however, as strong assumptions on the source structure are imposed in the choice of a simple $m = 1$ model for fitting the circular polarization. The m -ring results are discussed in more detail in F. Roelofs et al. (2023, in preparation, hereafter R23).

To test whether the observed inconsistency in the recovered Stokes V images is mitigated by a different calibration strategy, we produced images with a subset of methods using data on April

5 and April 11 reduced with the CASA-based `rPICARD` pipeline (Janssen et al. 2019). We made use of a new `rPICARD` calibration mode, where instrumental phase and delay offsets are solved initially to align the RR^* , RL^* , LR^* , and LL^* phases of the high and low bands. Subsequently, all fringe-fitting and phase calibration solutions are obtained from the combined data of the four correlation products and the two frequency bands. Additionally, unlike the data reduced with the EHT-HOPS pipeline used for Figures 4 and 5, where a multiday and multisource fit is employed, the CASA data utilized here are calibrated with $G_{R/L}$ assuming $V = 0$. The station-based amplitude gains are solved every few minutes and do not affect closures. We note that the $R - L$ gain calibration strategies employed for the HOPS/CASA data will likely underfit/overfit instrumental gain offsets. The images produced from the CASA data still show inconsistency across imaging methods and observing days and do not change our main conclusions based on the EHT-HOPS images, which we adopt as fiducial for the rest of the paper. We show full results of this test in Appendix F.

The general inconsistency among methods and across bands is in sharp contrast to the total intensity images presented in Paper IV and linear polarization images from Paper VII. This inconsistency is a result of the severe difficulties in recovering resolved circular polarization from sparse 2017 EHT observations with low S/N and $G_{R/L}$ calibration uncertainties.

4.2. Combining Days and Bands

Given the lack of consistency in the Stokes V reconstructions presented in Figures 4 and 5 for individual EHT data sets, it is natural to wonder whether by averaging data across frequency bands and in time we may increase S/N enough to more confidently recover circular polarization structure. A subset of our imaging methods tested this hypothesis. Figure 6 presents imaging results from DMC, THEMIS, and `DIFMAP` on data sets combining high- and low-band EHT observations of M87* on two pairs of days: 2017 April 5 and 6 (top row) and April 10 and 11 (bottom

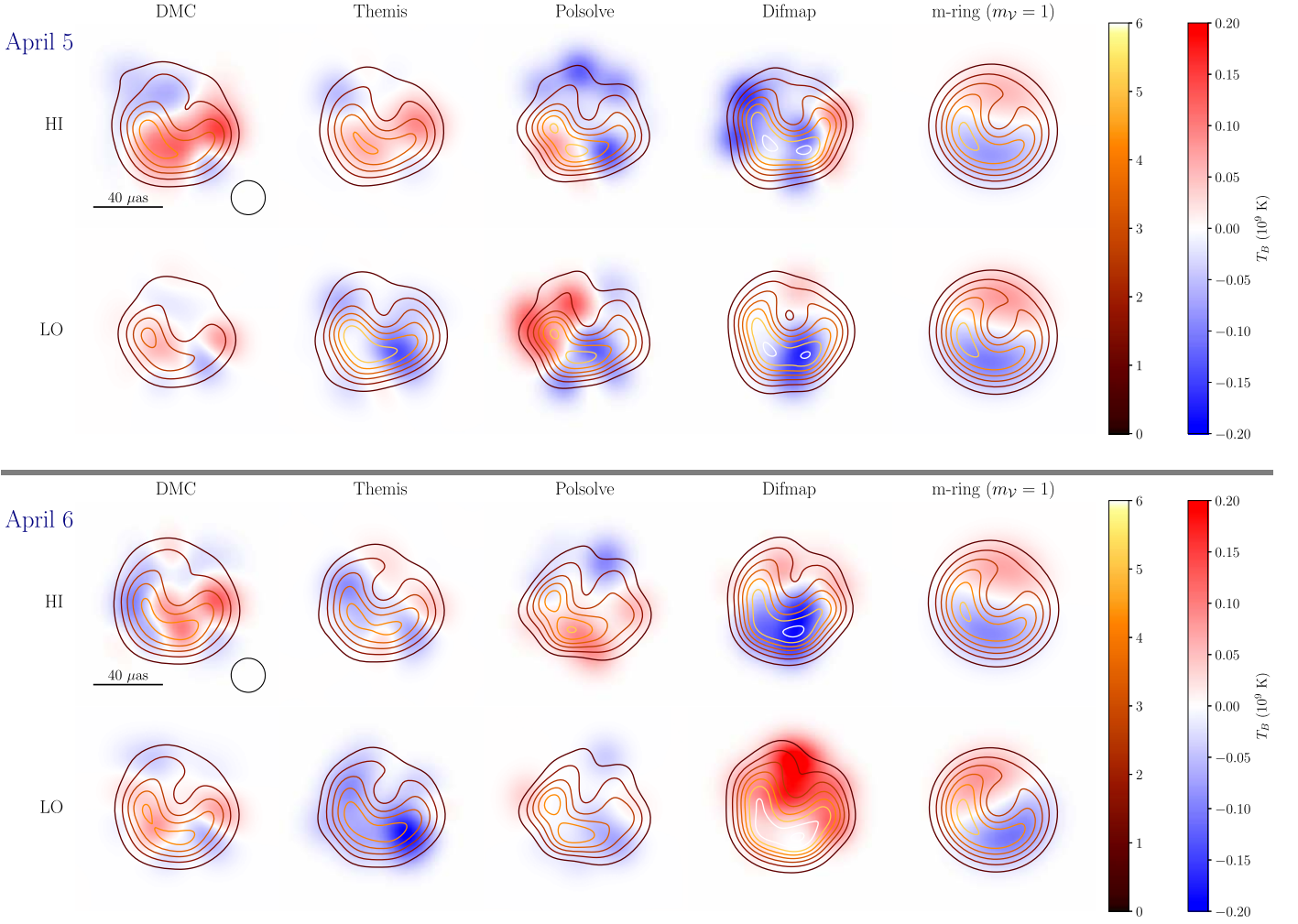


Figure 4. Circular polarization imaging results from 2017 EHT observations of M87* on April 5 (top two rows) and April 6 (bottom two rows). Images of circular polarization on these consecutive days are expected to be nearly identical, as is seen in total intensity and linear polarization. The top/bottom row in each pair shows results from imaging the high/low-band data. In each panel, total intensity is indicated in the colored linear-scale contours, and the Stokes V brightness is indicated in the diverging color map, with red/blue indicating a positive/negative sign. The color bar ranges are fixed in both plots (and in Figure 5). For posterior exploration methods (DMC, THEMIS, m -ring fitting), the posterior-average image is shown. All images are blurred with the same $20 \mu\text{as}$ FWHM Gaussian, shown with the black inset circle in the upper left panels.

row). The source structure in M87* evolved slightly over the week of observations in 2017 but was stable on each pair of days combined here (Paper IV, Paper VII). As in Figures 4 and 5, we find no consistency in the reconstructions from the band- and day-averaged data. While the S/N is improved by a factor of two by averaging four data sets in each combined reconstruction, the low intrinsic circular polarization and residual $G_{R/L}$ uncertainty are still too large and result in inconsistent image reconstructions from the combined data. We focus on results derived from the individual day/band images in the rest of this work.

4.3. Upper Limit on the Resolved Circular Polarization Fraction

We quantify the average circular polarization fraction in resolved EHT images by the image-averaged fractional circular polarization magnitude $|\mathcal{V}/\mathcal{I}|$, weighted by the Stokes \mathcal{I} intensity:

$$\langle |v| \rangle = \frac{\int |\mathcal{V}/\mathcal{I}| \mathcal{I} dA}{\int \mathcal{I} dA}, \quad (5)$$

where the integral is over the whole area of the image \mathcal{A} . The definition of $\langle |v| \rangle$ in Equation (5) is in close analogy with the average resolved linear polarization fraction $\langle |m| \rangle$ used in Papers VII and VIII.

We contrast this average polarization fraction on EHT scales with the unresolved circular polarization fraction v_{net} ,

$$v_{\text{net}} = \frac{\int \mathcal{V} dA}{\int \mathcal{I} dA}. \quad (6)$$

By definition, $\langle |v| \rangle$ depends on the image resolution, while v_{net} is invariant to convolution with a blurring kernel. The definition of $\langle |v| \rangle$ naturally down-weights contributions from regions where the total intensity image \mathcal{I} is dim and noisy. However, $\langle |v| \rangle$ is by definition constrained to be positive and thus cannot indicate the predominant sign of circular polarization in an image. For weakly polarized/noisy images, $\langle |v| \rangle$ will also be biased to nonzero values (see Appendix G for a discussion).

We chose to use $\langle |v| \rangle$ instead of alternative metrics like \mathcal{V}/\mathcal{I} at the peak brightness location because our image

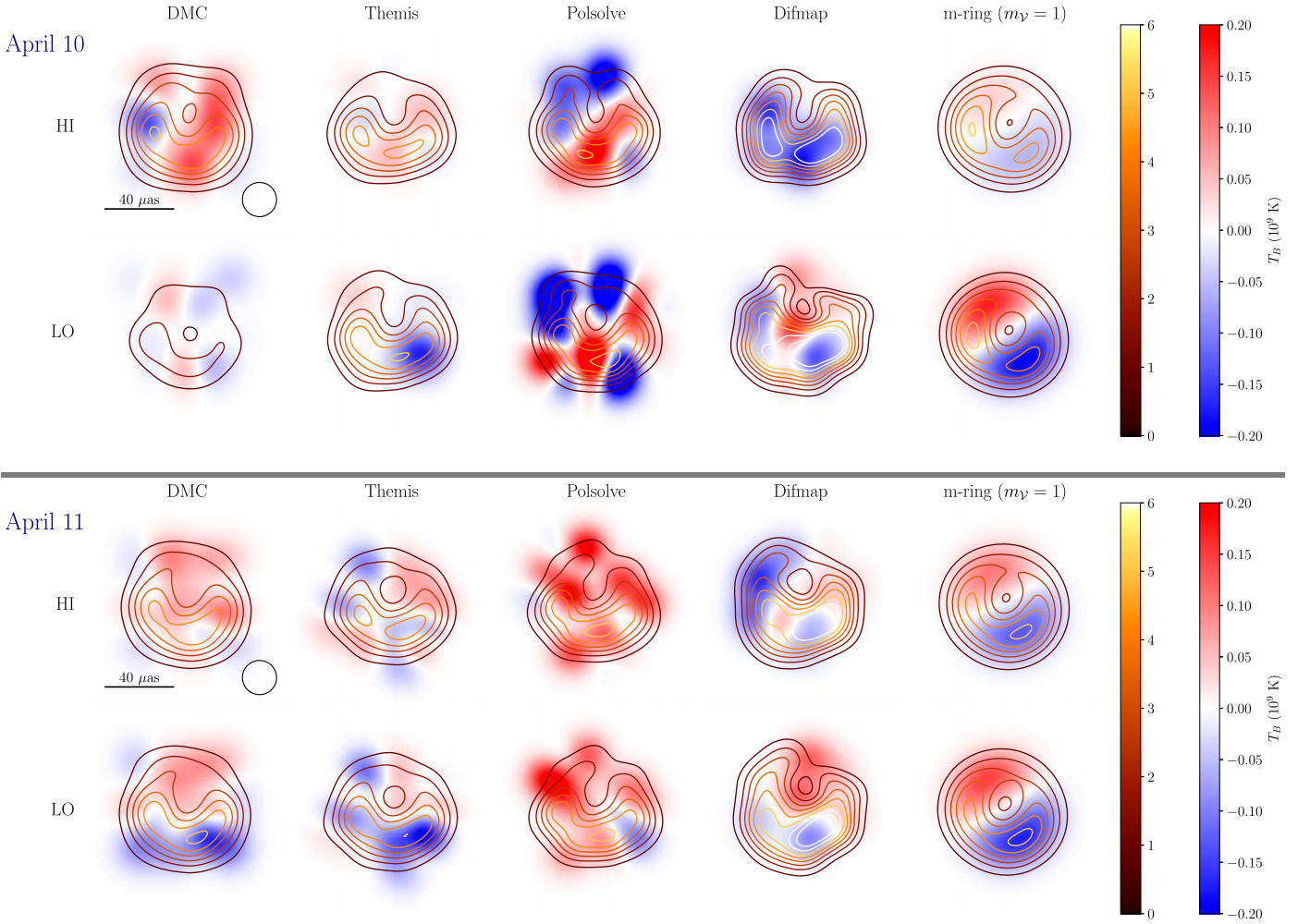


Figure 5. The same as Figure 4, but for 2017 EHT observations of M87* on April 10 (top two rows) and April 11 (bottom two rows).

reconstructions and GRMHD synthetic images often do not show a single peak in circular polarized intensity. Furthermore, image-integrated metrics are preferable for computing sensible posterior means and error bars from the results of Bayesian methods like DMC, THEMIS, and m -ring model fitting.

Figure 7 shows the image-integrated and resolved circular polarization fractions for all reconstruction methods, days, and bands of EHT 2017 M87* data. For the posterior exploration methods (DMC, THEMIS, and m -ring modeling), we plot the posterior median of each quantity and error bars corresponding to the 2.27th and 97.7th percentile of the posterior distribution (corresponding to 2σ error bars if the posterior were Gaussian). For DIFMAP and *polsolve*, we plot the single value corresponding to the image results in Figures 4 and 5; we derive approximate 2σ error bars for these methods based on the measured off-source residuals in \mathcal{V} and \mathcal{I} using standard Gaussian error propagation. DMC has particularly large error bars on fractional circular polarization because of its permissive priors on the relative gain ratios $G_{R/L}$.

The values of the integrated circular polarization fraction v_{net} in Figure 7 recovered by each imaging method are typically within the upper limits reported by Goddi et al. (2021) from ALMA observations, though some methods produce anomalously larger integrated polarization fractions on certain data sets (e.g., *polsolve* on both April 11 data sets, DIFMAP and THEMIS on April 6 low-band data). Our methods do not recover

a consistent sign of the integrated circular polarization fraction v_{net} .

The values of the resolved circular polarization fraction at $20 \mu\text{as}$ resolution $\langle |v| \rangle$ for most reconstructed images in the right panel of Figure 7 are typically less than 4%. The synthetic data results in Figure 20 indicate that for 2017 EHT coverage and sensitivity most methods will always produce $\langle |v| \rangle \gtrsim 1\%$, even when the actual value is lower, as a consequence of uncertainties in the reconstruction in the low-S/N regime, uncertainty in the residual gain ratios $G_{R/L}$, and an upward bias on the quantity $\langle |v| \rangle$ (Appendix G). As a result, and because of the lack of agreement among methods in recovering a consistent source structure in \mathcal{V} in Figures 4 and 5, we use these results only to obtain a conservative upper limit for the resolved circular polarization fraction.

We estimate a combined upper limit on $\langle |v| \rangle$ in M87* from each method’s eight measurements across the 4 days and two observing bands. We average the measurements of $\langle |v| \rangle$ within each method across these eight data sets using standard inverse-variance weighting and then compute the 99th percentile of the resulting distribution (assumed to be Gaussian). These 99% upper limits for each method are reported in Table 2.

We adopt the DMC results as the most conservative upper limit on $\langle |v| \rangle$. The upper limit computed from *polsolve* is slightly higher than that from DMC (3.8% vs. 3.7%). Nonetheless, we adopt the DMC value as our fiducial upper

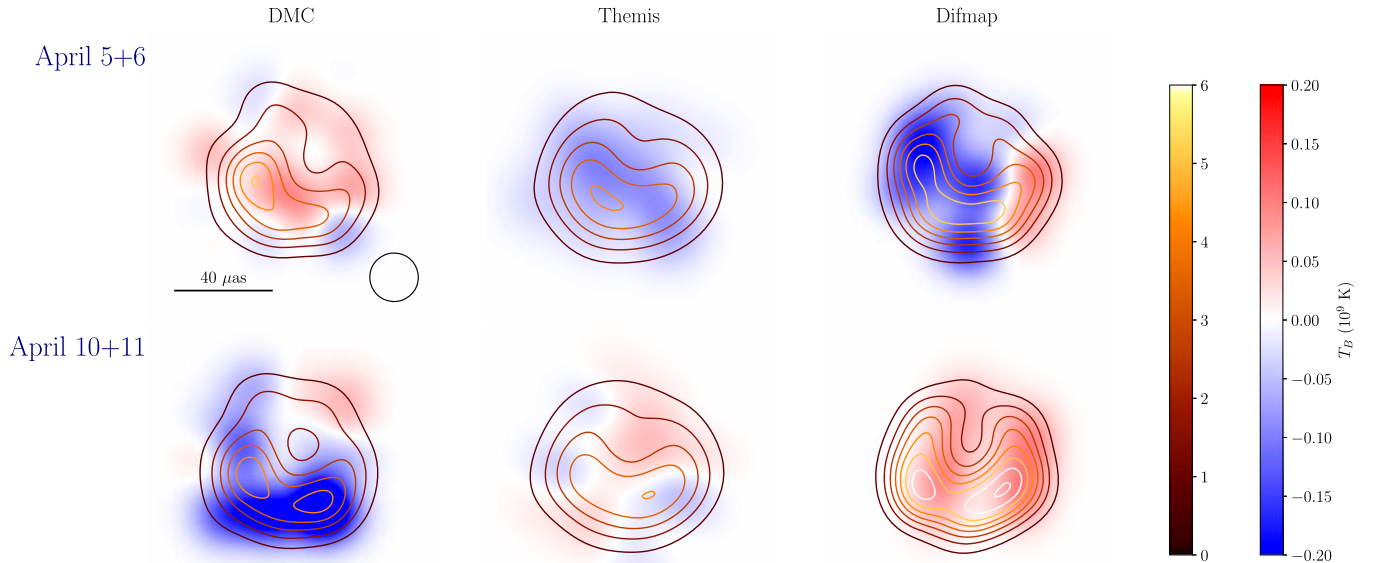


Figure 6. M87* imaging results combining days and bands. The top row shows results from three methods on a data set combining April 5 and 6 observations, both low and high band. The bottom row shows corresponding results from combining April 10 and 11 observations, low and high band. For THEMIS and DMC we show the posterior-average image. Images are plotted in the same manner as in Figure 4.

limit in this work. DMC performs full Bayesian posterior exploration over the image intensity values and station gains. In contrast, `polsolve` computes only a single image with error bars computed using idealized Gaussian error propagation. None of our interpretation in Section 5 is affected by choosing 3.8% versus 3.7% as our fiducial upper limit. For these reasons, we adopt the DMC value $\langle |v| \rangle < 3.7\%$ going forward as our conservative upper limit on the average resolved horizon-scale circular polarization fraction in M87* at 230 GHz.

5. Theoretical Interpretation

On event horizon scales, circularly polarized images of hot accretion flows can encode valuable information about the plasma, including its magnetic field geometry and composition. While our imaging methods are unable to unambiguously determine the horizon-scale structure of circular polarization, in Section 4 we establish an upper limit on the magnitude of the fractional circular polarization magnitude on scales of the EHT beam: $\langle |v| \rangle < 3.7\%$. This upper limit, combined with the existing limit on the unresolved, source-integrated circular polarization fraction from ALMA ($|v_{\text{net}}| < 0.8\%$), can be used to constrain models of the emitting plasma and accretion flow around M87*, *even without* additional information on the structure of the circularly polarized emission.

Recent work has revealed that intrinsically circularly polarized synchrotron emission, Faraday conversion and rotation, and twisted field geometries are all important in generating the CP image that we observe in millimeter wavelengths (Tsunetoe et al. 2020, 2021, 2022; Mościbrodzka et al. 2021; Ricarte et al. 2021). For the plasma parameters of interest in M87*, millimeter synchrotron emission is intrinsically circularly polarized only at the $\sim 1\%$ level. However, Faraday conversion also plays a role in exchanging linear and circular polarization states, and in fact dominates the Stokes \mathcal{V} production in many models. Stokes \mathcal{V} generated by Faraday conversion is understood to be sensitive to the magnetic field geometry, which has been utilized to infer helical field structure

in jets (e.g., Wardle & Homan 2003; Gabuzda et al. 2008). On event horizon scales, the connection between Stokes \mathcal{V} and the magnetic field geometry is complicated by the geometries probed by geodesics that take nontrivial paths through the spacetime, leading in particular to Stokes \mathcal{V} sign flips of successive subimages in some models (Mościbrodzka et al. 2021; Ricarte et al. 2021).

As with the previous papers in this series, our theoretical interpretation centers on ray-traced GRMHD simulations, which self-consistently generate the plasma that performs emission, absorption, and Faraday effects. Since images of total intensity and linear polarization were studied in detail in Papers V and VIII, respectively, here we focus mainly on the astrophysics governing the generation of circular polarization.

5.1. One-zone Models

In Papers V and VIII, we used one-zone isothermal sphere models to derive order-of-magnitude estimates on the parameters of the synchrotron emitting plasma in M87*. In particular, applying constraints on the linear polarization, Paper VIII found that one-zone models for M87* imply that the dimensionless electron temperature $\Theta_e = k_B T_e / m_e c^2$ lies in a mildly relativistic regime, $2 < \Theta_e < 20$, the magnetic field is in the range $1 \text{ G} \lesssim |B| \lesssim 20 \text{ G}$, and the number density lies in the range $10^4 \text{ cm}^{-3} \lesssim n_e \lesssim 10^7 \text{ cm}^{-3}$. While one-zone models neglect the critical effects of fluid velocity, gravitational redshift, and a nonuniform emitting region, they favor plasma parameters that are in general agreement with those found in favored GRMHD simulations (Paper VIII).

Here we explore some implications for circular polarization in the one-zone model space selected in Paper VIII. We use the same model of a uniform $R = 5r_g$ isothermal sphere¹⁵⁷ with a magnetic field at a pitch angle $\theta_B = \pi/3$ to the line of sight. We

¹⁵⁷ For M87*, adopting $M_{\text{BH}} = 6.5 \times 10^9 M_\odot$, the gravitational radius $r_g = GM_{\text{BH}}/c^2 = 9.6 \times 10^{14} \text{ cm} = 64 \text{ au}$.

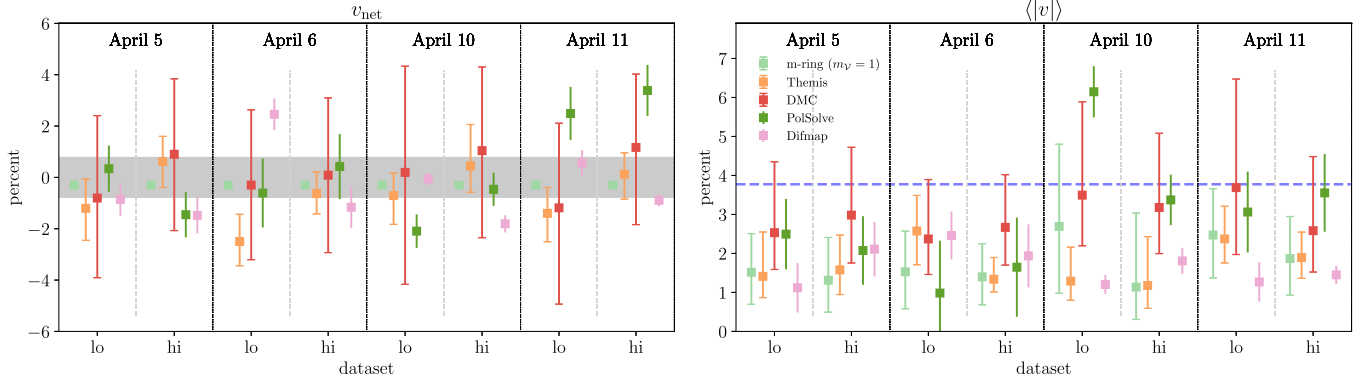


Figure 7. Image-integrated statistics from M87* images. The left panel shows the net circular polarization fraction v_{net} computed from each method for the eight EHT M87* data sets, and the right panel shows the average resolved circular polarization fraction $\langle |v| \rangle$ computed after blurring each image with a $20 \mu\text{as}$ kernel. The results from the posterior exploration methods are presented with the median value and 2σ error bars (note that *m*-ring modeling strictly enforces $|v_{\text{net}}| = 0$). The DIFMAP and *polsolve* results are derived from the single fitted image and standard error propagation from measurements of the off-source image rms in \mathcal{V} and \mathcal{I} . In the left panel, the limits on v_{net} from ALMA observations reported in Goddi et al. (2021) (used to constrain GRMHD models in Paper VIII) are indicated in the gray shaded region. In the right panel, the upper limit on $\langle |v| \rangle$ derived in this work ($\langle |v| \rangle < 3.7\%$, Table 2) is indicated by the dashed horizontal line.

Table 2
99th Percentile $\langle |v| \rangle$ Upper Limits from Each Method

Method	Combined Upper Limit on $\langle v \rangle$
DMC	3.7%
THEMIS	2.0%
<i>polsolve</i>	3.8%
DIFMAP	1.9%
<i>m</i> -ring	2.2%

Note. We derive each upper limit using inverse-variance weighted averaging of the results in the right panel of Figure 7. We adopt the conservative upper limit, $\langle |v| \rangle < 3.7\%$ from DMC (in bold), as our fiducial value. Note that in computing the upper limit from *polsolve*, we exclude the anomalously highly polarized April 10 low-band result, as it suffers from an anomalously large overall offset in the recovered $G_{R/L}$.

determine polarized emissivities $j_{\mathcal{I}}, j_{\mathcal{Q}}, j_{\mathcal{V}}$; absorption coefficients $\alpha_{\mathcal{I}}, \alpha_{\mathcal{Q}}, \alpha_{\mathcal{U}}$; and Faraday rotation/conversion coefficients $\rho_{\mathcal{V}}, \rho_{\mathcal{Q}}$ using the fitting functions for relativistic thermal electron distributions in Dexter (2016).

First, we consider the importance of Faraday conversion in producing the observed circular polarization in these models. The left panel of Figure 8 shows the distribution of the Faraday conversion optical depth $\tau_{\rho_{\mathcal{Q}}} = 2R \rho_{\mathcal{Q}}$ in one-zone models that pass the Paper VIII constraints. In all cases $\tau_{\rho_{\mathcal{Q}}} > 1$, indicating that Faraday conversion dominates intrinsic emission in producing circular polarization.

In the middle panel of Figure 8, we consider the ratio of the Faraday rotation to conversion optical depth, $\tau_{\rho_{\mathcal{V}}}/\tau_{\rho_{\mathcal{Q}}}$. In all passing models, this ratio is greater than unity, as we expect for a mildly relativistic plasma. As a consequence, these one-zone models typically produce *more* circular polarization than linear polarization. This is because in one-zone models where Faraday effects are significant ($\tau_{\rho_{\mathcal{Q}}} > 1, \tau_{\rho_{\mathcal{V}}} > 1$) but absorption is insignificant ($\tau_{\mathcal{I}} < 1$), the ratio of circular to linear polarized intensity in the limit of large Faraday depth approaches the ratio of the rotativities: $|\mathcal{V}|/|\mathcal{P}| \rightarrow |\rho_{\mathcal{V}}|/|\rho_{\mathcal{Q}}|$ (see Appendix C of Dexter 2016, for exact one-zone solutions in this limit). As a result, our one-zone models with uniform magnetic field orientations overproduce circular polarization relative to linear

polarization, and we must invoke effects not considered in the one-zone models to explain the observed $|\mathcal{V}| < |\mathcal{P}|$.

We consider the effects of one such complication—a spatially twisted magnetic field—in the right panel of Figure 8. Faraday conversion only converts linear polarization in the \mathcal{U} Stokes parameter into \mathcal{V} , so a spatial rotation of the projected *B*-field can work with or against Faraday rotation to enhance or suppress the produced circular polarization. We add a constant rotation of the magnetic field direction at a spatial frequency ω to our one-zone models, and we solve for the final circular polarization fraction \mathcal{V}/\mathcal{I} as a function of ω in three models: one with a low conversion depth ($\tau_{\rho_{\mathcal{Q}}} \approx 1$: $|B| = 11$ G, $\Theta_e = 7$, $n_e = 10^4 \text{ cm}^{-3}$), one with a median conversion depth ($\tau_{\rho_{\mathcal{Q}}} \approx 5$: $|B| = 5$ G, $\Theta_e = 8$, $n_e = 10^5 \text{ cm}^{-3}$), and one with a large conversion depth ($\tau_{\rho_{\mathcal{Q}}} \approx 200$: $|B| = 9$ G, $\Theta_e = 3$, $n_e = 4 \times 10^6 \text{ cm}^{-3}$).

For moderate rates of field rotation in the plane of the sky, only the model with the highest Faraday depth produces a constant circular polarization fraction at the values of ω considered. The other models produce \mathcal{V}/\mathcal{I} that varies rapidly and changes sign as the rotation rate ω passes through the critical frequency $\omega_{\text{crit}} = \rho_{\mathcal{V}}/2$. Thus, field twist through an inhomogeneous emission region can have a significant impact on both the magnitude and sign of the observed circular polarization (e.g., Mościbrodzka et al. 2021; Ricarte et al. 2021). We also show results for the same models artificially setting the intrinsic circular polarization emission to zero ($j_{\mathcal{V}} = 0$, dashed lines). In all three cases, this change has a minor effect on the produced circular polarization fraction, again confirming that Faraday conversion dominates the production of circular polarization for plasma parameters appropriate for M87*.

The fact that one-zone models overproduce \mathcal{V} relative to linear polarization and the fact that a changing field geometry in the emission region has significant impact on \mathcal{V}/\mathcal{I} motivate consideration of more complex models (see, e.g., Goddi et al. 2021, for a two-zone model of M87*'s Faraday rotation). We proceed next to consider circular polarization in GRMHD simulation images, which self-consistently include the effects of an inhomogeneous emission region, field twist, and special and general relativistic redshift and parallel transport.

