



JWST Photometric Time-delay and Magnification Measurements for the Triply Imaged Type Ia “SN H0pe” at $z = 1.78$

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Abstract

Supernova (SN) SN H0pe is a gravitationally lensed, triply imaged, Type Ia SN (SN Ia) discovered in James Webb Space Telescope imaging of the PLCK G165.7+67.0 cluster of galaxies. Well-observed multiply imaged SNe provide a rare opportunity to constrain the Hubble constant (H_0), by measuring the relative time delay between the images and modeling the foreground mass distribution. SN H0pe is located at $z = 1.783$ and is the first SN Ia with sufficient light-curve sampling and long enough time delays for an H_0 inference. Here we present photometric time-delay measurements and SN properties of SN H0pe. Using JWST/NIRCam photometry, we measure time

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delays of $\Delta t_{\text{ab}} = -116.6^{+10.8}_{-9.3}$ observer-frame days and $\Delta t_{\text{cb}} = -48.6^{+3.6}_{-4.0}$ observer-frame days relative to the last image to arrive (image 2b; all uncertainties are 1σ), which corresponds to a $\sim 5.6\%$ uncertainty contribution for H_0 assuming $70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. We also constrain the absolute magnification of each image to $\mu_a = 4.3^{+1.6}_{-1.8}$, $\mu_b = 7.6^{+3.6}_{-2.6}$, $\mu_c = 6.4^{+1.6}_{-1.5}$ by comparing the observed peak near-IR magnitude of SN H0pe to the nonlensed population of SNe Ia.

Unified Astronomy Thesaurus concepts: Cosmology (343); Type Ia supernovae (1728); Gravitational lensing (670); Galaxy clusters (584)

1. Introduction

Strong gravitational lensing can cause multiple images of a background source to appear, as light propagating along different paths is focused by a foreground galaxy or galaxy cluster (e.g., Narayan & Bartelmann 1997). Such a phenomenon requires chance alignment between the observer, the background source, and the foreground galaxy. If the multiply imaged source has variable brightness, then, depending on the relative geometrical and gravitational potential differences of each path, the source images will appear delayed by hours to weeks (for galaxy-scale lenses) or months to years (for cluster-scale lenses) (e.g., Oguri 2019).

Precise measurements of this “time delay” yield a direct distance measurement to the lens system that constrains the Hubble constant (H_0) in a single step (e.g., Refsdal 1964; Paraficz & Hjorth 2009; Linder 2011; Treu & Marshall 2016; Grillo et al. 2018, 2020; Birrer et al. 2022b; Treu et al. 2022; Suyu et al. 2024). Quasars have historically been used for this purpose (e.g., Vuissoz et al. 2008; Suyu et al. 2010; Tewes et al. 2013; Bonvin et al. 2017, 2018, 2019a; Birrer et al. 2019; Wong et al. 2020), since supernovae (SNe) exploding in multiply imaged galaxies are relatively rare, with only a handful discovered thus far (Kelly et al. 2015; Goobar et al. 2017, 2023; Rodney et al. 2021; Chen et al. 2022; Kelly et al. 2022; Pierel et al. 2023). Nevertheless, SNe have several advantages over quasars that are attractive for cosmological measurements: (1) SNe fade, enabling predictive experiments on the delayed appearance of trailing images and more accurate photometry, as the SN (or quasar), lens, and host fluxes are otherwise highly blended (Ding et al. 2021); (2) SNe generally have predictable behavior compared to quasars, simplifying time-delay measurements and requiring much shorter observing campaigns; and (3) the impact of a “microlensing time delay” is significantly mitigated owing to the more compact source size (Tie & Kochanek 2018; Bonvin et al. 2019b).

SNe of Type Ia (SNe Ia), known as “standardizable candles,” can be used to measure cosmological parameters by way of luminosity distances and the cosmic distance ladder (e.g., Garnavich et al. 1998; Riess et al. 1998; Perlmutter et al. 1999; Scolnic et al. 2018; Brout et al. 2022). SNe Ia are of particular value when strongly lensed, as their standardizable absolute brightness can provide additional leverage for lens modeling by limiting the uncertainty caused by the mass sheet degeneracy (Falco et al. 1985; Kolatt & Bartelmann 1998; Holz 2001; Oguri & Kawano 2003; Nordin et al. 2014; Patel et al. 2014; Rodney et al. 2015; Xu et al. 2016; Foxley-Marrable et al. 2018; Birrer et al. 2022a), though only in cases where microlensing and microlensing are not extreme (see Goobar et al. 2017; Yahalomi et al. 2017; Foxley-Marrable et al. 2018; Dhawan et al. 2020). Additionally, SNe Ia have well-understood models of light-curve evolution (Hsiao et al. 2007; Guy et al. 2010; Saunders et al. 2018; Leget et al. 2020; Kenworthy et al. 2021; Pierel et al. 2022; Mandel et al. 2022) that enable

precise time-delay measurements using color curves, which removes the effects of macro- and achromatic microlensing (e.g., Pierel & Rodney 2019; Huber et al. 2021; Rodney et al. 2021; Pierel et al. 2023). The first two spectroscopically confirmed multiply imaged SNe Ia were iPTF16geu (Goobar et al. 2017) and SN Zwicky (Goobar et al. 2023; Pierel et al. 2023), but both had very short time delays of ~ 0.25 – 1.5 days that precluded competitive H_0 measurements (Dhawan et al. 2020; Pierel et al. 2023).

Overall, while the advantages of using SNe for time-delay cosmography relative to other probes have been well documented (e.g., Refsdal 1964; Kelly et al. 2015; Goobar et al. 2017; Goldstein et al. 2018; Huber et al. 2019; Pierel & Rodney 2019; Suyu et al. 2020; Pierel et al. 2021; Rodney et al. 2021; Suyu et al. 2024), only a single strongly lensed SN (the core-collapse SN “Refsdal”; Kelly et al. 2015) has had time delays long enough (~ 1 yr) and sufficient light-curve sampling for precise time-delay measurements ($\sim 1.5\%$; Kelly et al. 2023b), which yielded a $\sim 6\%$ measurement of H_0 ($64.8^{+4.4}_{-4.3}$ or $66.6^{+4.1}_{-3.3} \text{ km s}^{-1} \text{ Mpc}^{-1}$, depending on lens model weights; Kelly et al. 2023a). A new multiply imaged SN Ia, SN H0pe (Frye et al. 2024, hereafter F24), was discovered in 2023 March by the James Webb Space Telescope (JWST) program “Prime Extragalactic Areas for Reionization and Lensing Science” (PEARLS; PID 1176, Windhorst et al. 2023) in NIRCcam observations of the PLCK G165.7+67.0 cluster of galaxies (Figure 1). Previous observations of the PLCK G165.7+67.0 cluster are described in Cañameras et al. (2015), Harrington et al. (2016), Cañameras et al. (2018), Frye et al. (2019), and Pascale et al. (2022). This work is part of a series of papers analyzing the SN H0pe system, which, together with M. Pascale et al. (2024, in preparation), will provide the first measurement of H_0 with a multiply imaged SN Ia. The discovery of SN H0pe is described in F24, and time delays derived only from spectroscopy of the system are presented in Chen et al. (2024). In this paper we focus on imaging of the SN itself, including the photometric time-delay measurement critical to that forthcoming H_0 inference. These measurements are made using NIRCcam photometry and a well-vetted model for SN Ia light-curve evolution, and we produce a series of simulations to ensure that the uncertainties are well characterized. We also compare the observed brightness of SN H0pe to the general SN Ia population to infer absolute magnifications for the system.

The outline of this paper is as follows: Section 2 gives an overview of SN H0pe and summarizes the JWST observation characteristics, which are described in more detail in F24. Section 3 details the photometric measurements used in the remainder of the paper, with analysis of time delays and magnifications reported in Section 4. In Section 5 we create a suite of simulations to understand various systematic uncertainties associated with the results of Section 4, and we

