

PAPER

JUNO sensitivity to the annihilation of MeV dark matter in the galactic halo

To cite this article: Angel Abusleme *et al* JCAP09(2023)001

View the [article online](#) for updates and enhancements.

You may also like

- [Exploring the \$^{10}\text{Li}\$ structure by the \$d\(^7\text{Li},p\)^{10}\text{Li}\$ transfer reaction](#)
M Cavallaro, M De Napoli, F Cappuzzello et al.
- [Reaction mechanisms in the scattering of exotic nuclei](#)
Angela Bonaccorso
- [Unified description of structure and reactions: implementing the nuclear field theory program](#)
R A Broglia, P F Bortignon, F Barranco et al.

JUNO sensitivity to the annihilation of MeV dark matter in the galactic halo



The JUNO collaboration

E-mail: Juno_pub_comm@juno.ihep.ac.cn

Received June 21, 2023

Accepted August 12, 2023

Published September 1, 2023

Abstract. We discuss JUNO sensitivity to the annihilation of MeV dark matter in the galactic halo via detecting inverse beta decay reactions of electron anti-neutrinos resulting from the annihilation. We study possible backgrounds to the signature, including the reactor neutrinos, diffuse supernova neutrino background, charged- and neutral-current interactions of atmospheric neutrinos, backgrounds from muon-induced fast neutrons and cosmogenic isotopes. A fiducial volume cut, as well as the pulse shape discrimination and the muon veto are applied to suppress the above backgrounds. It is shown that JUNO sensitivity to the thermally averaged dark matter annihilation rate in 10 years of exposure would be significantly better than the present-day best limit set by Super-Kamiokande and would be comparable to that expected by Hyper-Kamiokande.

Keywords: neutrino astronomy, neutrino detectors, neutrino experiments

ArXiv ePrint: [2306.09567](https://arxiv.org/abs/2306.09567)

Contents

1	Introduction	1
2	JUNO detector	2
3	Neutrino signature from DM annihilation in the Milky Way	3
4	Backgrounds	7
4.1	Reactor $\bar{\nu}_e$	7
4.2	DSNB	7
4.3	Atmospheric ν charged current background	8
4.4	Cosmogenic isotopes	9
4.5	Muon induced fast neutrons	9
4.6	Atmospheric ν neutral current background	10
5	Pulse Shape Discrimination	11
6	Sensitivity	14
6.1	Total spectrum in JUNO	14
6.2	Two approaches to the sensitivity	14
7	Conclusion	19
	The JUNO collaboration	24

1 Introduction

The existence of non-baryonic Dark Matter (DM) in the Universe has been well established by astronomical observations. For most spiral galaxies, the rotation curve of stars or gases far from the galactic center does not decline with increasing distance but rather stays as a constant. This strongly indicates the existence of a massive dark halo which contains the galactic disk and extends well beyond the size of the visible part of the galaxy [1]. One promising DM candidate is the Weakly Interacting Massive Particle (WIMP) [2], which predicts a correct relic density based on the weak interaction annihilation cross section. In the WIMP scenario, the thermally averaged self-annihilation rate $\langle\sigma v\rangle$ of DM is predicted to be $3 \times 10^{-26} \text{cm}^3 \text{s}^{-1}$ regardless of the annihilation channel.

Quarks and leptons could be produced through DM annihilation and therefore provide foundations for numerous DM indirect search experiments [3–7]. In this paper, we focus on the scenario in which neutrinos are produced directly from DM annihilation, $\chi\chi \rightarrow \nu\bar{\nu}$. For DM candidates lighter than the muon, this is the only neutrino production channel. For heavier DM candidates, neutrinos can also arise from decays of hadronic or leptonic final states produced by the annihilation.