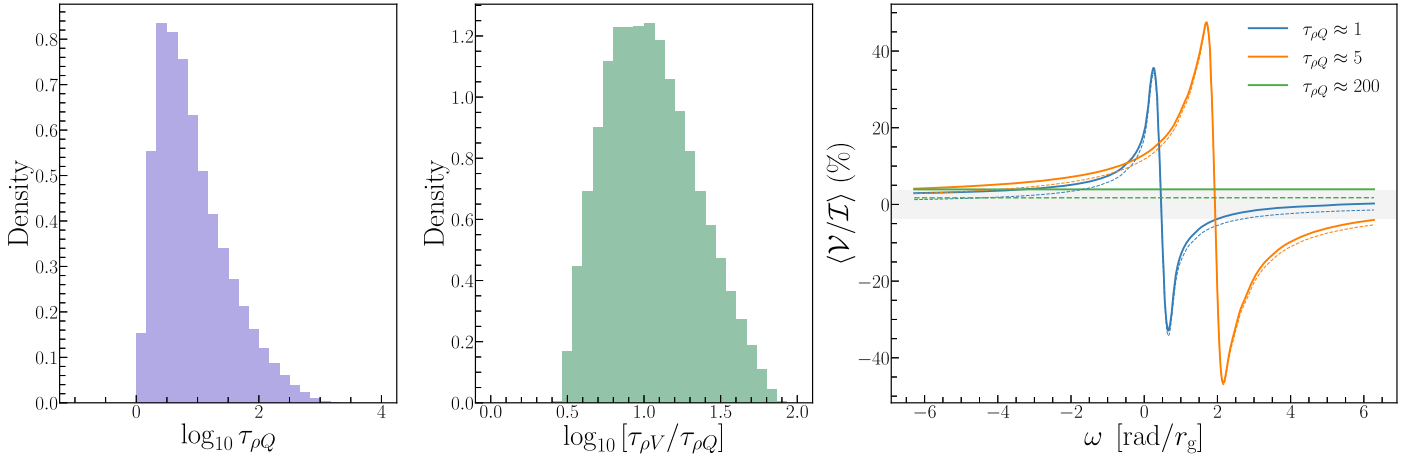


Figure 8. Circular polarization properties of passing Paper VIII one-zone models. Left: distribution of the Faraday conversion optical depth $\tau_{\rho Q}$ in passing models. In all passing models $\tau_{\rho Q} > 1$, indicating that most circular polarization is likely produced by Faraday conversion. Middle: distribution of the ratio of the Faraday rotation to Faraday conversion optical depths. In all cases, $\tau_{\rho V} > \tau_{\rho Q}$, indicating that with a constant field orientation in the emission region, circular polarization will dominate over linear polarization in these models. Right: the average fractional circular polarization between $5r_g$ and $10r_g$ in one-zone models with a rotating magnetic field direction along the line of sight, as a function of the angular rotation frequency ω . We show three different models: a model with low Faraday conversion depth (blue), a model with median conversion depth (orange), and a model with high conversion depth (green). Dashed lines show corresponding results for one-zone models with no intrinsic emission of circular polarization, $j_V = 0$.

5.2. GRMHD Image Libraries

The 3D GRMHD simulations used in this work were first considered in Papers V and VIII. We use the code IPOLE (Noble et al. 2007; Mościbrodzka & Gammie 2018) to perform general relativistic radiative transfer (GRRT) following the methodology outlined in Paper VIII and Wong et al. (2022). We briefly summarize our library generation here, and we refer to Appendix H and these previous works for more details.

All images assume a fixed observing frequency of 230 GHz, BH mass of $M_{\text{BH}} = 6.2 \times 10^9 M_\odot$, and distance of 16.9 Mpc.¹⁵⁸ The fluid density in the simulations is scaled to reproduce an average flux density of $F_\nu = 0.5$ Jy for each model, following the observed compact flux density in 2017 April (Paper IV). The observing inclination is tilted such that the approaching jet (parallel to the BH spin axis) is inclined at 17° with respect to our line of sight.

Our simulation library probes five free parameters: (i) the magnetic field state, either a strong field “magnetically arrested disk” (MAD) (BisnovatyiKogan & Ruzmaikin 1974; Igumenshchev et al. 2003; Narayan et al. 2003) or weak field “standard and normal evolution” (SANE; Narayan et al. 2012; Sądowski et al. 2013); (ii) the BH spin $a_* \in \{-0.94, -0.5, 0, 0.5, 0.94\}$, where a negative sign denotes retrograde accretion; (iii)–(iv) $R_{\text{high}} \in \{1, 10, 20, 40, 80, 160\}$ and $R_{\text{low}} \in \{1, 10\}$, which modulate the ion-to-electron temperature ratio in different regions; and (v) the magnetic field polarity, which is either aligned or anti-aligned with respect to the disk angular momentum on large scales. Each model is imaged at a cadence of $5 GM/c^3$ for a total duration ranging from 2500 to $5000 GM/c^3$ depending on the model.

In total, we compute 184,796 image snapshots from 10 GRMHD simulations. For each snapshot we compute a set of

polarimetric observables to compare with the data and score our models. Figure 9 displays a random selection of snapshots visualized in circular polarization as a function of spin, magnetic field state, and R_{high} . All images have been rotated such that the oncoming jet is projected at a position angle of 288° east of north. In the first three rows, we visualize the images in Stokes \mathcal{V} at high resolution in symmetric logarithmic scale with three decades in dynamic range. We find a wide variety of morphologies in near-horizon Stokes \mathcal{V} images, including sign flips in almost every snapshot. Some images, such as the SANE $a_* = 0$ models, exhibit “noise-like” regions with rapid sign flips among adjacent pixels. This is equivalent to randomly oriented ticks in linear polarization: circular polarization in these regions is scrambled owing to either large Faraday rotation¹⁵⁹ or Faraday conversion depths. In the bottom three rows, we visualize these snapshots in linear scale with Stokes \mathcal{I} contours, blurred with a Gaussian with a $20 \mu\text{as}$ FWHM. At EHT resolution, GRMHD simulations again predict a wide variety of morphologies, including ubiquitous sign flips.

5.3. The Importance of Magnetic Field Polarity

This is the first paper in this series to consider the magnetic field polarity as a free parameter. We flip the magnetic field polarity in postprocessing using the same GRMHD snapshots, since the equations of ideal GRMHD are invariant to a sign flip in the magnetic field vector. The equations of polarized GRRT are *not* invariant to sign of the magnetic field vector, however. Reversing the polarity of the magnetic field reverses the sign of ρ_V (Faraday rotation) and j_V (circularly polarized emission), but it does not affect ρ_Q (Faraday conversion). Consequently, flipping the magnetic field direction does not necessarily simply reverse the sign of circular polarization across the image.

¹⁵⁹ Although Faraday rotation does not affect circular polarization directly, it can scramble the EVPA of the linear polarization, which is then converted into circular.

¹⁵⁸ Note that the values of the M87* BH mass and distance used in the GRMHD library ($M_{\text{BH}} = 6.2 \times 10^9 M_\odot$, $D = 16.9$ Mpc, following Papers V, VIII) are slightly different from the EHT’s measured value from the ring diameter ($M_{\text{BH}} = 6.5 \times 10^9 M_\odot$, $D = 16.8$ Mpc, Paper VI), which we adopt in other sections of this paper. Because we do not rely on the image size as a constraint on our models, we do not expect the 5% difference in mass and 0.5% difference in distance between the values adopted in our GRMHD simulations and Paper VI measurements to affect our interpretation.

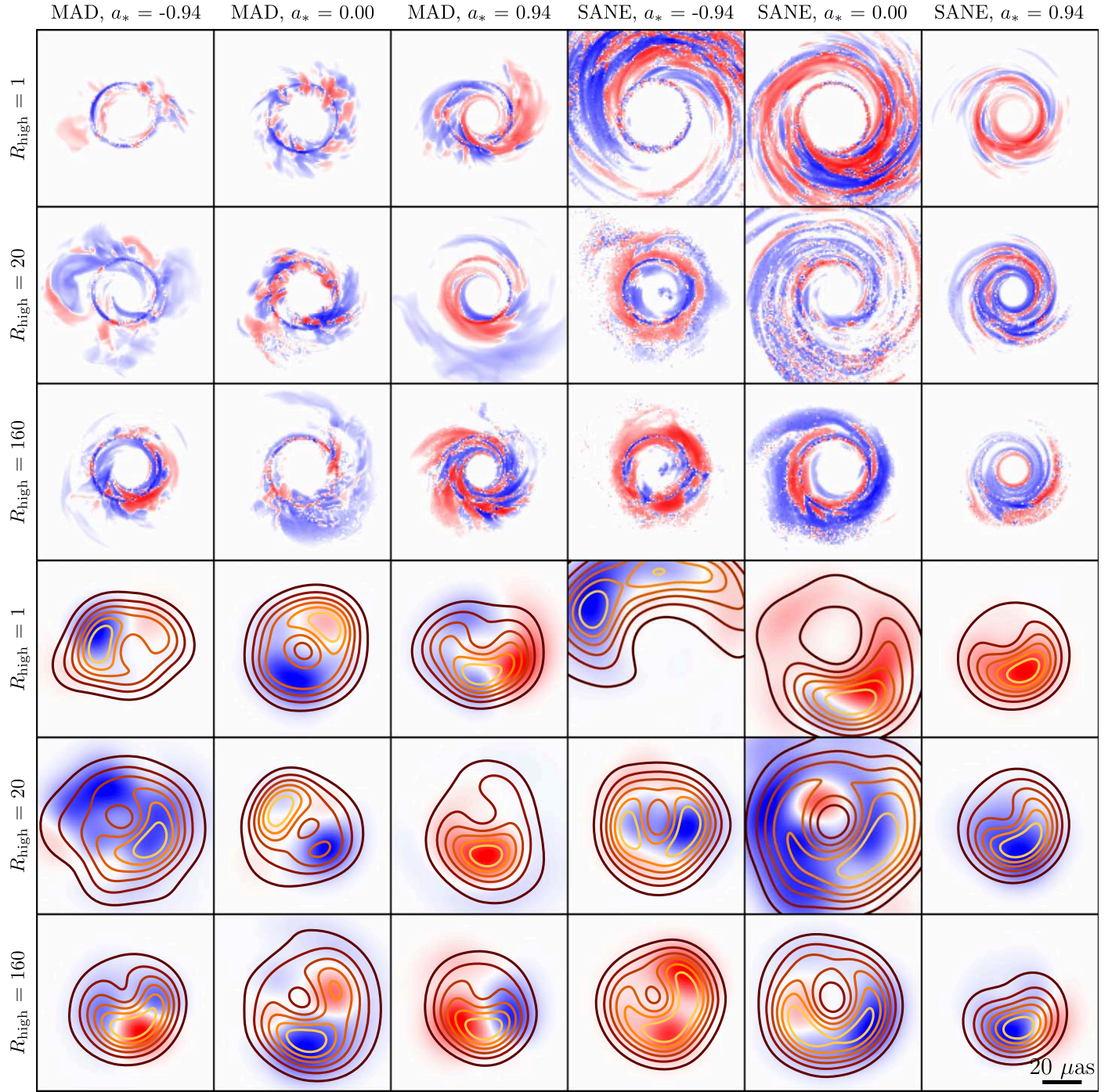


Figure 9. A random selection of representative snapshots from our GRMHD image library. The color scales for each snapshot are normalized individually. The first three rows are presented at native resolution in symmetric logarithmic scale with three decades in dynamic range shown to better visualize faint features. The bottom three rows plot Stokes \mathcal{I} in contours and Stokes \mathcal{V} in color after blurring with a $20 \mu\text{as}$ FWHM Gaussian, both in linear scale. Models exhibit a wide variety of morphologies and almost always show sign reversals, at both perfect resolution and EHT resolution.

Throughout, we describe the magnetic field as “aligned” if its polarity is parallel to the angular momentum of the disk on large scales, or “reversed” if this polarity is antiparallel. Note that the magnetic field structure can be complicated and turbulent in the near-horizon emission region, especially in retrograde disks, so a magnetic field aligned with the angular momentum on large scales is not necessarily trivially aligned on event horizon scales.

In Figure 10, we visualize a snapshot of the MAD $a_* = -0.94$ $R_{\text{low}} = 10$ $R_{\text{high}} = 160$ model with both magnetic

field polarities. The top row is presented in linear scale blurred with a $20 \mu\text{as}$ FWHM Gaussian kernel, while the bottom row is presented in symmetric logarithmic scale with two decades of dynamic range in intensity. With a polarity aligned with the angular momentum of the disk as in previous work, the unresolved circular polarization fraction $v_{\text{net}} = 2.7\%$ vastly exceeds the upper limit of $|v_{\text{net}}| < 0.8\%$ from ALMA observations (Goddi et al. 2021). However, flipping the magnetic field polarity reverses the sign of circular polarization in a significant portion of the image, reducing v_{net} to 0.7% and allowing the

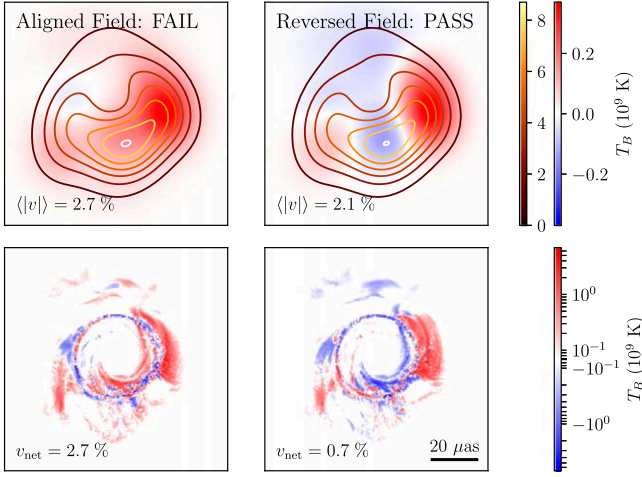


Figure 10. Example GRMHD snapshot ($MAD\ a_* = -0.94$ $R_{low} = 10$ $R_{high} = 160$) plotted with both magnetic field configurations, aligned (left) and reversed (right). The top panels show the images blurred to EHT resolution, and the bottom panels show the images at their native resolution. As shown in the left panels, this snapshot fails simultaneous polarimetric constraints with the aligned-field configuration, overproducing v_{net} . However, as shown in the right panels, flipping the magnetic field polarity produces some oppositely signed regions that reduce $|v_{net}|$. Flipping the field has no effect on the total intensity image.

model to pass. In fact, this snapshot simultaneously passes all polarimetric constraints considered in this work with the reversed magnetic field configuration.

Figure 10 illustrates that it is not easy to predict which regions of a given image change upon a reversal of the magnetic field direction. We expect that regions dominated by intrinsic synchrotron emission should trivially flip sign, while regions dominated by Faraday conversion may remain unchanged unless Faraday rotation is significant along those geodesics. We further explore the effect of flipping the magnetic field polarity on linear polarization in Appendix I. While there are noticeable differences in the distribution of the $\angle\beta_2$ parameter (Palumbo et al. 2020) across all GRMHD models, the effect is less dramatic for linear polarization metrics than it is for circular.

5.4. GRMHD Simulation Scoring

In Paper VIII, five observational metrics were used to score GRMHD models of M87*’s accretion flow: (1) the unresolved linear polarization fraction $|m|_{net}$, (2) the unresolved circular polarization fraction $|v|_{net}$, (3) the image-average linear polarization fraction $\langle|m|\rangle$, (4) the amplitude $|\beta_2|$, and (5) the phase $\angle\beta_2$ of the second azimuthal Fourier coefficient of the linear polarization pattern (see Palumbo et al. 2020; Paper VIII). In this paper, we add one additional constraint to this set from our observational results in Section 4: an upper limit on the circular polarization fraction on EHT scales: $\langle|v|\rangle < 3.7\%$. We summarize the observational constraints used to score GRMHD models in Table 3.¹⁶⁰

In Paper VIII, we found that many GRMHD models that could satisfy both resolved and unresolved linear polarization constraints could also self-consistently satisfy the upper limit on the unresolved circular polarization, $|v_{net}| < 0.8\%$. In the left

Table 3
Observational Constraints Applied to Our GRMHD Image Library

Parameter	Minimum	Maximum
m_{net}	1.0%	3.7%
v_{net}	−0.8%	0.8%
$\langle m \rangle$	5.7%	10.7%
$ \beta_2 $	0.04	0.07
$\angle\beta_2$	−163°	−129°
$\langle v \rangle$ (This Work)	0	3.7%

Note. Most of these constraints are inherited from Paper VII and were previously used to constrain models in Paper VIII. This work adds the new upper limit on $\langle|v|\rangle$.

panel of Figure 11, we plot histograms of v_{net} for our models, now including flipped magnetic field configurations, which were not included in the original analysis of Paper VIII (their Figure 8). Both MAD and SANE models are capable of passing the upper limit on v_{net} , but SANE models are more likely to fail. In the right panel of Figure 11, we plot histograms of the spatially resolved circular polarization $\langle|v|\rangle$, and we overplot our allowed region from the imaging results in Section 4: $\langle|v|\rangle < 3.7\%$. SANE models produce the largest values of $\langle|v|\rangle$, which extend to nearly 10% in some cases. We find that 87% of the images that fail our new upper limit on $\langle|v|\rangle$ are SANE.

We use two different methods for applying observational constraints to our GRMHD images, as in Paper VIII:

1. *Simultaneous* scoring: If a single model snapshot simultaneously passes every polarimetric constraint, then it passes. Otherwise, it is rejected. This method is relatively strict, as for a model to pass at least one image must satisfy all constraints. See Section 5.2 of Paper VIII for more details on the simultaneous scoring procedure.
2. *Joint* scoring: We compute $\chi^2_{j,k}$ statistics for each of the six metrics j for each GRMHD snapshot k (Paper VIII, Equation (17)). We then compute a likelihood \mathcal{L}_j for each metric by calculating the fraction of snapshots where $\chi^2_{j,k} > \chi^2_{j,data}$. The final model likelihood is the product of the individual likelihoods from each metric: $\mathcal{L} = \prod_j \mathcal{L}_j$. This method is relatively lenient, as no single image is required to satisfy all constraints simultaneously. See Section 5.3 of Paper VIII for more details on the joint scoring procedure.

In Paper VIII, these methods produced slightly different results, but both methods favored MAD models.

Our new model scoring results from both methods are shown in Figure 12. This figure is an update to Figure 12 of Paper VIII, updated for both new observational results (the addition of the upper limit on $\langle|v|\rangle$) and new theoretical models (the set of reversed-polarity GRMHD snapshots). As in Paper VIII, we find that MAD models are strongly favored over their SANE counterparts.

In the simultaneous scoring results (left panel of Figure 12), we find that a slight majority of the passing snapshots have aligned magnetic polarities with respect to the disk angular momentum vector. Furthermore, we find that 83% of passing snapshots have magnetic field polarities aligned with the spin of M87*, pointed away from us. This is due not to Stokes \mathcal{V} constraints, but rather to $\angle\beta_2$, for which the distributions shift slightly upon flipping the magnetic field polarity (see Appendix I). While potentially interesting, especially for

¹⁶⁰ Note that in computing these image metrics from GRMHD simulation snapshots, we first blur the model images with a $20\ \mu\text{as}$ FWHM circular Gaussian beam.

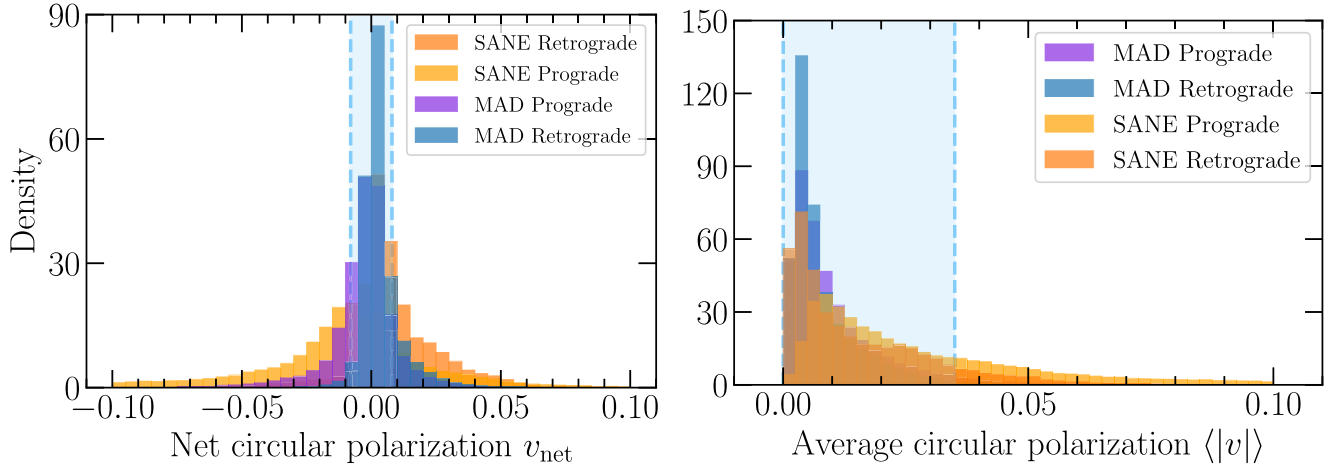


Figure 11. Number density distribution of net circular polarization v_{net} fraction (left panel) and image-averaged fractional circular polarization $\langle |v| \rangle$ (right panel) with an aligned and anti-aligned magnetic field respecting all images in the M87* library. Allowed inferred ranges for ALMA-only data (v_{net}) and measured values of reconstructed polarimetric images of M87* reconstructions ($\langle |v| \rangle$) are limited by the dashed lines.

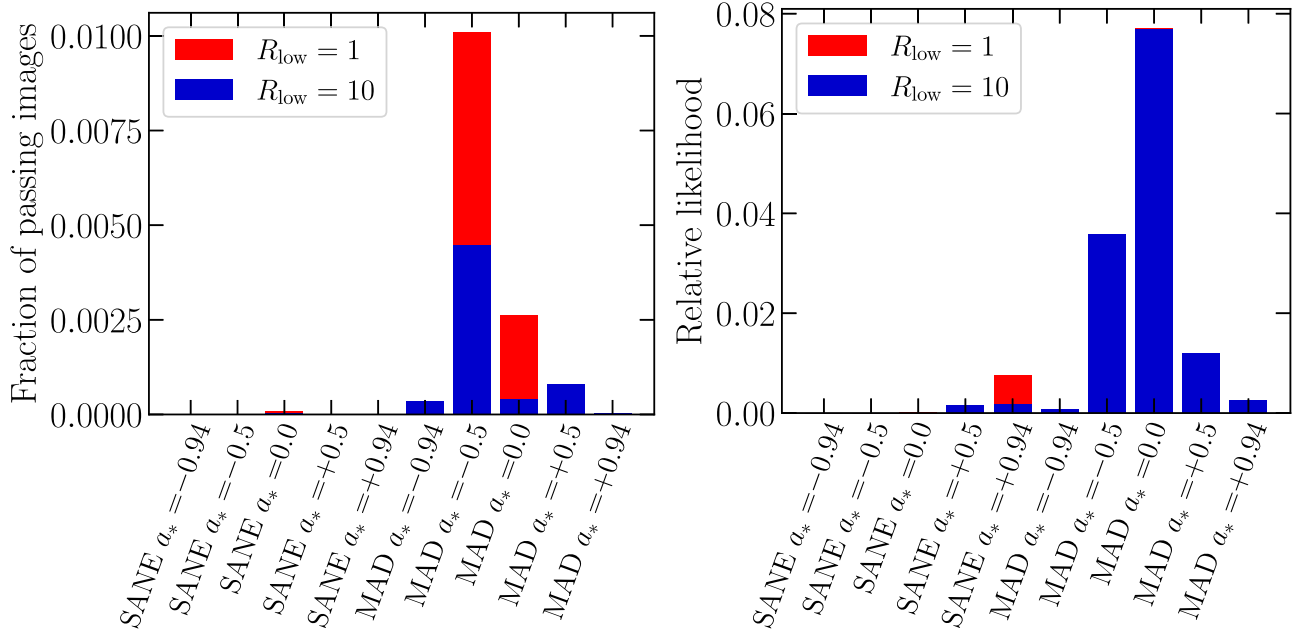


Figure 12. Scoring results for all the models using the simultaneous scoring method (left panel) and the joint distribution method (right panel). Passing fraction or relative likelihoods are summed over R_{high} and \mathbf{B} -field alignment. MADs remain favored across both methods.

constraining the origins of the magnetic field (e.g., Contopoulos & Kazanas 1998; Contopoulos et al. 2022), upon examining Figure 26 in Appendix I, it appears that this preference for an aligned field arises from the fact that $\angle \beta_2$ happens to be Faraday rotated out of the observed range more often in reversed-field models than in aligned-field models for the few spins that we sampled. It is possible that this effect may disappear if spin is sampled more densely, as $\angle \beta_2$ depends strongly on spin in GRMHD models (Palumbo et al. 2020; Emami et al. 2023; Qiu et al. 2023).

Only a small improvement on the upper limit on $\langle |v| \rangle$ would have been necessary to rule out some currently passing snapshots. Among the passing snapshots, we find $0.30\% \leq \langle |v| \rangle \leq 2.8\%$, with a median value of 0.48%. If the upper limit had instead been $\langle |v| \rangle < 1\%$, then 204 snapshots would have passed, down from 288. If the upper limit had been

$\langle |v| \rangle < 0.5\%$, then only 35 snapshots would have passed, all of which would have been MAD.

In Figure 13, we visualize in circular polarization five of the snapshots that pass simultaneous constraints. Each of them is a MAD model, of increasing spin, described in the figure caption. Some, but not all, of the Stokes \mathcal{V} morphologies are dominated by a dipole at EHT resolution. Perfect-resolution images reveal ubiquitous sign reversals within the EHT beam and a rich morphology. Some snapshots feature an inverted photon ring, discussed in detail in Mościbrodzka et al. (2021) and Ricarte et al. (2021).

In Appendix J, we discuss the scoring results for both scoring strategies in more detail. We find that the new constraint on $\langle |v| \rangle$ has no effect on the results of the simultaneous scoring method and only a slight effect on the results from the joint scoring method. The largest driver of

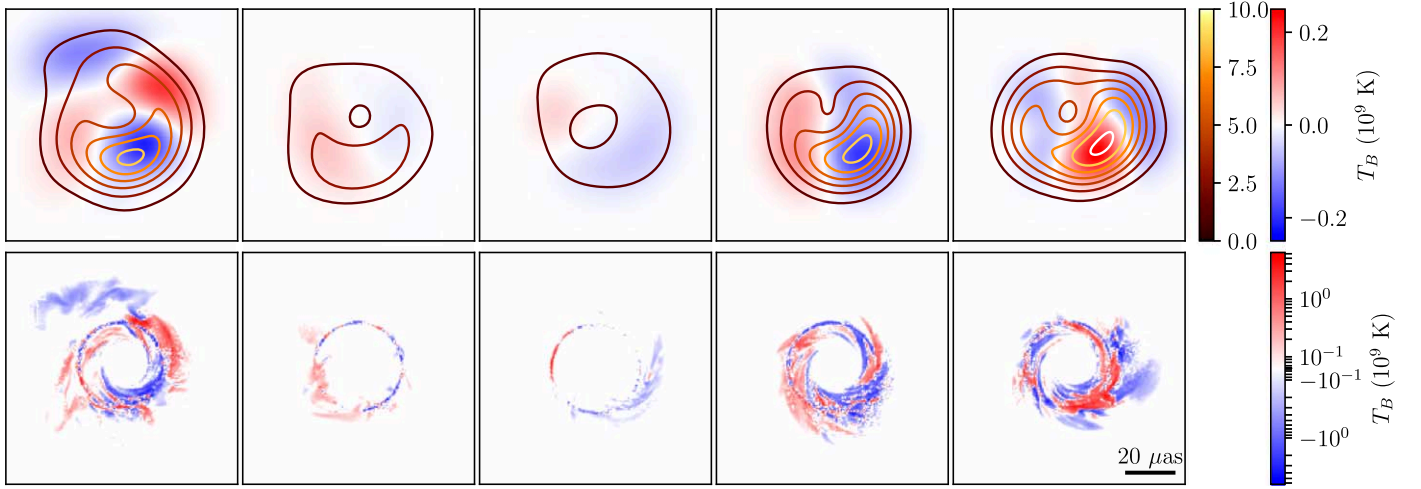


Figure 13. Five passing snapshots visualized in circular polarization at EHT resolution in the top row, and at native simulation resolution in the bottom row in a symmetric logarithmic scale. From left to right, these models correspond to (MAD $a_* = -0.94 R_{\text{high}} = 160 R_{\text{low}} = 10$ reversed field), (MAD $a_* = 0 R_{\text{high}} = 80 R_{\text{low}} = 1$ aligned field), (MAD $a_* = +0.5 R_{\text{high}} = 160 R_{\text{low}} = 10$ aligned field), and (MAD $a_* = +0.94 R_{\text{high}} = 160 R_{\text{low}} = 10$ aligned field). Sign reversals are ubiquitous in these models, even at spatial scales smaller than the EHT beam. Some, but not all, Stokes \mathcal{V} images exhibit a clear dipolar structure.

quantitative differences in the scoring results from Paper VIII is the inclusion of reversed magnetic field models, while updated radiative transfer coefficients and slight differences in the snapshots included in ray-tracing also play a minor role.

5.5. Origins of Stokes \mathcal{V} in Passing Models

To better understand the mechanisms by which Stokes \mathcal{V} is generated in passing GRMHD models, we conducted a series of tests using our radiative transfer code IPOLE. By artificially switching off certain radiative transfer coefficients, we can isolate their effects on the results on the resulting images.

We present these tests in detail in Appendix K. In summary, we find that Faraday conversion is the dominant production mechanism for Stokes \mathcal{V} in passing GRMHD models. Faraday rotation also plays a critical role, as the linear polarization that is later converted into circular polarization is scrambled on scales smaller than the EHT beam, leading to beam depolarization in both the linear and circular polarization images. Intrinsic or direct emission of circular polarization is subdominant to Faraday conversion, though it is not negligible in all models.

5.6. Extensions to GRMHD Models

We caution that, as with the previous papers in this series, constraints on astrophysical models from EHT results come with systematic uncertainties due to limits on our parameter space. As with previous theoretical interpretation papers in this series, we have limited our scoring to ideal GRMHD simulations with either perfectly aligned or anti-aligned disks, populated with ion–electron plasma with purely thermal distribution functions.

While our R – β prescription for setting the electron-to-ion temperature ratio broadly describes general trends as a function of plasma β , simulations with electron heating exhibit substantial zone-to-zone scatter and additional features that our prescription cannot reproduce (Mizuno et al. 2021; Dihingia et al. 2023; Jiang et al. 2023). On the other hand, simulations including radiative cooling produce denser and

cooler accretion disk midplanes, which may help produce Faraday rotation (Yoon et al. 2020).

Tilted disks are known to imprint signatures on the total intensity image (Chatterjee et al. 2020), and additional signatures may occur in polarization. Circularly polarized emission is also sensitive to the content of the plasma and its distribution function, which we briefly explore in Appendix L. While beyond the scope of our analysis, we expect that the polarized data published in this series will continue to constrain models probing additional aspects of BH and plasma astrophysics in future work.

6. Summary and Conclusion

We have presented an analysis of the circular polarization in M87* at 1.3 mm wavelength from 2017 EHT observations, the first analysis of circular polarization in any source using EHT data. By examining the difference between EHT closure phases in right and left circular polarization measurements, we find firm evidence for a weak, astrophysical (noninstrumental) circular polarization signal in M87* on event horizon scales.