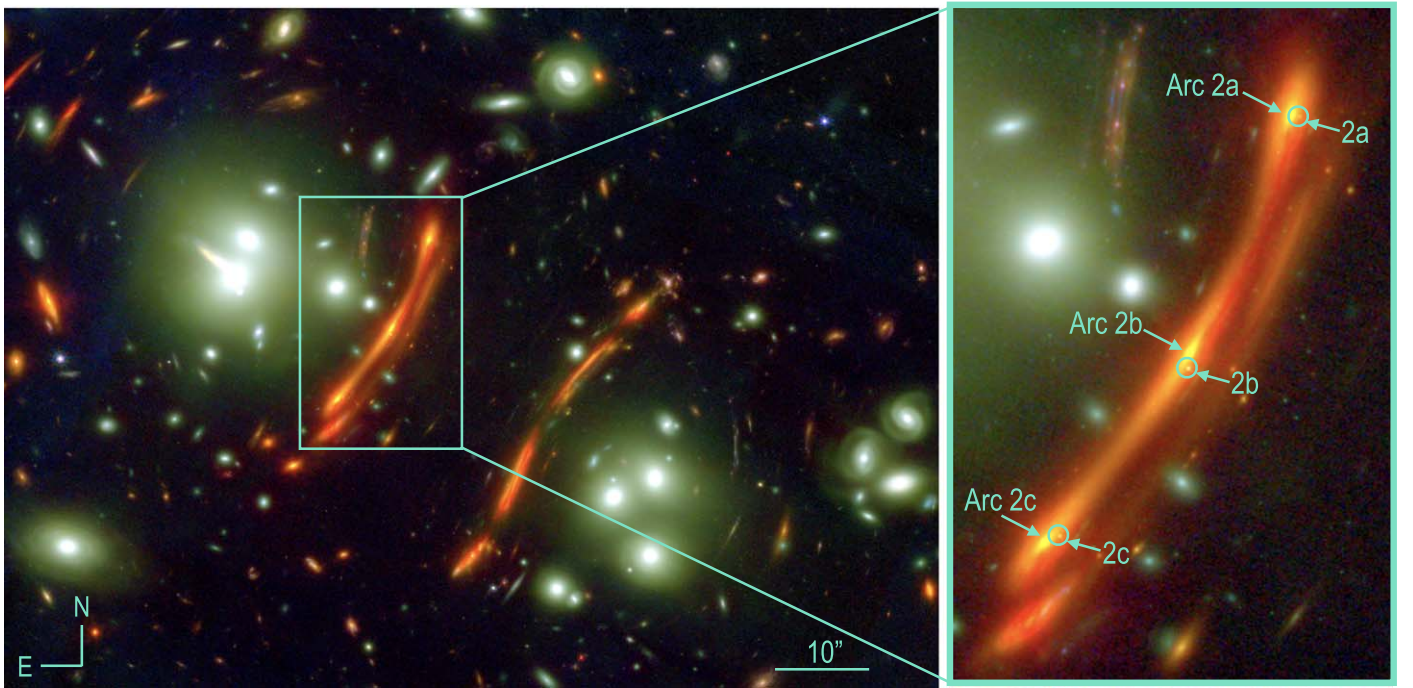


Figure 1. Left: JWST/NIRCam color image in the central region of PLCK G165.7+67.0 from PEARLS (epoch 1). The filters used are F090W (blue), F200W (green), and F444W (red). The images were drizzled to $0''.03 \text{ pixel}^{-1}$, and the image scale and orientation are as shown. Right: close-up of the boxed region depicting the three images of the galaxy Arc 2, as labeled. SN H0pe (circled) appears in all three images and flips parity between images 2c and 2b and between images 2b and 2a, as predicted by lensing theory.

conclude with a discussion of the implications for these measurements in Section 6.

2. Summary of JWST Observations

The description of PEARLS and Space Telescope Science Institute’s Director’s Discretionary Time (DDT) observations, SN discovery, and initial analysis of the SN H0pe lensing system are presented by F24. Briefly, the SN was discovered by the PEARLS team in (eight-filter) NIRCam imaging taken on 2023 March 30 (MJD 60033) by comparing the F150W data with an archival Hubble Space Telescope (HST) WFC3/IR F160W image (2016 May 30, MJD 57538; Frye et al. 2023). These two filters are extremely well matched in wavelength and transmission, and as the source was very bright ($\lesssim 24 \text{ mag AB}$) in F150W but absent in F160W, its transient nature was confirmed. The initial photometric estimate of the host galaxy redshift was $z \sim 2.3$ from HST and Spitzer $3.6 \mu\text{m}/4.5 \mu\text{m}$ photometry (Pascale et al. 2022), but the poorly matched resolution between HST and Spitzer led to this being an overestimate. Updated NIRCam photometry provided a much more robust estimate of $z \sim 1.8$, benefiting from the improved wavelength range and resolution. Adopting predicted time delays and magnifications from the Frye et al. (2019) lens model updated with the NIRCam imaging, initial photometric classification suggested that the SN was likely of Type Ia. A JWST DDT proposal was subsequently submitted and accepted (PID 4446), which provided NIRSpec Prism, G140M, and G235M spectroscopy of the multiple images and two epochs of additional NIRCam imaging in six filters (see Section 3.4 for exposure times).

The spectroscopic analysis and updated lens modeling are described in companion papers, but the NIRSpec data unambiguously confirmed that the SN was of Type Ia at a

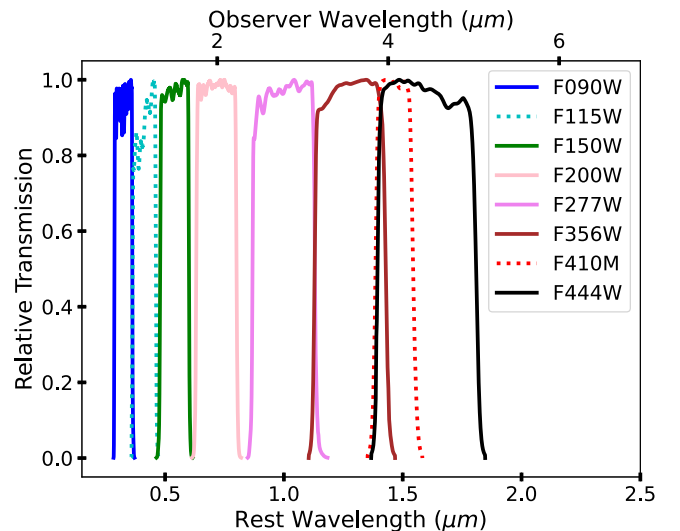


Figure 2. The JWST filters used to observe SN H0pe, with rest-frame wavelength on the lower axis and observer-frame wavelength on the upper axis. Filters with dashed lines (F115W and F410M) were only observed in epoch 1, while the others were observed in all three epochs, but overall rest-frame wavelength coverage is excellent.

refined spectroscopic redshift of $z = 1.783$ (Chen et al. 2024; F24), whose value matches the redshift of the SN host galaxy of $z = 1.7833 \pm 0.0005$ (F24; Polletta et al. 2023). The two NIRCam imaging epochs occurred on 2023 April 22 (MJD 60056) and 2023 May 9 (MJD 60073). This meant that the light curve for each image contained three observations each separated by ~ 3 observer-frame weeks, or ~ 1 rest-frame week. The combined data set effectively produces an SN Ia light curve with nine epochs and rest-frame UV–near-IR (NIR) wavelength coverage, which is sufficient to accurately measure

phases of each image and therefore the relative time delays (e.g., Pierel & Rodney 2019). The wavelength coverage of the NIRCcam imaging is shown in Figure 2.

3. Measuring Photometry

In most analyses involving dithered exposures, the individual exposures for each filter are drizzled together to create a single (level 3) image. When one is primarily concerned with precisely measuring the position and brightness of an SN, point-spread function (PSF) fitting on the (level 2) NIRCcam “CAL” images is the best-understood method (e.g., Rigby et al. 2023), which preserves the PSF structure compared to drizzled images. CALs are individual exposures that have been calibrated using the STScI JWST Pipeline³⁹ (Bushouse et al. 2022) and have been bias-subtracted, dark-subtracted, and flat-fielded but not yet corrected for geometric distortion. The NIRCcam pixel area map for the corresponding short-wavelength (SW) and long-wavelength (LW) modules containing the SN must also be applied to each exposure to correct for pixel area variations across the images.

Due to the increased brightness of the host galaxy Arc 2 and diminished NIRCcam resolution at longer wavelengths, the signal-to-noise ratio of SN H0pe is only high enough to rely on individual exposures in the SW channel, and most reliably in the F150W filter, where SN H0pe is brightest. We will therefore measure photometry in the F150W filter on CAL images as a baseline, since this well-established method should be robust (see Section 3.1), but in order to be consistent across all filters, this is only used for comparison and not in the final analysis. Instead, we implement a PSF fitting routine suitable for drizzled images in Section 3.2, where we test that the results are consistent with the CAL measurements. Finally, to enable accurate photometry at all wavelengths without a template image, we remove Arc 2 light in the drizzled images using a host surface brightness modeling technique described in Section 3.3. We use injected PSF models to ensure that the host subtraction method does not introduce biases, and we extend the method to all filters for our final photometry in Section 3.4. In all cases we use the PSF models provided by WebbPSF (version 1.2.1)⁴⁰ to represent the (level 2) PSF, which have been updated to better match the observed PSF in each filter, and take into account temporal and spatial variation across the detector.

3.1. PSF Fitting on F150W CAL Images

For each epoch in the F150W filter, we use the Python package `space_phot`⁴¹ to simultaneously constrain the (common) SN flux in all CALs (four, one for each dither position) for all three SN images. Each PSF was fit to the multiple SN images within a 3×3 pixel square in an attempt to limit the contamination of the Arc 2 halo. The final measured flux is the integral of each full fitted PSF model, which includes a correction to the infinite aperture flux. These total fluxes, which are in units of MJy sr^{-1} , are converted to AB magnitudes using the native pixel scale of each image ($0''.03 \text{ pixel}^{-1}$ for SW, $0''.06 \text{ pixel}^{-1}$ for LW). A constant background is removed from the fitting region, which is small enough that a constant

background is sufficient. The background is estimated as the median value in an annulus centered at the SN position with an inner radius of 5 pixels and an outer radius of 7 pixels. In this case the method of background estimation is not critical, as long as it is reasonable and the same method is used in the following section, as the final photometry uses the host galaxy surface brightness modeling method (Section 3.3). The measured photometry from this section is only used as a baseline to compare with the F150W photometry measured on drizzled images in the next section, and it is not used in the final analysis.

