



# Measurements of the Crab Pulsar’s Giant Radio Pulse Amplitude Power-law Index Using Low-frequency Arecibo and Green Bank Telescope Observations

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## Abstract

We report two low-frequency measurements of the power-law index for the amplitudes of giant radio pulses from the Crab pulsar. The two observations were taken with the Arecibo and Green Bank radio telescopes at center frequencies of 327 MHz and 350 MHz, respectively. We find best-fit values for the differential power-law index  $\beta$  (where  $dN/dS \propto S^\beta$  and  $S$  is the pulse amplitude) of  $-2.63 \pm 0.05$  and  $-3.6 \pm 0.5$  from the Arecibo and Green Bank data sets, respectively. Both values are broadly consistent with other values previously measured for the Crab pulsar at low radio frequencies. These reported values may be useful in future giant pulse studies of the Crab pulsar.

*Unified Astronomy Thesaurus concepts:* Pulsars (1306); Radio transient sources (2008)

## 1. Introduction

Since the discovery of the Crab pulsar as a source of individual dispersed radio pulses (Staelin & Reifenstein 1968), it has been studied as an emitter of giant pulses (see, e.g., Lewandowska 2015 for a review). Giant radio pulses from pulsars have been defined as pulses having energies much greater than the mean value and having amplitudes that follow a power-law distribution (Johnston & Romani 2004; Geyer et al. 2021). A departure from this behavior has been observed for the Crab pulsar by Cordes et al. (2004) and Mickaliger et al. (2012), where they saw a slight excess at very large amplitudes that might be explained by rare “supergiant” pulses. However, such pulses would not necessarily be expected to repeat if, for example, they were due to lensing phenomena, which is possible given the role that filaments play in producing multiple images. Subsequent observations by Bera & Chengalur (2019) did not show evidence for supergiant pulses in a longer set of observations, so the observed excess may be a statistical fluke (as mentioned by Cordes et al. 2004). Here, we present two new measurements of the differential amplitude power-law index  $\beta$  for Crab pulsar giant pulses, where  $\beta$  is defined according to  $dN/dS \propto S^\beta$  and where  $S$  is the pulse amplitude. The measurements were obtained from two low-frequency observations taken with the Arecibo and the Green Bank telescopes.

## 2. Observations and Analysis

The Crab pulsar was observed with the Arecibo 305 m telescope in a 5 minute diagnostic observation on 2014 April 13 (MJD 56,760) as part of the Arecibo 327 MHz Drift-Scan