Searches for neutrinos originating from DM annihilation have been proposed for astrophysical neutrino observatories [8], accelerator neutrino experiments [9, 10] and solar neutrino measurements [11]. Several experiments including KamLAND [12, 13], Super-Kamiokande (SuperK) [14], IceCube [15] and ANTARES [16] have searched for neutrino signatures from

DM annihilation. Among them, KamLAND [13] obtained the updated 90% confidence level upper limit $\langle\sigma v\rangle = (1\text{--}11)\times 10^{-26}\text{cm}^3\text{s}^{-1}$ for DM mass in the range of (9–21) MeV with a nominal angular-averaged intensity of the galactic DM profile. More stringent limits on $\langle\sigma v\rangle$ were deduced from the data of SuperK [17, 18], while IceCube high-energy cosmic neutrino data can be used to perform the imaging of galactic DM [19]. Finally, the expected sensitivity to $\langle\sigma v\rangle$ is also discussed for Hyper- Kamiokande (HyperK) [20].

JUNO [21, 22], equipped with a central acrylic sphere containing 20 kt of liquid scintillator (LS), will significantly improve the sensitivity to $\langle\sigma v\rangle$. In this paper, we discuss the JUNO sensitivity to the neutrino flux from DM annihilation $\chi\chi \rightarrow \nu\bar{\nu}$ in the galactic halo [23] where final state neutrinos are monochromatic i.e. $E_\nu = m_\chi$. We specifically focus on the DM mass range of (15–100) MeV where the lower mass limit is set to avoid the reactor neutrino background while the upper mass limit allows us to consider only the direct annihilation channel for neutrino productions. We note that JUNO is also capable of measuring more energetic neutrino events, such as those arising from heavier DMs through more complicated neutrino production mechanisms or atmospheric neutrino events in the energy range between 0.1 GeV and 10 GeV [24].

We shall first evaluate the inverse beta decay (IBD) signature in JUNO arising from DM annihilation $\chi\chi \rightarrow \nu\bar{\nu}$ in the galactic halo. Backgrounds dominating or comparable to the signal will be investigated hereafter. Among them, the events induced by atmospheric neutrinos interacting with ^{12}C nuclei through neutral-current (Atm- ν NC) are dominant. We perform pulse shape discrimination (PSD) to reject such backgrounds. Considering all backgrounds and the overall event selection efficiency, we find that the JUNO sensitivity to $\langle\sigma v\rangle$ can reach $\sim 1 \times 10^{-25}\text{cm}^3\text{s}^{-1}$ with a 10 years of exposure time, which will be competitive with any other existing or upcoming detectors in the near future.

The paper is organized as follows: In section 2 we give a brief introduction of the JUNO detector. In section 3 we present predictions on the DM annihilation rate in the galactic halo and the IBD event rate in the JUNO detector. The background studies are presented in section 4, with efficiencies of various veto methods evaluated. Section 5 focuses on PSD methods employed for separating IBD and non-IBD events. This is particularly useful for suppressing backgrounds from Atm- ν NC events mentioned before. In section 6 we present the JUNO sensitivity to $\langle\sigma v\rangle$ for the DM mass range of (15–100) MeV. We summarize and conclude in section 7.

2 JUNO detector

JUNO is a multi-purpose underground liquid scintillator (LS) detector that aims to decipher the neutrino mass ordering as the primary goal. The low muon rate, $0.004\text{s}^{-1}\text{m}^{-2}$, is due to the 700 m rock overburden that shields the detector from the flux of cosmic muons with 207 GeV of average energies, and therefore makes the detector suitable for searching exotic sources of neutrinos.