3.2. PSF Fitting on F150W Drizzled Images

In this section we use `space_phot` to measure PSF photometry on level 3 F150W drizzled images as a test of our method. `space_phot` creates files with the same header and WCS information as the individual CAL files, but containing only a PSF model at the SN location. The PSF/CAL files are then drizzled using the same method used to produce the drizzled data files, which builds a PSF model suitable for measuring photometry only in that single drizzled image and at the SN position. We then fit the SN flux in the same 3×3 pixel square with the same background estimation method as the previous section, and we convert to AB magnitudes in the same manner. The resulting magnitudes agree with the F150W CAL file measurements from the previous section to within $\sigma_{\text{PSF}} = 0.04$ mag. This is within the statistical uncertainty of our level 3 PSF fitting for most filters, but we include it as a possible systematic in our final photometry and proceed with measuring photometry on level 3 drizzled images for the remainder of this work.

3.3. Host Model-subtracted Images

As shown in Figure 1, Arc 2 is bright ($m_{\text{F200W}} = 20.3$) and red. This makes extraction of the SN light difficult without a template used for difference imaging, particularly at longer wavelengths. To achieve an efficient subtraction of the contaminant light, we need to perform a surface brightness reconstruction of the host/background galaxy accounting for the lensing effect. Such additional distortions of the arc image are not well represented by analytical profiles, and we therefore use the Gravitational Lens Efficient Explorer (GLEE) software because of its capability to perform detailed surface brightness reconstruction in lensing scenarios (Suyu & Halkola 2010; Suyu et al. 2012). The detailed galaxy surface brightness and lens modeling will be published in G. Caminha et al. (2024, in preparation), and here we give a short description of the main steps. We first use the positions of substructures along Arc 2 and other multiply imaged galaxies to model the gravitational lensing deflection adopting a parametric approach for the cluster total mass distribution. Our total mass parameterization consists of two main cluster-scale components plus a dual pseudo-isothermal elliptical (dPIE) density profile for each securely identified cluster member using the F150W magnitudes for the scaling relation, which is similar to previous works (see, e.g., Pascale et al. 2022). In a second step we use the extended light of Arc 2 to model the cluster lens properties and the source light simultaneously, using the best-fit mass model obtained with the multiple image positions as a starting point to find the final solution. The light distribution of Arc 2 in the source plane is described by a grid of pixels but with a

³⁹ <https://github.com/spacetelescope/jwst>

⁴⁰ <https://www.stsci.edu/jwst/science-planning/proposal-planning-toolbox/psf-simulation-tool>

⁴¹ space-phot.readthedocs.io

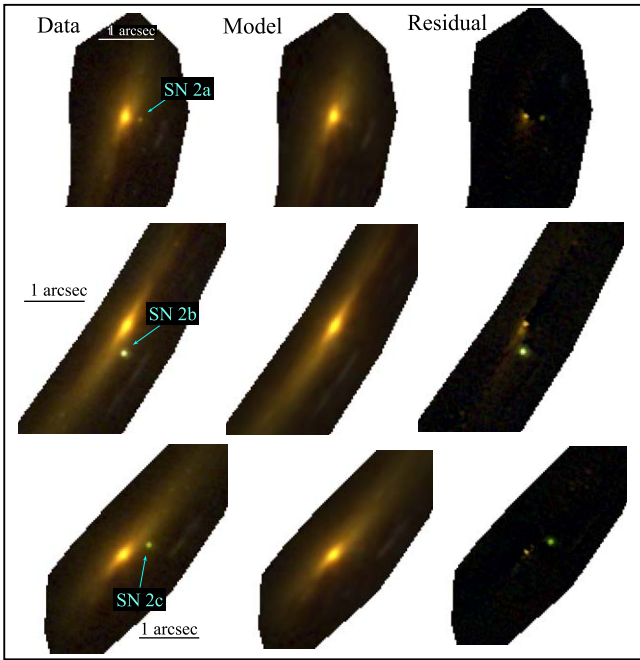


Figure 3. Host light subtraction for Arc 2. Color-composite images of epoch 1 are shown using the filters F090W, F150W, and F200W as blue, green and red, respectively. Left panels show the original data, middle panels show the surface brightness distribution predicted by our strong-lensing total mass model, and right panels show the residuals (i.e., data–model) with the same scaling. Note that the SN is not present in the model, and so it is expected to be present in the residuals. The orientation is the same as in Figure 1.

curvature regularization to avoid high pixel-to-pixel variations (see Suyu et al. 2006, for more details). The light distribution is then mapped to the image plane and convolved with the PSF estimated directly on each single band image interactively to minimize the differences between the observed data and model. In Figure 3, we show an example of subtraction using epoch 1 data. The process is repeated for all epochs and all filters, which are used for the final photometry measured in Section 3.4.

3.3.1. Injected Source Recovery Test

In order to test whether the host subtraction method above biases our SN photometry measurements, we plant three point sources representing SNe around each image of Arc 2 for all epochs and filters. Each injected point source is roughly the same brightness as the corresponding image of SN H0pe, and together the three fake sources and SN H0pe form a rectangle around each Arc 2 image. The locations were chosen to obtain similar contamination from the host galaxy in image 2a, and then those plant positions were translated to positions in 2b and 2c with the lens model. After subtracting the host light derived in the previous section, we measure the flux of each injected PSF with the same method described above. We use the resulting rms in each filter as an additional systematic uncertainty on our final measured photometry (Table 1), described in the following section. Note that F356W and F444W are absent, as we found that the source reconstruction method was not reliable at these longest wavelengths when recovering injected PSF models. We therefore remove these filters from the current analysis until a template is obtained for the system in JWST Cycle 3 (PID 4744) for traditional difference imaging.

Table 1

Systematic Photometry Uncertainties Measured by Recovering Injected PSF Models

Filter	σ_{plant} (mag)
F090W	0.02
F115W	0.05
F150W	0.06
F200W	0.07
F277W	0.14

Table 2

Photometry Measured for Each Image of SN H0pe in Each Filter and for Each Observing Epoch

MJD	Filter	Exp. Time (s)	m_{2a}	m_{2b}	m_{2c}
60033	F090W	2490	30.55 ± 0.47	25.53 ± 0.05	28.71 ± 0.11
60033	F115W	2490	28.35 ± 0.11	24.69 ± 0.06	26.16 ± 0.07
60033	F150W	1889	26.50 ± 0.08	23.93 ± 0.07	24.95 ± 0.07
60033	F200W	2104	26.45 ± 0.08	24.21 ± 0.08	25.25 ± 0.08
60033	F277W	2104	25.72 ± 0.14	24.93 ± 0.14	25.58 ± 0.14
60057	F090W	1245	>29.71	26.44 ± 0.07	30.23 ± 0.74
60057	F150W	858	27.06 ± 0.10	24.03 ± 0.07	25.45 ± 0.08
60057	F200W	1760	26.82 ± 0.09	24.35 ± 0.08	25.06 ± 0.08
60057	F277W	1760	26.07 ± 0.15	25.20 ± 0.14	25.37 ± 0.14
60074	F090W	1417	>29.48	27.21 ± 0.09	30.45 ± 0.81
60074	F150W	1245	27.37 ± 0.12	24.42 ± 0.07	25.94 ± 0.08
60074	F200W	1760	26.74 ± 0.09	24.81 ± 0.08	25.24 ± 0.08
60074	F277W	1760	26.28 ± 0.15	25.21 ± 0.14	25.18 ± 0.14

Note. Columns show Modified Julian Date, NIRCcam filter, exposure time in seconds, and photometry plus final uncertainty for SN image “ m_{2a} ,” “ m_{2b} ,” and “ m_{2c} ” in AB magnitudes. Upper limits are 3σ .

3.4. Final Photometry

We measure the final photometry on the host-subtracted image from the previous section in all filters and epochs. The final photometry measured here is reported in Table 2, and the uncertainties correspond to $\sigma_{\text{tot}} = \sqrt{\sigma_{\text{stat}}^2 + \sigma_{\text{plant}}^2 + \sigma_{\text{PSF}}^2}$, where σ_{stat} is the statistical uncertainty derived from the posterior of each PSF fitting routine, σ_{plant} is measured from injected source recovery in Section 3.3.1, and σ_{PSF} is the estimated additional uncertainty from performing PSF photometry on level 3 data (Section 3.2). A final source of photometric uncertainty is a systematic uncertainty on the zero-points, which is $\lesssim 0.01$ mag for all filters and is therefore subdominant to the uncertainties derived here (M. Boyer 2024, private communication; M. Boyer et al. 2024, in preparation).

4. Light-curve Fitting

We fit the photometry of the multiple images simultaneously using the SN Time Delays (SNTD) software package (Pierel & Rodney 2019). Milky Way dust extinction ($E(B - V) = 0.019$, $R_V = 3.1$) was based on the maps of Schlafly & Finkbeiner (2011) and extinction curve of Fitzpatrick (1999). We used a phase-extended version of the BayeSN SN Ia spectral energy distribution (SED) model (Mandel et al. 2022; Grayling et al.