To further analyze these data, we developed five different approaches to image analysis, testing each method on a suite of synthetic data from GRMHD simulations. When applied to the observations of M87*, these methods all find a moderate degree of image circular polarization ($\lesssim 4\%$), which is consistent with expectations from the low degree of circular polarization ($\lesssim 1\%$) seen in unresolved observations of M87* with other facilities such as ALMA. The details of our reconstructed images vary considerably among the five methods, indicating that the circular polarization structure is sensitive to choices in the imaging and calibration. Overall, we find that the structure cannot be reliably inferred without additional data or stronger assumptions. From our set of image reconstructions, we establish an upper limit on the (unsigned) resolved circular polarization fraction at the EHT’s resolution of $20 \mu\text{as}$: $\langle |\nu| \rangle < 3.7\%$.

Our results provide new constraints for models of the central supermassive BH in M87 and its environment. We apply these constraints to a large library of images produced from ray-

traced GRMHD simulations, which span a broad range of values for the BH spin and weak (SANE) and strong (MAD) magnetic fields, updating our constraints relative to Paper VIII. We consider both prograde and retrograde accretion flows and both aligned and anti-aligned poloidal magnetic fields relative to the angular momentum of the accretion disk.

We again find that strongly magnetized MAD models remain favored over weakly magnetized SANEs. As in Paper VIII and Qiu et al. (2023), our scoring prefers spin values roughly around -0.5 and 0.0 , driven largely by our constraints on $\angle\beta_2$. That is, the toroidal morphology of the linear polarization ticks favors models with substantial poloidal fields. Our scoring also favors models with larger ion-to-electron temperature ratio R_{high} . That is, the relatively low polarization fraction favors models that contain relatively dense and cool electrons to perform Faraday depolarization. For the model snapshots that simultaneously pass observational constraints, we find that Faraday conversion is typically more important than the circular polarization inherent to synchrotron emission in generating the circular polarization that we observe. This is consistent with the conclusions of parsec-scale studies of other active galactic nuclei at lower frequencies (e.g., Jones 1988; Wardle et al. 1998; Bower & Falcke 1999; Gabuzda et al. 2008). Faraday rotation also serves an indirect role in limiting the circular polarization fraction in many models, scrambling much of the linear polarization that could otherwise be converted into circular.

Theoretical models of M87* exhibit significant variability. In Appendix M, we explore this variability in circular polarization metrics in passing GRMHD models. Future M87* observations may present more favorable conditions for Stokes \mathcal{V} imaging (Figure 31). Thus, continued polarimetric monitoring of M87* on horizon scales will allow us to place better constraints on physical models of the 230 GHz emitting region. Sagittarius A*, the BH in the galactic center recently imaged by the EHT (Event Horizon Telescope Collaboration et al. 2022a), exhibits a much larger $v_{\text{net}} \approx -1\%$ that may make prospects of Stokes \mathcal{V} imaging more promising for this source (Bower et al. 2018; Goddi et al. 2021; Wielgus et al. 2022b), although its much more rapid time variability presents separate challenges (Event Horizon Telescope Collaboration et al. 2022b, 2022c; Wielgus et al. 2022a).

Our results complete the analysis of EHT observations of M87* in 2017, revealing the 230 GHz emission encircling the apparent shadow of the BH in all four Stokes parameters. More recent EHT observations, with additional telescopes, significantly improved sensitivity, and higher observing frequencies, will provide crucial improvements to reveal unambiguous structure in the circular polarization and to characterize its variability.

Acknowledgments

The Event Horizon Telescope Collaboration thanks the following organizations and programs: the Academia Sinica; the Academy of Finland (AKA, projects 274477, 284495, 312496, 315721); the Agencia Nacional de Investigación y Desarrollo (ANID), Chile via NCN19_058 (TITANs), Fondecyt 1221421 and BASAL FB210003; the Alexander von Humboldt Stiftung; an Alfred P. Sloan Research Fellowship; Allegro, the European ALMA Regional Centre node in the Netherlands, the NL astronomy research network NOVA, and the astronomy institutes of the University of Amsterdam,

Leiden University, and Radboud University; the ALMA North America Development Fund; the Astrophysics and High Energy Physics program by MCIN (with funding from European Union NextGenerationEU, PRTR-C17I1); the Black Hole Initiative, which is funded by grants from the John Templeton Foundation and the Gordon and Betty Moore Foundation (although the opinions expressed in this work are those of the author(s) and do not necessarily reflect the views of these Foundations); the Brinson Foundation; “la Caixa” Foundation (ID 100010434) through fellowship codes LCF/BQ/DI22/11940027 and LCF/BQ/DI22/11940030; Chandra DD7-18089X and TM6-17006X; the China Scholarship Council; the China Postdoctoral Science Foundation fellowships (2020M671266, 2022M712084); Consejo Nacional de Ciencia y Tecnología (CONACYT, Mexico, projects U0004-246083, U0004-259839, F0003-272050, M0037-279006, F0003-281692, 104497, 275201, 263356); the Colfuturo Scholarship; the Consejería de Economía, Conocimiento, Empresas y Universidad of the Junta de Andalucía (grant P18-FR-1769), the Consejo Superior de Investigaciones Científicas (grant 2019AEP112); the Delaney Family via the Delaney Family John A. Wheeler Chair at Perimeter Institute; Dirección General de Asuntos del Personal Académico-Universidad Nacional Autónoma de México (DGAPA-UNAM, projects IN112417 and IN112820); the Dutch Organization for Scientific Research (NWO) for the VICI award (grant 639.043.513), grant OCENW.KLEIN.113, and the Dutch Black Hole Consortium (with project No. NWA 1292.19.202) of the research program the National Science Agenda; the Dutch National Supercomputers, Cartesius and Snellius (NWO grant 2021.013); the EACOA Fellowship awarded by the East Asia Core Observatories Association, which consists of the Academia Sinica Institute of Astronomy and Astrophysics, the National Astronomical Observatory of Japan, Center for Astronomical Mega-Science, Chinese Academy of Sciences, and the Korea Astronomy and Space Science Institute; the European Research Council (ERC) Synergy Grant “BlackHoleCam: Imaging the Event Horizon of Black Holes” (grant 610058); the European Union Horizon 2020 research and innovation program under grant agreements RadioNet (No. 730562) and M2FINDERS (No. 101018682); the Horizon ERC Grants 2021 program under grant agreement No. 101040021; the Generalitat Valenciana (grants APOSTD/2018/177 and ASFAE/2022/018) and GenT Program (project CIDEAGENT/2018/021); MICINN Research Project PID2019-108995GB-C22; the European Research Council for advanced grant “JETSET: Launching, propagation and emission of relativistic jets from binary mergers and across mass scales” (grant No. 884631); the FAPESP (Fundação de Amparo à Pesquisa do Estado de São Paulo) under grant 2021/01183-8; the Institute for Advanced Study; the Istituto Nazionale di Fisica Nucleare (INFN) sezione di Napoli, iniziative specifiche TEONGRAV; the International Max Planck Research School for Astronomy and Astrophysics at the Universities of Bonn and Cologne; DFG research grant “Jet physics on horizon scales and beyond” (grant No. FR 4069/2-1); Joint Columbia/Flatiron Postdoctoral Fellowship (research at the Flatiron Institute is supported by the Simons Foundation); the Japan Ministry of Education, Culture, Sports, Science and Technology (MEXT; grant JPMXP1020200109); the Japan Society for the Promotion of Science (JSPS) Grant-in-Aid for JSPS Research Fellowship (JP17J08829); the Joint Institute for

Computational Fundamental Science, Japan; the Key Research Program of Frontier Sciences, Chinese Academy of Sciences (CAS, grants QYZDJ-SSW-SLH057, QYZDJ-SSW-SYS008, ZDBS-LY-SLH011); the Leverhulme Trust Early Career Research Fellowship; the Max-Planck-Gesellschaft (MPG); the Max Planck Partner Group of the MPG and the CAS; the MEXT/JSPS KAKENHI (grants 18KK0090, JP21H01137, JP18H03721, JP18K13594, 18K03709, JP19K14761, 18H01245, 25120007, 23K03453); the Malaysian Fundamental Research Grant Scheme (FRGS) FRGS/1/2019/STG02/UM/02/6; the MIT International Science and Technology Initiatives (MISTI) Funds; the Ministry of Science and Technology (MOST) of Taiwan (103-2119-M-001-010-MY2, 105-2112-M-001-025-MY3, 105-2119-M-001-042, 106-2112-M-001-011, 106-2119-M-001-013, 106-2119-M-001-027, 106-2923-M-001-005, 107-2119-M-001-017, 107-2119-M-001-020, 107-2119-M-001-041, 107-2119-M-110-005, 107-2923-M-001-009, 108-2112-M-001-048, 108-2112-M-001-051, 108-2923-M-001-002, 109-2112-M-001-025, 109-2124-M-001-005, 109-2923-M-001-001, 110-2112-M-003-007-MY2, 110-2112-M-001-033, 110-2124-M-001-007, and 110-2923-M-001-001); the Ministry of Education (MoE) of Taiwan Yushan Young Scholar Program; the Physics Division, National Center for Theoretical Sciences of Taiwan; the National Aeronautics and Space Administration (NASA, Fermi Guest Investigator grant 80NSSC20K1567, NASA Astrophysics Theory Program grant 80NSSC20K0527, NASA NuSTAR award 80NSSC20K0645); NASA Hubble Fellowship grants HST-HF2-51431.001-A and HST-HF2-51482.001-A awarded by the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., for NASA, under contract NAS5-26555; the National Institute of Natural Sciences (NINS) of Japan; the National Key Research and Development Program of China (grants 2016YFA0400704, 2017YFA0402703, 2016YFA0400702); the National Science Foundation (NSF, grants AST-0096454, AST-0352953, AST-0521233, AST-0705062, AST-0905844, AST-0922984, AST-1126433, AST-1140030, DGE-1144085, AST-1207704, AST-1207730, AST-1207752, MRI-1228509, OPP-1248097, AST-1310896, AST-1440254, AST-1555365, AST-1614868, AST-1615796, AST-1715061, AST-1716327, OISE-1743747, AST-1816420, AST-1935980, AST-2034306, AST-2307887); NSF Astronomy and Astrophysics Postdoctoral Fellowship (AST-1903847); the Natural Science Foundation of China (grants 11650110427, 10625314, 11721303, 11725312, 11873028, 11933007, 11991052, 11991053, 12192220, 12192223, 12273022); the Natural Sciences and Engineering Research Council of Canada (NSERC, including a Discovery Grant and the NSERC Alexander Graham Bell Canada Graduate Scholarships-Doctoral Program); the National Youth Thousand Talents Program of China; the National Research Foundation of Korea (the Global PhD Fellowship Grant: grants NRF-2015H1A2A1033752; the Korea Research Fellowship Program: NRF-2015H1D3A1066561; Brain Pool Program: 2019H1D3A1A01102564; Basic Research Support Grant 2019R1F1A1059721, 2021R1A6A3A01086420, 2022R1C1C1005255); Netherlands Research School for Astronomy (NOVA) Virtual Institute of Accretion (VIA) postdoctoral fellowships; Onsala Space Observatory (OSO) national infrastructure, for the provisioning of its facilities/observational support (OSO receives funding through the Swedish Research Council under grant 2017-00648); the Perimeter Institute for

Theoretical Physics (research at Perimeter Institute is supported by the Government of Canada through the Department of Innovation, Science and Economic Development and by the Province of Ontario through the Ministry of Research, Innovation and Science); the Princeton Gravity Initiative; the Spanish Ministerio de Ciencia e Innovación (grants PGC2018-098915-B-C21, AYA2016-80889-P, PID2019-108995GB-C21, PID2020-117404GB-C21); the University of Pretoria for financial aid in the provision of the new Cluster Server nodes and SuperMicro (USA) for a SEEDING GRANT approved toward these nodes in 2020; the Shanghai Municipality orientation program of basic research for international scientists (grant No. 22JC1410600); the Shanghai Pilot Program for Basic Research, Chinese Academy of Science, Shanghai Branch (JCYJ-SHFY-2021-013); the State Agency for Research of the Spanish MCIU through the “Center of Excellence Severo Ochoa” award for the Instituto de Astrofísica de Andalucía (SEV-2017-0709); the Spanish Ministry for Science and Innovation grant CEX2021-001131-S funded by MCIN/AEI/10.13039/501100011033; the Spinoza Prize SPI 78-409; the South African Research Chairs Initiative, through the South African Radio Astronomy Observatory (SARAO, grant ID 77948), which is a facility of the National Research Foundation (NRF), an agency of the Department of Science and Innovation (DSI) of South Africa; the Toray Science Foundation; the Swedish Research Council (VR); the US Department of Energy (USDOE) through the Los Alamos National Laboratory (operated by Triad National Security, LLC, for the National Nuclear Security Administration of the USDOE, contract 89233218CNA000001); and the YCAA Prize Postdoctoral Fellowship.

We thank the staff at the participating observatories, correlation centers, and institutions for their enthusiastic support. This paper makes use of the following ALMA data: ADS/JAO.ALMA#2016.1.01154.V. ALMA is a partnership of the European Southern Observatory (ESO; Europe, representing its member states), NSF, and National Institutes of Natural Sciences of Japan, together with National Research Council (Canada), Ministry of Science and Technology (MOST; Taiwan), Academia Sinica Institute of Astronomy and Astrophysics (ASIAA; Taiwan), and Korea Astronomy and Space Science Institute (KASI; Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, Associated Universities, Inc. (AUI)/NRAO, and the National Astronomical Observatory of Japan (NAOJ). The NRAO is a facility of the NSF operated under cooperative agreement by AUI. This research used resources of the Oak Ridge Leadership Computing Facility at the Oak Ridge National Laboratory, which is supported by the Office of Science of the U.S. Department of Energy under contract No. DE-AC05-00OR22725; the ASTROVIVES FEDER infrastructure, with project code IDIFEDER-2021-086; and the computing cluster of Shanghai VLBI correlator supported by the Special Fund for Astronomy from the Ministry of Finance in China. We also thank the Center for Computational Astrophysics, National Astronomical Observatory of Japan. This work was supported by FAPESP (Fundação de Amparo a Pesquisa do Estado de São Paulo) under grant 2021/01183-8.

APEX is a collaboration between the Max-Planck-Institut für Radioastronomie (Germany), ESO, and the Onsala Space Observatory (Sweden). The SMA is a joint project between the SAO and ASIAA and is funded by the Smithsonian Institution

and the Academia Sinica. The JCMT is operated by the East Asian Observatory on behalf of the NAOJ, ASIAA, and KASI, as well as the Ministry of Finance of China, Chinese Academy of Sciences, and the National Key Research and Development Program (No. 2017YFA0402700) of China and Natural Science Foundation of China grant 11873028. Additional funding support for the JCMT is provided by the Science and Technologies Facility Council (UK) and participating universities in the UK and Canada. The LMT is a project operated by the Instituto Nacional de Astr f sica,  ptica, y Electr nica (Mexico) and the University of Massachusetts at Amherst (USA). The IRAM 30 m telescope on Pico Veleta, Spain, is operated by IRAM and supported by CNRS (Centre National de la Recherche Scientifique, France), MPG (Max-Planck-Gesellschaft, Germany), and IGN (Instituto Geogr fico Nacional, Spain). The SMT is operated by the Arizona Radio Observatory, a part of the Steward Observatory of the University of Arizona, with financial support of operations from the State of Arizona and financial support for instrumentation development from the NSF. Support for SPT participation in the EHT is provided by the National Science Foundation through award OPP-1852617 to the University of Chicago. Partial support is also provided by the Kavli Institute of Cosmological Physics at the University of Chicago. The SPT hydrogen maser was provided on loan from the GLT, courtesy of ASIAA.

This work used the Extreme Science and Engineering Discovery Environment (XSEDE), supported by NSF grant ACI-1548562, and CyVerse, supported by NSF grants DBI-0735191, DBI-1265383, and DBI-1743442. The XSEDE Stampede2 resource at TACC was allocated through TG-AST170024 and TG-AST080026N. The XSEDE JetStream resource at PTI and TACC was allocated through AST170028. This research is part of the Frontera computing project at the Texas Advanced Computing Center through the Frontera Large-Scale Community Partnerships allocation AST20023. Frontera is made possible by National Science Foundation award OAC-1818253. This research was done using services provided by the OSG Consortium (Pordes et al. 2007; Sfiligoi et al. 2009), which is supported by the National Science Foundation award Nos. 2030508 and 1836650. Additional work used ABACUS2.0, which is part of the eScience center at Southern Denmark University. Simulations were also performed on the SuperMUC cluster at the LRZ in Garching, on the LOEWE cluster in CSC in Frankfurt, on the HazelHen cluster at the HLRS in Stuttgart, and on the Pi2.0 and Siyuan Mark-I at Shanghai Jiao Tong University. The computer resources of the Finnish IT Center for Science (CSC) and the Finnish Computing Competence Infrastructure (FCCI) project are acknowledged. This research was enabled in part by support provided by Compute Ontario (<http://computeontario.ca>), Calcul Quebec (<http://www.calculquebec.ca>), and Compute Canada (<http://www.computeCanada.ca>).

The EHTC has received generous donations of FPGA chips from Xilinx Inc., under the Xilinx University Program. The EHTC has benefited from technology shared under open-source license by the Collaboration for Astronomy Signal Processing and Electronics Research (CASPER). The EHT project is grateful to T4Science and Microsemi for their assistance with hydrogen masers. This research has made use of NASA’s Astrophysics Data System. We gratefully acknowledge the support provided by the extended staff of the ALMA,

from the inception of the ALMA Phasing Project through the observational campaigns of 2017 and 2018. We would like to thank A. Deller and W. Bricken for EHT-specific support with the use of DiFX. We thank Martin Shepherd for the addition of extra features in the Difmap software that were used for the CLEAN imaging results presented in this paper. We acknowledge the significance that Maunakea, where the SMA and JCMT EHT stations are located, has for the indigenous Hawaiian people.

Facilities: EHT, ALMA, APEX, IRAM:30m, JCMT, LMT, SMA, ARO:SMT, SPT.

Software: AIPS (Greisen 2003), ParselTongue (Kettenis et al. 2006), GNU Parallel (Tange 2011), eht-imaging (Chael et al. 2016), DIFMAP (Shepherd 2011), Numpy (Harris et al. 2020), Scipy (Virtanen et al. 2020), Pandas (McKinney 2010), Astropy (The Astropy Collaboration et al. 2013, 2018), Jupyter (Kluyver et al. 2016), Matplotlib (Hunter 2007), THEMIS (Broderick et al. 2020a), DMC (Pesce 2021), polysolve (Mart -Vidal et al. 2021), HARM (Gammie et al. 2003; Noble et al. 2006), IPOLE (Noble et al. 2007; Mo cibrodzka & Gammie 2018).

Appendix A Strategies for Polarimetric Gain Calibration

The fiducial multisource procedure for polarimetric gain calibration, described in Section 2.3, was employed for the whole series of EHT papers on the M87* observations performed in 2017, as well as the multiyear M87* variability study presented in Wielgus et al. (2020) and the 3C 279 quasar study (Kim et al. 2020).

Figure 14 shows an example of how we estimated the phase of $G_{R/L}$ for the LMT using this fiducial multisource approach. The figure shows visibility phase differences between RR^* and LL^* components recorded on the ALMA-LMT baseline, following the a priori correction for the field angle rotation. A near constant in time phase residual of about -157° is clearly present, consistent between observing days and for multiple observed sources. This residual is interpreted as the negative phase of the LMT instrumental gain ratio $G_{R/L}$ and subsequently corrected in the data. Similarly, $G_{R/L}$ gains

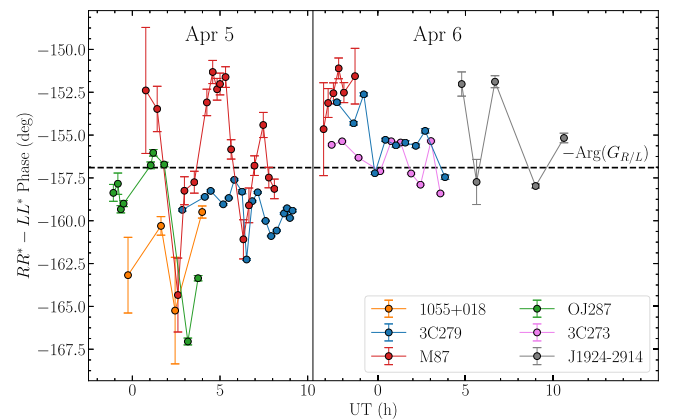


Figure 14. Example of a multisource, multiday estimation of the $G_{R/L}$ phase for the LMT, using RR^* and LL^* high-band visibilities measured on the ALMA-LMT baseline. For the actual estimate of a constant $G_{R/L}$ phase, 5 days and 10 sources were used. The origin of the small residual phases may be instrumental; however, they do not exhibit clear source-independent features. Alternatively, residuals may be caused by the presence of a small, source-specific circular polarization.

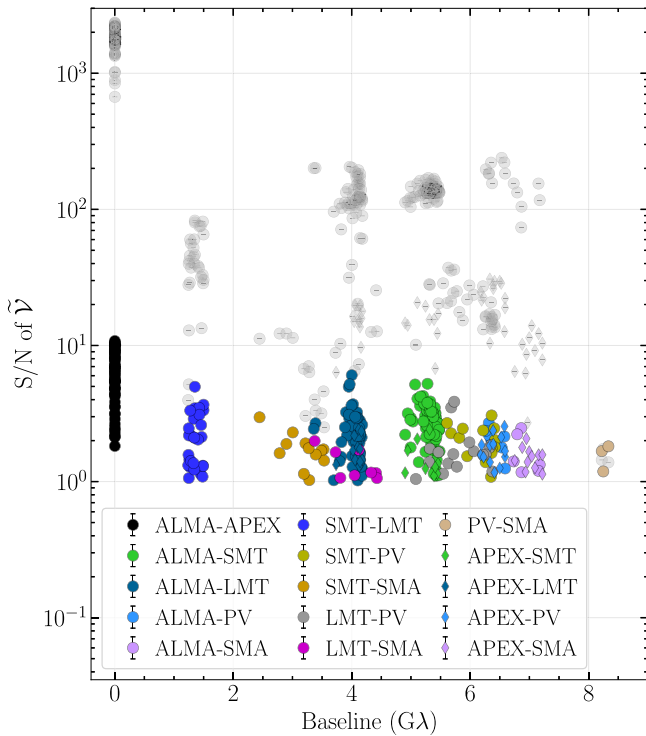


Figure 15. S/N of the $\tilde{\mathcal{V}}$ observations as a function of the projected baseline length, for the data set self-calibrated to $\tilde{\mathcal{V}} = 0$. Gray points in the background correspond to the S/N of Stokes $\tilde{\mathcal{I}}$ detections.

corresponding to other sites were corrected using the baseline from LMT to ALMA on multiple days and sources.

An alternative approach is based on calibrating complex $G_{R/L}$ to $RR^* = LL^*$, that is, to $\tilde{\mathcal{V}} = 0$. The latter strategy, which is simpler but more aggressive, was used in the CASA-based `rPICARD` pipeline (Janssen et al. 2019), where the phase difference between RR^* and LL^* visibilities is minimized through a polarimetric gain calibration. The $\tilde{\mathcal{V}} = 0$ self-calibration constitutes a very robust approach that remains valid for the recovery of the remaining Stokes parameters as long as $\tilde{\mathcal{I}} \gg \tilde{\mathcal{V}}$. In addition to its use in the `rPICARD` pipeline, we have made a choice to employ this calibration variant for the HOPS data in the EHT papers on Sagittarius A* (Event Horizon Telescope Collaboration et al. 2022a), Centaurus A

(Janssen et al. 2021), J1924–2914 (Issaoun et al. 2022), and NRAO 530 (Jorstad et al. 2023).

Note that self-calibration (multiplicative gain calibration) to $\tilde{\mathcal{V}} = 0$ is not equivalent to removing the baseline signatures of $\tilde{\mathcal{V}} = 0$. In particular, this alternative calibration does not affect closure signatures, such as those presented in Figure 2. This is different from the EHT M87* data set released with Paper III in 2019,¹⁶¹ where all signatures of $\tilde{\mathcal{V}}$ were deliberately wiped out by replacing RR^* and LL^* on all baselines with their average, only retaining the Stokes $\tilde{\mathcal{I}}$ total intensity information intact.

Despite certain advantages, self-calibration to $\tilde{\mathcal{V}} = 0$ may render circular polarization more difficult to recover, and for that reason the multisource polarimetric gain calibration procedure was selected as better suited for the conservative upper limits presented in this paper. In Figure 15 we present a version of Figure 1, but for the $\tilde{\mathcal{V}} = 0$ self-calibration data set. Significant decrease of $\tilde{\mathcal{V}}$ S/N is clearly visible, indicating that high-S/N detections seen in Figure 1 could be related to systematic errors of the polarimetric gain calibration, rather than to robust signatures of circular polarization. S/N of the Stokes $\tilde{\mathcal{I}}$, on the other hand, is mostly insensitive to the $G_{R/L}$ calibration scheme.

Appendix B

Closure Phase Differences on ALMA Triangles

Figure 2 in the main text shows the closure phase differences for the triangle ALMA-SMT-PV. For this triangle, an average deviation from zero at the level of $\sim 5\sigma$ is detected, which we interpret as evidence of the presence of a spatially resolved (and asymmetric) Stokes \mathcal{V} brightness distribution in M87*.

There are, though, a total of 10 antenna triangles where ALMA is included, which are related to the closure phases with highest S/N. In Figure 16, we show the closure phase differences of scan-averaged visibilities for all these triangles. There are hints of departures from zero in some cases and for some time ranges (e.g., ALMA-SMT-LMT toward the end of the observations, as well as ALMA-APEX-LMT and ALMA-LMT-PV), but there is no clear detection above 3σ – 4σ . The detection of Stokes \mathcal{V} from all the ALMA-related antenna triangles is, therefore, only tentative if we use the closure phase differences.

¹⁶¹ https://datacommons.cyverse.org/browse/iplant/home/shared/commons_repo/curated/EHTC_FirstM87Results_Apr2019

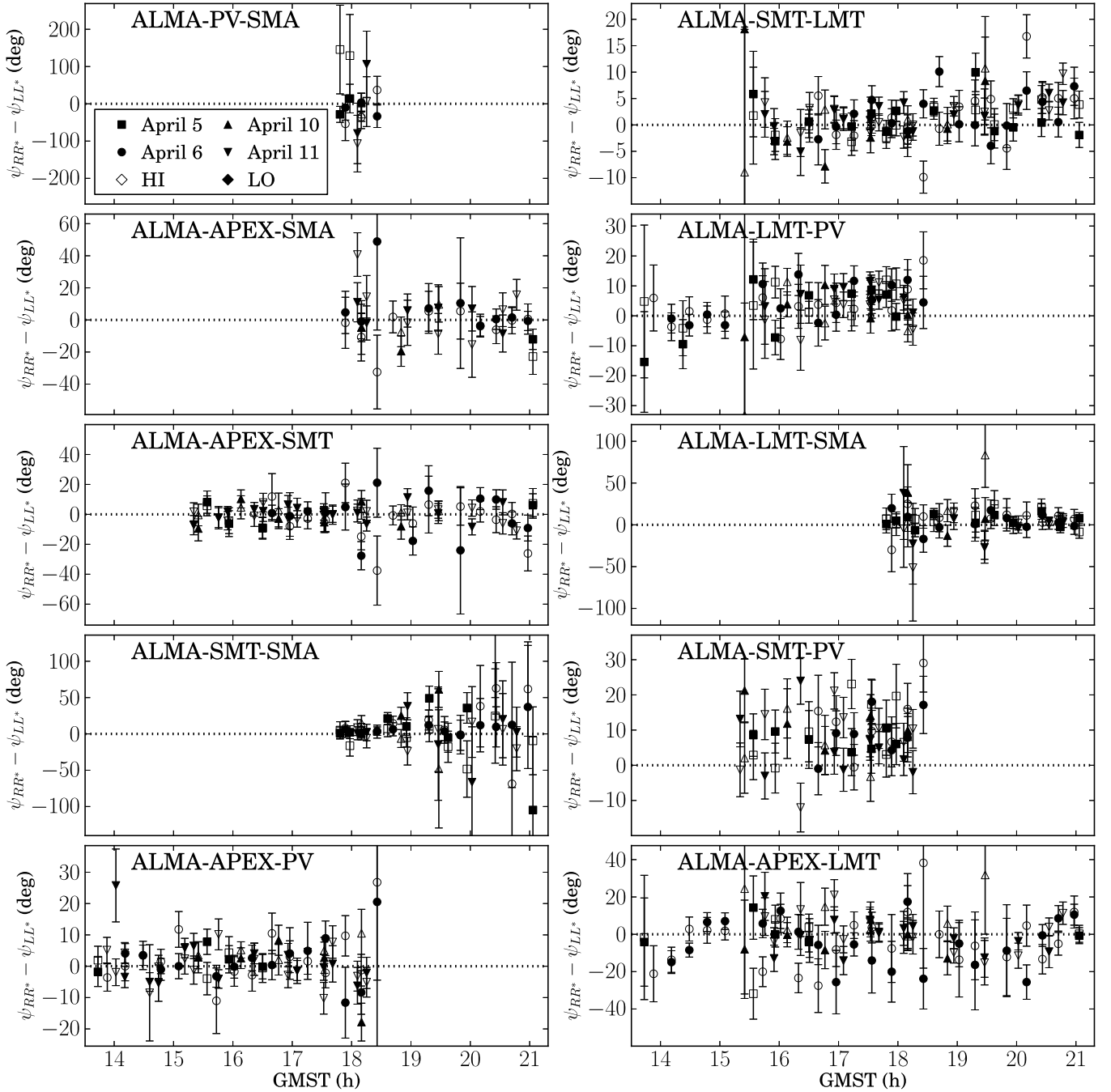


Figure 16. Closure phase differences between scan-averaged RR^* and LL^* visibilities, for all the antenna triangles related to ALMA. All epochs and both bands (open symbols for high band; filled symbols for low band) are shown. The zero line (i.e., no detection of spatially resolved Stokes \mathcal{V}) is shown as a dotted line.

Appendix C Closure Traces

Closure traces are a set of complex closure quantities defined on quadrangles that are insensitive to station-based corruptions (see Broderick & Pesce 2020). Of particular relevance to the current discussion of Stokes \mathcal{V} is the fact that this includes the right- and left-hand complex station gains and leakage terms. From these, a combination can be constructed on a single quadrangle, the conjugate closure trace product that differs from unity only in the presence of nontrivial polarization structure (see Broderick & Pesce 2020). For this reason, these conjugate closure trace products provide a robust direct test for source polarization.