JUNO comprises central and veto detectors as illustrated in figure 1 [22]. The central detector is a spherically shaped acrylic shell with a 17.7 m radius filled with 20 kt of liquid scintillator. There are 17,612 20-inch and 25,600 3-inch PMTs mounted outside the acrylic ball, which provide around 78% photocathode coverage and excellent $3\%/\sqrt{E(\text{MeV})}$ energy resolution [25]. The veto detector is composed of a top tracker and a water Cherenkov detector. The top tracker [26] covers half of the top surface of the water pool as shown in figure 1. It is composed of 3-layer plastic scintillators originally used in the OPERA [27, 28] experiment

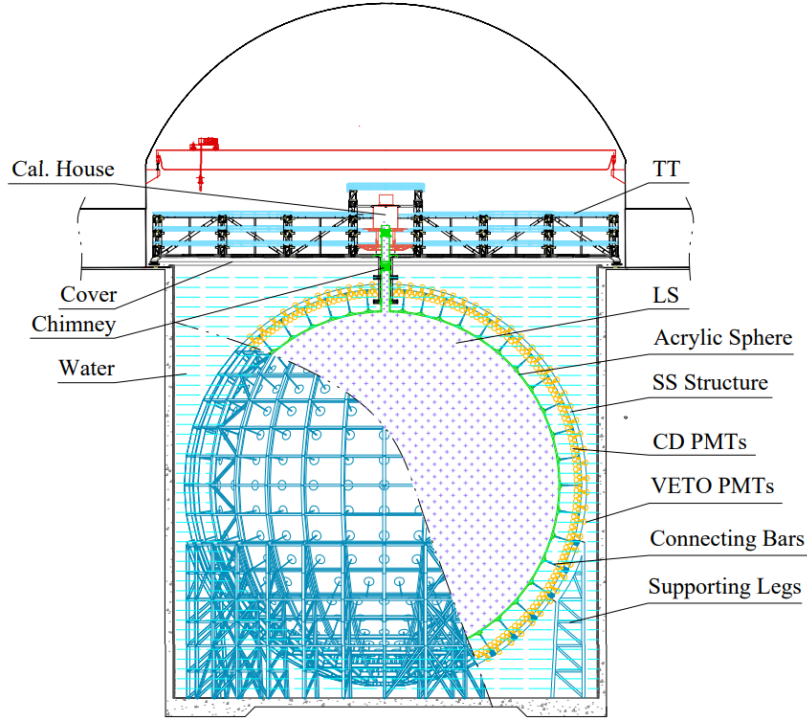


Figure 1. Scheme of the JUNO detector.

and re-utilized by JUNO. The water Cherenkov detector is a cylindrical water pool with 43.5 m in both diameter and height. It is filled with 30 kt of ultrapure water and maintained by a circulation system. There are 2400 20-inch PMTs mounted on the stainless steel frame to detect the Cherenkov light of cosmic muons. Together with veto and central detectors, the tracks of cosmic muons can be precisely reconstructed and therefore a high efficiency (99.5%) on muon tagging is achieved. The low muon rate and high muon tagging efficiency enable the detector to perform rare event searches.

3 Neutrino signature from DM annihilation in the Milky Way

The estimation of the DM-induced neutrino event rate is based on two assumptions. First, we assume that DM annihilates 100% into neutrino-antineutrino pairs, which sets an upper limit for the neutrino flux resulting from DM annihilation in the galactic halo. Second, we only consider the DM mass range of (15–100) MeV which meets the selection criteria we set for DM search in JUNO. The entire galactic halo is chosen as the target for maximizing the neutrino flux and at the same time minimizing the impact of profile uncertainties around the galactic center (GC).

The $\bar{\nu}_e$ neutrino flux spectrum resulting from DM annihilation is given by [23]

$$\frac{d\phi_{\bar{\nu}_e}^{\text{DM}}}{dE_{\bar{\nu}_e}} = \frac{1}{2} \langle \sigma v \rangle J_{\text{avg}} \frac{R_{\text{sc}} \rho_0^2}{m_\chi^2} \frac{1}{3} \delta(E_\nu - m_\chi), \quad (3.1)$$

where m_χ is the DM mass, $R_{\text{sc}} \sim 8.5 \text{ kpc}$ the distance between the GC and the solar system, and $\rho_0 \equiv \rho(R_{\text{sc}})$ the DM density in the local universe. The factor 1/2 comes from the