Table 3
Summary of the BayeSN Parameters Used in Fitting SN H0pe

Parameter	Description	Free/Fixed	Bounds	Value
z	Redshift	Fixed	...	1.783
t_{pk} (MJD)	Image 2b time of maximum	Free	$[-20, 20] + 60033$	$60037.77^{+2.12}_{-2.01}$
θ	Light-curve shape	Free	$[-2, 2]$	$-0.74^{+0.30}_{-0.27}$
A_V	V-band dust extinction	Free	$[0, 3]$	$1.21^{+0.11}_{-0.11}$
R_V	$A_V/E(B - V)$	Free	$[1, 6]$	$3.10^{+0.40}_{-0.34}$

2024), which covers the full plausible rest-frame phase range of SN H0pe, to fit the measured photometry. The version of BayeSN we used is a variant of the one presented by Ward et al. (2023) and uses the training set detailed therein. The model is trained to cover rest-frame phases as late as 50 days, with linear extrapolation in the space of the logarithm of the SED beyond that. We included u -band data in the training, enabling coverage of the F090W photometry of SN H0pe. This is the only model for SN Ia light-curve evolution with the necessary flexibility and sufficient wavelength and temporal coverage for this data set (particularly for image 2a), making it difficult to estimate systematic uncertainties associated with a single choice of SN model. We therefore produced a suite of simulations (described in Section 5) to quantify the final time-delay uncertainties. In addition to the base template and Milky Way dust noted above, BayeSN includes a rest-frame host galaxy dust component (parameterized by the V-band extinction A_V and ratio of total to selective extinction R_V) and a light-curve shape parameter θ . (This has a strong negative correlation with SALT2’s stretch parameter, x_1 , meaning that positive θ implies a faster decline rate; see, e.g., Figure 4 of Mandel et al. 2022.) The free parameters in the model and their status in the fitting process are summarized by Table 3.

4.1. Time Delays

With the knowledge that SN H0pe is an SN Ia, we are able to leverage the BayeSN SED model (Mandel et al. 2022) to measure time delays with the SNTD Color method (Pierel & Rodney 2019; Pierel et al. 2023). This method finds common values for the BayeSN light-curve parameters (i.e., shape θ and dust extinction parameters R_V and A_V) among all SN H0pe images, while measuring time delays relative to the value of t_{pk} corresponding to image 2b (chosen as the brightest image and last to arrive; Figure 1). The Color method directly fits using color curves (i.e., observed SN colors over time), which removes the need for normalization and relative magnifications. The primary advantage of the Color method is that it is insensitive to achromatic microlensing, as any such effect is identically present in all bands (The impacts of chromatic microlensing are discussed in Section 5.2.2.) Degeneracies between magnification ratios and time delays are also removed, but the cost of color curves is fewer clear inflection points compared to light curves in certain filter combinations and phases. The SN H0pe light curve has sufficient features in the color curves to use this method, which optimizes a joint likelihood function including both the SN Ia light-curve model parameters and the time delays with bounds of ± 200 observer-frame days. Figure 4 shows the measured photometry and model with best-fit parameters and time delays applied, and the results are summarized in Tables 3 and 4. Note that the measured value for θ , -0.74 , is equivalent to a normal decline

Table 4
Measured Time Delays and Magnifications

Image	$\Delta t (t_i - t_b)$	μ
	(Days)	
2a	$-116.6^{+10.8}_{-9.3}$	$4.3^{+1.6}_{-1.8}$
2b	...	$7.6^{+3.6}_{-2.6}$
2c	$-48.6^{+3.6}_{-4.0}$	$6.4^{+1.6}_{-1.5}$

rate of $\Delta m_{15}(B) = 1.13$ (Phillips 1993). The measured value of Δt in Table 4 includes corrections based on the systematic uncertainties derived in Section 5 following the methods of Kelly et al. (2023b). While some of the points appear discrepant from the model at $\sim 1\sigma$ – 2σ (e.g., the first epoch for image 2a in F150W–F200W), this has been propagated through to the final time-delay and magnification uncertainties by way of our simulations in Section 5, and further improvements will require improved photometry from difference imaging and an improved late-time SN Ia model (see Section 6). The posterior distributions from the fitting are shown in Figure 5. The distributions display an expected covariance between (R_V, A_V) , as well as some covariances between $(\theta, \Delta t_{bc})$ and t_{pk} and the time delays. The covariances with t_{pk} are also to be expected, while the $(\theta, \Delta t_{bc})$ covariance is likely because the second infrared maximum constrains Δt_{bc} and is directly related to θ . While we cannot fit all three images of SN H0pe with existing models other than BayeSN, for consistency we do measure time delays for images 2b and 2c with the SALT3-NIR model. We find agreement within 1σ and proceed with our BayeSN result, as it can be applied to all three images.

4.2. Magnifications

Relative magnifications between the multiple images have been used as a model-independent means of weighting lens models while inferring H_0 (i.e., Kelly et al. 2023a), but absolute magnifications provide a more stringent constraint when measurable. Given their standardizability, SNe Ia provide a rare opportunity to measure an absolute magnification, increasing their utility for time-delay cosmology (e.g., Birrer et al. 2022a). As mentioned in the previous section, the SNTD Color method does not measure an overall normalization. We therefore fixed the intrinsic model parameters and time delays to those measured in the previous section and then fit for the image 2b normalization and magnification ratios of images 2a and 2c relative to image 2b. We converted the magnification ratios to absolute magnifications by assuming that SN H0pe is standardizable in the rest-frame NIR, where the impacts from dust are mitigated relative to the rest-frame optical. Specifically, we used the peak F277W magnitude (\sim rest-frame Y

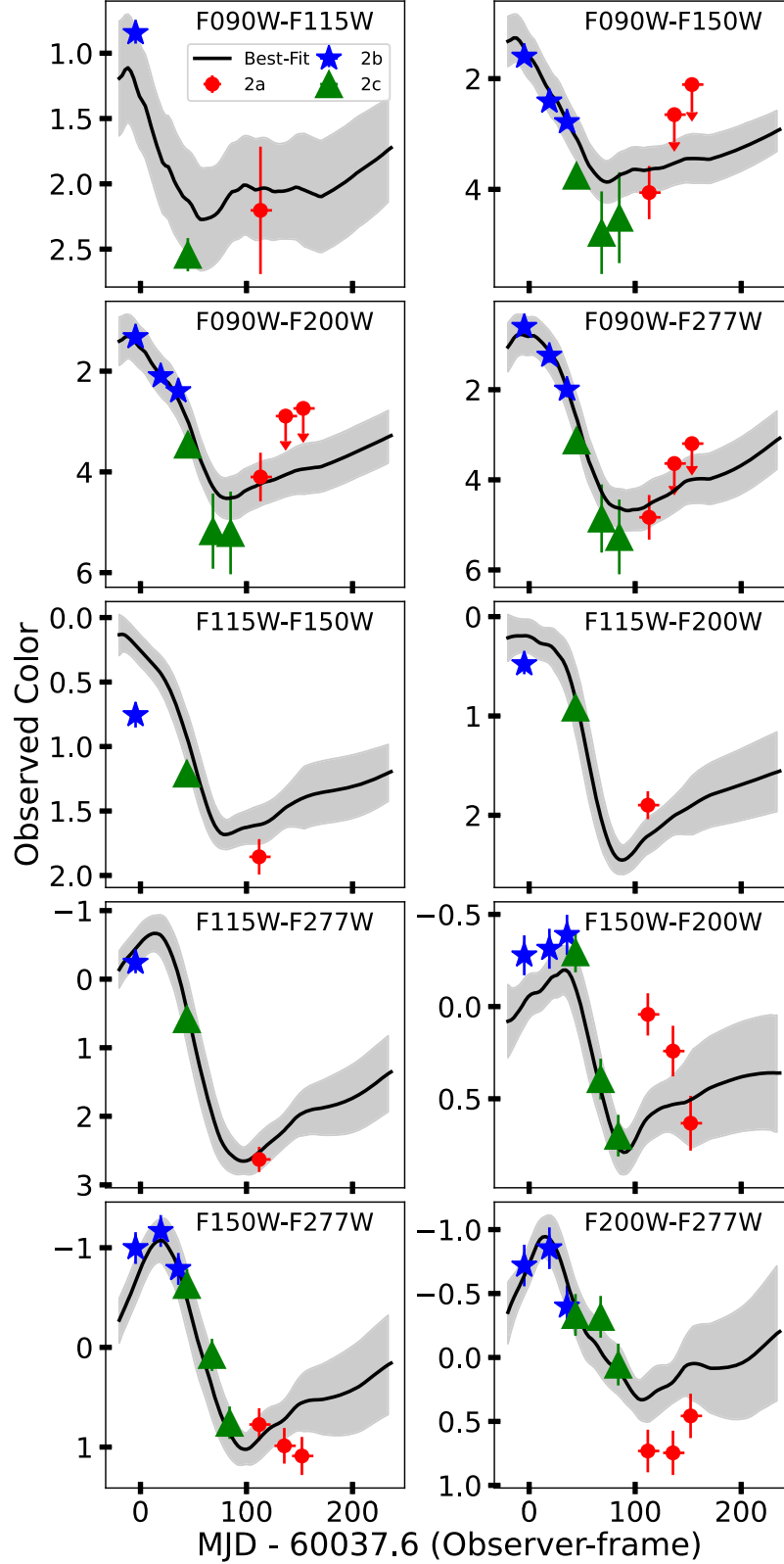


Figure 4. Reconstructed color curves for SN H0pe, after applying the best-fit time delays. The vertical error bars are the photometric precision, while the horizontal error bars are the 16th and 84th percentile simulation recovery time-delay posterior for images 2a and 2c (see Table 5), and t_{pk} for image 2b. The gray shaded region is uncertainty on the best-fit model from the SNTD Color method (black solid).

band) measured here from BayeSN including the final measurement uncertainty, which previous work has shown to be an excellent estimator for luminosity distance while fitting for shape and color for normal SNe Ia such as SN H0pe (e.g.,

Pierel et al. 2022). We fit the complete sample of spectroscopic SN Ia peak m_B measurements (Scolnic et al. 2018; Brout et al. 2022) with three separate methods to obtain the predicted value of m_B at $z = 1.783$ (Figure 6). The three methods are a simple