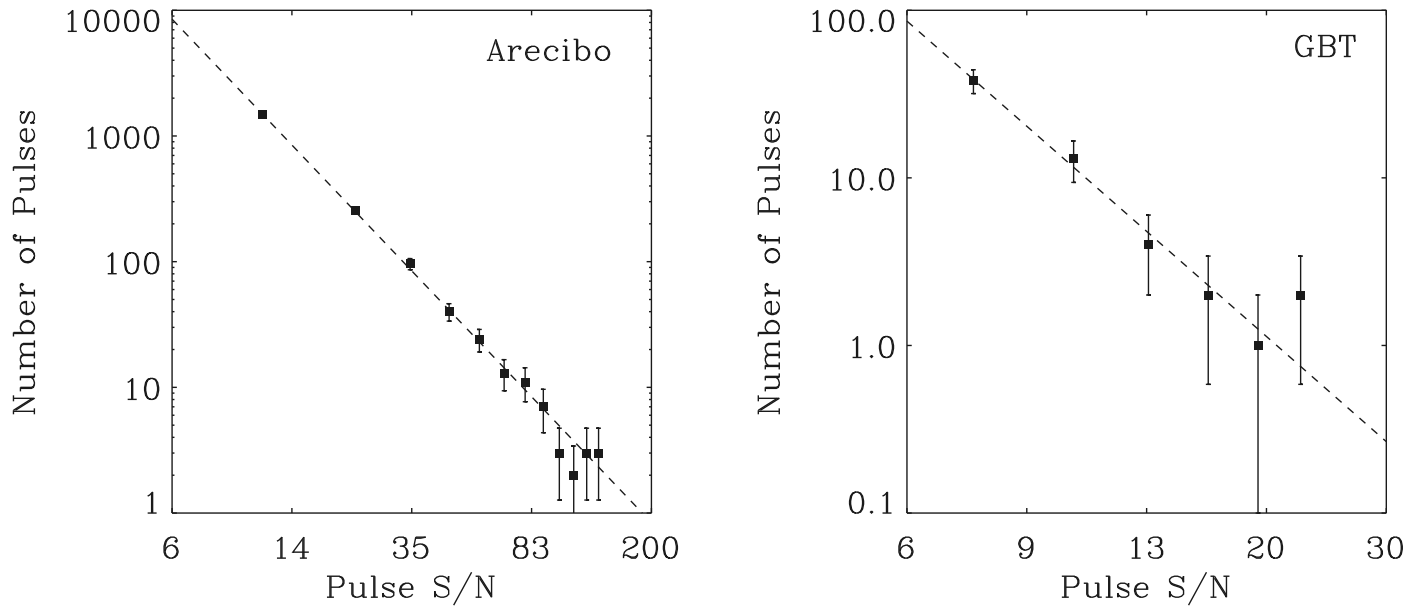
Pulsar Survey (Deneva et al. 2013). This observation used an effective bandwidth of 68.75 MHz divided into 2816 channels sampled at 81.92  $\mu$ s. The Crab pulsar was also observed with the Green Bank Telescope (GBT) on 2019 October 22 (MJD 58,778) as part of the Green Bank Northern Celestial Cap (GBNCC) survey (Stovall et al. 2014). This survey used a bandwidth of 100 MHz centered on 350 MHz that was split into 4096 channels. This was the only survey beam from the GBNCC survey that overlapped with the position of the Crab pulsar. This beam had a position offset of 0°.25 from the Crab pulsar position and had an integration time of two minutes. This position offset is close to (but still within) the edge of the  $\sim 0.3$  deg beam radius at 350 MHz for the GBT. The native sampling time of 81.92  $\mu$ s of the GBNCC survey beam was increased by a factor of two in our analysis, leading to an effective sampling time of 163.84  $\mu$ s. At these observing frequencies, the amount of dispersion smearing experienced by Crab pulses within the frequency channels in both observations is  $\sim 0.3$  ms, which significantly exceeds the sampling times. Thus, the different effective sampling times used in the two different observations had no effect on the Crab pulse detection rate.

In both observations, the data were searched blindly for single pulses at a range of dispersion measures (DMs) encompassing the Crab pulsar’s DM. We used the HEIMDALL single-pulse detection package in the search (Barsdell 2012; Barsdell et al. 2012).<sup>7</sup> All of the pulses detected by HEIMDALL that were within 0.3 pc cm<sup>−3</sup> the Crab’s nominal DM of 56.77 pc cm<sup>−3</sup> (Bilous et al. 2016) were retained. This DM window is consistent with the DM variability observed for the Crab’s giant pulses, which does not exceed this range (e.g., Lewandowska et al. 2022). We note that all of the pulses that were discarded as radio frequency interference (RFI) had DMs far from the Crab’s DM value (at least 20 pc cm<sup>−3</sup> away), so no



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<sup>7</sup> <https://sourceforge.net/projects/heimdall-astro>



**Figure 1.** Histograms of pulse amplitudes (S/N) for detected Crab pulses in observations taken with Arecibo (left) and the GBT (right). A total of 1943 pulses were detected with Arecibo in a 5 minute diagnostic observation at 327 MHz. For the GBT, 60 pulses were detected in a 2 minute observation that was part of the GBNCC survey at 350 MHz. In each case, the error bars were computed as the square root of the number of pulses in each bin. The best-fit values from power-law fits to the distributions are also plotted; they yielded differential power-law index values of  $-2.63 \pm 0.05$  and  $-3.6 \pm 0.5$  for the Arecibo and GBT data sets, respectively.

Crab pulses were accidentally eliminated. We measured the signal-to-noise ratio (S/N) of each detected pulse with HEIMDALL. The S/N is proportional to the pulse energy, given that the (narrow) Crab pulses are temporally unresolved due to the dominance of the DM smearing, and they therefore will have similar observed widths.

A total of 1943 and 60 Crab pulses were detected with an S/N above 6 in the Arecibo and the GBT observations, respectively. We did not make any distinction between main pulses (MPs) and interpulses (IPs) in our detected sample, even though IPs are expected to constitute a large fraction (about a third) of the giant pulses detected at these low frequencies (Cordes et al. 2004). We do not expect a significant difference between these two classes of pulses at low frequencies: Mikami et al. (2016) showed that, at 325 MHz, the average power-law index values for MPs and IPs are quite close (see their Table 3). This is supported by other studies (Cordes et al. 2004; Hankins et al. 2016; Lin et al. 2023) that have found that Crab MPs and IPs have similar properties at frequencies below a few gigahertz. In contrast to this, Hankins et al. (2015, 2016) have shown that, at high frequencies ( $\gtrsim 5$  GHz), the Crab MP and IP properties are quite different and the number of detected IPs exceeds the number of MPs by more than an order of magnitude.