Halo profiles	α	β	γ	r_s [kpc]	$\rho(R_{sc})$ [GeV/cm ³]	J_{avg}
NFW [29]	1	3	1	20	0.3	3
MQGSL [31]	1.5	3	1.5	28	0.27	8
KKBP [30]	2	3	0.4	10	0.37	2.6
Canonical						5.0 [23]

Table 1. Summary of parameters for different DM halo profiles.

assumption that DM is a Majorana particle while the factor $1/3$ arises from the assumption that DM candidates annihilate into all three neutrino flavors with an equal probability. The canonical value for the thermally averaged DM annihilation rate, $\langle\sigma v\rangle$, is $3 \times 10^{-26} \text{cm}^3 \text{s}^{-1}$. In the case that DM candidates annihilate predominantly into one particular flavor, neutrino flavor transitions occurring between the production and detection points would generate the other flavors with comparable intensities. Since we focus on detecting the annihilation channel $\chi + \chi \rightarrow \nu + \bar{\nu}$, the neutrino energy is equal to the DM mass and this is implemented by the delta function $\delta(E_\nu - m_\chi)$.

The angular-averaged intensity, J_{avg} , is an integration over the square of the DM density along the line of sight and normalized by the square of the local DM density,

$$J_{\text{avg}} = \frac{1}{2R_{sc}\rho_0^2} \int_{-1}^1 d\cos\psi \int_0^{\ell_{\text{max}}} \rho^2(r(l, \psi)) dl, \quad (3.2)$$

where ρ is the DM density at the specific location described by the coordinate (l, ψ) with l the distance between the DM and the Earth while ψ the direction of the DM viewed from the Earth with $\psi = 0$ corresponding to the direction of GC. The distance between DM and GC is given by $r = \sqrt{R_{sc}^2 - 2\ell R_{sc} \cos\psi + \ell^2}$ while $\ell_{\text{max}} = \sqrt{r_s^2 - \sin^2\psi R_{sc}^2} + R_{sc} \cos\psi$ with r_s the radius of the galactic halo. The integration runs from $\psi = 0^\circ$ to $\psi = 180^\circ$ since it is challenging to precisely determine the neutrino direction with the JUNO detector at these energies. In the galactic halo, one assumes a spherically symmetric DM density profile with isotropic velocity dispersion. Hence the DM profile can be parametrized as:

$$\rho(r) = \rho_0 \left(\frac{R_{sc}}{r} \right)^\gamma \left[\frac{1 + (R_{sc}/r_s)^\alpha}{1 + (r/r_s)^\alpha} \right]^{(\beta-\gamma)/\alpha}. \quad (3.3)$$

For the variable set $(\alpha, \beta, \gamma, r_s)$, α determines the profile shape around r_s , β is the slope in the limit $r \rightarrow \infty$, γ is the inner cusp index, and r_s is the halo radius. In table 1, we summarize the parameter values corresponding to three commonly used profiles, the Navarro-Frenk-White (NFW) [29], Kravtsov [30] and Moore [31] profiles. We adopt the benchmark case of $J_{\text{avg}} = 5.0$ [23] for presenting our results.

Finally, the number of neutrino events in JUNO is given by

$$\frac{dN_S(E_{\bar{\nu}_e})}{dE_{\bar{\nu}_e}} = \sigma_{\text{IBD}}(E_{\bar{\nu}_e}) \cdot \frac{d\phi_{\bar{\nu}_e}^{\text{DM}}(E_{\bar{\nu}_e})}{dE_{\bar{\nu}_e}} \cdot N_{\text{target}} \cdot t \cdot \epsilon. \quad (3.4)$$

Here σ_{IBD} is the cross section for the IBD reaction, $\nu_e + p \rightarrow e^+ + n$, which leads to a prompt signal from positron-electron annihilation and a delayed signal from neutron capture. The value for $\sigma_{\text{IBD}}(E_{\bar{\nu}_e})$ is taken from [32]. The flux spectrum $d\phi_{\bar{\nu}_e}^{\text{DM}}/dE_{\bar{\nu}_e}$ is given by eq. (3.1). The number of free protons inside the JUNO central detector, N_{target} , is about $1.45 \cdot 10^{33}$ [21].