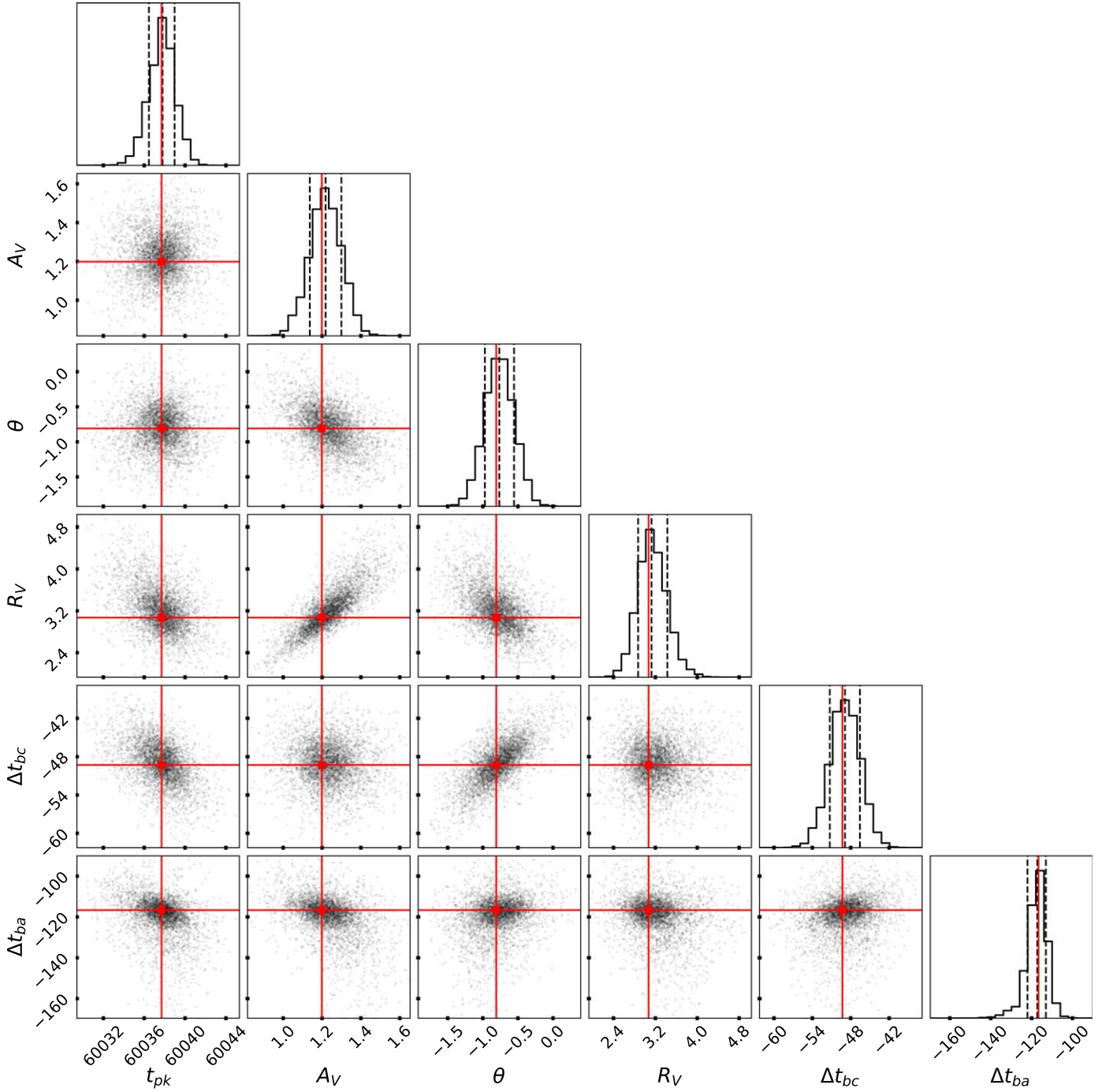


Figure 5. Posterior distributions of the SNTD Color method fitting of HST photometry. The dashed vertical lines correspond to the distribution 16th, 50th, and 84th percentiles. The reported measurement values (solid red lines) are the maximum likelihood estimates from these distributions.

second-order polynomial, a Gaussian process, and a kinematic expansion (Riess et al. 2022, Equation (4) with $j_0 = 1$). We take the average of these methods (26.13 AB mag) and propagate the standard deviation to the magnification uncertainties (0.14 mag). We then scaled the best-fit BayeSN model to this B -band magnitude and predict a nonlensed F277W AB magnitude of 26.98. We compared the best-fit model-predicted F277W to this reference magnitude for each image to give absolute magnification measurements. This would be fully model independent given a set of distance measurements using F277W at this redshift, but this is not possible at the current time. However, as the m_B measurements we used as a reference

are independent of cosmology, our magnification measurements are still independent of H_0 . We combined the uncertainties on each measured magnification (Section 5) with systematic uncertainties based on the intrinsic scatter of SN Ia absolute magnitudes measured by Scolnic et al. (2018; 0.1 mag). They further report zero-point uncertainties of <0.01 mag, and we therefore ignore this additional systematic. The measured time delays and magnifications are shown in Table 4, and the final reconstructed light curve is shown in Figure 7, where all observations are consistent with the model within 1σ . Once again, we attempt to measure magnifications in the same manner using SALT3-NIR for images 2b and 2c,

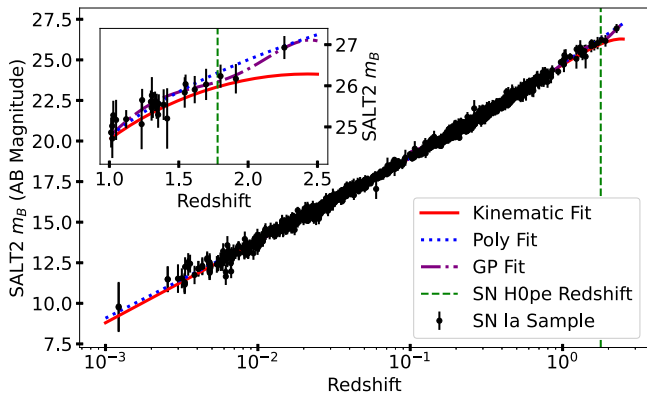


Figure 6. A visual representation of finding the predicted value of m_B for an SN Ia at $z = 1.783$, using the SNe Ia from Scolnic et al. (2018) and Brout et al. (2022) (black). The data are fit with a second-degree polynomial (blue dotted), a Gaussian process (purple dashed-dotted), and a kinematic expansion (red solid). The redshift of SN H0pe is shown as a dashed green line.

and we find agreement within 1σ – 2σ . While this is in slight tension, without more SNe Ia at high z this method must remain model dependent, and we therefore proceed with the BayeSN result, as it can be applied to all three SN images.

5. Simulations

We simulated variations of the measured color curves both to test the robustness of our fitting procedure and to estimate the statistical and systematic uncertainties of the time-delay and magnification measurements. We created two distinct sets of simulations (photometric variability, Section 5.1; SED variability, Section 5.2) that test different aspects of the fitting uncertainties, described below. We then varied the second set further by introducing the effects of chromatic microlensing (Section 5.2.2). These simulations serve to illuminate the relative sources of uncertainty for this analysis, and then we create a single final set of simulations that include all of the effects discussed in Sections 5.1–5.2.2 that define our combined uncertainties (Section 5.3). These are added in quadrature with three external (not related to the SED or light curve) sources of systematic uncertainty, namely millilensing (Section 5.4), the intrinsic scatter of SNe Ia, and the SN Ia m_B population interpolation uncertainty (the last two discussed in Section 4.2). The results of this section are summarized in Table 5.

5.1. Random Photometric Realizations

We created mock observations drawn from a normal distribution centered at each measured flux, with σ corresponding to the measured photometric uncertainty (Table 2). We created simulated color curves and generated 100 realizations of each SN image. We ran the identical SNTD Color method and magnification measurement used in the previous section on each of the realizations, and Table 5 shows the results. Due to the relatively large systematic uncertainties in the photometry (Table 1), varying the photometry has quite a large impact on the time-delay measurements. The effect is most significant for image 2a, where the bulk of the time-delay inference comes from the F277W photometry, which has the largest uncertainties. The template imaging planned for SN H0pe in JWST Cycle 3 will both improve the photometry in F277W and add

the F356W/F444W filters, likely reducing this uncertainty significantly.

5.2. SED-based Simulations

5.2.1. Baseline Simulations

Mock color curves were also created using the BayeSN model to explore model- and/or wavelength-dependent systematics. These simulations include realizations of residual intrinsic light-curve variation (beyond the principal effect of light-curve shape θ) drawn from BayeSN’s learned distribution of intrinsic residuals (ϵ ; see Mandel et al. 2022, Section 5.2.2 for details). In this way, our simulations should capture the intrinsic diversity of SN Ia color curves. The simulations include a range of values for the BayeSN light-curve shape (θ) and dust extinction (A_V) values based on independent constraints from the NIRSpect data (Chen et al. 2024). The spectra show SN H0pe to be a normal SN Ia, and therefore we simulated SNe Ia over the reasonable population values for θ and A_V of $-2 \leq \theta \leq 2$ and $0 \leq A_V \leq 1.5$ given $R_V = 3.1$ (Mandel et al. 2022). This is the best-fit value of R_V from our light-curve fitting, and we fix it here, as it is highly covariant with A_V and has little impact on measured time delays or magnifications (Figure 5). We then simulated 100 color curves with cadence and photometric precision matching those described above but with random offsets to the relative time delays and magnifications between images. Again, the results of running the identical SNTD Color method and magnification measurement used in Section 4 on each of the realizations are summarized in Table 5. The results from fitting this set of simulations, which essentially includes both the model uncertainty and intrinsic variability of SNe Ia as a function of wavelength, are the largest source of uncertainty for both time delays and magnifications. The most promising means of improving this uncertainty would be to add the F356W and F444W filters, which provide additional light-curve information particularly at late times, and to continue improving SN Ia SED models in the NIR and at late times (beyond 50 rest-frame days after peak brightness).