We produced histograms of the pulse amplitudes using the Sturges criterion for the binning, where the number of bins  $k$  is determined by the total number of events  $n$  in the histogram, according to  $k = 1 + \log_2 n$ . Error bars were calculated by taking the square root of the number of events in each bin, in accordance with Poisson statistics. Each data set was separately fit using a power law of the form  $dN/dS \propto S^\beta$ , where both a coefficient and a power-law index were fit as free parameters. Figure 1 shows these two histograms and the best fits.

### 3. Results and Discussion

Scaling the GBT integration from 2 to 5 minutes (the integration time of the Arecibo observation) would result in about 150 detected pulses in the GBT data set. This number is

still much smaller than the almost 2000 pulses detected with Arecibo. The telescope gain difference is not a significant factor in this difference, because the Crab Nebula dominates the system temperature in both observations, as shown below. The Crab Nebula’s flux density at 350 MHz is  $S \approx 1270$  Jy (derived from the relation  $S = 955f^{-0.27}$  Jy, where  $f$  is the observing frequency in GHz; Bietenholz et al. 1997), while Arecibo’s system equivalent flux density (SEFD) is 10 Jy and the GBT’s is 35 Jy. The Crab Nebula (with a characteristic diameter  $5''.5$ ; Cordes et al. 2004) is also unresolved in both observations: the beam diameters of Arecibo and the GBT at these frequencies are  $15'$  and  $36'$ , respectively. Thus, all of the flux from the Crab Nebula was received by the telescope in each observation (see, e.g., Cordes et al. 2004). In both cases, the telescope SEFD is more than an order of magnitude smaller than the Crab Nebula’s flux density and does not significantly increase the system noise. Therefore, despite their difference in raw sensitivity, both Arecibo and the GBT have the same effective sensitivity to Crab giant pulses. The small difference in the central observing frequencies would also not affect the detection rates to this degree, nor would the sampling rate difference in the observations (see above). The RFI in each observation was also minimal (less than 0.1% of the data was masked in each case). However, the two observations were conducted at epochs separated by about 5.5 yr. Lundgren et al. (1995) showed that refractive interstellar scintillation (RISS) produces day-to-day variability for Crab giant pulses that affects the detection rate, so RISS may account for most of this difference in the detection rates.

Our best fits to the two histograms yielded differential power-law index values  $\beta$  of  $-2.63 \pm 0.05$  and  $-3.6 \pm 0.5$  for the Arecibo and GBT data, respectively. An alternate method of estimating the power-law index using integrated number counts in a maximum likelihood estimate has been outlined by Crawford et al. (1970) and James et al. (2019). This method avoids binning of the data and is considered to be less biased. We fit our two data sets using this approach as a check, and we

obtained differential power-law indices (derived from the cumulative power-law index) of  $-2.45 \pm 0.03$  and  $-3.7 \pm 0.4$  for the Arecibo and GBT data sets, respectively. These values are similar to those we obtained with binning, suggesting that both approaches produce consistent results. Our best-fit power-law values from the two data sets do not overlap within the stated ( $1\sigma$ ) errors. This difference may be attributable to the large amount of time separating the two observations (5.5 yr). Variability in the power-law index has been previously observed on long timescales (e.g., Rudnitskii et al. 2017).