Figure 17 compares the phases of the conjugate closure trace product generated on the APEX-ALMA-LMT-SMT and APEX-SMT-LMT-ALMA quadrangles from the image reconstructions in Figure 3 directly to the 2017 April 11 low-band data.¹⁶² All reconstructions are broadly consistent with the conjugate closure trace products, similar to the results found in Appendix B of Paper VII.

Because the closure traces are invariant to rotations on the Poincaré sphere, the conjugate closure trace products cannot

¹⁶² The insensitivity of the closure traces to the gain ratios $|G_R/L|$ results in similar width 2σ bands in Figure 17 for THEMIS and DMC despite their different treatments of the station gains, and thus very different error bars in Figure 7.

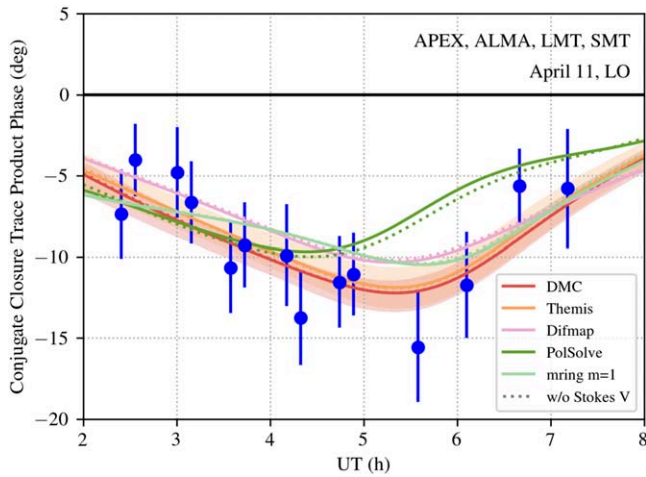


Figure 17. Phases of the conjugate closure trace product constructed on the APEX-ALMA-LMT-SMT and APEX-SMT-LMT-ALMA quadrangles from the 2017 April 11 low-band observations. Colored lines show the same quantity computed from the mean images in Figure 3, with the 2σ regions for THEMIS and DMC indicated by bands. For comparison, the conjugate closure trace products when the Stokes \mathcal{V} maps are ignored are shown by the corresponding dotted lines for each method.

isolate the contribution from the Stokes \mathcal{V} maps. However, we also show by dotted lines the conjugate closure trace products for each reconstruction in Figure 3 after setting the Stokes \mathcal{V} to zero, indicating the gross magnitude of the impact of circular polarization. For all reconstruction methods, the difference attributable to the Stokes \mathcal{V} map is small. For THEMIS and DMC, the difference is small in comparison to the range of conjugate closure trace product phases spanned by the image posterior, implying that reconstruction of the Stokes \mathcal{V} maps is strongly dependent on the assumed calibration priors, consistent with what is found in Section 4.

Appendix D Method Summaries

In this appendix, we describe each image reconstruction method introduced in Section 3 and summarized in Table 1.

D.1. DIFMAP

The CLEAN algorithm (Högbom 1974) is a traditional VLBI imaging method that performs imaging via inverse modeling in all four Stokes parameters, \mathcal{I} , \mathcal{Q} , \mathcal{U} , and \mathcal{V} for visibility data with dual polarization. The method is implemented in the DIFMAP (Shepherd 2011) and Astronomical Imaging Process System (AIPS Greisen 2003) software packages. Imaging in DIFMAP is an iterative process using a set of CLEAN windows until the residual χ^2 between data and a model reach a minimum value. For total intensity imaging, the process involves self-calibration of phase and amplitude of visibilities with the model at different steps of iterations, which results in calibrated visibility data and a file containing a model of \mathcal{I} delta functions characterized by the amplitude and position on the image. For linear and circular polarization imaging, no additional self-calibration procedure is performed, and DIFMAP generates models of \mathcal{Q} , \mathcal{U} , and \mathcal{V} delta functions.

For EHT data, to obtain a Stokes \mathcal{V} image, we first perform an amplitude and phase calibration of R and L visibility data with the Stokes \mathcal{I} image via the AIPS task CALIB, assuming that $(RR^* + LL^*)/2 = \tilde{\mathcal{I}}$. Then, we apply the D -term

correction to the visibility data and repeat the amplitude and phase calibration with the image assuming again $(RR^* + LL^*)/2 = \tilde{\mathcal{I}}$. For this study we use the M87* D -terms, which are specified in Tables 3 and 5 of Appendix A in Paper VII, applying them through an antenna table via AIPS tasks TBOU and TBIN. At this stage, the calibrated data files are written out to DIFMAP for visual inspection and further editing, if needed.

After reading the edited visibility data back to AIPS, we proceed with the final R/L gain calibration, which is necessary to remove any residual instrumental circular polarization. To do that, we apply an additional amplitude and phase calibration of R and L visibility data *separately* with the \mathcal{I} image using the CALIB task, assuming that $RR^* = \tilde{\mathcal{I}}$ and $LL^* = \tilde{\mathcal{I}}$. This is necessary because DIFMAP does not calibrate R and L polarization hands separately. The resulting solution (SN) table ideally includes only a small residual correction for R and L that represents their gain offset for each antenna in the array. This SN table is processed outside AIPS to edit the R and L amplitude and phases with the assumption that the R/L gains of the ALMA station equal unity and hence any R/L ratio seen by ALMA is ideally intrinsic to the source. Under this assumption, the R and L visibility data for all antennas are modified accordingly so that the average R/L ratios better match the R/L ratio seen by ALMA. We note that this approach corrects only for an R/L offset for each antenna and is insensitive to possible time variations of the R/L gains. The modified SN table is read back to AIPS with the task TBIN and used to produce the final data files (uvfits), using tasks CLCAL and SPLIT, which should be calibrated for both circular and linear polarization. Such calibrated uv-data are reloaded in DIFMAP and imaged in Stokes \mathcal{V} using the set of windows employed for the Stokes \mathcal{I} image. This is the procedure that we have used to produce Stokes \mathcal{V} images for the 2017 M87* data and synthetic data analyzed in this work. We have used the noise level (rms) multiplied by a factor of 3 of Stokes \mathcal{I} and \mathcal{V} images as the uncertainties of the total and circular polarized intensities, respectively, to derive uncertainties of parameters presented in Appendix E and Section 4.

D.2. polsolve

The polsolve algorithm is also based on CLEAN, but for polarimetric calibration it uses the full measurement equation (e.g., Smirnov 2011), including (second-order) nonlinear corrections to the instrumental polarization (see Martí-Vidal et al. 2021, for full details). While the original polsolve algorithm was not designed to include a Stokes \mathcal{V} source model in the fitting of the instrumental polarization, the flexibility of CASA makes it possible to account for circular polarization by manually filling the *model* column of the measurement set and asking polsolve to use it in the fitting. This strategy corresponds to the “polarimetric self-calibration” approach described in detail in Martí-Vidal et al. (2021).

For the recovery of the Stokes \mathcal{V} image with polsolve, we have followed this iterative polarimetric self-calibration approach, starting from the data calibrated as described in Paper VII. First, we have run the CASA-based CLEAN on the four Stokes parameters, by using a common mask based on the Stokes \mathcal{I} image. The parameters used in this CLEAN are the same as those summarized for polsolve in Paper VII. Each run of (full-polarization) CASA-based CLEAN was followed by a (polarization-independent) self-calibration (i.e., gain

solutions were computed based on Stokes $\tilde{\mathcal{I}}$ only, using RR^* and LL^*). Then, a D -term fitting with `polsolve` was performed (this time, the full visibility matrix was used). This procedure was iterated until convergence (i.e., until the changes in the antenna gains and D -terms between consecutive iterations were negligible). After convergence, a final CLEAN image was generated. This approach was the same for all the data presented in this publication (i.e., all the synthetic data tests and all the EHT observations).

D.3. DMC and THEMIS

THEMIS (Broderick et al. 2020b) and DMC (Pesce 2021) produced Stokes \mathcal{V} maps concurrently with the linear polarization maps presented in Paper VII, and details about the analyses are provided there. Here we briefly summarize both methods and emphasize the differences that are most relevant for imaging circular polarization.

THEMIS and DMC formulate the imaging problem as one of Bayesian posterior exploration over the combined space of both image structures and station-based calibration quantities (i.e., complex gains and leakage terms). The output of both codes is a set of Markov Chain Monte Carlo samples from the joint posterior distribution over both the full-Stokes image structure and the calibration quantities.

For the purposes of Stokes \mathcal{V} reconstruction, the most salient difference between THEMIS and DMC is in their treatment of the relative right- and left-handed station gain quantities. Whereas THEMIS holds that the right- and left-handed station gains are identical, DMC explores potentially large differences between them. With the exception of a reference station—chosen here to be ALMA—that has both right- and left-hand gain phases fixed to be zero-valued at all times, DMC independently models the right- and left-hand complex gain quantities for each scan and each station.

D.4. Geometric Modeling

Geometric modeling differs from imaging methods because the reconstructed source structure is restricted to the space defined by a family of simple geometric models. Geometric modeling has been used to infer total intensity source structure parameters for M87* and Sgr A*, generally showing excellent consistency with the results of imaging (Event Horizon Telescope Collaboration et al. 2019f, 2022c). The restricted parameter space in geometric modeling is generally easier to constrain than that of many-parameter image reconstructions and their associated regularizers. Hence, geometric modeling provides a useful complement to imaging, especially to analyze observations with sparse uv -sampling. However, geometric modeling results will not be accurate if the true source structure is not represented in the model space, so care must be taken when modeling sources where the source structure is unknown and unconstrained from other methods.

In this work, we fit a full-Stokes m -ring model (R23) to EHT M87* data from 2017. The total intensity m -ring model (Johnson et al. 2020; Event Horizon Telescope Collaboration et al. 2022c) consists of a ring with diameter d , with azimuthal brightness variations that are decomposed in Fourier modes. The image is given by

$$\mathcal{I}(\rho, \varphi) = \frac{F}{\pi d} \delta\left(\rho - \frac{d}{2}\right) \sum_{k=-m_{\mathcal{I}}}^{m_{\mathcal{I}}} \beta_k e^{ik\varphi}. \quad (\text{D1})$$

Here ρ and φ are polar sky coordinates centered on the ring, β_k are the complex Fourier coefficients, F is the total flux density (we set $\beta_0 \equiv 1$), and the total intensity m -ring order $m_{\mathcal{I}}$ is the maximum nonzero Fourier coefficient. By setting a higher $m_{\mathcal{I}}$, the total number of model parameters increases and an increasingly complex azimuthal structure can be modeled. Because the total intensity is real, $\beta_k = -\beta_k^*$. The m -ring is given a finite width by convolving $\mathcal{I}(\rho, \varphi)$ with a circular Gaussian with FWHM α . A useful property of this model is that both the image and its corresponding visibility function are straightforward to calculate analytically (see, e.g., R23).

The m -ring model can be straightforwardly generalized to include all Stokes parameters. In linear polarization, the azimuthal variations are set by $\{\beta_{\mathcal{P},k}\}$. Because the linear polarization image is complex, $\beta_{\mathcal{P},k}$ and $\beta_{\mathcal{P},k}^*$ are not related by conjugation symmetry. In circular polarization, the azimuthal variations are set by $\{\beta_{\mathcal{V},k}\}$. Because the circular polarization image is real, $\beta_{\mathcal{V},k} = -\beta_{\mathcal{V},k}^*$.

For our m -ring fits to synthetic and real EHT data, we use the geometric modeling module in `eht-imaging` (Chael et al. 2016, 2018). In this work, we fit an ($m_{\mathcal{I}} = 3$, $m_{\mathcal{P}} = 3$, $m_{\mathcal{V}} = 1$) m -ring to the parallel-hand closure amplitudes and closure phases. Posterior exploration is done with the `dynesty` sampler (Speagle 2020). We first fit the total intensity and linear polarization structure, after which we fix the linear polarization parameters and fit the total intensity and circular polarization structure simultaneously. Because the closure products do not constrain the net circular polarization $\beta_{\mathcal{V},0} \equiv v_{\text{net}}$, we fix this parameter to the ALMA-only estimate of -0.3% (Goddi et al. 2021) for the EHT M87* data and to the ground-truth values for the synthetic data sets. Our fitting procedure and assumptions are motivated in more detail in R23, where additional tests are presented, including varying the circular polarization m -ring order $m_{\mathcal{V}}$, varying the assumed net circular polarization fraction ($m_c \equiv \beta_{\mathcal{V},0}$), and varying the fitted data products (e.g., using closure quantities vs. using parallel-hand complex visibility ratios).

Appendix E Synthetic Data Tests

In this appendix, we test the ability of our image reconstruction and geometric modeling methods to recover resolved circular polarization from an idealized, high circular polarization “convention test,” as well as representative images from GRMHD simulations of M87*.

E.1. Synthetic Data Generation

As in Paper VII, we generated synthetic observations on the M87* 2017 April 11 low-band (u, v) coverage using the `eht-imaging` software package. After sampling the image Fourier transforms on EHT baselines, the synthetic data were corrupted with time-variable, station-based systematic phase and gain offsets, time-stable polarimetric leakage terms, and thermal noise.

Unlike in the Paper VII synthetic data, in this paper we do not assume that the complex gains in the right and left circular polarizations are equal: $G_R \neq G_L$. Instead, we introduce time-variable offsets in both the amplitude and phase of the ratio G_R/L . The magnitude of these offsets and their characteristic timescales were motivated by the a priori limits described in

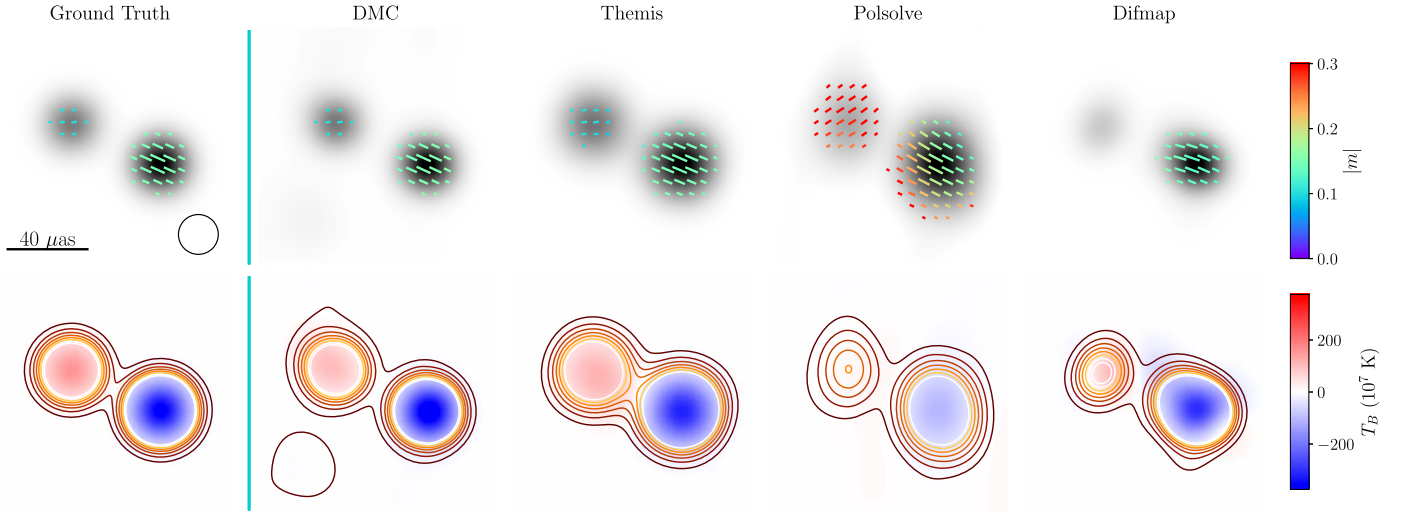


Figure 18. Imaging results from the convention test data set, consisting of a highly polarized ($V_{\text{pk}}/I \approx 10\%$) double source. Images are plotted in the same manner as in Figure 19.

Section 2.3. For most sites, we sample the R/L amplitude gain offset from a normal distribution centered at unity with a standard deviation of 20% and the $R - L$ phase offset from a zero-centered normal distribution with a standard deviation of 10° . The exceptions are ALMA (where $G_R = G_L$) and the SMA, where a priori limits on the phase offset motivate the use of a 40° standard deviation. The phase and amplitude gain offsets are correlated in time (as are the overall amplitude and phase gain terms), and we set the correlation timescales for the offsets to be larger than the correlation timescales for the overall amplitude (2 hr vs. 1 hr) and phase gains (24 hr vs. 0.5 hr).

E.2. Convention Test Results

Before imaging the GRMHD models presented in Appendix E.3, we first tested our methods on an optimistic “convention test,” a strongly circularly polarized ($v \approx 10\%$) double Gaussian source with 1.2 Jy total flux density. In this test, instead of a GRMHD snapshot, a synthetic source image was constructed from two Gaussians with FWHM $\approx 20 \mu\text{as}$ separated by a distance $d \approx 50 \mu\text{as}$. The total flux density of the double source was 1.2 Jy, and the fractional circular polarization of each Gaussian component was high, $V_{\text{pk}}/I \approx 10\%$. We show the results of this initial test in Figure 18.

All participating methods (the m -ring modeling group did not participate in this test) were able to recover the two-component structure, as well as the approximate degree of circular polarization in each component, even with systematic gain offsets $G_{R/L}$ present in the synthetic data at similar levels to the GRMHD data sets discussed in the following section. This test demonstrates that our analysis methods would have been able to recover consistent, accurate resolved Stokes \mathcal{V} images from 2017 EHT observations of M87* if the intrinsic circular polarization brightness was ≈ 5 –10 times higher than observed in 2017.

E.3. GRMHD Results

To test our imaging methods on more realistic source models for M87*, we selected three GRMHD images from Paper VIII that pass all the observational constraints from Paper VII. We sorted these images by the value of $\langle |v| \rangle$ after blurring to $20 \mu\text{as}$

Table 4
Summary of Synthetic Data Models

Model	B -field	a^*	v_{net} (%)	$\langle v \rangle$ (%)
01	MAD	−0.5	−0.05	0.5
02	MAD	0.5	+0.25	1.75
03	SANE	0	+0.15	4.1

resolution. We then chose three snapshot images from across this distribution to represent weak, average, and near-maximal $\langle |v| \rangle$ from the distribution of passing Paper VIII images. All of these images have low image-integrated circular polarization $|v_{\text{net}}| < 0.3\%$; images that have stronger polarization at $20 \mu\text{as}$ scales have significant regions of both positive and negative circular polarization that cancel in unresolved observations.

Our three models are the “weak polarization” case (model 01), with an average circular polarization fraction at $20 \mu\text{as}$ resolution of $\langle |v| \rangle \approx 0.5\%$; the “moderate polarization” case (model 02), with $\langle |v| \rangle \approx 2\%$; and the “strong polarization” case (model 03), with $\langle |v| \rangle \approx 4\%$. In Table 4, we summarize these three models, including the magnetic field configuration and spin of their underlying GRMHD simulation, the net 230 GHz circular polarization fraction v_{net} , and the average resolved fraction $\langle |v| \rangle$.

Figure 19 shows the GRMHD ground-truth models and reconstructions from our three synthetic data sets (Appendix E), for each of the five imaging or modeling methods described in Section 3. Because all three GRMHD synthetic data models were chosen from the set of passing Paper VIII models, all feature relatively similar total linear polarization fractions and EVPA structure, similar to those observed in M87* by the EHT in Paper VII.

For all models, the total intensity and linear polarization structure (top rows) are generally recovered well by all methods. Because the synthetic data included both large overall amplitude gains G on certain stations, the total flux density is not completely constrained by these data sets (as is the case for the real EHT 2017 M87* data). The Stokes \mathcal{V} reconstructions (bottom rows) show significant variation between the methods, and most methods have difficulty reconstructing the ground-truth features in circular polarization.

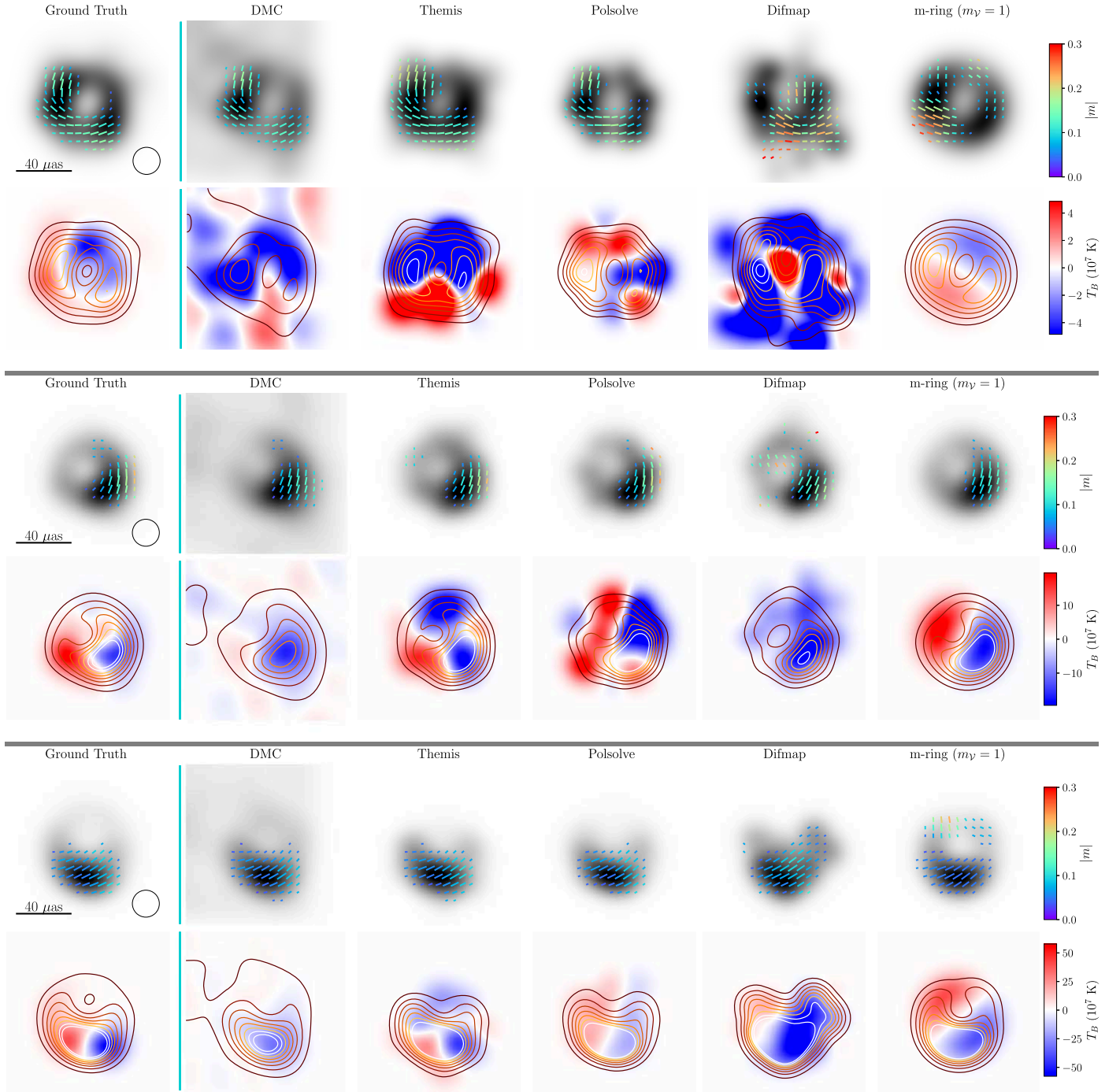


Figure 19. Imaging results from the three synthetic data sets in Table 4. The first row shows the total intensity in gray scale and the fractional linear polarization in the colored ticks. The left column is the ground-truth simulation image. From left to right, the next columns show the posterior-average images from DMC and THEMIS, the final CLEAN images from `polsolve` and `DIFMAP`, and the posterior average of the m -ring model fits. All images are blurred with the same $20 \mu\text{as}$ FWHM circular Gaussian beam. The second row shows the total intensity image with eight contour levels on a linear scale and the Stokes \mathcal{V} structure in a diverging color map. The first and second rows show the results for model 01 (low resolved circular polarization), the third and fourth rows show the results for model 02 (moderate resolved circular polarization), and the fifth and sixth rows show the results for model 03 (high resolved circular polarization). Note that the color bar for \mathcal{V} has different maximum values in the second, fourth, and sixth rows as the GRMHD simulation images become more polarized.

All methods struggle the most with model 01, the weakly polarized case ($\langle |\mathcal{V}| \rangle \approx 0.5\%$). In this model, no method correctly recovers the spatial distribution of the sign of circular polarization, and due to the presence of residual relative gains $G_{R/L}$ in the data, some methods significantly overproduce the total Stokes \mathcal{V} brightness in the image. In contrast, for the “high-polarization” model 03, ($\langle |\mathcal{V}| \rangle \approx 4\%$), all methods are able to recover the resolved distribution of \mathcal{V} somewhat more

accurately. THEMIS, `polsolve`, and m -ring modeling all recover the correct dipolar Stokes \mathcal{V} structure of the image; DMC and `DIFMAP` produce mostly one sign of \mathcal{V} corresponding to the more strongly polarized negative component in the image. The results of model 02 are more mixed, with some methods (DMC, THEMIS, and `DIFMAP`) accurately recovering the sign and approximate magnitude of \mathcal{V} at the total intensity peak, while m -ring modeling and `polsolve` see

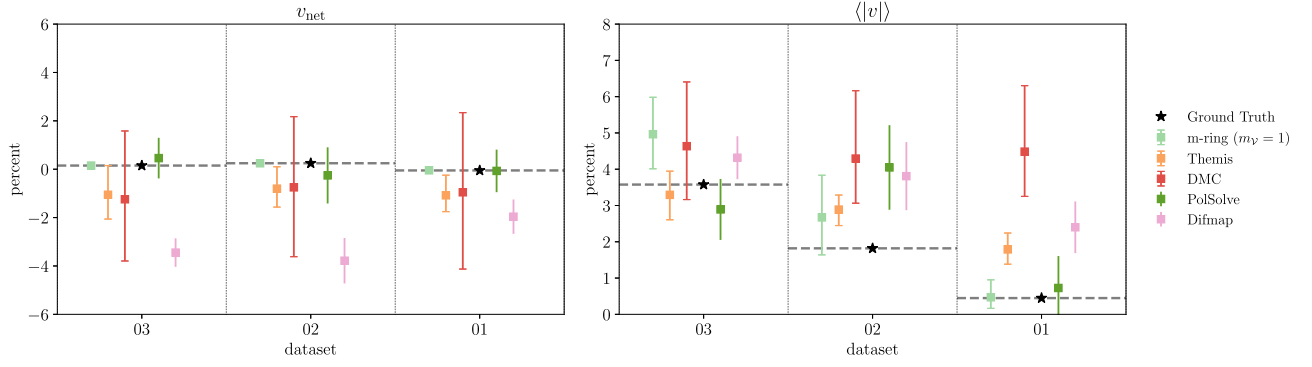


Figure 20. Image-integrated statistics from the results of the synthetic data tests presented in Figure 19. The left panel shows the net circular polarization fraction v_{net} computed from each method for the three synthetic data sets, and the right panel shows the average resolved circular polarization fraction $\langle |v| \rangle$ computed after blurring each image with a $20 \mu\text{as}$ kernel. The results from the posterior exploration methods are presented with the median value and 2σ error bars (note that m -ring modeling strictly enforces that v_{net} is equal to the ground truth). The CLEAN-based methods produce only one image, and so only one value of each metric is reported, but 2σ error bars for DIFMAP and `polSolve` are shown from standard error propagation using the off-source residuals in \mathcal{V} and \mathcal{I} . The ground-truth value of each statistic is indicated with the black star and horizontal dashed line.

an overall dipole in the same orientation as the ground truth. No image reconstruction accurately captures the weak circular polarization in the northern half of the image.

In all three synthetic models, m -ring modeling does the best job of consistently recovering the level of \mathcal{V} in the image and the orientation of the dipolar structure across the ring. The limits on structural complexity baked into the $m_{\mathcal{V}} = 1$ model likely help the method lock on to the overall orientation of the source in these synthetic data tests, which all feature an underlying dipolar structure in \mathcal{V} . While many GRMHD images show this dipolar structure at EHT resolution, some show more complicated patterns (see Section 5). More details and other model choices applied to m -ring modeling of these synthetic data sources are presented in R23.

Figure 20 shows the recovery of the image-integrated and resolved circular polarization fractions for all data sets and methods. The low image-integrated circular polarization fraction is recovered decently well by all methods on the three sources. An exception is DIFMAP, which recovers a relatively large negative integrated circular polarization fraction. The m -ring modeling method is constrained by construction to produce the ground-truth v_{net} . The recovered resolved circular polarization $\langle |v| \rangle$ shows larger differences between methods and larger deviations from the ground truth. For the high circular polarization case 03, both DMC and THEMIS have the ground-truth value of $\langle |v| \rangle$ within their 2σ error bars, while the `polSolve` and DIFMAP single-image results are off by $\approx 1\%$. In the moderate and low polarization models 02 and 01, all

methods except m -ring modeling produce results biased significantly upward from the ground truth. This could indicate that residual uncertainty in the $G_{R/L}$ gains is entering into overproducing the resolved circular polarization in these methods with weak intrinsic \mathcal{V} and low \mathcal{V} S/N in the data, even after self-calibration during imaging. Furthermore, $\langle |v| \rangle$ is biased upward by uncertainty in the underlying image reconstructions (see Appendix G). Interestingly, m -ring modeling performs the best on data sets 01 and 02, accurately recovering $\langle |v| \rangle$ in both cases, while overproducing $\langle |v| \rangle$ in the “easier” case 03. This difference in recovery accuracy could be driven by the limited freedom of the m -ring model in the source structures it can recover; the ground-truth structure for model 03 is less dipolar and ring-like than for models 01 and 02, with both total intensity and polarized flux concentrated in the south.

Figure 21 shows the recovered amplitude gain ratios $|G_{R/L}|$ for selected stations on the three synthetic data sets. Note that THEMIS does not fit for these gain ratios, and m -ring modeling does not fit for any gains, as it directly fits to closure quantities. The trends of $|G_{R/L}|$ as a function of time are generally recovered by all methods, although systematic offsets up to $\sim 10\%$ do occur. DMC is the only method that returns uncertainty estimates on these gain ratios; the DMC results are nearly all within 2σ of the ground truth. The inability of all methods to accurately recover these gain ratios even in the model with the most circular polarization (03) indicates that residual uncertainty in the solution for $G_{R/L}$ may significantly affect our \mathcal{V} imaging results for M87*.