5.2.2. The Impacts of Microlensing

Microlensing from small perturbers such as stars in the lens plane is well-known to potentially impact time-delay measurements (e.g., Dobler & Keeton 2006; Goldstein et al. 2018; Huber et al. 2019; Pierel & Rodney 2019; Huber et al. 2022). The effects of microlensing are not generally achromatic because the specific intensity profiles for SNe vary for different filters. Therefore, the expanding SN shell interacts with the microlensing caustics, and the resulting magnification is chromatic (Foxley-Marrable et al. 2018; Goldstein et al. 2018; Huber et al. 2019). For SNe Ia, microlensing is essentially achromatic in the first ~ 3 rest-frame weeks after explosion (i.e., the “achromatic phase”; Goldstein et al. 2018; Huber et al. 2021). The method of measuring time delays with color curves is therefore valuable because it removes the uncertainty of macrolensing and achromatic microlensing from the fit. The observations of (at least) image 2b of SN H0pe are likely in the achromatic phase, but it is plausible that the remaining observations are impacted by chromatic microlensing.

We produced a magnification map for each SN image and convolved the SN Ia light profiles from four theoretical models

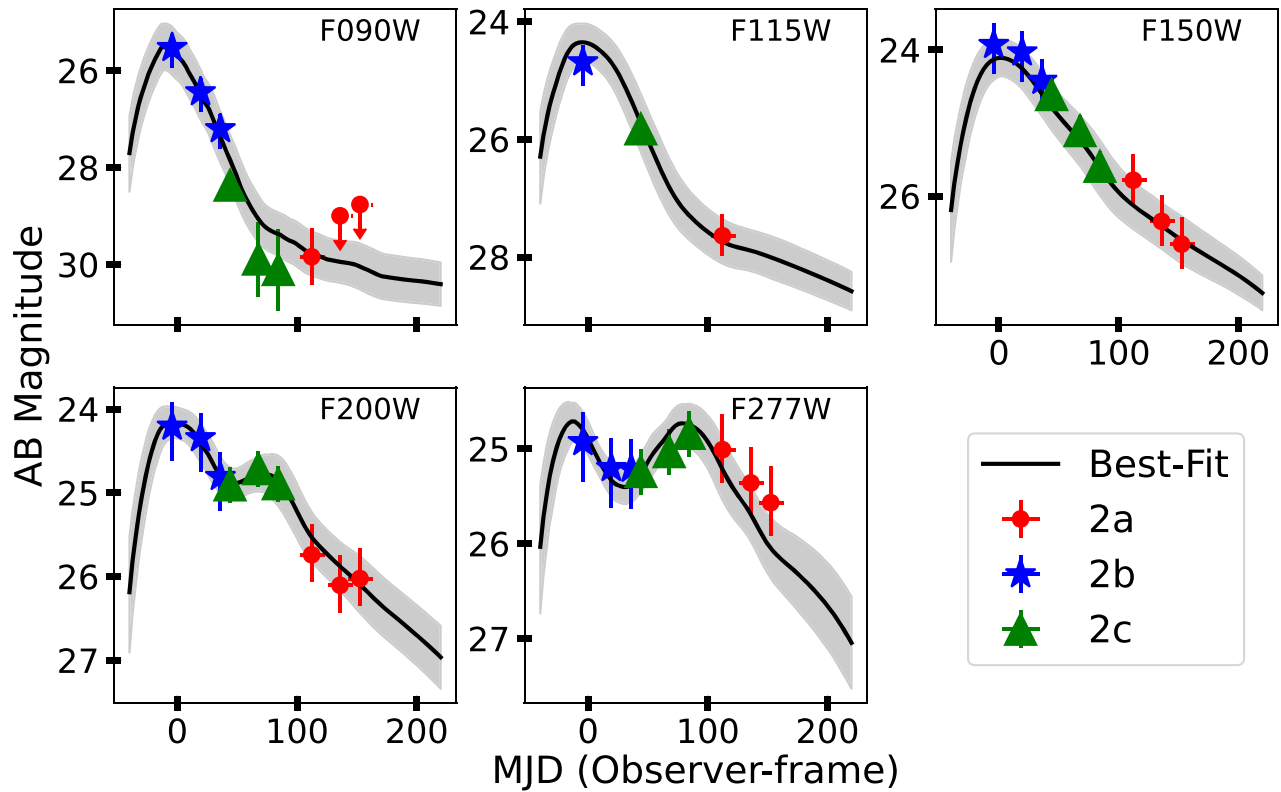


Figure 7. Reconstructed light curve for SN H0pe after applying the best-fit time delays and magnifications relative to image 2b. The vertical error bars are the photometric precision based on the work in Section 3 combined with an additional magnification uncertainty from Section 5, while the horizontal error bars are the 16th and 84th percentiles for the time-delay posteriors from Figure 5. The gray shaded region is the best-fit model from the SNTD Color method scaled to image 2b.

Table 5
The Variation in Measured Time Delay and Magnification of Each Simulation Scenario from Section 5, Relative to the Measurements of Section 4

Sim Type (Section)	Image A		Image B		Image C	
	δt (days)	$\delta\mu$	δt (days)	$\delta\mu$	δt (days)	$\delta\mu$
Photometric (5.1)	$0.55^{+8.74}_{-7.13}$	$-0.36^{+1.07}_{-1.36}$...	$-0.26^{+1.05}_{-0.86}$	$-0.57^{+2.76}_{-3.94}$	$-0.22^{+0.79}_{-0.70}$
SED (5.2.1)	$-1.24^{+4.80}_{-4.95}$	$0.08^{+0.44}_{-1.61}$...	$-0.47^{+1.26}_{-0.59}$	$-0.45^{+1.91}_{-2.11}$	$-0.01^{+0.70}_{-1.36}$
Microlensing (5.2.2)	$-0.20^{+0.95}_{-0.82}$	$0.22^{+0.19}_{-0.20}$...	$-0.63^{+2.51}_{-1.49}$	$-0.04^{+0.15}_{-0.16}$	$0.30^{+0.20}_{-0.47}$
Combined (5.3)	$-0.97^{+10.83}_{-9.29}$	$0.03^{+1.38}_{-1.61}$...	$-0.32^{+3.36}_{-2.21}$	$-0.67^{+3.58}_{-3.97}$	$0.28^{+0.95}_{-1.07}$
Millilensing (5.4)	...	$0.06^{+0.43}_{-0.28}$...	$-0.12^{+0.52}_{-0.65}$...	$0.12^{+0.74}_{-0.43}$
Intrinsic SN Ia Scatter (4.2)	...	$0.00^{+0.40}_{-0.40}$...	$0.00^{+0.76}_{-0.76}$...	$0.00^{+0.59}_{-0.59}$
m_B Interpolation (4.2)	...	$0.00^{+0.56}_{-0.56}$...	$0.00^{+0.99}_{-0.83}$...	$0.00^{+0.83}_{-0.83}$
Final Uncertainty	$-0.97^{+10.83}_{-9.29}$	$0.09^{+1.60}_{-1.77}$...	$-0.44^{+3.61}_{-2.61}$	$-0.67^{+3.58}_{-3.97}$	$0.40^{+1.57}_{-1.54}$

(Suyu et al. 2020; Huber et al. 2021) with each magnification map (1000 random positions per model for a total of 4000 convolutions per map) in the manner of Huber et al. (2019). The parameters used to generate the magnification maps are the lens surface mass density scaled to the critical value (κ), the image shear (γ), and the smooth matter fraction (s). Here κ and γ at each SN image position are the median of six independent lens model predictions from M. Pascale et al. (2024, in preparation), and we assume a value of 0.99 for s across all maps, consistent with Kelly et al. (2023b). Table 6 shows the κ and γ values used for each SN image. We selected 40 random chromatic microlensing curves for baseline simulation from the previous section and varied the flux in each band by the phase- and wavelength-dependent magnification curves. Our total number of microlensed simulations was therefore 4000.

Table 6
Parameters Used to Create Microlensing Magnification Maps for Each SN Image

Image	κ	γ
2a	0.563	0.193
2b	0.699	0.445
2c	0.576	0.246

Finally, we ran the SNTD Color method and magnification measurement used in Section 4 one more time on each of the realizations. Table 5 summarizes the results.