We set the total exposure time as 10 years. The parameter ϵ is the IBD detection efficiency optimized for the DM search, which is obtained from the official JUNO offline simulation and analysis frameworks. The simulation procedure includes the event generator, detector response simulation, electronics simulation, energy reconstruction, and vertex reconstruction. We apply the same simulation and analysis frameworks to evaluate the backgrounds. The final efficiency, ϵ , is the product of IBD signal selection, muon veto and PSD cut efficiencies. We note that the cut efficiencies hereafter always represent the event survival probability after the cut, regardless of whether these events are signal or backgrounds.

The criteria for IBD signal selection are composed of the following:

- (1) cut on the time difference between prompt and delayed signals, $\Delta T < 1$ ms,
- (2) $1.8 \text{ MeV} < \text{deposited energy of the delayed signal } (E_d) < 2.6 \text{ MeV}$,
- (3) multiplicity cut condition, $N_{\text{mult}} = 1$,
- (4) cut on the root mean square of time residual profile,¹ $\sigma(T_{\text{res}}) < 77$ ns,
- (5) fiducial volume cut, $R < 16$ m,
- (6) prompt signal deposited energy (E_p) cut, and
- (7) prompt-delay distance cut.

The efficiencies for (1), (2), (3), and (4) are 99.6%, 98.2%, 99.9%, and 99.9%, respectively. The efficiency for (5) is 73.9% owing to the uniform distribution of DM events. The efficiency of (6) varies between 92.0% and 99.8% for the DM mass range of (15–100) MeV. The efficiency of (7) is maintained at around 99.5% by varying the distance cut parameter with m_χ . We summarize the efficiencies and the uncertainties in table 2. The muon veto cut efficiency is well studied in JUNO [22]. In the case that the muon is detected by all detectors except the scintillator detector or when its track is not successfully reconstructed, any event occurring within 0.2 seconds after the muon detection is discarded. On the other hand, if the muon is also detected by the central detector and its track is successfully reconstructed, only events within a 3 m radius cylindrical region around the muon track and occurring within 0.2 seconds after the muon are rejected.

We present in figure 2 the visible energy spectra of neutrino events induced by DM annihilation. The non-Gaussian appearance of the peaks in figure 2 can be attributed to the IBD kinematics and its differential cross section. In our interested energy range, the positron energy is related to its direction by $E_{e^+} \approx (E_\nu - 1.30 \text{ MeV}) \cdot (1 - E_\nu(1 - \cos \theta)/M)$ [33] with θ the positron angle relative to the neutrino direction and M the nucleon mass. The maximal energy of the positron occurs at $\theta = 0$ with $E_{e^+}^{\text{max}} = (E_\nu - 1.30 \text{ MeV})$ while the minimal energy

¹To separate atmospheric $\bar{\nu}_\mu$ charged-current ($\bar{\nu}_\mu$ CC) events from atmospheric $\bar{\nu}_e$ charged-current ($\bar{\nu}_e$ CC) events and IBD events of other sources (see section 4.3 for further details), we define the time residual T_{res} for each hit on the i -th 3-inch PMT as $T_{\text{res}}^i = t_{\text{hit}}^i - n \cdot R_V^i/c$ with t_{hit}^i the hit time on the i -th PMT, n the refraction index of JUNO liquid scintillator (LS) and R_V^i the distance between the reconstructed vertex position and the i -th PMT. The time residual profile of the scintillation light emitted by $\bar{\nu}_\mu$ CC events is different from that of general $\bar{\nu}_e$ CC events including IBD, since $\mu^\pm$ from the former takes a longer time to deposit its energy to LS than $e^\pm$ from the latter does. Hence the root mean square of the T_{res} distribution over the fired 3-inch PMTs, denoted as $\sigma(T_{\text{res}})$, is a useful parameter for event selections [24]. The cut $\sigma(T_{\text{res}}) < 77$ ns can effectively reject $\bar{\nu}_\mu$ CC events.