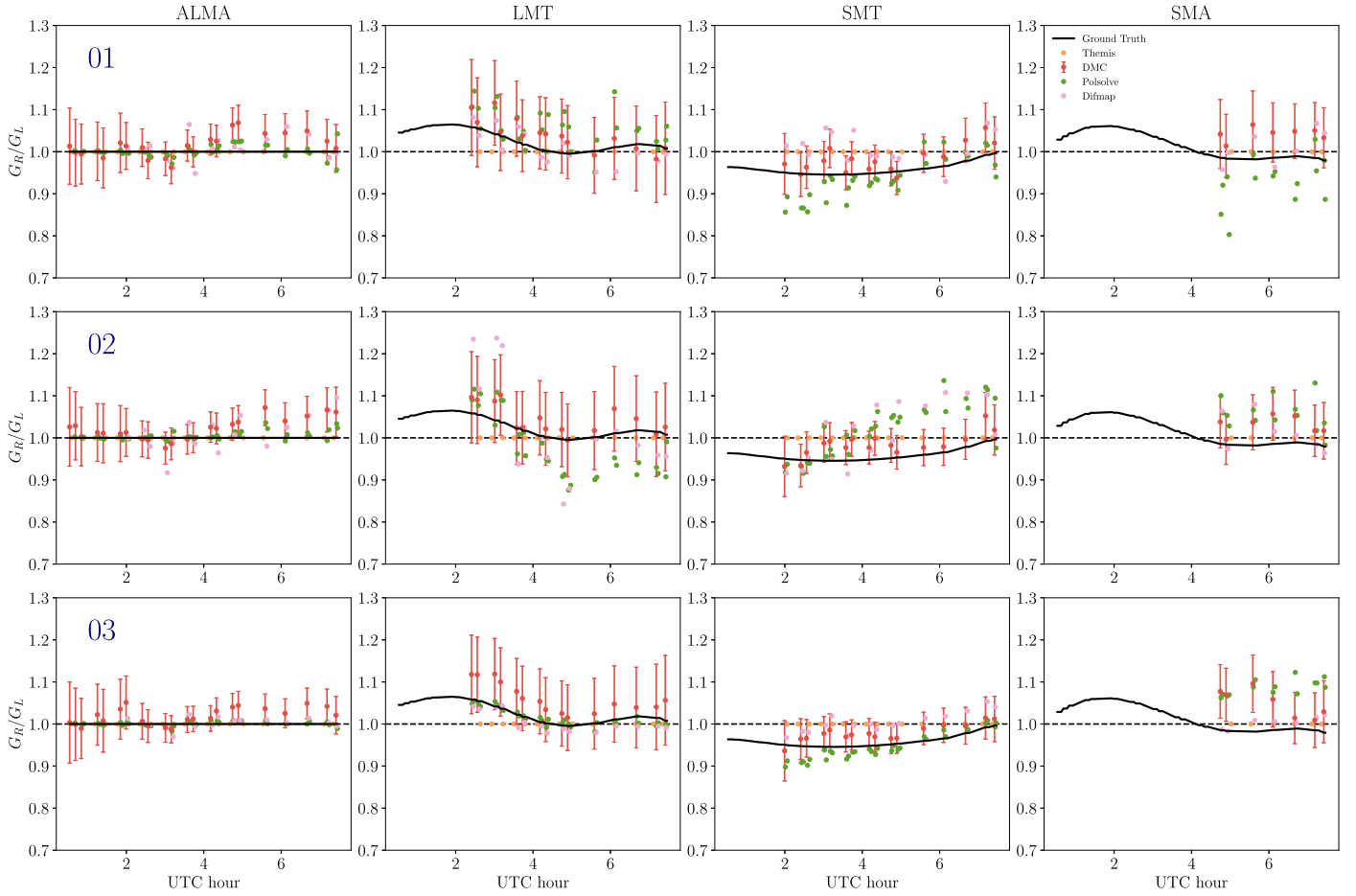


Figure 21. Derived gain ratios $|G_R|/|G_L|$ from different imaging methods applied to synthetic data set 01 (top), 02 (middle), and 03 (bottom). From left to right, gain ratios are reported for ALMA, the LMT, the SMT, and the SMA sites. The actual applied gain ratios in the synthetic data set (motivated by the limits discussed in Section 2.3) are displayed in black. DMC gains are reported with 2σ error bars.

Appendix F M87* Results from CASA-calibrated Data

Figure 22 presents Stokes \mathcal{V} imaging results using data reduced with the CASA `rPICARD` pipeline (Janssen et al. 2019). The CASA-reduced data calibrated the amplitude gain ratios $G_{R/L}$ assuming $\mathcal{V} = 0$, while the fiducial HOPS-reduced data used to produce the results in Figures 4 and 5 used a multiday and multisource fit to calibrate $G_{R/L}$. By producing images with these data, we test whether our main conclusions regarding the level of the resolved \mathcal{V} and (in)consistency of the resolved structure in different reconstructions depend strongly on the choices and assumptions made in data reduction, particularly in the way $G_{R/L}$ are calibrated before imaging.

Three of the imaging pipelines used in this paper (DIFMAP, THEMIS, and m -ring model fitting) produced Stokes \mathcal{V} images

using the CASA data for both high and low band on April 5 and April 11. The results presented in Figure 22 show that the resulting Stokes \mathcal{V} images from the CASA-calibrated data are still inconsistent across reconstruction methods and observing days. Interestingly, the reconstructions using CASA data show somewhat more consistency across the high- and low-band images than in the reconstructions using the fiducial EHT-HOPS data (Figures 4, 5). This may be a result of the fact that the `rPICARD` pipeline combines information across the bands for the signal stabilization (Janssen et al. 2022) of the data, while the bands are treated entirely separately in the EHT-HOPS reduction. R23 presents an analysis of the CASA data m -ring reconstructions compared to the EHT-HOPS data reconstructions in more detail and without any $G_{R/L}$ amplitude gain calibration.

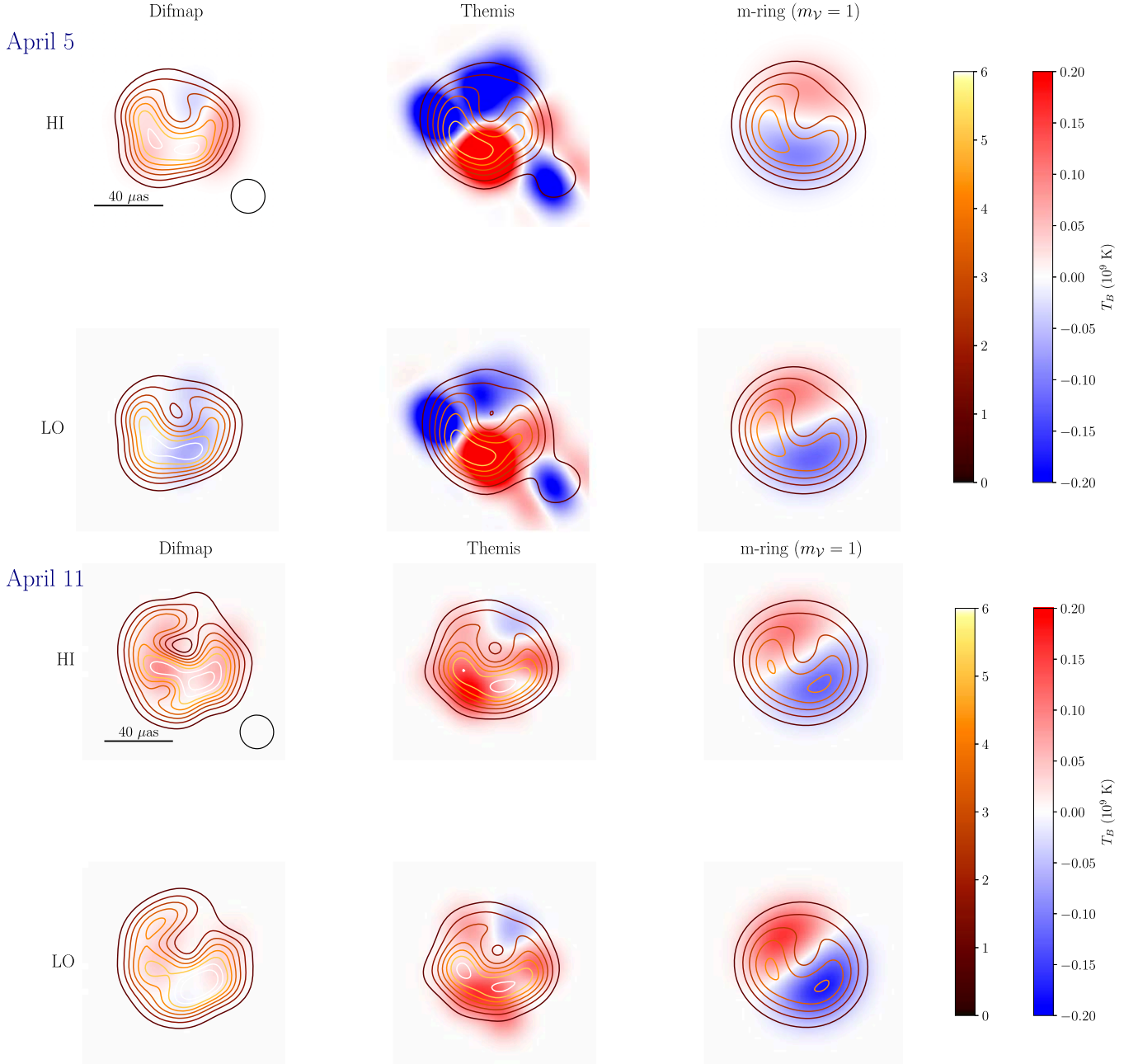


Figure 22. M87* Stokes \mathcal{V} imaging results using data reduced with the CASA-based `rPICARD` pipeline (Janssen et al. 2019) and precalibrated with the assumption $\mathcal{V} = 0$. The top two rows show results from three methods (DIFMAP, THEMIS, and m -ring geometric modeling) on April 5 observations, both low and high band. The bottom row shows corresponding results from April 11 observations. For THEMIS and m -ring modeling we show the posterior-average image. Images are plotted in the same manner as in Figure 4.

Appendix G

Biases in Resolved Circular Polarization Fraction

For an image composed of N_{res} pixels or resolution elements, each of which has angular area ΔA , Equation (5) can be expressed as a sum

$$\langle |v| \rangle = \frac{\Delta A}{F} \sum_{i=1}^{N_{\text{res}}} \left| \frac{\mathcal{V}_i}{\mathcal{I}_i} \right| \mathcal{I}_i, \quad (\text{G1})$$

where F is the total integrated Stokes \mathcal{I} flux density in the image and i is a subscript that indicates a particular pixel or resolution element. If \mathcal{I} is everywhere positive in the image, then $|\mathcal{V}_i/\mathcal{I}_i| \mathcal{I}_i = |\mathcal{V}_i|$ and we can simplify the expression for

$\langle |v| \rangle$ to be

$$\langle |v| \rangle = \frac{\Delta A}{F} \sum_{i=1}^{N_{\text{res}}} |\mathcal{V}_i|. \quad (\text{G2})$$

If \mathcal{V}_i is purely noise-like and distributed according to a normal distribution with mean zero and variance σ_i^2 , then the probability distribution function of $|\mathcal{V}|$ is given by

$$|\mathcal{V}_i| \sim \begin{cases} \sqrt{\frac{2}{\pi \sigma_i^2}} \exp\left(-\frac{\mathcal{V}_i^2}{2\sigma_i^2}\right), & |\mathcal{V}_i| \geq 0 \\ 0, & |\mathcal{V}_i| < 0. \end{cases} \quad (\text{G3})$$

The mean of this distribution is $\mu_i = \sigma_i \sqrt{2/\pi}$, which is nonzero. The mean of $\langle |v| \rangle$ will then be given by

$$E[\langle |v| \rangle] = \frac{\Delta A}{F} \sqrt{\frac{2}{\pi}} \sum_{i=1}^{N_{\text{res}}} \sigma_i. \quad (\text{G4})$$

If each pixel or resolution element has a similar noise variance σ_i^2 , then we can see that $E[\langle |v| \rangle]$ will be proportional to N_{res} . Precisely, the value of the resolved circular polarization fraction depends on the number of pixels or resolution elements contained in the image, which means that it will be sensitive to things such as the size of a blurring kernel or the field of view of the image.

Appendix H

Additional GRMHD Library Generation Details

The GRMHD models in this paper were computed using the code HARM (Gammie et al. 2003; Noble et al. 2006) and were first presented in Paper V. Simulations were initialized with an idealized Fishbone–Moncrief torus of gas in hydrostatic equilibrium (Fishbone & Moncrief 1976), within which the magnetorotational instability (Balbus & Hawley 1991, 1998) naturally develops, driving accretion. Our GRMHD simulations are mainly characterized by two parameters. The first is the BH spin a_* ranging from $-1 < a_* < 1$, with positive/negative values representing alignment/anti-alignment between the engines' angular momenta. We only consider models where the BH angular momentum and the disk angular momentum on large scales are either perfectly aligned or anti-aligned, where anti-alignment is denoted with a negative sign. The second is the absolute magnetic flux in its dimensionless form $\phi \equiv \Phi_{\text{BH}} (\dot{M} r_g^2 c)^{-1/2}$, where Φ_{BH} is the magnetic flux's magnitude traversing one hemisphere of the event horizon (see, e.g., Porth et al. 2019) and the mass accretion rate \dot{M} set over the event horizon. The relative strength of the magnetic flux near the horizon is determined by the dimensionless quantity ϕ and affects the dynamics of the accretion flow solutions. This led to a division of the solutions into two categories: the MAD state, which highly affects the accretion dynamics $\phi \geq 50$ (e.g., Tchekhovskoy et al. 2011; McKinney et al. 2012; Narayan et al. 2022), and the SANE state $\phi \approx 1 - 5$ (in Gaussian units). SANE simulations were produced on a $288 \times 128 \times 128$ grid, a fluid adiabatic index $\gamma = 4/3$, and a boundary of $r_{\text{out}} = 50 r_g$. MAD simulations adopt a $384 \times 192 \times 192$ resolution, $\gamma = 13/9$, and a domain of $r_{\text{out}} = 10^3 r_g$.

To produce images via GRRT, we use the code IPOLE (Mościbrodzka & Gammie 2018). First, null geodesics are calculated backward from the camera through the source. Then, radiative transfer is calculated forward accounting for polarized synchrotron emission and self-absorption, as well as Faraday rotation and conversion. We use radiative transfer coefficients appropriate for a thermal relativistic plasma (Leung et al. 2011), with the updated fitting functions from Marszewski et al. (2021). The fast-light approximation is made during ray-tracing.

In ideal GRMHD, the fluid density, magnetic field strength, and internal energy can be rescaled during postprocessing, and we do so to achieve an average flux density of $F_\nu \approx 0.5$ Jy for our images at $\nu = 230$ GHz (e.g., Event Horizon Telescope Collaboration et al. 2019e). In addition to the fluid scaling and BH parameters, there are three other free parameters we explore at this stage. Two of them, R_{low} and R_{high} , are associated with the ion-to-electron temperature ratio, where the electron

temperature is defined as

$$T_e = \frac{2m_e u_{\text{gas}}}{3\rho k_B (2 + R)}. \quad (\text{H1})$$

In order for the description in Equation (H1) to be valid, the ions are assumed to be nonrelativistic while electrons are assumed to be relativistic. Adopting a prescription in Mościbrodzka et al. (2016), the ion-to-electron temperature ratio R is given by

$$R = \frac{T_i}{T_e} = R_{\text{high}} \frac{\beta^2}{1 + \beta^2} + R_{\text{low}} \frac{1}{1 + \beta^2}, \quad (\text{H2})$$

where the components describe the local plasma $\beta = p_{\text{gas}}/p_{\text{mag}}$ with $p_{\text{gas}} = (\gamma - 1)u_{\text{gas}}$, $p_{\text{mag}} = B^2/8\pi$, and further u_{gas} being the total internal energy density. To regulate the ratio in Equation (H2), our models explore $R_{\text{low}} = 1, 10$ and $R_{\text{high}} = 1, 10, 20, 40, 80, 160$ as in Paper VIII.

Such a parameterization to treat the model as thermal flow is motivated by the fact that mean free path is much larger than the Debye length and Larmor radius of the jet's plasma. This assumption, together with Equation (H2), mimics collisionless plasma properties and allows us to associate the electron heating with its magnetic properties (see Paper VIII).

Appendix I

Field Reversal and Linear Polarization

As we discuss in Section 5.2, reversing the polarity of the magnetic field has significant effects on circular polarization, on both resolved and unresolved scales. Our previous studies including linear polarization of this source only included models wherein the poloidal magnetic field and the angular momentum of the disk are aligned. Here we briefly explore the effect of flipping the magnetic field on linear polarization metrics. Overall, we find that the effect is minor and does not have a strong impact on linear polarization constraints.

In Figures 23–26, we plot histograms of observable quantities considered in Paper VIII for the aligned case (as in previous work) on the left and for the anti-aligned case on the right. As expected, most of the linear polarization metrics are insensitive to this flip. The quantity that shows the most significant, albeit minor, effect is $\angle\beta_2$. As explored in more detail in Emami et al. (2023), on average, Faraday rotation imprints a small shift in the EVPA pattern corresponding to the line-of-sight magnetic field direction. For example, since we are viewing M87* nearly face-on, a poloidal field pointed toward the observer will on average Faraday rotate EVPA ticks in the counterclockwise direction, and β_2 is also rotated counterclockwise. Similarly, a poloidal field pointed away from the observer will on average rotate ticks β_2 clockwise. More complicated behavior can even arise if differences in the Faraday depth among different emitting regions cause summation or cancellation depending on the degree of Faraday rotation.

Ultimately, although flipping the polarity of the magnetic field has an effect on the distributions of $\angle\beta_2$, it has little effect on our linear polarization constraints. SANE models do not appear to change significantly in this metric. The MAD models appear to shift slightly in $\angle\beta_2$ when the field polarity is reversed. Prograde models have a slight increase, whereas retrograde models have a slight decrease in $\angle\beta_2$. This does not affect the simultaneous snapshot scoring model results but

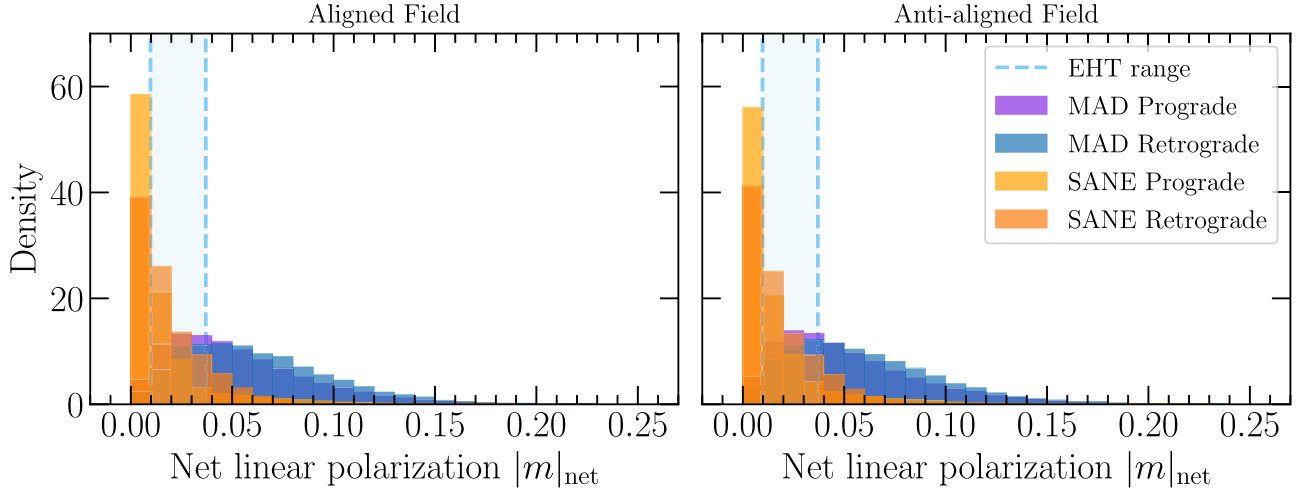


Figure 23. Distribution of image-integrated net linear polarization fraction with an aligned (left panel) and anti-aligned (right panel) magnetic field respecting all images in the M87* library. Allowed inferred ranges of EHT image reconstructions are limited by the dashed lines.

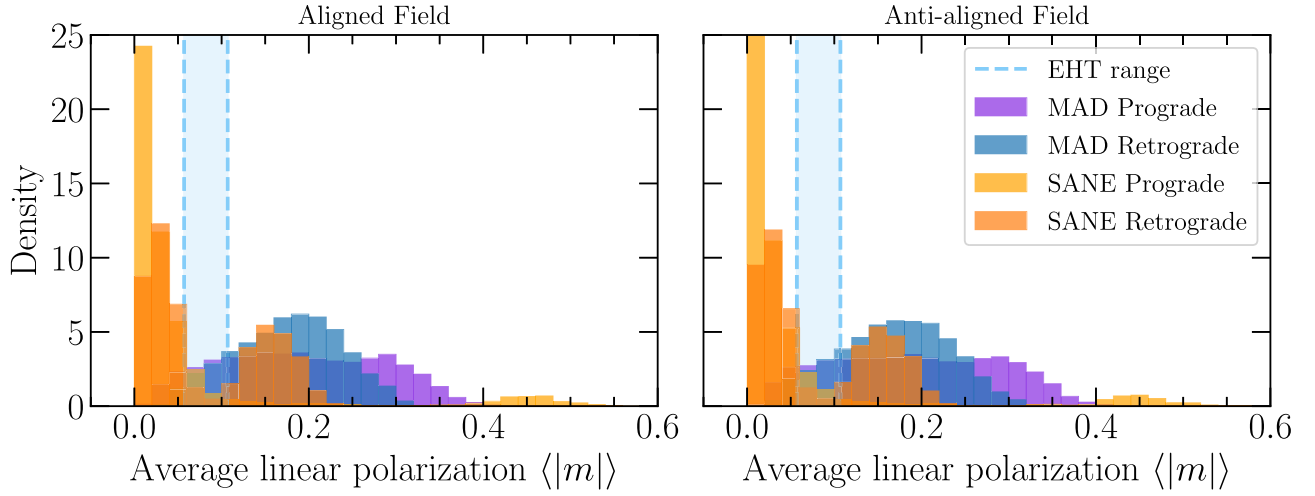


Figure 24. Distribution of image-averaged fractional linear polarization $\langle |m| \rangle$ with an aligned (left panel) and anti-aligned (right panel) magnetic field respecting all images in the M87* library blurred with a $20 \mu\text{s}$ FWHM Gaussian beam. Allowed inferred ranges of EHT image reconstructions are limited by the dashed lines.

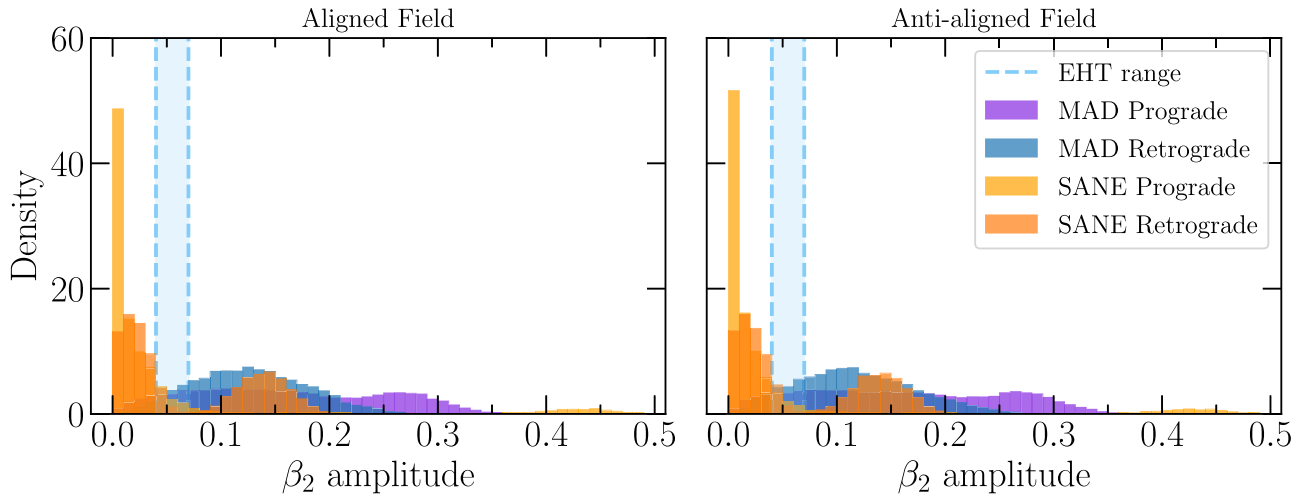


Figure 25. Distribution of β_2 amplitude for an aligned (left panel) and anti-aligned (right panel) magnetic field representing values taken from all images in the EHT M87* library blurred with a $20 \mu\text{s}$ beam. Allowed inferred ranges of EHT image reconstructions are limited by the dashed lines.

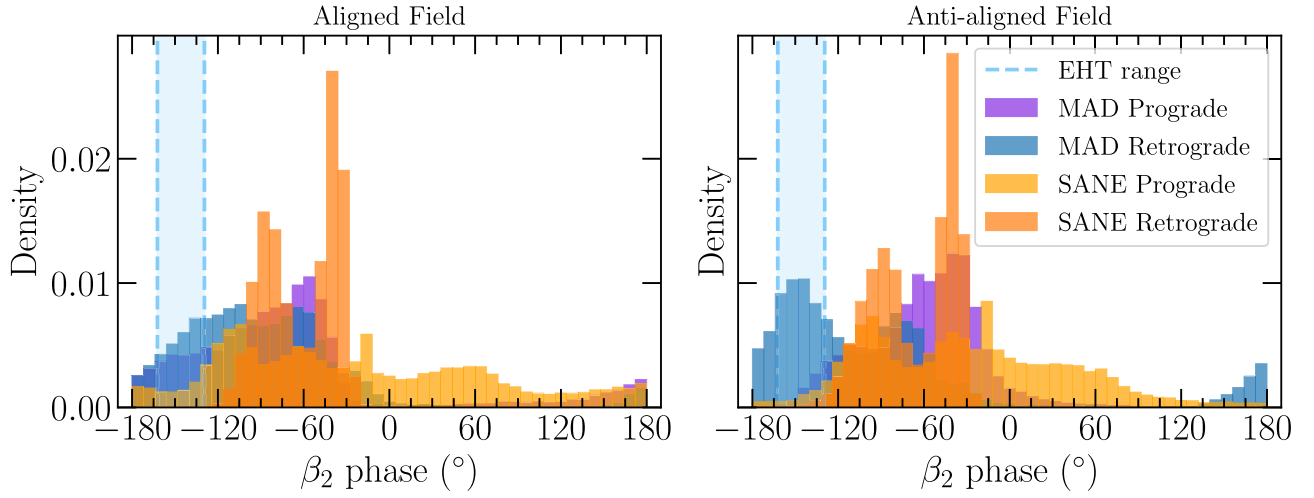


Figure 26. Distribution of β_2 phase for an aligned (left panel) and anti-aligned (right panel) magnetic field representing values taken from all images in the EHT M87* library blurred with a 20 μ as beam. Allowed inferred ranges of EHT image reconstructions are limited by the dashed lines.

causes an increased preference for retrograde MAD models in the joint scoring method. With multifrequency data, a sign flip in the magnetic field polarity should naturally lead to a sign flip in the rotation measure (e.g., Contopoulos et al. 2022), but we do not consider rotation measure in this work.

Appendix J More Details on GRMHD Scoring Results

J.1. Simultaneous Snapshot Scoring

For simultaneous scoring, the upper limit (3.7%) for $\langle|v|\rangle$ reported in this paper does not provide any additional information to constrain the GRMHD image parameter space. That is, all models that pass the other polarimetric constraints already naturally pass this constraint. This is perhaps not surprising, since the upper limit on v_{net} reported by Goddi et al. (2021) and already considered in Paper VIII is more stringent than that on $\langle|v|\rangle$ obtained in this work. Only 261/184,796 snapshots (0.14%) fail the $\langle|v|\rangle$ constraint but pass the v_{net} constraint. Producing a low v_{net} but a high $\langle|v|\rangle$ requires regions of high resolved circular polarization fraction of both positive and negative sign to nearly cancel across an image. Ultimately when combined with the linear polarization constraints, $\langle|v|\rangle$ adds no discerning power. Consequently, model predictions for M87* remain consistent with Paper VIII.

The minor differences in the simultaneous scoring histogram (left panel of Figure 12) compared to Paper VIII can be attributed not to the additional constraint on $\langle|v|\rangle$ but rather to the additional parameter probed in this paper, the magnetic field polarity, as well as slight differences in time sampling of the images and updates to radiative transfer coefficients.

To better understand how Stokes \mathcal{V} alone affects our model scoring, Figure 27 plots the fraction of snapshots of a given model that pass Stokes \mathcal{V} constraints without considering the others. We find that Stokes \mathcal{V} constraints are more likely to rule out SANE models than MADs, since SANE models are more likely to overproduce v_{net} . The $\langle|v|\rangle$ constraint alone does not rule out many models. Passing fractions appear similar for both field configurations, and thus they have been combined in the figure.

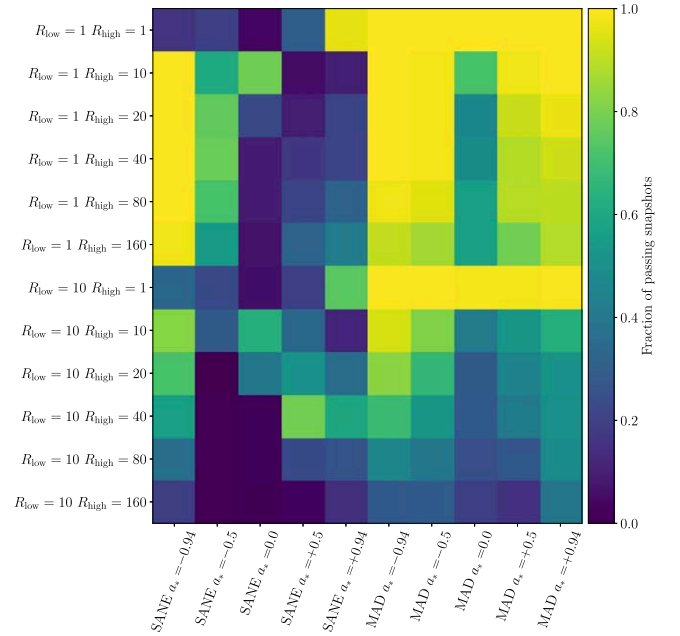


Figure 27. Passing fraction of the v_{net} and $\langle|v|\rangle$ constraints for all the models in the simulation library. Aligned and anti-aligned models are scored combined since the individual passing fractions are similar.