For completeness, we checked for the level of chromatic microlensing effects that could be impacting our observed

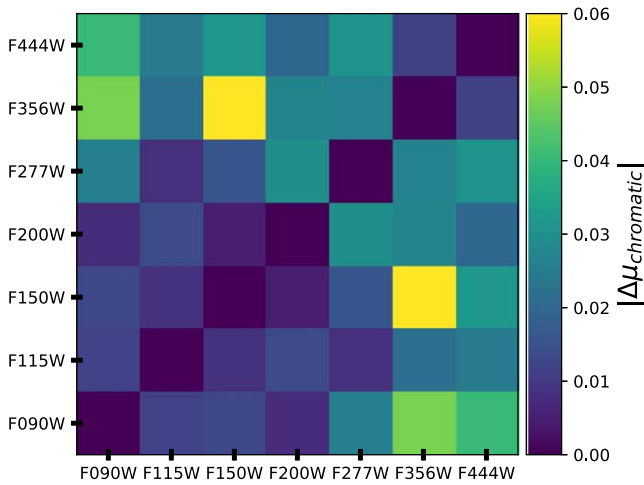


Figure 8. The level of chromatic microlensing expected from our simulations in the case of SN H0pe (across all three SN images) showing all the possible color combinations. The color bar corresponds to the shift in magnitudes. For example, the cell associated with F150W and F200W corresponds to the F150W–F200W color.

photometry by examining the maximum chromatic deviation due to microlensing over the course of the full light curve for all filter combinations. The results are summarized in Figure 8, which suggests a maximum of ~ 0.06 mag of chromatic microlensing. The impact overall is relatively low, reaching a maximum of $\gtrsim 0.05$ mag in the F356W and F444W filters, and is generally < 0.02 mag in the filters used in this analysis. This is consistent with the small additional uncertainty due to microlensing found by fitting the simulated light curves. If a future time-delay analysis includes the F356W and F444W filters, particular care should be taken with respect to chromatic microlensing, but the impacts are small for the present analysis. This is consistent with the results of fitting the simulations, which found the extra uncertainty due to microlensing to be quite small compared to the other sources of uncertainty explored here.

5.3. Combined Simulation

We produce one final set of 1000 simulated versions of SN H0pe, which include all the sources of variability discussed above, ensuring that the full covariance of these uncertainties is correctly included. Our final uncertainties are the result of fitting these simulations, which contain all sources of uncertainty we are able to simulate. We combine this with an additional systematic uncertainty from millilensing, discussed in the following section, along with systematics of 0.1 and 0.03 mag due to the intrinsic scatter of SNe Ia and the uncertainty in our m_B interpolation (see Section 4.2), with the results shown in Table 5. The final row in Table 5 includes systematic shifts, which were used to correct each actual measurement from Sections 4.1 and 4.2.

5.4. The Impacts of Millilensing

Kelly et al. (2023b) found that millilensing, which is additional lensing caused by dark matter halos (referring to subhalos associated with the cluster and halos along the line of sight), caused an additional $\sim 10\%$ uncertainty in the measured magnifications for SN Refsdal. The impact of millilensing is achromatic and therefore will not impact our time-delay

measurements but could plausibly impact the image magnifications at a similar level. Using the ray-tracing technique presented by Gilman et al. (2019), we computed the expected millilensing due to dark matter cluster subhalos and line-of-sight halos for a projected mass density in substructure bracketing the range of expected values based on theoretical and observational arguments (Gilman et al. 2020). This corresponds to a 1%–10% fraction of mass in dark matter subhalos with masses in the range 10^6 – $10^{10} M_\odot$. Halos more massive than $10^{10} M_\odot$ would likely contain a visible counterpart, and halos less massive than $10^6 M_\odot$ are too small to significantly alter the image magnifications. We modeled the line-of-sight structure using a Sheth–Tormen halo mass function (Sheth et al. 2001), placed uniformly (in 2D) around each SN image and along a cylindrical line-of-sight volume with diameter equal to the comoving distance to the cluster times $1''$. We report the additional uncertainty in magnification due to millilensing in Table 5, where we assume the median scenario of a 5% dark matter contribution from subhalos, finding that magnification due to millilensing ranges from a $\sim 4\%$ to $\sim 8\%$ effect among the three SN images.

6. Conclusions

SN H0pe is the first multiply imaged SN Ia and second multiply imaged SN capable of delivering a competitive H_0 measurement. Time delays (relative to the last image to arrive, 2b) are $\Delta t_{ab} = -116.6^{+10.8}_{-9.3}$ observer-frame days and $\Delta t_{cb} = -48.6^{+3.6}_{-4.0}$ observer-frame days. Leveraging the fact that SN H0pe has a standardizable light curve, absolute magnification estimates are $\mu_a = 4.3^{+1.6}_{-1.8}$, $\mu_b = 7.6^{+3.6}_{-2.6}$, $\mu_c = 6.4^{+1.6}_{-1.5}$. These time-delay measurements will be included in an upcoming inference of H_0 (M. Pascale et al. 2024, in preparation). While the inferred value of H_0 is not shown here, if we assume a fiducial result of $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, then the uncertainty from this time-delay measurement alone would be $\sim 3.9 \text{ km s}^{-1} \text{ Mpc}^{-1}$, which is then added in quadrature to the lens model uncertainty.

While we are unable to fully explore possible systematic uncertainties introduced by using a specific model to analyze SN H0pe, we attempted to mitigate this by using simulations to estimate the uncertainties, including a model of intrinsic SN Ia variability in the generative model to account for possible shortcomings in the model. We modeled the possible magnification uncertainty due to millilensing from dark matter subhalos, which does not impact time-delay measurements but contributes an additional magnification uncertainty of $\delta\mu_a = -0.06^{+0.43}_{-0.28}$, $\delta\mu_b = 0.12^{+0.52}_{-0.65}$, $\delta\mu_c = -0.12^{+0.74}_{-0.43}$. We also considered the possible impacts of chromatic microlensing, which are relatively small ($\delta t_{ab} = -0.20^{+0.95}_{-0.82}$, $\delta t_{cb} = -0.04^{+0.15}_{-0.16}$ days and $\delta\mu_a = 0.22^{+0.19}_{-0.20}$, $\delta\mu_b = -0.63^{+2.51}_{-1.49}$, $\delta\mu_c = 0.30^{+0.20}_{-0.47}$), as the SN is well offset from cluster-member galaxies, and the observations of image 2b are in the achromatic phase of SN Ia microlensing variability, which limits the impact for time-delay measurements (Goldstein et al. 2018; Huber et al. 2021). The exception is the magnification measurement for image 2b, which, based on our simulations, may be significantly impacted by microlensing, as it has higher values for both κ and γ relative to images 2a and 2c.

The most significant sources of uncertainty in this analysis are the photometry ($\delta t_{ab} = 0.55^{+8.74}_{-7.13}$, $\delta t_{cb} = -0.57^{+2.76}_{-3.94}$ days and $\delta\mu_a = -0.36^{+1.07}_{-1.36}$, $\delta\mu_b = -0.26^{+1.05}_{-0.86}$, $\delta\mu_c = -0.22^{+0.79}_{-0.70}$),

due to the brightness of Arc 2 (particularly at longer wavelengths), and uncertainties from the SN Ia SED model BayeSN ($\delta t_{ab} = -1.24^{+4.80}_{-4.95}$, $\delta t_{cb} = -0.45^{+1.91}_{-2.11}$ days and $\delta \mu_a = 0.08^{+0.44}_{-1.61}$, $\delta \mu_b = -0.47^{+1.26}_{-0.59}$, $\delta \mu_c = -0.01^{+0.70}_{-1.36}$). The latter issue is likely to be resolved with time, as more NIR data are collected from (in particular) the Nancy Grace Roman Space Telescope. While we implemented a method to remove host galaxy light using a novel lens modeling technique, large residuals remained at the longest wavelengths, and we were unable to detect the SN despite predicted brightnesses that were above our detection threshold. A set of JWST/NIRCam template images in the same filters is planned for Cycle 3 (PID 4744), which will greatly improve the LW channel photometry and thereby the time-delay measurement precision. For now, we used planted PSF models to estimate the photometric precision on our final SN H0pe photometry.



Upcoming surveys such as the Vera C. Rubin Observatory Legacy Survey of Space and Time (LSST; Ivezić et al. 2019) and the Nancy Grace Roman Space Telescope High Latitude Time Domain Survey (HLTDS; Pierel et al. 2021; Rose et al. 2021) are expected to deliver dozens of useful galaxy-scale multiply imaged SNe, but the number of expected cluster-lensed SNe similar to SN H0pe and SN Refsdal is uncertain. Roman in particular, with its $\sim 2000 \text{ deg}^2$ High Latitude Survey (HLS) and $\sim 20 \text{ deg}^2$ HLTDS, will observe $\sim 23,000$ massive galaxy clusters (Eifler et al. 2021). A subset of these clusters have already been accurately modeled using observations from programs such as the Reionization Lensing Cluster Survey (RELICS; Coe et al. 2019), Hubble Frontier Fields (HFF; Lotz et al. 2017), and JWST UNCOVER (Weaver et al. 2024). This allows an estimate of the multiply imaged SN yields in the area strongly lensed by the clusters (M. Bronikowski et al. 2024, in preparation). While the Roman HLTDS observing fields have not yet been selected, it is important to explore which of the massive clusters provides the optimal prospects for discovering strongly lensed SNe in order to maximize the potential for a precise measurement of H_0 , facilitated by the fact that these systems offer long time delays and wide image separations. As a result, cluster-lensed SNe are likely to continue as an extremely valuable subset of the future cosmological sample.