Our power-law fit values are broadly consistent with previous low-frequency measurements from several different telescopes. Oronsaye et al. (2015) measured a power-law index of  $-3.35 \pm 0.35$  for the fluence distribution using the Murchison Widefield Array at a center frequency of 193 MHz (see their Figure 2). Observations taken with the Low Frequency Array at a similarly low center frequency of 150 MHz by van Leeuwen et al. (2020) showed an amplitude index for the fluence of  $-3.04 \pm 0.03$  (see their Figure 3). From measurements at 325 MHz with the Iitate Planetary Radio Telescope (IPRT) that were taken just a few months after our Arecibo observation, Mikami et al. (2016) measured  $\beta = -2.61^{+0.13}_{-0.15}$  for the Crab main pulse (see their Figure 4 and Table 3). This is remarkably close (within the formal uncertainties) to the value of  $-2.63 \pm 0.05$  we obtained from our Arecibo observation at essentially the same central observing frequency (327 MHz). The Crab Nebula’s flux density slightly exceeds the SEFD of the IPRT (Mikami et al. 2016; see their Table 3), so the IPRT system noise is comparable to the Crab Nebula contribution. This makes it slightly less sensitive to Crab giant pulses compared to Arecibo, where the SEFD is negligible relative to the Crab Nebula contribution (see above). The similarity of the Arecibo power-law index value to the IPRT value indicates that, even though the two observations were separated by a few months, a single power-law dependence extends to a slightly lower fluence level than what was measured with the IPRT. Table 4 of Mickaliger et al. (2012) listed four previously published low-frequency measurements (taken between 150 and 430 MHz) for the Crab differential power-law index (Argyle & Gower 1972; Cordes et al. 2004; Bhat et al. 2007; Smirnova & Logvinenko 2009) plus one new 330 MHz measurement using the Green Bank 43 m telescope. Combining these five measurements from this table yields an average power-law index of  $-2.9 \pm 0.5$  for the low-frequency regime ( $<500$  MHz). Table 1 shows a summary of these and other low-frequency measurements.

At higher observing frequencies, various studies have measured the power-law index. Mickaliger et al. (2012) (Table 4) listed measured values ranging from  $-2.1$  to  $-4.1$  for frequencies between 600 and 4850 MHz. Bera & Chengalur (2019) measured a value of  $-2.81 \pm 0.05$  at 1330 MHz, consistent with this range. Some of this variability in values can be attributed to RISS (both Lundgren et al. 1995 and Mickaliger et al. 2012 observed day-to-day variability in the measured power-law index at 812 MHz and 1.2 GHz, respectively). We note, however, that Bera & Chengalur (2019) did not see variability in the power-law index on timescales of a few days in observations taken at 1330 MHz, though on much longer timescales (a few years) it has been observed to vary (Rudnitskii et al. 2017). Mickaliger et al. (2012) do not see a clear trend in how the power-law index changes with frequency, but Cordes et al. (2004) report a possible steepening of the index from 430 MHz to 8.8 GHz.

**Table 1**  
Measured Crab Giant Pulse Power-law Indices at Low Radio Frequencies

Frequency Range (MHz)	Power-law Index	Telescope	References
20–84	...	LWA1	Eftekhari et al. (2016)
20–84	...	LWA1	Ellingson et al. (2013)
23	...	UTR-2	Popov et al. (2006)
111	...	LPA	Popov et al. (2006)
600	...	RT-64	Popov et al. (2006)
110–180	$-1.73 \pm 0.45$	WSRT	Karuppusamy et al. (2012)
185–200	$-3.35 \pm 0.35$	MWA	Oronsaye et al. (2015) (Figure 2)
111–189	$-3.04 \pm 0.03$	LOFAR	van Leeuwen et al. (2020) (Figure 3)
325	$-2.61^{+0.13}_{-0.15}$	IPRT	Mikami et al. (2016) (Figure 4 and Table 3)
112–430	$-2.9 \pm 0.5$	multiple	Mickaliger et al. (2012) (Table 4)
293–361	$-2.63 \pm 0.05$	Arecibo	this work
300–400	$-3.6 \pm 0.5$	GBT	this work

**Notes.** The table lists differential (not cumulative) power-law indices. Ellingson et al. (2013) do not provide an estimated power-law index, due to concerns about calibration and correcting for the flux density of the Crab Nebula. Eftekhari et al. (2016) do not provide an estimated power-law index, due to concerns about the small number of giant pulses detected. Popov et al. (2006) do not report a power-law index. Telescope abbreviations: LWA1 = Long Wavelength Array Station 1, UTR-2 = Ukrainian T-shaped Radio Telescope (Second Modification), LPA = Large Phased Array (Pushchino), MWA = Murchison Widefield Array, WSRT = Westerbork Synthesis Radio Telescope, RT-64 = Kalyazan 64 m Radio Telescope, LOFAR = Low Frequency Array, and IPRT = Iitate Planetary Radio Telescope.

In conclusion, we have measured power-law index values for the Crab giant pulse amplitude distribution using two separate low-frequency observations taken with Arecibo and the GBT. The best-fit values are broadly consistent with values previously measured at low frequencies with different telescopes. These measurements may be useful in future giant pulse studies of the Crab pulsar.

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*Facilities:* Arecibo, GBT.

*Software:* HEIMDALL (Barsdell 2012; Barsdell et al. 2012).

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