J.2. Joint Scoring

As shown in Figure 12, the joint scoring method favors MAD $R_{\text{low}} = 10$ models that are not highly spinning. The inclusion of reversed-field images greatly increases the relative likelihood for MAD spin -0.5 models, particularly for $\angle\beta_{2,\text{LP}}$. The joint scoring results appear qualitatively different from that given in Paper VIII, which is likely due to differences in the image library sets rather than just the $\langle|v|\rangle$ constraint.

The inclusion of the $\langle|v|\rangle$ constraint removes small likelihoods for MAD spin 0, -0.5 , $R_{\text{low}} = 1$ models but otherwise does not qualitatively change the likelihoods. Note that the midpoint for the $\langle|v|\rangle$ allowed range is 1.85%, or half of the conservative upper limit. The results from the joint scoring method are more sensitive to this choice than is the case in simultaneous scoring. Using an aggressive upper limit on $\langle|v|\rangle$ of 2% from THEMIS (Table 2), we find that qualitatively the

joint scoring results remain largely the same, but some $R_{\text{low}} = 1$ models now are included, as those models have lower $\langle |v| \rangle$. A lower limit on $\langle |v| \rangle$ at 1% tends to increase the likelihood of the MAD spin +0.5, $R_{\text{low}} = 10$ models.

Comparisons between the image library sets of this paper and Paper VIII can yield insight into the differences in the joint scoring method. For example, the relatively lower likelihoods for the MAD spin +0.5 models in Figure 12 can in part be attributed to only scoring images at 17° or 163° inclination, while Paper VIII included models at inclinations $\pm 5^\circ$ from 17° or 163° . Models from Paper VIII scored using only 17° or 163° inclination cause the relative likelihood for MAD spin +0.5 models to decrease. The time sampling in the library sets can also influence the model scoring. In this paper, we sample uniformly for each model at 5M, while in Paper VIII the number of snapshots per model is kept fixed to 200, with the sampling cadence varying. Using the sampling in Paper VIII causes a slight increase to the likelihoods for the prograde MAD models. Moreover, the library set in Paper VIII excludes emission from within 5% of the event horizon, while the current image library consists of emission from further inside. While this only affects models that are very optically thin (MAD models), it can change the EVPA structure and thus the $\beta_{2,\text{LP}}$ modes. The $\beta_{2,\text{LP}}$ modes appear to be the most discrepant between the two library sets, and most linear and circular polarization constraints otherwise remain consistent.

Appendix K Stokes \mathcal{V} Origin in Simulations

Here we present in more detail the test summarized in Section 5.5 to determine the physical origin of Stokes \mathcal{V} emission in physical models. We investigate the full set of 288 snapshots that simultaneously pass all polarimetric constraints. In IPOLE, our GRRT code, it is straightforward to probe the effects of each physical effect by setting the appropriate radiative transfer coefficient to 0. Here the most relevant coefficients are intrinsic

circularly polarized emission (j_v), Faraday conversion (ρ_Q), and Faraday rotation (ρ_V). For details of these coefficients, we refer the reader to Mościbrodzka & Gammie (2018).

In this experiment, we reimage all the passing snapshots, but each time we turn off one of these coefficients. Selecting one representative snapshot per passing model, we compare both v_{net} and $\langle |v| \rangle$ against the full radiative transfer solution and plot the results in Figure 28. Since significant cancellation occurs in the image plane for Stokes \mathcal{V} , large differences to the image do not necessarily correspond to large changes in v_{net} . Therefore, $\langle |v| \rangle$ is somewhat more informative and robust to cancellations, although cancellations within the Gaussian kernel of $20 \mu\text{as}$ still occur.

From Figure 28 we note three main results:

1. Turning off intrinsic Stokes \mathcal{V} emission (j_v) does not appear to greatly influence either v_{net} or $\langle |v| \rangle$ (first column of Figure 28). While this is not true for every snapshot in every model, for most of the passing models we find that intrinsic emission is not the dominant origin of Stokes \mathcal{V} .
2. On the other hand, turning off Faraday conversion ($\rho_Q = 0$) greatly suppresses $\langle |v| \rangle$ (second row, second column in Figure 28), suggesting that Faraday conversion is the dominant production mechanism for Stokes \mathcal{V} in passing models.
3. Interestingly, turning off Faraday rotation increases $\langle |v| \rangle$ for some of the passing models. This implies that scrambling via Faraday rotation is important not only for reducing the linear polarization fraction but also for indirectly reducing the circular polarization fraction. In these models, Faraday rotation partly scrambles the linear polarization that is converted into circular.

In summary, Faraday effects are critical for Stokes \mathcal{V} generation in our models, while intrinsic emission plays a subdominant, though nonnegligible, role.

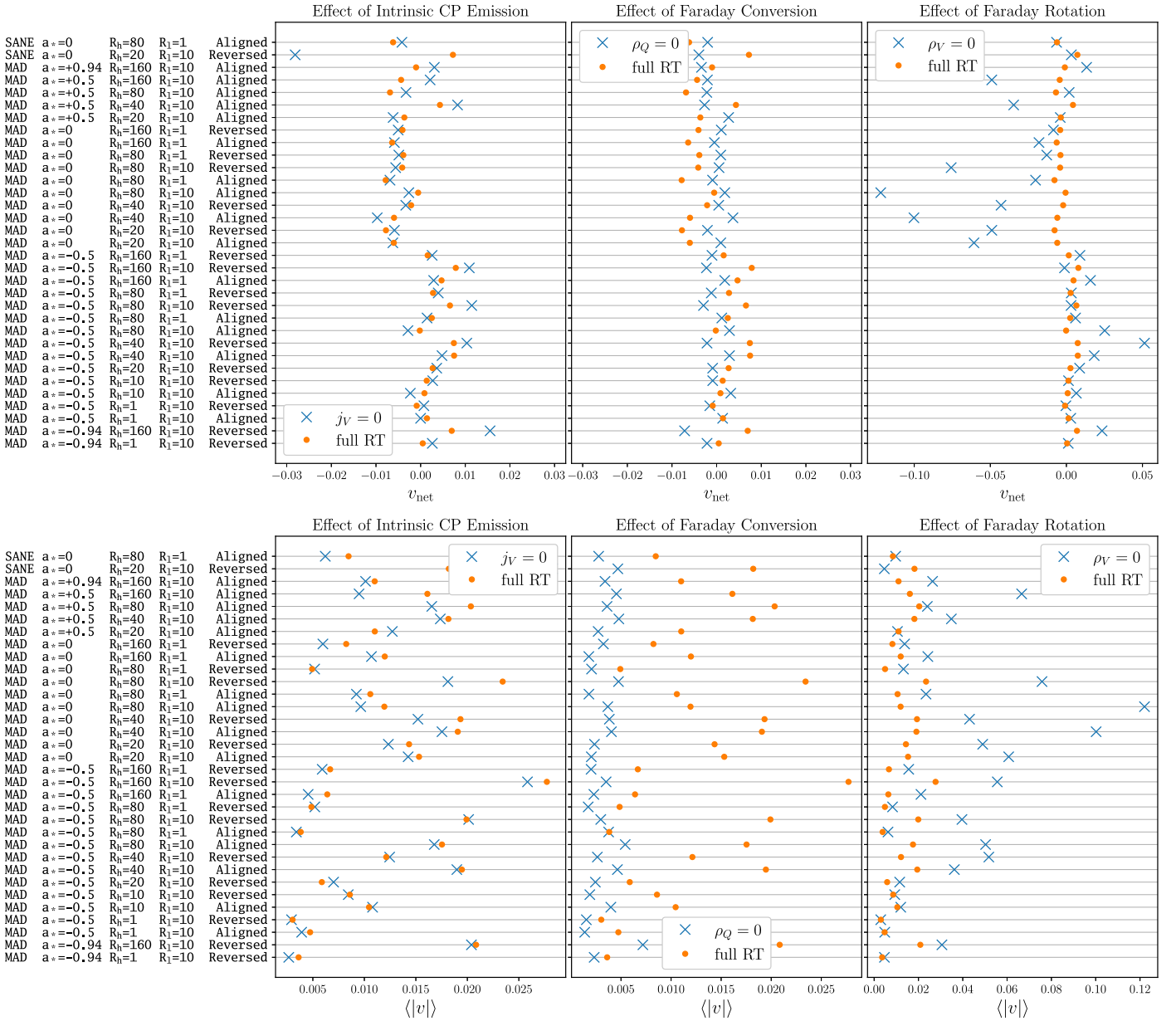


Figure 28. Exploration of the effects of various CP production mechanisms on v_{net} (top panels) and $\langle |v| \rangle$ (bottom panels). Each column corresponds to a comparison between the full radiative transfer (RT) simulation and the same calculation with one coefficient turned off: $j_V = 0$ (intrinsic emission) for the first column, $\rho_Q = 0$ (Faraday conversion) for the second column, and $\rho_V = 0$ (Faraday rotation) for the third column. Each tick on the y-axis corresponds to a model for which at least one snapshot passes all the polarimetric constraints. A single snapshot from each of these models has been plotted at each tick.

Appendix L Alternative Plasma Models

Throughout this work, we have modeled a pure ion–electron plasma with thermal electron distributions. Here we explore our sensitivity to the details of the plasma by considering the existence of electron–positron pairs and nonthermal electron distribution functions. In brief, our Stokes \mathcal{V} images are indeed sensitive to these details, motivating future studies in these areas.

L.1. Ionic versus Pair Plasmas

Unlike an ionic plasma (assumed throughout this paper), a pure pair plasma produces no circular polarization via synchrotron emission and has no Faraday rotation but can still perform Faraday conversion (Jones & O’Dell 1977; Jones 1988; Wardle et al. 1998). For ionic plasma models, all of the radiative transfer

coefficients are believed to be important for producing the Stokes \mathcal{V} image. On EHT scales, the few recent studies on this topic find that increasing the pair fraction can alter both resolved Stokes \mathcal{V} images at 230 GHz and the evolution as a function of frequency (Anantua et al. 2020; Emami et al. 2021, 2023).

In Figure 29, we test a single snapshot of the MAD $a = +0.5 R_{\text{high}} = 80 R_{\text{low}} = 10$ aligned-field model, performing polarized ray-tracing with an increasing positron-to-electron number density ratio, denoted as f in each panel. To obtain a given value of f , we inject electron–positron pairs into each cell, representing a simplistic scenario in which a large fraction of pairs are produced at number densities directly proportional to the number densities of preexisting electrons. Since this would drastically increase the number density of emitting leptons in each successive panel, we also find a new value of

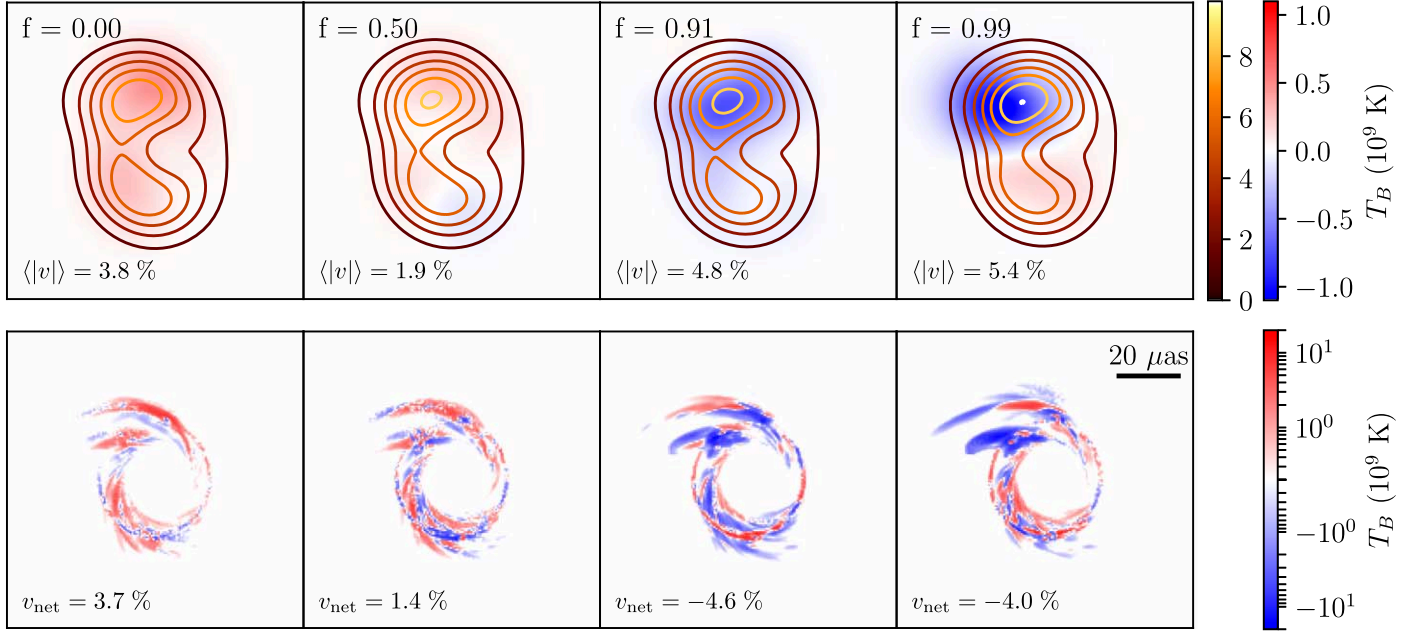


Figure 29. Test of plasma content, where we ray-trace a single snapshot of the MAD $a = +0.5$ $R_{\text{high}} = 80$ $R_{\text{low}} = 10$ aligned-field model with an increasing positron-to-electron ratio, denoted as f in the upper left corner of each panel. The Stokes \mathcal{V} structure clearly evolves as f increases, but we do not observe a clear discriminant of plasma content.

the density normalization for each value of f , such that the total flux is always 0.5 Jy as in the original snapshot.

For a given value of f , the radiative transfer coefficients are modified via (e.g., MacDonald & Marscher 2018; Emami et al. 2021)

$$\begin{aligned}
 j_{I,Q,U} &\rightarrow (1 + f)j_{I,Q,U}, \\
 j_V &\rightarrow (1 - f)j_V, \\
 \alpha_{I,Q,U} &\rightarrow (1 + f)\alpha_{I,Q,U}, \\
 \alpha_V &\rightarrow (1 - f)\alpha_V, \\
 \rho_{Q,U} &\rightarrow (1 + f)\rho_{Q,U}, \\
 \rho_V &\rightarrow (1 - f)\rho_V,
 \end{aligned} \tag{L1}$$

capturing the effects described above. Figure 29 illustrates that the circularly polarized morphology of this snapshot clearly and strongly evolves with f , consistent with previous works. However, given the inherent diversity of Stokes \mathcal{V} structures among the library, there is no known signature in Stokes \mathcal{V} that clearly indicates the presence of pairs. Interestingly, we find that $\langle |v| \rangle$ increases monotonically in the perfect-resolution images, which can be attributed to drastically falling Faraday rotation depth (to be discussed in R. Anantua et al., in preparation). However, due to cancellations within the EHT beam, $\langle |v| \rangle$ does *not* necessarily increase monotonically at EHT resolution. In summary, we are sensitive to the content of the emitting plasma, but we are unaware of a clear signature distinguishing a pair plasma from an ionic plasma on event horizon scales where both intrinsic emission and Faraday conversion are important. While not explored here, images originating from GRMHD are also surprisingly sensitive to the atomic composition of a given ionic plasma (Wong & Gammie 2022). Identifying clear signatures to distinguish different models of plasma content on event horizon scales remains a topic of ongoing research.

L.2. Nonthermal Electrons

Throughout this work, we consider only models with thermal electron distribution functions (eDFs). However, it is believed that nonthermal particle acceleration can occur as a result of magnetic reconnection, MHD turbulence, and collective plasma modes. At present, a consensus model for the presence of nonthermal electrons has not been found, but here we explore two physically motivated implementations. For a single snapshot, we compare their thermal eDF images to images produced assuming “kappa” models, which feature a thermal core and a high-energy nonthermal tail (Vasyliunas 1968). Such distributions have been observed in the solar wind (Decker & Krimigis 2003; Pierrard & Lazar 2010) and occur naturally in particle-in-cell (PIC) simulations of particle acceleration in magnetized plasmas (Kunz et al. 2016). We try both a constant $\kappa = 5$ model, where $\kappa - 1 = 4$ is the power-law index of the high-energy tail, and a variable kappa model where $\kappa = \kappa(\sigma, \beta)$ as fit by the PIC simulations of Ball et al. (2016). Due to limitations of the fitting functions used, any time the prescription would assign $\kappa > 7$, the code instead uses a radiative transfer coefficient appropriate for a thermal eDF.

In Figure 30, we spot-check our sensitivity to our assumption of thermal electrons by ray-tracing a GRMHD snapshot with three different assumptions for the electron distribution function. A new plasma density scale is found for each image to match the total flux of the image with thermal electrons, 0.7 Jy. This snapshot corresponds to the MAD $a = +0.5$ $R_{\text{high}} = 80$ $R_{\text{low}} = 10$ aligned-field model. In the top row, we plot images blurred to EHT resolution, while in the bottom row we plot perfect-resolution images of Stokes \mathcal{V} in symmetric logarithmic scale. Compared to the leftmost image assuming thermal electrons, we find very little changes in the variable kappa model shown in the middle panel. This is broadly consistent with our findings in Event Horizon Telescope Collaboration et al. (2022d) for Sgr A* in total

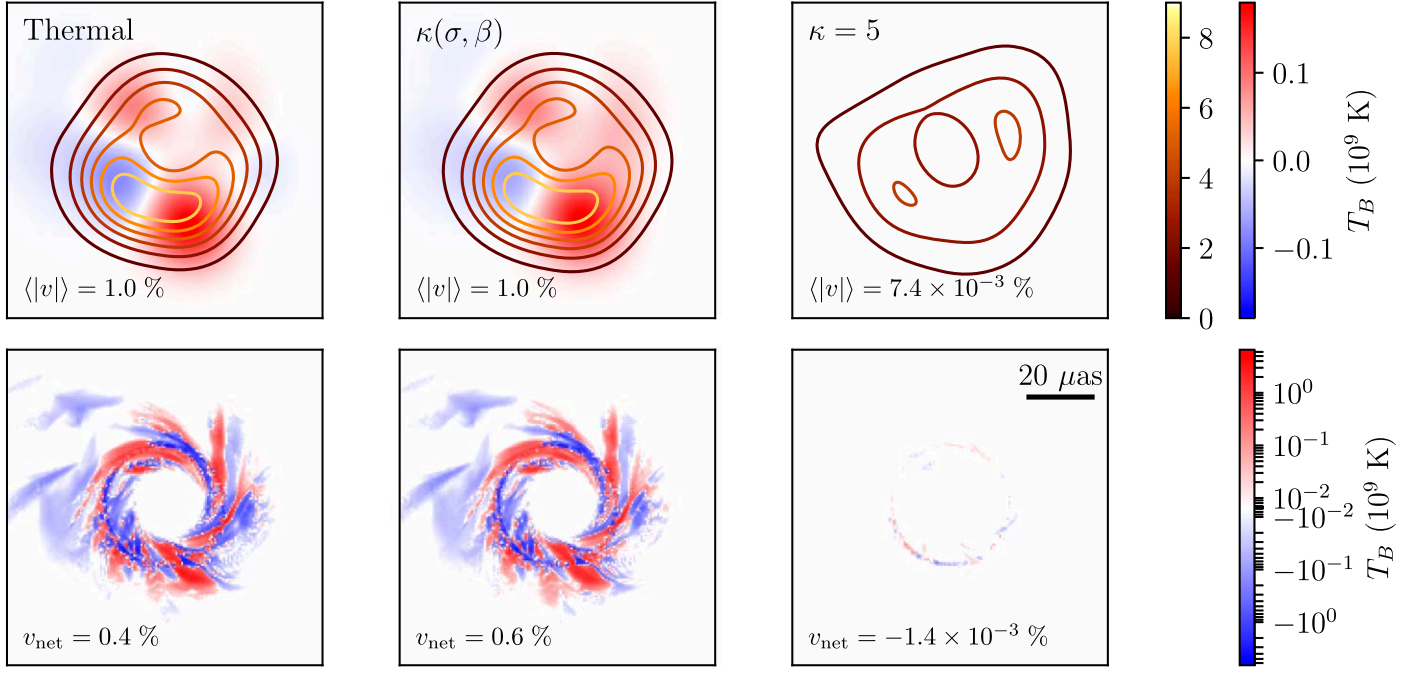


Figure 30. A single GRMHD snapshot (MAD $a = +0.5$ $R_{\text{high}} = 80$ $R_{\text{low}} = 10$ aligned field) ray-traced with three different electron distribution functions. The physically motivated variable kappa model produces both a Stokes \mathcal{I} and Stokes \mathcal{V} image very similar to the thermal model. However, a model with fixed $\kappa = 5$ produces a much more diffuse Stokes \mathcal{I} and extremely little Stokes \mathcal{V} .

intensity. Thus, at least for this snapshot, exchanging a thermal distribution for a physically motivated nonthermal electron distribution has very little effect on Stokes \mathcal{V} . However, we report dramatic differences when switching to a constant kappa model with a value of $\kappa = 5$. As found in many other studies, nonthermal electrons make the image noticeably larger and more diffuse (e.g., Özel 2000; Mao et al. 2017; Fromm et al. 2022; Ricarte et al. 2023). Intriguingly, although the plasma density is decreased only by a factor of 4 compared to the thermal model, the Stokes \mathcal{V} signal almost entirely vanishes. This may be due to the fact that Faraday conversion is caused by the coldest relativistic electrons, which occur at a smaller fraction in κ models by definition.

Appendix M Stokes \mathcal{V} Variability in Passing Models

In Figure 31, we consider each model's distribution of v_{net} over time and plot its standard deviation, which we denote as $\sigma_{v_{\text{net}}}$. We average together the standard deviations for each magnetic field polarity, since their distributions can be disjoint. We find that models with larger R_{low} are more variable in v_{net} , with a similar but weaker trend in R_{high} . Although not shown, we find that the variability of $\langle |v| \rangle$ is qualitatively similar. Some models have $\sigma_{v_{\text{net}}}$ in excess of our upper limit of 0.008, suggesting that future observations during more favorable conditions may result in higher-S/N detections of Stokes \mathcal{V} on EHT baselines and hence more robust image reconstructions of the source.

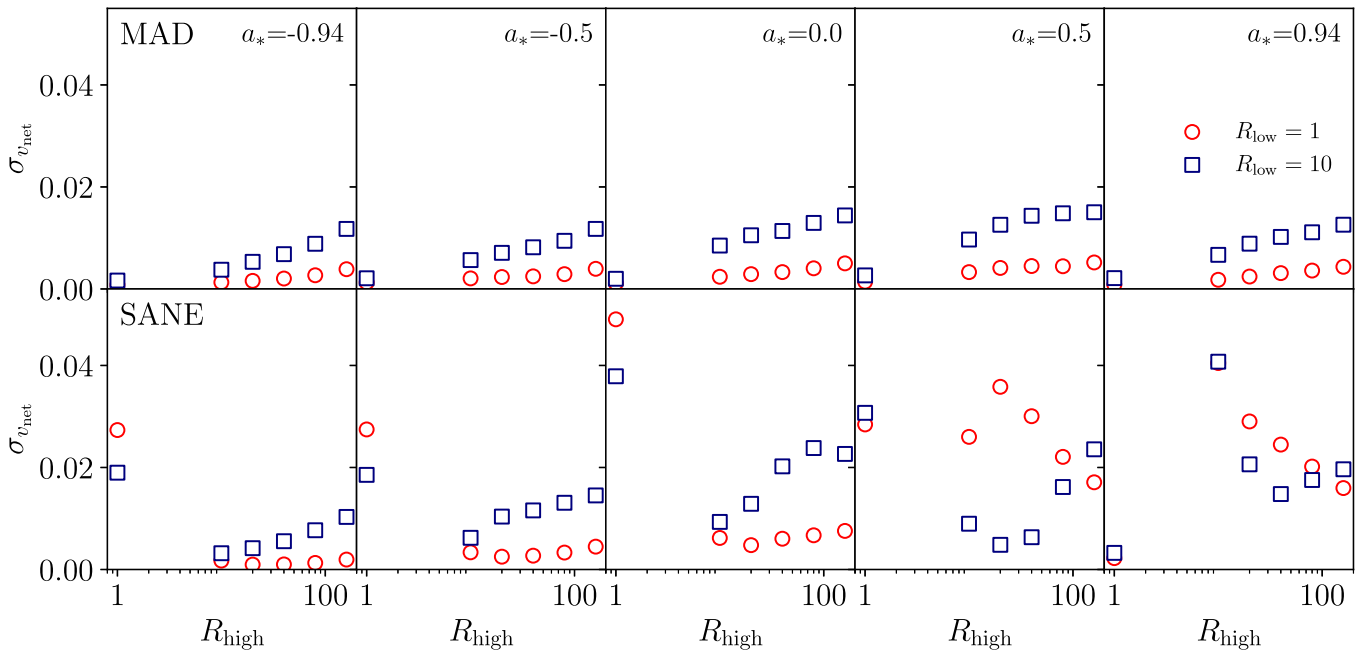




































































































Figure 31. Standard deviations of the distributions of v_{net} in GRMHD models over time. The standard deviations of models with different magnetic field polarities have been averaged together. Different magnetic field states are plotted in different rows, different spins are plotted in different columns, R_{high} is plotted on the x-axis of each panel, and the color and shape of the markers encode R_{low} . In particular, models with $R_{\text{low}} = 10$ are more variable than those with $R_{\text{low}} = 1$, and a similar but weaker trend occurs with R_{high} .

ORCID iDs

Kazunori Akiyama <https://orcid.org/0000-0002-9475-4254>
 Antxon Alberdi <https://orcid.org/0000-0002-9371-1033>
 Juan Carlos Algaba <https://orcid.org/0000-0001-6993-1696>
 Richard Anantua <https://orcid.org/0000-0003-3457-7660>
 Keiichi Asada <https://orcid.org/0000-0001-6988-8763>
 Rebecca Azulay <https://orcid.org/0000-0002-2200-5393>
 Uwe Bach <https://orcid.org/0000-0002-7722-8412>
 Anne-Kathrin Baczko <https://orcid.org/0000-0003-3090-3975>
 Mislav Baloković <https://orcid.org/0000-0003-0476-6647>
 John Barrett <https://orcid.org/0000-0002-9290-0764>
 Michi Bauböck <https://orcid.org/0000-0002-5518-2812>
 Bradford A. Benson <https://orcid.org/0000-0002-5108-6823>
 Lindy Blackburn <https://orcid.org/0000-0002-9030-642X>
 Raymond Blundell <https://orcid.org/0000-0002-5929-5857>
 Katherine L. Bouman <https://orcid.org/0000-0003-0077-4367>
 Geoffrey C. Bower <https://orcid.org/0000-0003-4056-9982>
 Hope Boyce <https://orcid.org/0000-0002-6530-5783>
 Christiaan D. Brinkerink <https://orcid.org/0000-0002-2322-0749>
 Roger Brissenden <https://orcid.org/0000-0002-2556-0894>
 Silke Britzen <https://orcid.org/0000-0001-9240-6734>
 Avery E. Broderick <https://orcid.org/0000-0002-3351-760X>
 Dominique Brogiere <https://orcid.org/0000-0001-9151-6683>
 Thomas Bronzwaer <https://orcid.org/0000-0003-1151-3971>
 Sandra Bustamante <https://orcid.org/0000-0001-6169-1894>
 Do-Young Byun <https://orcid.org/0000-0003-1157-4109>
 John E. Carlstrom <https://orcid.org/0000-0002-2044-7665>
 Chiara Ceccobello <https://orcid.org/0000-0002-4767-9925>
 Andrew Chael <https://orcid.org/0000-0003-2966-6220>

Chi-kwan Chan <https://orcid.org/0000-0001-6337-6126>
 Dominic O. Chang <https://orcid.org/0000-0001-9939-5257>
 Koushik Chatterjee <https://orcid.org/0000-0002-2825-3590>
 Shami Chatterjee <https://orcid.org/0000-0002-2878-1502>
 Ming-Tang Chen <https://orcid.org/0000-0001-6573-3318>
 Yongjun Chen (陈永军) <https://orcid.org/0000-0001-5650-6770>
 Xiaopeng Cheng <https://orcid.org/0000-0003-4407-9868>
 Ilje Cho <https://orcid.org/0000-0001-6083-7521>
 Pierre Christian <https://orcid.org/0000-0001-6820-9941>
 Nicholas S. Conroy <https://orcid.org/0000-0003-2886-2377>
 John E. Conway <https://orcid.org/0000-0003-2448-9181>
 James M. Cordes <https://orcid.org/0000-0002-4049-1882>
 Thomas M. Crawford <https://orcid.org/0000-0001-9000-5013>
 Geoffrey B. Crew <https://orcid.org/0000-0002-2079-3189>
 Alejandro Cruz-Orsorio <https://orcid.org/0000-0002-3945-6342>
 Yuzhu Cui (崔玉竹) <https://orcid.org/0000-0001-6311-4345>
 Rohan Dahale <https://orcid.org/0000-0001-6982-9034>
 Jordy Davelaar <https://orcid.org/0000-0002-2685-2434>
 Mariafelicia De Laurentis <https://orcid.org/0000-0002-9945-682X>
 Roger Deane <https://orcid.org/0000-0003-1027-5043>
 Jessica Dempsey <https://orcid.org/0000-0003-1269-9667>
 Gregory Desvignes <https://orcid.org/0000-0003-3922-4055>
 Jason Dexter <https://orcid.org/0000-0003-3903-0373>
 Vedant Dhruv <https://orcid.org/0000-0001-6765-877X>
 Sheperd S. Doeleman <https://orcid.org/0000-0002-9031-0904>
 Sean Dougal <https://orcid.org/0000-0002-3769-1314>
 Sergio A. Dzib <https://orcid.org/0000-0001-6010-6200>
 Ralph P. Eatough <https://orcid.org/0000-0001-6196-4135>
 Razieh Emami <https://orcid.org/0000-0002-2791-5011>