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References

- Birrer, S., Dhawan, S., & Shajib, A. J. 2022a, *ApJ*, 924, 2
 Birrer, S., Millon, M., Sluse, D., et al. 2022b, arXiv:2210.10833
 Birrer, S., Treu, T., Rusu, C. E., et al. 2019, *MNRAS*, 484, 4726
 Bonvin, V., Chan, J. H. H., Millon, M., et al. 2018, *A&A*, 616, A183
 Bonvin, V., Courbin, F., Suyu, S. H., et al. 2017, *MNRAS*, 465, 4914
 Bonvin, V., Millon, M., Chan, J. H.-H., et al. 2019a, *A&A*, 629, A97
 Bonvin, V., Tihhonova, O., Millon, M., et al. 2019b, *A&A*, 621, A55
 Brout, D., Scolnic, D., Popovic, B., et al. 2022, *ApJ*, 938, 110
 Bushouse, H., Eisenhamer, J., Dencheva, N., et al. 2022, JWST Calibration Pipeline, v1.12.5, Zenodo, doi:10.5281/zenodo.7325378
 Cañameras, R., Nesvadba, N. P. H., Guery, D., et al. 2015, *A&A*, 581, A105
 Cañameras, R., Nesvadba, N. P. H., Limousin, M., et al. 2018, *A&A*, 620, A60
 Chen, W., Kelly, P., Castellano, M., et al. 2022, Transient Name Server Discovery Report, 2022-3517
 Chen, W., Kelly, P. L., Frye, B. L., et al. 2024, arXiv:2403.19029
 Coe, D., Salmon, B., Bradač, M., et al. 2019, *ApJ*, 884, 85
 Dhawan, S., Johansson, J., Goobar, A., et al. 2020, *MNRAS*, 491, 2639
 Ding, X., Liao, K., Birrer, S., et al. 2021, *MNRAS*, 504, 5621
 Dobler, G., & Keeton, C. R. 2006, *ApJ*, 653, 1391
 Eifler, T., Miyatake, H., Krause, E., et al. 2021, *MNRAS*, 507, 1746
 Falco, E. E., Gorenstein, M. V., & Shapiro, I. I. 1985, *ApJL*, 289, L1
 Fitzpatrick, E. 1999, *PASP*, 111, 63
 Foxley-Marrable, M., Collett, T. E., Vernardos, G., Goldstein, D. A., & Bacon, D. 2018, *MNRAS*, 478, 5081
 Frye, B., Pascale, M., Cohen, S., et al. 2023, TNSAN, 96, 1
 Frye, B. L., Pascale, M., Pierel, J., et al. 2024, *ApJ*, 961, 171
 Frye, B. L., Pascale, M., Qin, Y., et al. 2019, *ApJ*, 871, 51
 Garnavich, P. M., Jha, S., Challis, P., et al. 1998, *ApJ*, 509, 74
 Gilman, D., Birrer, S., Nierenberg, A., et al. 2020, *MNRAS*, 491, 6077
 Gilman, D., Birrer, S., Treu, T., Nierenberg, A., & Benson, A. 2019, *MNRAS*, 487, 5721
 Goldstein, D. A., Nugent, P. E., Kasen, D. N., & Collett, T. E. 2018, *ApJ*, 855, 22
 Goobar, A., Amanullah, R., Kulkarni, S. R., et al. 2017, *Sci*, 356, 291
 Goobar, A., Johansson, J., Schulze, S., et al. 2023, *NatAs*, 7, 1098
 Grayling, M., Thorp, S., Mandel, K. S., et al. 2024, arXiv:2401.08755
 Grillo, C., Rosati, P., Suyu, S. H., et al. 2018, *ApJ*, 860, 94
 Grillo, C., Rosati, P., Suyu, S. H., et al. 2020, *ApJ*, 898, 87
 Guy, J., Sullivan, M., Conley, A., et al. 2010, *A&A*, 523, A7
 Harrington, K. C., Yun, M. S., Cybulski, R., et al. 2016, *MNRAS*, 458, 4383
 Holz, D. E. 2001, *ApJL*, 556, L71
 Hsiao, E. Y., Conley, A., Howell, D. A., et al. 2007, *ApJ*, 663, 1187
 Huber, S., Suyu, S. H., Ghoshdastidar, D., et al. 2022, *A&A*, 658, A157
 Huber, S., Suyu, S. H., Noebauer, U. M., et al. 2019, *A&A*, 631, A161
 Huber, S., Suyu, S. H., Noebauer, U. M., et al. 2021, *A&A*, 646, A110
 Ivezić, Z., Kahn, S. M., Tyson, J. A., et al. 2019, *ApJ*, 873, 111
 Kelly, P., Zitrin, A., Oguri, M., et al. 2022, TNSAN, 169, 1
 Kelly, P. L., Rodney, S., Treu, T., et al. 2023a, *Sci*, 380, abh1322
 Kelly, P. L., Rodney, S., Treu, T., et al. 2023b, *ApJ*, 948, 93
 Kelly, P. L., Rodney, S. A., Treu, T., et al. 2015, *Sci*, 347, 1123
 Kenworthy, W. D., Jones, D. O., Dai, M., et al. 2021, *ApJ*, 923, 265
 Kolatt, T. S., & Bartelmann, M. 1998, *MNRAS*, 296, 763
 Leget, P.-F., Gangler, E., Mondon, F., et al. 2020, *A&A*, 636, A46
 Linder, E. V. 2011, *PhRvD*, 84, 123529
 Lotz, J. M., Koekemoer, A., Coe, D., et al. 2017, *ApJ*, 837, 97
 Mandel, K. S., Thorp, S., Narayan, G., Friedman, A. S., & Avelino, A. 2022, *MNRAS*, 510, 3939
 Narayan, R., & Bartelmann, M. 1997, arXiv:astro-ph/9606001
 Nordin, J., Rubin, D., Richard, J., et al. 2014, *MNRAS*, 440, 2742
 Oguri, M. 2019, *RPPH*, 82, 126901
 Oguri, M., & Kawano, Y. 2003, *MNRAS*, 338, L25
 Paraficz, D., & Hjorth, J. 2009, *A&A*, 507, L49
 Pascale, M., Frye, B. L., Dai, L., et al. 2022, *ApJ*, 932, 85
 Patel, B., McCully, C., Jha, S. W., et al. 2014, *ApJ*, 786, 9
 Perlmutter, S., Aldering, G., Goldhaber, G., et al. 1999, *ApJ*, 517, 565
 Phillips, M. M. 1993, *ApJS*, 413, L105
 Pierel, J. D. R., Arendse, N., Ertl, S., et al. 2023, *ApJ*, 948, 115
 Pierel, J. D. R., Jones, D. O., Kenworthy, W. D., et al. 2022, *ApJ*, 939, 11
 Pierel, J. D. R., & Rodney, S. 2019, *ApJ*, 876, 107
 Pierel, J. D. R., Rodney, S., Vernardos, G., et al. 2021, *ApJ*, 908, 190
 Polletta, M., Nonino, M., Frye, B., et al. 2023, *A&A*, 675, L4
 Refsdal, S. 1964, *MNRAS*, 128, 307
 Riess, A. G., Filippenko, A. V., Challis, P., et al. 1998, *AJ*, 116, 1009
 Riess, A. G., Yuan, W., Macri, L. M., et al. 2022, *ApJL*, 934, L7
 Rigby, J., Perrin, M., McElwain, M., et al. 2023, *PASP*, 135, 048001
 Rodney, S. A., Brammer, G. B., Pierel, J. D. R., et al. 2021, *NatAs*, 5, 1118
 Rodney, S. A., Patel, B., Scolnic, D., et al. 2015, *ApJ*, 811, 70
 Rose, B. M., Baltay, C., Hounsell, R., et al. 2021, arXiv:2111.03081
 Saunders, C., Aldering, G., Antilogus, P., et al. 2018, *ApJ*, 869, 167
 Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, 737, 103
 Scolnic, D. M., Jones, D. O., Rest, A., et al. 2018, *ApJ*, 859, 101
 Sheth, R. K., Mo, H. J., & Tormen, G. 2001, *MNRAS*, 323, 1
 Suyu, S. H., Goobar, A., Collett, T., More, A., & Vernardos, G. 2024, *SSRv*, 220, 13
 Suyu, S. H., & Halkola, A. 2010, *A&A*, 524, A94
 Suyu, S. H., Hensel, S. W., McKean, J. P., et al. 2012, *ApJ*, 750, 10

- Suyu, S. H., Huber, S., Cañameras, R., et al. 2020, [A&A](#), **644**, A162
- Suyu, S. H., Marshall, P. J., Auger, M. W., et al. 2010, [ApJ](#), **711**, 201
- Suyu, S. H., Marshall, P. J., Hobson, M. P., & Blandford, R. D. 2006, [MNRAS](#), **371**, 983
- Tewes, M., Courbin, F., Meylan, G., et al. 2013, [A&A](#), **556**, A22
- Tie, S. S., & Kochanek, C. S. 2018, [MNRAS](#), **473**, 80
- Treu, T., & Marshall, P. J. 2016, [A&ARv](#), **24**, 11
- Treu, T., Suyu, S. H., & Marshall, P. J. 2022, [A&ARv](#), **30**, 8
- Vuissoz, C., Courbin, F., Sluse, D., et al. 2008, [A&A](#), **488**, 481
- Ward, S. M., Thorp, S., Mandel, K. S., et al. 2023, [ApJ](#), **956**, 111
- Weaver, J. R., Cutler, S. E., Pan, R., et al. 2024, [ApJS](#), **270**, 7
- Windhorst, R. A., Cohen, S. H., Jansen, R. A., et al. 2023, [AJ](#), **165**, 13
- Wong, K. C., Suyu, S. H., Chen, G. C. F., et al. 2020, [MNRAS](#), **498**, 1420
- Xu, D., Sluse, D., Schneider, P., et al. 2016, [MNRAS](#), **456**, 739
- Yahalom, D. A., Schechter, P. L., & Wambsganss, J. 2017, arXiv:1711.07919