IBD signal selection - m_χ independent		cut efficiency	MC uncertainty
(1) ΔT cut,	$\Delta T < 1$ ms	$\sim 99.6\%$	1.0%
(2) E_d cut,	$1.8 \text{ MeV} < E_d < 2.6 \text{ MeV}$	98.2%	1.0%
(3) Multiplicity cut,	$N_{\text{mult}} = 1$	99.9 %	1.0%
(4) T_{res} cut,	$\sigma(T_{\text{res}}) < 77$ ns	99.9%	1.0%
(5) Fiducial volume cut,	$R_{\text{prompt}} < 16$ m	73.9 %	0.8%
IBD signal selection - m_χ dependent		cut efficiency, $15 \leq m_\chi/\text{MeV} \leq 100$	MC uncertainty
(6) E_p cut	$(0.75 \cdot m_\chi + 2) \text{ MeV} < E_p < (0.97 \cdot m_\chi - 0.06) \text{ MeV}$	$\sim (99 - 0.04 \cdot m_\chi/\text{MeV})\%$	0.8%
(7) ΔD cut	$0 \text{ mm} < \Delta D < (7.7 \cdot m_\chi/\text{MeV} + 226) \text{ mm}$	$\sim 99.5\%$	1.0%
Muon veto cut		$\sim 97.5\%$	1.0%
PSD cut [15–100] MeV		[30.4% - 99.9%]	5.0%
Total [15–100] MeV		[22.5% - 68.8%]	$\sim 4\%$

Table 2. Summary of IBD signal selection cuts, the muon veto and the PSD cut.

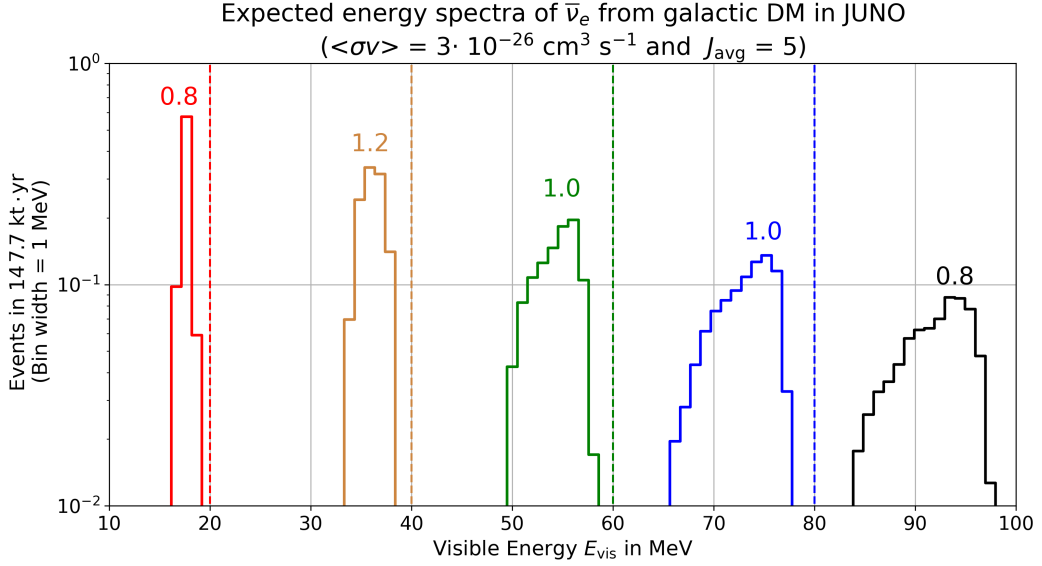


Figure 2. Visible energy spectra in the JUNO detector with 10 years of exposure time and 14.77 kt of fiducial mass after applying all except the PSD cuts for $J_{\text{avg}} = 5$ and $\langle\sigma v\rangle = 3 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}$. The distributions are shown for $m_\chi = 20$ MeV (red), 40 MeV (brown), 60 MeV (green), 80 MeV (blue) and 100 MeV (black). The total number of signal events is also displayed above the respective peaks.