- Heino Falcke  <https://orcid.org/0000-0002-2526-6724>
 Joseph Farah  <https://orcid.org/0000-0003-4914-5625>
 Vincent L. Fish  <https://orcid.org/0000-0002-7128-9345>
 Ed Fomalont  <https://orcid.org/0000-0002-9036-2747>
 H. Alyson Ford  <https://orcid.org/0000-0002-9797-0972>
 Marianna Foschi  <https://orcid.org/0000-0001-8147-4993>
 Raquel Fraga-Encinas  <https://orcid.org/0000-0002-5222-1361>
 Per Friberg  <https://orcid.org/0000-0002-8010-8454>
 Christian M. Fromm  <https://orcid.org/0000-0002-1827-1656>
 Antonio Fuentes  <https://orcid.org/0000-0002-8773-4933>
 Peter Galison  <https://orcid.org/0000-0002-6429-3872>
 Charles F. Gammie  <https://orcid.org/0000-0001-7451-8935>
 Roberto García  <https://orcid.org/0000-0002-6584-7443>
 Olivier Gentaz  <https://orcid.org/0000-0002-0115-4605>
 Boris Georgiev  <https://orcid.org/0000-0002-3586-6424>
 Ciriaco Goddi  <https://orcid.org/0000-0002-2542-7743>
 Roman Gold  <https://orcid.org/0000-0003-2492-1966>
 Arturo I. Gómez-Ruiz  <https://orcid.org/0000-0001-9395-1670>
 José L. Gómez  <https://orcid.org/0000-0003-4190-7613>
 Minfeng Gu (顾敏峰)  <https://orcid.org/0000-0002-4455-6946>
 Mark Gurwell  <https://orcid.org/0000-0003-0685-3621>
 Kazuhiro Hada  <https://orcid.org/0000-0001-6906-772X>
 Daryl Haggard  <https://orcid.org/0000-0001-6803-2138>
 Michael H. Hecht  <https://orcid.org/0000-0002-4114-4583>
 Ronald Hesper  <https://orcid.org/0000-0003-1918-6098>
 Dirk Heumann  <https://orcid.org/0000-0002-7671-0047>
 Luis C. Ho (何子山)  <https://orcid.org/0000-0001-6947-5846>
 Paul Ho  <https://orcid.org/0000-0002-3412-4306>
 Mareki Honma  <https://orcid.org/0000-0003-4058-9000>
 Chih-Wei L. Huang  <https://orcid.org/0000-0001-5641-3953>
 Lei Huang (黄磊)  <https://orcid.org/0000-0002-1923-227X>
 Shiro Ikeda  <https://orcid.org/0000-0002-2462-1448>
 C. M. Violette Impellizzeri  <https://orcid.org/0000-0002-3443-2472>
 Makoto Inoue  <https://orcid.org/0000-0001-5037-3989>
 Sara Issaoun  <https://orcid.org/0000-0002-5297-921X>
 David J. James  <https://orcid.org/0000-0001-5160-4486>
 Buell T. Jannuzi  <https://orcid.org/0000-0002-1578-6582>
 Michael Janssen  <https://orcid.org/0000-0001-8685-6544>
 Britton Jeter  <https://orcid.org/0000-0003-2847-1712>
 Wu Jiang (江悟)  <https://orcid.org/0000-0001-7369-3539>
 Alejandra Jiménez-Rosales  <https://orcid.org/0000-0002-2662-3754>
 Michael D. Johnson  <https://orcid.org/0000-0002-4120-3029>
 Svetlana Jorstad  <https://orcid.org/0000-0001-6158-1708>
 Abhishek V. Joshi  <https://orcid.org/0000-0002-2514-5965>
 Taehyun Jung  <https://orcid.org/0000-0001-7003-8643>
 Mansour Karami  <https://orcid.org/0000-0001-7387-9333>
 Ramesh Karuppusamy  <https://orcid.org/0000-0002-5307-2919>
 Tomohisa Kawashima  <https://orcid.org/0000-0001-8527-0496>
 Garrett K. Keating  <https://orcid.org/0000-0002-3490-146X>
 Mark Kettenis  <https://orcid.org/0000-0002-6156-5617>
 Dong-Jin Kim  <https://orcid.org/0000-0002-7038-2118>
 Jae-Young Kim  <https://orcid.org/0000-0001-8229-7183>
 Jongsoo Kim  <https://orcid.org/0000-0002-1229-0426>
 Junhan Kim  <https://orcid.org/0000-0002-4274-9373>
 Motoki Kino  <https://orcid.org/0000-0002-2709-7338>
 Jun Yi Koay  <https://orcid.org/0000-0002-7029-6658>
 Prashant Kocherlakota  <https://orcid.org/0000-0001-7386-7439>
 Patrick M. Koch  <https://orcid.org/0000-0003-2777-5861>
 Shoko Koyama  <https://orcid.org/0000-0002-3723-3372>
 Carsten Kramer  <https://orcid.org/0000-0002-4908-4925>
 Joana A. Kramer  <https://orcid.org/0009-0003-3011-0454>
 Michael Kramer  <https://orcid.org/0000-0002-4175-2271>
 Thomas P. Krichbaum  <https://orcid.org/0000-0002-4892-9586>
 Cheng-Yu Kuo  <https://orcid.org/0000-0001-6211-5581>
 Noemi La Bella  <https://orcid.org/0000-0002-8116-9427>
 Tod R. Lauer  <https://orcid.org/0000-0003-3234-7247>
 Daeyoung Lee  <https://orcid.org/0000-0002-3350-5588>
 Sang-Sung Lee  <https://orcid.org/0000-0002-6269-594X>
 Po Kin Leung  <https://orcid.org/0000-0002-8802-8256>
 Aviad Levis  <https://orcid.org/0000-0001-7307-632X>
 Zhiyuan Li (李志远)  <https://orcid.org/0000-0003-0355-6437>
 Rocco Lico  <https://orcid.org/0000-0001-7361-2460>
 Greg Lindahl <https://orcid.org/0000-0002-6100-4772>
 Michael Lindqvist <https://orcid.org/0000-0002-3669-0715>
 Mikhail Lisakov <https://orcid.org/0000-0001-6088-3819>
 Jun Liu (刘俊) <https://orcid.org/0000-0002-7615-7499>
 Kuo Liu <https://orcid.org/0000-0002-2953-7376>
 Elisabetta Liuzzo <https://orcid.org/0000-0003-0995-5201>
 Wen-Ping Lo <https://orcid.org/0000-0003-1869-2503>
 Andrei P. Lobanov <https://orcid.org/0000-0003-1622-1484>
 Laurent Loinard <https://orcid.org/0000-0002-5635-3345>
 Colin J. Lonsdale <https://orcid.org/0000-0003-4062-4654>
 Amy E. Lowitz <https://orcid.org/0000-0002-4747-4276>
 Ru-Sen Lu (路如森) <https://orcid.org/0000-0002-7692-7967>
 Nicholas R. MacDonald <https://orcid.org/0000-0002-6684-8691>
 Jirong Mao (毛基荣) <https://orcid.org/0000-0002-7077-7195>
 Nicola Marchili <https://orcid.org/0000-0002-5523-7588>
 Sera Markoff <https://orcid.org/0000-0001-9564-0876>
 Daniel P. Marrone <https://orcid.org/0000-0002-2367-1080>
 Alan P. Marscher <https://orcid.org/0000-0001-7396-3332>
 Iván Martí-Vidal <https://orcid.org/0000-0003-3708-9611>
 Satoki Matsushita <https://orcid.org/0000-0002-2127-7880>
 Lynn D. Matthews <https://orcid.org/0000-0002-3728-8082>
 Lia Medeiros <https://orcid.org/0000-0003-2342-6728>
 Karl M. Menten <https://orcid.org/0000-0001-6459-0669>
 Daniel Michalik <https://orcid.org/0000-0002-7618-6556>
 Izumi Mizuno <https://orcid.org/0000-0002-7210-6264>
 Yosuke Mizuno <https://orcid.org/0000-0002-8131-6730>
 James M. Moran <https://orcid.org/0000-0002-3882-4414>
 Kotaro Moriyama <https://orcid.org/0000-0003-1364-3761>
 Monika Moscibrodzka <https://orcid.org/0000-0002-4661-6332>
 Wanga Mulaudzi <https://orcid.org/0000-0003-4514-625X>
 Cornelia Müller <https://orcid.org/0000-0002-2739-2994>
 Hendrik Müller <https://orcid.org/0000-0002-9250-0197>
 Alejandro Mus <https://orcid.org/0000-0003-0329-6874>
 Gibwa Musoke <https://orcid.org/0000-0003-1984-189X>
 Ioannis Myserlis <https://orcid.org/0000-0003-3025-9497>
 Andrew Nadolski <https://orcid.org/0000-0001-9479-9957>

- Hiroshi Nagai  <https://orcid.org/0000-0003-0292-3645>
 Neil M. Nagar  <https://orcid.org/0000-0001-6920-662X>
 Masanori Nakamura  <https://orcid.org/0000-0001-6081-2420>
 Ramesh Narayan  <https://orcid.org/0000-0002-1919-2730>
 Gopal Narayanan  <https://orcid.org/0000-0002-4723-6569>
 Iniyan Natarajan  <https://orcid.org/0000-0001-8242-4373>
 Joey Neilsen  <https://orcid.org/0000-0002-8247-786X>
 Roberto Neri  <https://orcid.org/0000-0002-7176-4046>
 Chunchong Ni  <https://orcid.org/0000-0003-1361-5699>
 Aristeidis Noutsos  <https://orcid.org/0000-0002-4151-3860>
 Michael A. Nowak  <https://orcid.org/0000-0001-6923-1315>
 Junghwan Oh  <https://orcid.org/0000-0002-4991-9638>
 Hiroki Okino  <https://orcid.org/0000-0003-3779-2016>
 Héctor Olivares  <https://orcid.org/0000-0001-6833-7580>
 Gisela N. Ortiz-León  <https://orcid.org/0000-0002-2863-676X>
 Tomoaki Oyama  <https://orcid.org/0000-0003-4046-2923>
 Feryal Özel  <https://orcid.org/0000-0003-4413-1523>
 Daniel C. M. Palumbo  <https://orcid.org/0000-0002-7179-3816>
 Georgios Filippas Paraschos  <https://orcid.org/0000-0001-6757-3098>
 Jongho Park  <https://orcid.org/0000-0001-6558-9053>
 Harriet Parsons  <https://orcid.org/0000-0002-6327-3423>
 Nimesh Patel  <https://orcid.org/0000-0002-6021-9421>
 Ue-Li Pen  <https://orcid.org/0000-0003-2155-9578>
 Dominic W. Pesce  <https://orcid.org/0000-0002-5278-9221>
 Richard Plambeck  <https://orcid.org/0000-0001-6765-9609>
 Oliver Porth  <https://orcid.org/0000-0002-4584-2557>
 Felix M. Pötzel  <https://orcid.org/0000-0002-6579-8311>
 Ben Prather  <https://orcid.org/0000-0002-0393-7734>
 Jorge A. Preciado-López  <https://orcid.org/0000-0002-4146-0113>
 Dimitrios Psaltis  <https://orcid.org/0000-0003-1035-3240>
 Hung-Yi Pu  <https://orcid.org/0000-0001-9270-8812>
 Venkatesh Ramakrishnan  <https://orcid.org/0000-0002-9248-086X>
 Ramprasad Rao  <https://orcid.org/0000-0002-1407-7944>
 Mark G. Rawlings  <https://orcid.org/0000-0002-6529-202X>
 Alexander W. Raymond  <https://orcid.org/0000-0002-5779-4767>
 Luciano Rezzolla  <https://orcid.org/0000-0002-1330-7103>
 Angelo Ricarte  <https://orcid.org/0000-0001-5287-0452>
 Bart Ripperda  <https://orcid.org/0000-0002-7301-3908>
 Freek Roelofs  <https://orcid.org/0000-0001-5461-3687>
 Alan Rogers  <https://orcid.org/0000-0003-1941-7458>
 Cristina Romero-Cañizales  <https://orcid.org/0000-0001-6301-9073>
 Eduardo Ros  <https://orcid.org/0000-0001-9503-4892>
 Arash Roshanineshat  <https://orcid.org/0000-0002-8280-9238>
 Alan L. Roy  <https://orcid.org/0000-0002-1931-0135>
 Ignacio Ruiz  <https://orcid.org/0000-0002-0965-5463>
 Chet Ruszczyk  <https://orcid.org/0000-0001-7278-9707>
 Kazi L. J. Rygl  <https://orcid.org/0000-0003-4146-9043>
 Salvador Sánchez  <https://orcid.org/0000-0002-8042-5951>
 David Sánchez-Argüelles  <https://orcid.org/0000-0002-7344-9920>
 Miguel Sánchez-Portal  <https://orcid.org/0000-0003-0981-9664>
 Mahito Sasada  <https://orcid.org/0000-0001-5946-9960>
 Kaushik Satapathy  <https://orcid.org/0000-0003-0433-3585>
 Tuomas Savolainen  <https://orcid.org/0000-0001-6214-1085>
 Jonathan Schonfeld  <https://orcid.org/0000-0002-8909-2401>
 Karl-Friedrich Schuster  <https://orcid.org/0000-0003-2890-9454>
 Lijing Shao  <https://orcid.org/0000-0002-1334-8853>
 Zhiqiang Shen (沈志强)  <https://orcid.org/0000-0003-3540-8746>
 Des Small  <https://orcid.org/0000-0003-3723-5404>
 Bong Won Sohn  <https://orcid.org/0000-0002-4148-8378>
 Jason SooHoo  <https://orcid.org/0000-0003-1938-0720>
 León David Sosapanta Salas  <https://orcid.org/0000-0003-1979-6363>
 Kamal Souccar  <https://orcid.org/0000-0001-7915-5272>
 He Sun (孙赫)  <https://orcid.org/0000-0003-1526-6787>
 Fumie Tazaki  <https://orcid.org/0000-0003-0236-0600>
 Alexandra J. Tetarenko  <https://orcid.org/0000-0003-3906-4354>
 Paul Tiede  <https://orcid.org/0000-0003-3826-5648>
 Remo P. J. Tilanus  <https://orcid.org/0000-0002-6514-553X>
 Michael Titus  <https://orcid.org/0000-0001-9001-3275>
 Pablo Torne  <https://orcid.org/0000-0001-8700-6058>
 Teresa Toscano  <https://orcid.org/0000-0003-3658-7862>
 Efthalia Traianou  <https://orcid.org/0000-0002-1209-6500>
 Tyler Trent  <https://orcid.org/0009-0004-8116-3123>
 Sascha Trippe  <https://orcid.org/0000-0003-0465-1559>
 Matthew Turk  <https://orcid.org/0000-0002-5294-0198>
 Ilse van Bemmelen  <https://orcid.org/0000-0001-5473-2950>
 Huib Jan van Langevelde  <https://orcid.org/0000-0002-0230-5946>
 Daniel R. van Rossum  <https://orcid.org/0000-0001-7772-6131>
 Jesse Vos  <https://orcid.org/0000-0003-3349-7394>
 Jan Wagner  <https://orcid.org/0000-0003-1105-6109>
 Derek Ward-Thompson  <https://orcid.org/0000-0003-1140-2761>
 John Wardle  <https://orcid.org/0000-0002-8960-2942>
 Jasmin E. Washington  <https://orcid.org/0000-0002-7046-0470>
 Jonathan Weintraub  <https://orcid.org/0000-0002-4603-5204>
 Robert Wharton  <https://orcid.org/0000-0002-7416-5209>
 Maciek Wielgus  <https://orcid.org/0000-0002-8635-4242>
 Kaj Wiik  <https://orcid.org/0000-0002-0862-3398>
 Gunther Witzel  <https://orcid.org/0000-0003-2618-797X>
 Michael F. Wondrak  <https://orcid.org/0000-0002-6894-1072>
 George N. Wong  <https://orcid.org/0000-0001-6952-2147>
 Qingwen Wu (吴庆文)  <https://orcid.org/0000-0003-4773-4987>
 Nitika Yadlapalli  <https://orcid.org/0000-0003-3255-4617>
 Paul Yamaguchi  <https://orcid.org/0000-0002-6017-8199>
 Aristomenis Yfantis  <https://orcid.org/0000-0002-3244-7072>
 Doosoo Yoon  <https://orcid.org/0000-0001-8694-8166>
 André Young  <https://orcid.org/0000-0003-0000-2682>
 Ken Young  <https://orcid.org/0000-0002-3666-4920>
 Ziri Younsi  <https://orcid.org/0000-0001-9283-1191>
 Wei Yu (于威)  <https://orcid.org/0000-0002-5168-6052>
 Feng Yuan (袁峰)  <https://orcid.org/0000-0003-3564-6437>
 Ye-Fei Yuan (袁业飞)  <https://orcid.org/0000-0002-7330-4756>
 J. Anton Zensus  <https://orcid.org/0000-0001-7470-3321>

Shuo Zhang  <https://orcid.org/0000-0002-2967-790X>
 Guang-Yao Zhao  <https://orcid.org/0000-0002-4417-1659>
 Shan-Shan Zhao (赵杉杉)  <https://orcid.org/0000-0002-9774-3606>

References

- Anantua, R., Emami, R., Loeb, A., & Chael, A. 2020, *ApJ*, **896**, 30
 Balbus, S. A., & Hawley, J. F. 1991, *ApJ*, **376**, 214
 Balbus, S. A., & Hawley, J. F. 1998, *RvMP*, **70**, 1
 Ball, D., Özel, F., Psaltis, D., & Chan, C.-K. 2016, *ApJ*, **826**, 77
 BisnovatyiKogan, G. S., & Ruzmaikin, A. A. 1974, *Ap&SS*, **28**, 45
 Blackburn, L., Chan, C.-k., Crew, G. B., et al. 2019, *ApJ*, **882**, 23
 Blackburn, L., Pesce, D. W., Johnson, M. D., et al. 2020, *ApJ*, **894**, 31
 Bower, G. C., Broderick, A., Dexter, J., et al. 2018, *ApJ*, **868**, 101
 Bower, G. C., Falcke, H., & Backer, D. C. 1999, *ApJL*, **523**, L29
 Bower, G. C., Falcke, H., & Mellon, R. R. 2002, *ApJL*, **578**, L103
 Broderick, A. E., Gold, R., Karami, M., et al. 2020a, *ApJ*, **897**, 139
 Broderick, A. E., & Pesce, D. W. 2020, *ApJ*, **904**, 126
 Broderick, A. E., Pesce, D. W., Tiede, P., Pu, H.-Y., & Gold, R. 2020b, *ApJ*, **898**, 9
 Chael, A. A., Johnson, M. D., Bouman, K. L., et al. 2018, *ApJ*, **857**, 23
 Chael, A. A., Johnson, M. D., Narayan, R., et al. 2016, *ApJ*, **829**, 11
 Chatterjee, K., Younsi, Z., Liska, M., et al. 2020, *MNRAS*, **499**, 362
 Contopoulos, I., & Kazanas, D. 1998, *ApJ*, **508**, 859
 Contopoulos, I., Myserlis, I., Kazanas, D., & Nathanael, A. 2022, *Galax*, **10**, 80
 Decker, R. B., & Krimigis, S. M. 2003, *AdSpR*, **32**, 597
 Dexter, J. 2016, *MNRAS*, **462**, 115
 Dhiingia, I. K., Mizuno, Y., Fromm, C. M., & Rezzolla, L. 2023, *MNRAS*, **518**, 405
 Emami, R., Anantua, R., Chael, A. A., & Loeb, A. 2021, *ApJ*, **923**, 272
 Emami, R., Anantua, R., Ricarte, A., et al. 2023, *Galax*, **11**, 11
 Emami, R., Ricarte, A., Wong, G. N., et al. 2023, *ApJ*, **950**, 38
 Event Horizon Telescope Collaboration 2023, EHT M87 Polarized Data, CyVerse Data Commons, doi:10.25739/q46m-m857
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2019a, *ApJL*, **875**, L1
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2019b, *ApJL*, **875**, L2
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2019c, *ApJL*, **875**, L3
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2019d, *ApJL*, **875**, L4
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2019e, *ApJL*, **875**, L5
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2019f, *ApJL*, **875**, L6
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2021a, *ApJL*, **910**, L12
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2021b, *ApJL*, **910**, L13
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2022a, *ApJL*, **930**, L12
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2022b, *ApJL*, **930**, L14
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2022c, *ApJL*, **930**, L15
 Event Horizon Telescope Collaboration, Akiyama, K., Alberdi, A., et al. 2022d, *ApJL*, **930**, L16
 Fishbone, L. G., & Moncrief, V. 1976, *ApJ*, **207**, 962
 Fromm, C. M., Cruz-Orsio, A., Mizuno, Y., et al. 2022, *A&A*, **660**, A107
 Gabuzda, D. C., Vitrichchak, V. M., Mahmud, M., & O'Sullivan, S. P. 2008, *MNRAS*, **384**, 1003
 Gammie, C. F., McKinney, J. C., & Tóth, G. 2003, *ApJ*, **589**, 444
 Goddi, C., Martí-Vidal, I., Messias, H., et al. 2019, *PASP*, **131**, 075003
 Goddi, C., Martí-Vidal, I., Messias, H., et al. 2021, *ApJL*, **910**, L14
 Greisen, E. W. 2003, in *Information Handling in Astronomy—Historical Vistas*, ed. A. Heck (Dordrecht: Kluwer), 109
 Hamaker, J. P., Bregman, J. D., & Sault, R. J. 1996, *A&AS*, **117**, 137
 Harris, C. R., Millman, K. J., van der Walt, S. J., et al. 2020, *Natur*, **585**, 357
 Högbom, J. A. 1974, *A&AS*, **15**, 417
 Homan, D. C., & Lister, M. 2006, *AJ*, **131**, 1262
 Homan, D. C., & Wardle, J. F. C. 1999, *AJ*, **118**, 1942
 Hunter, J. D. 2007, *CSE*, **9**, 90
 Igumenshchev, I. V., Narayan, R., & Abramowicz, M. A. 2003, *ApJ*, **592**, 1042
 Issaoun, S., Wielgus, M., Jorstad, S., et al. 2022, *ApJ*, **934**, 145
 Janssen, M., Falcke, H., Kadler, M., et al. 2021, *NatAs*, **5**, 1017
 Janssen, M., Goddi, C., van Bemmell, I. M., et al. 2019, *A&A*, **626**, A75
 Janssen, M., Radcliffe, J. F., & Wagner, J. 2022, *Univ*, **8**, 527
 Jiang, H.-X., Mizuno, Y., Fromm, C. M., & Nathanael, A. 2023, *MNRAS*, **522**, 2307
 Johnson, M. D., Lupsasca, A., Strominger, A., et al. 2020, *SciA*, **6**, eaaz1310
 Jones, T. W. 1988, *ApJ*, **332**, 678
 Jones, T. W., & O'Dell, S. L. 1977, *ApJ*, **214**, 522
 Jorstad, S., Wielgus, M., Lico, R., et al. 2023, *ApJ*, **943**, 170
 Kettenis, M., van Langevelde, H. J., Reynolds, C., & Cotton, B. 2006, in *ASP Conf. Ser. 351, Astronomical Data Analysis Software and Systems XV*, ed. C. Gabriel et al. (San Francisco, CA: ASP), 497
 Kim, J.-Y., Krichbaum, T. P., Broderick, A. E., et al. 2020, *A&A*, **640**, A69
 Kluyver, T., Ragan-Kelley, B., Pérez, F., et al. 2016, in *Positioning and Power in Academic Publishing: Players, Agents and Agendas*, ed. F. Loizides & B. Schmidt (Amsterdam: IOS Press), 87
 Kunz, M. W., Stone, J. M., & Quataert, E. 2016, *PhRvL*, **117**, 235101
 Leung, P. K., Gammie, C. F., & Noble, S. C. 2011, *ApJ*, **737**, 21
 MacDonald, N. R., & Marscher, A. P. 2018, *ApJ*, **862**, 58
 Mao, S. A., Dexter, J., & Quataert, E. 2017, *MNRAS*, **466**, 4307
 Marszewski, A., Prather, B. S., Joshi, A. V., Pandya, A., & Gammie, C. F. 2021, *ApJ*, **921**, 17
 Martí-Vidal, I., & Marcaide, J. M. 2008, *A&A*, **480**, 289
 Martí-Vidal, I., Mus, A., Janssen, M., de Vicente, P., & González, J. 2021, *A&A*, **646**, A52
 Martí-Vidal, I., Roy, A., Conway, J., & Zensus, A. J. 2016, *A&A*, **587**, A143
 Matthews, L. D., Crew, G. B., Doleman, S. S., et al. 2018, *PASP*, **130**, 015002
 McKinney, J. C., Tchekhovskoy, A., & Blandford, R. D. 2012, *MNRAS*, **423**, 3083
 McKinney, W. 2010, in *Proc. 9th Python in Science Conf.*, ed. S. van der Walt & J. Millman, 56
 Mizuno, Y., Fromm, C. M., Younsi, Z., et al. 2021, *MNRAS*, **506**, 741
 Mościbrodzka, M., Falcke, H., & Shiokawa, H. 2016, *A&A*, **586**, A38
 Mościbrodzka, M., & Gammie, C. F. 2018, *MNRAS*, **475**, 43
 Mościbrodzka, M., Janiuk, A., & De Laurentis, M. 2021, *MNRAS*, **508**, 4282
 Myserlis, I., Angelakis, E., Kraus, A., et al. 2018, *A&A*, **609**, A68
 Narayan, R., Chael, A., Chatterjee, K., Ricarte, A., & Curd, B. 2022, *MNRAS*, **511**, 3795
 Narayan, R., Igumenshchev, I. V., & Abramowicz, M. A. 2003, *PASP*, **55**, L69
 Narayan, R., Sądowski, A., Penna, R. F., & Kulkarni, A. K. 2012, *MNRAS*, **426**, 3241
 Noble, S. C., Gammie, C. F., McKinney, J. C., & Del Zanna, L. 2006, *ApJ*, **641**, 626
 Noble, S. C., Leung, P. K., Gammie, C. F., & Book, L. G. 2007, *CQGra*, **24**, S259
 Özel, F., Psaltis, D., & Narayan, R. 2000, *ApJ*, **541**, 234
 Palumbo, D. C. M., Wong, G. N., & Prather, B. S. 2020, *ApJ*, **894**, 156
 Pesce, D. W. 2021, *AJ*, **161**, 178
 Pierrard, V., & Lazar, M. 2010, *SoPh*, **267**, 153
 Pordes, R., Petravick, D., Kramer, B., et al. 2007, *JPhCS*, **78**, 012057
 Porth, O., Chatterjee, K., Narayan, R., et al. 2019, *ApJS*, **243**, 26
 Qiu, R., Ricarte, A., Narayan, R., et al. 2023, *MNRAS*, **520**, 4867
 Ricarte, A., Gammie, C., Narayan, R., & Prather, B. S. 2023, *MNRAS*, **519**, 4203
 Ricarte, A., Qiu, R., & Narayan, R. 2021, *MNRAS*, **505**, 523
 Sądowski, A., Narayan, R., Penna, R., & Zhu, Y. 2013, *MNRAS*, **436**, 3856
 Sfiligoi, I., Bradley, D. C., Holzman, B., et al. 2009, in *2009 WRI World Congress on Computer Science and Information Engineering*, 2 (Piscataway, NJ: IEEE), 428
 Shepherd, M., 2011 Difmap: Synthesis Imaging of Visibility Data, Astrophysics Source Code Library, ascl:1103.001
 Smirnov, O. M. 2011, *A&A*, **527**, A106
 Speagle, J. S. 2020, *MNRAS*, **493**, 3132
 Steel, S., Wielgus, M., Blackburn, L., Issaoun, S., & Johnson, M. 2019, EHT Memo Series 2019-CE-03
 Tange, O. 2011, login: The USENIX Magazine, 36, 42, <https://www.usenix.org/system/files/login/articles/105438-Tange.pdf>
 Tchekhovskoy, A., Narayan, R., & McKinney, J. C. 2011, *MNRAS*, **418**, L79
 The Astropy Collaboration, Price-Whelan, A. M., Sipőcz, B. M., et al. 2018, *AJ*, **156**, 123
 The Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, **558**, A33
 Thompson, A. R., Moran, J. M., & Swenson, G. W., Jr. 2017, *Interferometry and Synthesis in Radio Astronomy* (3rd ed.; Berlin: Springer)
 Thum, C., Agudo, I., Molina, S. N., et al. 2018, *MNRAS*, **473**, 2506

- Tsunetoe, Y., Mineshige, S., Kawashima, T., et al. 2022, *ApJ*, 931, 25
- Tsunetoe, Y., Mineshige, S., Ohsuga, K., Kawashima, T., & Akiyama, K. 2020, *PASJ*, 72, 32
- Tsunetoe, Y., Mineshige, S., Ohsuga, K., Kawashima, T., & Akiyama, K. 2021, *PASJ*, 73, 912
- Vasyliunas, V. M. 1968, *JGR*, 73, 2839
- Virtanen, P., Gommers, R., Oliphant, T. E., et al. 2020, *NatMe*, 17, 261
- Vitrichchak, V. M., Gabuzda, D. C., Algaba, J. C., et al. 2008, *MNRAS*, 391, 124
- Wardle, J. F. C., & Homan, D. C. 2003, *Ap&SS*, 288, 143
- Wardle, J. F. C., Homan, D. C., Ojha, R., & Roberts, D. H. 1998, *Natur*, 395, 457
- Wielgus, M., Akiyama, K., Blackburn, L., et al. 2020, *ApJ*, 901, 67
- Wielgus, M., Marchili, N., Martí-Vidal, I., et al. 2022a, *ApJL*, 930, L19
- Wielgus, M., Moscibrodzka, M., Vos, J., et al. 2022b, *A&A*, 665, L6
- Wong, G. N., & Gammie, C. F. 2022, *ApJ*, 937, 60
- Wong, G. N., Prather, B. S., Dhruv, V., et al. 2022, *ApJS*, 259, 64
- Yoon, D., Chatterjee, K., Markoff, S. B., et al. 2020, *MNRAS*, 499, 3178

The Event Horizon Telescope Collaboration,

- Kazunori Akiyama^{1,2,3} , Antxon Alberdi⁴ , Walter Alef⁵ , Juan Carlos Algaba⁶ , Richard Anantua^{3,7,8} , Keiichi Asada⁹ ,
 Rebecca Azulay^{10,11,5} , Uwe Bach⁵ , Anne-Kathrin Baczko^{12,5} , David Ball¹³ , Mislav Baloković¹⁴ , John Barrett¹ ,
 Michi Bauböck¹⁵ , Bradford A. Benson^{16,17} , Dan Bintley^{18,19} , Lindy Blackburn^{3,7} , Raymond Blundell⁷ ,
 Katherine L. Bouman²⁰ , Geoffrey C. Bower^{21,22} , Hope Boyce^{23,24} , Michael Bremer²⁵ , Christiaan D. Brinkerink²⁶ ,
 Roger Brissenden^{3,7} , Silke Britzen⁵ , Avery E. Broderick^{27,28,29} , Dominique Broguiere²⁵ , Thomas Bronzwaer²⁶ ,
 Sandra Bustamante³⁰ , Do-Young Byun^{31,32} , John E. Carlstrom^{33,17,34,35} , Chiara Ceccobello¹² , Andrew Chael³⁶ ,
 Chi-kwan Chan^{13,37,38} , Dominic O. Chang^{3,7} , Koushik Chatterjee^{3,7} , Shami Chatterjee³⁹ , Ming-Tang Chen²¹ ,
 Yongjun Chen (陈永军)^{40,41} , Xiaopeng Cheng³¹ , Ilje Cho⁴ , Pierre Christian⁴² , Nicholas S. Conroy^{43,7} ,
 John E. Conway¹² , James M. Cordes³⁹ , Thomas M. Crawford^{17,33} , Geoffrey B. Crew¹ , Alejandro Cruz-Orsorio^{44,45} ,
 Yuzhu Cui (崔玉竹)^{46,47} , Rohan Dahale⁴ , Jordy Davelaar^{48,49,26} , Mariafelicia De Laurentis^{50,45,51} ,
 Roger Deane^{52,53,54} , Jessica Dempsey^{18,19,55} , Gregory Desvignes^{5,56} , Jason Dexter⁵⁷ , Vedant Dhruv¹⁵ ,
 Sheperd S. Doeleman^{3,7} , Sean Dougal¹³ , Sergio A. Dzib^{25,5} , Ralph P. Eatough^{58,5} , Razieh Emami⁷ , Heino Falcke²⁶ ,
 Joseph Farah^{59,60} , Vincent L. Fish¹ , Ed Fomalont⁶¹ , H. Alyson Ford¹³ , Marianna Foschi⁴ , Raquel Fraga-Encinas²⁶ ,
 William T. Freeman^{62,63} , Per Friberg^{18,19} , Christian M. Fromm^{64,45,5} , Antonio Fuentes⁴ , Peter Galison^{3,65,66} ,
 Charles F. Gammie^{15,43,67} , Roberto García²⁵ , Olivier Gentaz²⁵ , Boris Georgiev^{28,29,27} , Ciriaco Goddi^{68,69,70,71} ,
 Roman Gold⁷² , Arturo I. Gómez-Ruiz^{73,74} , José L. Gómez⁴ , Minfeng Gu (顾敏峰)^{40,75} , Mark Gurwell⁷ ,
 Kazuhiro Hada^{76,77} , Daryl Haggard^{23,24} , Kari Haworth⁷ , Michael H. Hecht¹ , Ronald Hesper⁷⁸ , Dirk Heumann¹³ ,
 Luis C. Ho (何子山)^{79,80} , Paul Ho^{9,19,18} , Mareki Honma^{76,77,81} , Chih-Wei L. Huang⁹ , Lei Huang (黄磊)^{40,75} ,
 David H. Hughes⁷³ , Shiro Ikeda^{2,82,83,84} , C. M. Violette Impellizzeri^{85,61} , Makoto Inoue⁹ , Sara Issaoun^{7,153} ,
 David J. James^{86,87} , Buell T. Jannuzi¹³ , Michael Janssen⁵ , Britton Jeter⁹ , Wu Jiang (江悟)⁴⁰ ,
 Alejandra Jiménez-Rosales²⁶ , Michael D. Johnson^{3,7} , Svetlana Jorstad⁸⁸ , Abhishek V. Joshi¹⁵ , Taehyun Jung^{31,32} ,
 Mansour Karami^{27,28} , Ramesh Karuppusamy⁵ , Tomohisa Kawashima⁸⁹ , Garrett K. Keating⁷ , Mark Kettenis⁹⁰ ,
 Dong-Jin Kim⁵ , Jae-Young Kim^{91,5} , Jongsoo Kim³¹ , Junhan Kim²⁰ , Motoki Kino^{2,92} , Jun Yi Koay⁹ ,
 Prashant Kocherlakota⁴⁵ , Yutaro Kofuji^{76,81} , Patrick M. Koch⁹ , Shoko Koyama^{9,93} , Carsten Kramer²⁵ ,
 Joana A. Kramer⁵ , Michael Kramer⁵ , Thomas P. Krichbaum⁵ , Cheng-Yu Kuo^{94,9} , Noemi La Bella²⁶ , Tod R. Lauer⁹⁵ ,
 Daeyoung Lee¹⁵ , Sang-Sung Lee³¹ , Po Kin Leung⁹⁶ , Aviad Levis²⁰ , Zhiyuan Li (李志远)^{97,98} , Rocco Lico^{99,4} ,
 Greg Lindahl⁷ , Michael Lindqvist¹² , Mikhail Lisakov⁵ , Jun Liu (刘俊)⁵ , Kuo Liu⁵ , Elisabetta Liuzzo¹⁰⁰ ,
 Wen-Ping Lo^{9,101} , Andrei P. Lobanov⁵ , Laurent Loinard^{102,44} , Colin J. Lonsdale¹ , Amy E. Lowitz¹³ , Ru-Sen Lu (路如森)^{40,41,5} ,
 Nicholas R. MacDonald⁵ , Jirong Mao (毛基荣)^{103,104,105} , Nicola Marchili^{100,5} , Sera Markoff^{106,107} ,
 Daniel P. Marrone¹³ , Alan P. Marscher⁸⁸ , Iván Martí-Vidal^{10,11} , Satoki Matsushita⁹ , Lynn D. Matthews¹ ,
 Lia Medeiros^{108,153} , Karl M. Menten⁵ , Daniel Michalik^{109,17} , Izumi Mizuno^{18,19} , Yosuke Mizuno^{47,111,45} ,
 James M. Moran^{3,7} , Kotaro Moriyama^{45,1,76} , Monika Moscibrodzka²⁶ , Wanga Muladzi¹⁰⁶ , Cornelia Müller^{5,26} ,
 Hendrik Müller⁵ , Alejandro Mus^{10,11} , Gibwa Musoke^{106,26} , Ioannis Myserlis¹¹⁰ , Andrew Nadolski⁴³ ,
 Hiroshi Nagai^{2,77} , Neil M. Nagar¹¹² , Masanori Nakamura^{113,9} , Ramesh Narayan^{3,7} , Gopal Narayanan³⁰ ,
 Iniyan Natarajan^{7,3} , Antonios Nathanail^{114,45} , Santiago Navarro Fuentes¹¹¹ , Joey Neilsen¹¹⁵ , Roberto Neri²⁵ ,
 Chunchong Ni^{28,29,27} , Aristeidis Noutsos⁵ , Michael A. Nowak¹¹⁶ , Junghwan Oh¹¹⁷ , Hiroki Okino^{76,81} ,
 Héctor Olivares²⁶ , Gisela N. Ortiz-León^{73,5} , Tomoaki Oyama⁷⁶ , Feryal Özel¹¹⁸ , Daniel C. M. Palumbo^{3,7} ,
 Georgios Filippou Paraschos⁵ , Jongho Park¹¹⁹ , Harriet Parsons^{18,19} , Nimesh Patel⁷ , Ue-Li Pen^{9,27,120,121,122} ,
 Dominic W. Pesce^{7,3} , Vincent Piétu²⁵ , Richard Plambeck¹²³ , Aleksandar PopStefanija³⁰ , Oliver Porth^{106,45} ,
 Felix M. Pötzel^{124,5} , Ben Prather¹⁵ , Jorge A. Preciado-López²⁷ , Dimitrios Psaltis¹¹⁸ , Hung-Yi Pu^{125,126,9} ,
 Venkatesh Ramakrishnan^{112,127,128} , Ramprasad Rao⁷ , Mark G. Rawlings^{129,18,19} , Alexander W. Raymond^{3,7} ,
 Luciano Rezzolla^{45,130,131} , Angelo Ricarte^{7,3} , Bart Ripperda^{120,121,132,27} , Freek Roelofs^{7,3,26} , Alan Rogers¹ ,
 Cristina Romero-Cañizales⁹ , Eduardo Ros⁵ , Arash Roshanineshat¹³ , Helge Rottmann⁵ , Alan L. Roy⁵ , Ignacio Ruiz¹¹¹ ,
 Chet Ruszczyk¹ , Kazi L. J. Rygl¹⁰⁰ , Salvador Sánchez¹¹¹ , David Sánchez-Argüelles^{73,74} , Miguel Sánchez-Portal¹¹¹ 

Mahito Sasada^{133,76,134} , Kaushik Satapathy¹³ , Tuomas Savolainen^{135,128,5} , F. Peter Schloerb³⁰, Jonathan Schonfeld⁷ , Karl-Friedrich Schuster²⁵ , Lijing Shao^{80,5} , Zhiqiang Shen (沈志强)^{40,41} , Des Small⁹⁰ , Bong Won Sohn^{31,32,136} , Jason SooHoo¹ , León David Sosapanta Salas¹⁰⁶ , Kamal Souccar³⁰ , He Sun (孙赫)^{137,138} , Fumie Tazaki⁷⁶ , Alexandra J. Tetarenko¹³⁹ , Paul Tiede^{7,3} , Remo P. J. Tilanus^{13,26,85,140} , Michael Titus¹ , Pablo Torne^{111,5} , Teresa Toscano⁴ , Efthalia Traianou^{4,5} , Tyler Trent¹³ , Sascha Trippe¹⁴¹ , Matthew Turk⁴³ , Ilse van Bemmelen⁹⁰ , Huib Jan van Langevelde^{90,85,142} , Daniel R. van Rossum²⁶ , Jesse Vos²⁶ , Jan Wagner⁵ , Derek Ward-Thompson¹⁴³ , John Wardle¹⁴⁴ , Jasmin E. Washington¹³ , Jonathan Weintraub^{3,7} , Robert Wharton⁵ , Maciek Wielgus⁵ , Kaj Wiik¹⁴⁵ , Gunther Witzel⁵ , Michael F. Wondrak^{26,146} , George N. Wong^{147,36} , Qingwen Wu (吴庆文)¹⁴⁸ , Nitika Yadlapalli²⁰ , Paul Yamaguchi⁷ , Aristomenis Yfantis²⁶ , Doosoo Yoon¹⁰⁶ , André Young²⁶ , Ken Young⁷ , Ziri Younsi^{149,45} , Wei Yu (于威)⁷ , Feng Yuan (袁峰)^{40,75,150} , Ye-Fei Yuan (袁业飞)¹⁵¹ , J. Anton Zensus⁵ , Shuo Zhang¹⁵² , Guang-Yao Zhao⁴ , and Shan-Shan Zhao (赵杉杉)⁴⁰ 

¹ Massachusetts Institute of Technology Haystack Observatory, 99 Millstone Road, Westford, MA 01886, USA

² National Astronomical Observatory of Japan, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan

³ Black Hole Initiative at Harvard University, 20 Garden Street, Cambridge, MA 02138, USA

⁴ Instituto de Astrofísica de Andalucía-CSIC, Glorieta de la Astronomía s/n, E-18008 Granada, Spain

⁵ Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, D-53121 Bonn, Germany

⁶ Department of Physics, Faculty of Science, Universiti Malaya, 50603 Kuala Lumpur, Malaysia

⁷ Center for Astrophysics | Harvard & Smithsonian, 60 Garden Street, Cambridge, MA 02138, USA

⁸ Department of Physics & Astronomy, The University of Texas at San Antonio, One UTSA Circle, San Antonio, TX 78249, USA

⁹ Institute of Astronomy and Astrophysics, Academia Sinica, 11F of Astronomy-Mathematics Building, AS/NTU No. 1, Sec. 4, Roosevelt Rd., Taipei 10617, Taiwan, R.O.C.

¹⁰ Departament d'Astronomia i Astrofísica, Universitat de València, C. Dr. Moliner 50, E-46100 Burjassot, València, Spain

¹¹ Observatori Astronòmic, Universitat de València, C. Catedrático José Beltrán 2, E-46980 Paterna, València, Spain

¹² Department of Space, Earth and Environment, Chalmers University of Technology, Onsala Space Observatory, SE-43992 Onsala, Sweden

¹³ Steward Observatory and Department of Astronomy, University of Arizona, 933 N. Cherry Ave., Tucson, AZ 85721, USA

¹⁴ Yale Center for Astronomy & Astrophysics, Yale University, 52 Hillhouse Avenue, New Haven, CT 06511, USA

¹⁵ Department of Physics, University of Illinois, 1110 West Green Street, Urbana, IL 61801, USA

¹⁶ Fermi National Accelerator Laboratory, MS209, P.O. Box 500, Batavia, IL 60510, USA

¹⁷ Department of Astronomy and Astrophysics, University of Chicago, 5640 South Ellis Avenue, Chicago, IL 60637, USA

¹⁸ East Asian Observatory, 660 N. A'ohoku Place, Hilo, HI 96720, USA

¹⁹ James Clerk Maxwell Telescope (JCMT), 660 N. A'ohoku Place, Hilo, HI 96720, USA

²⁰ California Institute of Technology, 1200 East California Boulevard, Pasadena, CA 91125, USA

²¹ Institute of Astronomy and Astrophysics, Academia Sinica, 645 N. A'ohoku Place, Hilo, HI 96720, USA

²² Department of Physics and Astronomy, University of Hawaii at Manoa, 2505 Correa Road, Honolulu, HI 96822, USA

²³ Department of Physics, McGill University, 3600 rue University, Montréal, QC H3A 2T8, Canada

²⁴ Trotter Space Institute at McGill, 3550 rue University, Montréal, QC H3A 2A7, Canada

²⁵ Institut de Radioastronomie Millimétrique (IRAM), 300 rue de la Piscine, F-38406 Saint Martin d'Hères, France

²⁶ Department of Astrophysics, Institute for Mathematics, Astrophysics and Particle Physics (IMAPP), Radboud University, P.O. Box 9010, 6500 GL Nijmegen, The Netherlands

²⁷ Perimeter Institute for Theoretical Physics, 31 Caroline Street North, Waterloo, ON N2L 2Y5, Canada

²⁸ Department of Physics and Astronomy, University of Waterloo, 200 University Avenue West, Waterloo, ON N2L 3G1, Canada

²⁹ Waterloo Centre for Astrophysics, University of Waterloo, Waterloo, ON N2L 3G1, Canada

³⁰ Department of Astronomy, University of Massachusetts, Amherst, MA 01003, USA

³¹ Korea Astronomy and Space Science Institute, Daedeok-daero 776, Yuseong-gu, Daejeon 34055, Republic of Korea

³² University of Science and Technology, Gajeong-ro 217, Yuseong-gu, Daejeon 34113, Republic of Korea

³³ Kavli Institute for Cosmological Physics, University of Chicago, 5640 South Ellis Avenue, Chicago, IL 60637, USA

³⁴ Department of Physics, University of Chicago, 5720 South Ellis Avenue, Chicago, IL 60637, USA

³⁵ Enrico Fermi Institute, University of Chicago, 5640 South Ellis Avenue, Chicago, IL 60637, USA

³⁶ Princeton Gravity Initiative, Jadwin Hall, Princeton University, Princeton, NJ 08544, USA

³⁷ Data Science Institute, University of Arizona, 1230 N. Cherry Ave., Tucson, AZ 85721, USA

³⁸ Program in Applied Mathematics, University of Arizona, 617 N. Santa Rita, Tucson, AZ 85721, USA

³⁹ Cornell Center for Astrophysics and Planetary Science, Cornell University, Ithaca, NY 14853, USA

⁴⁰ Shanghai Astronomical Observatory, Chinese Academy of Sciences, 80 Nandan Road, Shanghai 200030, People's Republic of China

⁴¹ Key Laboratory of Radio Astronomy, Chinese Academy of Sciences, Nanjing 210008, People's Republic of China

⁴² Physics Department, Fairfield University, 1073 North Benson Road, Fairfield, CT 06824, USA

⁴³ Department of Astronomy, University of Illinois at Urbana-Champaign, 1002 West Green Street, Urbana, IL 61801, USA

⁴⁴ Instituto de Astronomía, Universidad Nacional Autónoma de México (UNAM), Apdo Postal 70-264, Ciudad de México, México

⁴⁵ Institut für Theoretische Physik, Goethe-Universität Frankfurt, Max-von-Laue-Straße 1, D-60438 Frankfurt am Main, Germany

⁴⁶ Research Center for Intelligent Computing Platforms, Zhejiang Laboratory, Hangzhou 311100, People's Republic of China

⁴⁷ Tsung-Dao Lee Institute, Shanghai Jiao Tong University, Shengrong Road 520, Shanghai, 201210, People's Republic of China

⁴⁸ Department of Astronomy and Columbia Astrophysics Laboratory, Columbia University, 500 W. 120th Street, New York, NY 10027, USA

⁴⁹ Center for Computational Astrophysics, Flatiron Institute, 162 Fifth Avenue, New York, NY 10010, USA

⁵⁰ Dipartimento di Fisica "E. Pancini," Università di Napoli "Federico II," Compl. Univ. di Monte S. Angelo, Edificio G, Via Cinthia, I-80126, Napoli, Italy

⁵¹ INFN Sez. di Napoli, Compl. Univ. di Monte S. Angelo, Edificio G, Via Cinthia, I-80126, Napoli, Italy

⁵² Wits Centre for Astrophysics, University of the Witwatersrand, 1 Jan Smuts Avenue, Braamfontein, Johannesburg 2050, South Africa

⁵³ Department of Physics, University of Pretoria, Hatfield, Pretoria 0028, South Africa

⁵⁴ Centre for Radio Astronomy Techniques and Technologies, Department of Physics and Electronics, Rhodes University, Makhanda 6140, South Africa

⁵⁵ ASTRON, Oude Hoogeveensedijk 4, 7991 PD Dwingeloo, The Netherlands

⁵⁶ LESIA, Observatoire de Paris, Université PSL, CNRS, Sorbonne Université, Université de Paris, 5 place Jules Janssen, F-92195 Meudon, France

⁵⁷ JILA and Department of Astrophysical and Planetary Sciences, University of Colorado, Boulder, CO 80309, USA

⁵⁸ National Astronomical Observatories, Chinese Academy of Sciences, 20A Datun Road, Chaoyang District, Beijing 100101, People's Republic of China

- ⁵⁹ Las Cumbres Observatory, 6740 Cortona Drive, Suite 102, Goleta, CA 93117-5575, USA
- ⁶⁰ Department of Physics, University of California, Santa Barbara, CA 93106-9530, USA
- ⁶¹ National Radio Astronomy Observatory, 520 Edgemont Road, Charlottesville, VA 22903, USA
- ⁶² Department of Electrical Engineering and Computer Science, Massachusetts Institute of Technology, 32-D476, 77 Massachusetts Ave., Cambridge, MA 02142, USA
- ⁶³ Google Research, 355 Main St., Cambridge, MA 02142, USA
- ⁶⁴ Institut für Theoretische Physik und Astrophysik, Universität Würzburg, Emil-Fischer-Str. 31, D-97074 Würzburg, Germany
- ⁶⁵ Department of History of Science, Harvard University, Cambridge, MA 02138, USA
- ⁶⁶ Department of Physics, Harvard University, Cambridge, MA 02138, USA
- ⁶⁷ NCSA, University of Illinois, 1205 W. Clark St., Urbana, IL 61801, USA
- ⁶⁸ Instituto de Astronomia, Geofísica e Ciências Atmosféricas, Universidade de São Paulo, R. do Matão, 1226, São Paulo, SP 05508-090, Brazil
- ⁶⁹ Dipartimento di Fisica, Università degli Studi di Cagliari, SP Monserrato-Sestu km 0.7, I-09042 Monserrato (CA), Italy
- ⁷⁰ INAF—Osservatorio Astronomico di Cagliari, via della Scienza 5, I-09047 Selargius (CA), Italy
- ⁷¹ INFN, sezione di Cagliari, I-09042 Monserrato (CA), Italy
- ⁷² CP3-Origins, University of Southern Denmark, Campusvej 55, DK-5230 Odense M, Denmark
- ⁷³ Instituto Nacional de Astrofísica, Óptica y Electrónica. Apartado Postal 51 y 216, 72000. Puebla Pue., México
- ⁷⁴ Consejo Nacional de Ciencia y Tecnología, Av. Insurgentes Sur 1582, 03940, Ciudad de México, México
- ⁷⁵ Key Laboratory for Research in Galaxies and Cosmology, Chinese Academy of Sciences, Shanghai 200030, People's Republic of China
- ⁷⁶ Mizusawa VLBI Observatory, National Astronomical Observatory of Japan, 2-12 Hoshigaoka, Mizusawa, Oshu, Iwate 023-0861, Japan
- ⁷⁷ Department of Astronomical Science, The Graduate University for Advanced Studies (SOKENDAI), 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan
- ⁷⁸ NOVA Sub-mm Instrumentation Group, Kapteyn Astronomical Institute, University of Groningen, Landleven 12, 9747 AD Groningen, The Netherlands
- ⁷⁹ Department of Astronomy, School of Physics, Peking University, Beijing 100871, People's Republic of China
- ⁸⁰ Kavli Institute for Astronomy and Astrophysics, Peking University, Beijing 100871, People's Republic of China
- ⁸¹ Department of Astronomy, Graduate School of Science, The University of Tokyo, 7-3-1 Hongo, Bunkyo-ku, Tokyo 113-0033, Japan
- ⁸² The Institute of Statistical Mathematics, 10-3 Midori-cho, Tachikawa, Tokyo, 190-8562, Japan
- ⁸³ Department of Statistical Science, The Graduate University for Advanced Studies (SOKENDAI), 10-3 Midori-cho, Tachikawa, Tokyo 190-8562, Japan
- ⁸⁴ Kavli Institute for the Physics and Mathematics of the Universe, The University of Tokyo, 5-1-5 Kashiwanoha, Kashiwa, 277-8583, Japan
- ⁸⁵ Leiden Observatory, Leiden University, Postbus 2300, 9513 RA Leiden, The Netherlands
- ⁸⁶ ASTRAVEO LLC, P.O. Box 1668, Gloucester, MA 01931, USA
- ⁸⁷ Applied Materials Inc., 35 Dory Road, Gloucester, MA 01930, USA
- ⁸⁸ Institute for Astrophysical Research, Boston University, 725 Commonwealth Ave., Boston, MA 02215, USA
- ⁸⁹ Institute for Cosmic Ray Research, The University of Tokyo, 5-1-5 Kashiwanoha, Kashiwa, Chiba 277-8582, Japan
- ⁹⁰ Joint Institute for VLBI ERIC (JIVE), Oude Hoogeveensedijk 4, 7991 PD Dwingeloo, The Netherlands
- ⁹¹ Department of Astronomy and Atmospheric Sciences, Kyungpook National University, Daegu 702-701, Republic of Korea
- ⁹² Kogakuin University of Technology & Engineering, Academic Support Center, 2665-1 Nakano, Hachioji, Tokyo 192-0015, Japan
- ⁹³ Graduate School of Science and Technology, Niigata University, 8050 Ikarashi 2-no-cho, Nishi-ku, Niigata 950-2181, Japan
- ⁹⁴ Physics Department, National Sun Yat-Sen University, No. 70, Lien-Hai Road, Kaosiung City 80424, Taiwan, R.O.C.
- ⁹⁵ National Optical Astronomy Observatory, 950 N. Cherry Ave., Tucson, AZ 85719, USA
- ⁹⁶ Department of Physics, The Chinese University of Hong Kong, Shatin, N.T., Hong Kong
- ⁹⁷ School of Astronomy and Space Science, Nanjing University, Nanjing 210023, People's Republic of China
- ⁹⁸ Key Laboratory of Modern Astronomy and Astrophysics, Nanjing University, Nanjing 210023, People's Republic of China
- ⁹⁹ INAF-Istituto di Radioastronomia, Via P. Gobetti 101, I-40129 Bologna, Italy
- ¹⁰⁰ INAF-Istituto di Radioastronomia & Italian ALMA Regional Centre, Via P. Gobetti 101, I-40129 Bologna, Italy
- ¹⁰¹ Department of Physics, National Taiwan University, No. 1, Sec. 4, Roosevelt Rd., Taipei 10617, Taiwan, R.O.C
- ¹⁰² Instituto de Radioastronomía y Astrofísica, Universidad Nacional Autónoma de México, Morelia 58089, México
- ¹⁰³ Yunnan Observatories, Chinese Academy of Sciences, 650011 Kunming, Yunnan Province, People's Republic of China
- ¹⁰⁴ Center for Astronomical Mega-Science, Chinese Academy of Sciences, 20A Datun Road, Chaoyang District, Beijing, 100012, People's Republic of China
- ¹⁰⁵ Key Laboratory for the Structure and Evolution of Celestial Objects, Chinese Academy of Sciences, 650011 Kunming, People's Republic of China
- ¹⁰⁶ Anton Pannekoek Institute for Astronomy, University of Amsterdam, Science Park 904, 1098 XH, Amsterdam, The Netherlands
- ¹⁰⁷ Gravitation and Astroparticle Physics Amsterdam (GRAPPA) Institute, University of Amsterdam, Science Park 904, 1098 XH Amsterdam, The Netherlands
- ¹⁰⁸ Department of Astrophysical Sciences, Peyton Hall, Princeton University, Princeton, NJ 08544, USA
- ¹⁰⁹ Science Support Office, Directorate of Science, European Space Research and Technology Centre (ESA/ESTEC), Keplerlaan 1, 2201 AZ Noordwijk, The Netherlands
- ¹¹⁰ School of Physics and Astronomy, Shanghai Jiao Tong University, 800 Dongchuan Road, Shanghai, 200240, People's Republic of China
- ¹¹¹ Institut de Radioastronomie Millimétrique (IRAM), Avenida Divina Pastora 7, Local 20, E-18012, Granada, Spain
- ¹¹² Astronomy Department, Universidad de Concepción, Casilla 160-C, Concepción, Chile
- ¹¹³ National Institute of Technology, Hachinohe College, 16-1 Uwanotai, Tamonoki, Hachinohe City, Aomori 039-1192, Japan
- ¹¹⁴ Research Center for Astronomy, Academy of Athens, Soranou Efessiou 4, 115 27 Athens, Greece
- ¹¹⁵ Department of Physics, Villanova University, 800 Lancaster Avenue, Villanova, PA 19085, USA
- ¹¹⁶ Physics Department, Washington University, CB 1105, St. Louis, MO 63130, USA
- ¹¹⁷ Sejong University, 209 Neungdong-ro, Gwangjin-gu, Seoul, Republic of Korea
- ¹¹⁸ School of Physics, Georgia Institute of Technology, 837 State St. NW, Atlanta, GA 30332, USA
- ¹¹⁹ Department of Astronomy and Space Science, Kyung Hee University, 1732, Deogyong-daero, Giheung-gu, Yongin-si, Gyeonggi-do 17104, Republic of Korea
- ¹²⁰ Canadian Institute for Theoretical Astrophysics, University of Toronto, 60 St. George Street, Toronto, ON M5S 3H8, Canada
- ¹²¹ Dunlap Institute for Astronomy and Astrophysics, University of Toronto, 50 St. George Street, Toronto, ON M5S 3H4, Canada
- ¹²² Canadian Institute for Advanced Research, 180 Dundas St. W, Toronto, ON M5G 1Z8, Canada
- ¹²³ Radio Astronomy Laboratory, University of California, Berkeley, CA 94720, USA
- ¹²⁴ Institute of Astrophysics, Foundation for Research and Technology—Hellas, Voutes, 7110 Heraklion, Greece
- ¹²⁵ Department of Physics, National Taiwan Normal University, No. 88, Sec. 4, Tingzhou Rd., Taipei 116, Taiwan, R.O.C.
- ¹²⁶ Center of Astronomy and Gravitation, National Taiwan Normal University, No. 88, Sec. 4, Tingzhou Road, Taipei 116, Taiwan, R.O.C.
- ¹²⁷ Finnish Centre for Astronomy with ESO, FI-20014 University of Turku, Finland
- ¹²⁸ Aalto University Metsähovi Radio Observatory, Metsähovintie 114, FI-02540 Kylmälahti, Finland
- ¹²⁹ Gemini Observatory/NSF NOIRLab, 670 N. A'ohōkū Place, Hilo, HI 96720, USA
- ¹³⁰ Frankfurt Institute for Advanced Studies, Ruth-Moufang-Strasse 1, D-60438 Frankfurt, Germany
- ¹³¹ School of Mathematics, Trinity College, Dublin 2, Ireland

- ¹³² Department of Physics, University of Toronto, 60 St. George Street, Toronto, ON M5S 1A7, Canada
- ¹³³ Department of Physics, Tokyo Institute of Technology, 2-12-1 Ookayama, Meguro-ku, Tokyo 152-8551, Japan
- ¹³⁴ Hiroshima Astrophysical Science Center, Hiroshima University, 1-3-1 Kagamiyama, Higashi-Hiroshima, Hiroshima 739-8526, Japan
- ¹³⁵ Aalto University Department of Electronics and Nanoengineering, PL 15500, FI-00076 Aalto, Finland
- ¹³⁶ Department of Astronomy, Yonsei University, Yonsei-ro 50, Seodaemun-gu, 03722 Seoul, Republic of Korea
- ¹³⁷ National Biomedical Imaging Center, Peking University, Beijing 100871, People's Republic of China
- ¹³⁸ College of Future Technology, Peking University, Beijing 100871, People's Republic of China
- ¹³⁹ Department of Physics and Astronomy, University of Lethbridge, Lethbridge, AB T1K 3M4, Canada
- ¹⁴⁰ Netherlands Organisation for Scientific Research (NWO), Postbus 93138, 2509 AC Den Haag, The Netherlands
- ¹⁴¹ Department of Physics and Astronomy, Seoul National University, Gwanak-gu, Seoul 08826, Republic of Korea
- ¹⁴² University of New Mexico, Department of Physics and Astronomy, Albuquerque, NM 87131, USA
- ¹⁴³ Jeremiah Horrocks Institute, University of Central Lancashire, Preston PR1 2HE, UK
- ¹⁴⁴ Physics Department, Brandeis University, 415 South Street, Waltham, MA 02453, USA
- ¹⁴⁵ Tuorla Observatory, Department of Physics and Astronomy, University of Turku, Finland
- ¹⁴⁶ Radboud Excellence Fellow of Radboud University, Nijmegen, The Netherlands
- ¹⁴⁷ School of Natural Sciences, Institute for Advanced Study, 1 Einstein Drive, Princeton, NJ 08540, USA
- ¹⁴⁸ School of Physics, Huazhong University of Science and Technology, Wuhan, Hubei, 430074, People's Republic of China
- ¹⁴⁹ Mullard Space Science Laboratory, University College London, Holmbury St. Mary, Dorking, Surrey, RH5 6NT, UK
- ¹⁵⁰ School of Astronomy and Space Sciences, University of Chinese Academy of Sciences, No. 19A Yuquan Road, Beijing 100049, People's Republic of China
- ¹⁵¹ Astronomy Department, University of Science and Technology of China, Hefei 230026, People's Republic of China
- ¹⁵² Department of Physics and Astronomy, Michigan State University, 567 Wilson Rd., East Lansing, MI 48824, USA