corresponds to $\theta = \pi$ so that $E_{e^+}^{\text{min}} = (1 - 2E_\nu/M) \cdot E_{e^+}^{\text{max}}$. The relation between $E_{e^+}^{\text{max}}$ and $E_{e^+}^{\text{min}}$ explains why the DM event spectrum becomes broader as m_χ (E_ν) increases. Furthermore the positron average energy is given by $\langle E_{e^+} \rangle = (E_{e^+}^{\text{max}} + E_{e^+}^{\text{min}})/2 + E_{e^+}^{\text{max}} E_\nu \langle \cos \theta \rangle / M$ with $\langle \cos \theta \rangle \approx 2.4 \cdot (E_\nu - 13 \text{ MeV})/M$ [33]. Clearly, as m_χ (E_ν) increases, $\langle E_{e^+} \rangle$ moves farther away from $(E_{e^+}^{\text{max}} + E_{e^+}^{\text{min}})/2$. This is reflected by the increasingly asymmetrical shape of DM event spectrum as m_χ increases. In section 5, we shall discuss the PSD method used for further background suppression. The efficiency for the PSD cut will be evaluated.

Despite a low event number, the unprecedented energy resolution of JUNO provides a great advantage for observing monochromatic neutrino signatures from DM annihilation as we shall see in section 6.

4 Backgrounds

In this section, we discuss backgrounds to the indirect DM signature. We divide the backgrounds into two categories: IBD backgrounds and non-IBD backgrounds.

IBD backgrounds come from other $\bar{\nu}_e$ sources, which are intrinsically indistinguishable from $\bar{\nu}_e$ produced by DM annihilation. The reactor neutrinos, charged current interactions of atmospheric neutrinos (Atm- ν CC) and diffuse supernova neutrino background (DSNB) act as the neutrino floor for the indirect DM search in JUNO. We focus on the energy range beginning at 12 MeV since there are a large number of reactor neutrino events for energies less than this, which overwhelm the DM signature by several orders of magnitudes.

Non-IBD backgrounds mimic the IBD coincidence, including fast neutrons (FN) induced by muons passing through the surrounding rock, radionuclides (^{11}Li and ^{14}B) induced by muon spallation on carbon nuclei, and neutral current interactions of atmospheric neutrinos. These events can be suppressed by proper veto strategies with good efficiencies. Backgrounds from FN are reduced by the fiducial volume cut while the muon spallation backgrounds are reduced by the 12 MeV cut on E_p and the muon veto cut customized for JUNO. The neutral current interactions from atmospheric neutrinos are suppressed to an acceptable level with the PSD cut.

4.1 Reactor $\bar{\nu}_e$

The JUNO detector is located 52.5 km away from the eight reactors of the Taishan and Yangjiang nuclear power plants with a total thermal power of 26.6 GW. The $\bar{\nu}_e$ flux is generated through β decays of ^{235}U , ^{238}U , ^{239}Pu and ^{241}Pu . We follow [21] to simulate IBD rate and spectrum from reactor $\bar{\nu}_e$. The total number of IBD events from the reactors is estimated to be 223,736 after taking into account the oscillation effect for 147.7 kt · year exposure. The spectral shape of the reactor neutrino background is derived from Huber-Muller model [34, 35]. The event spectrum drops rapidly beyond 12 MeV, resulting in only a few events. To avoid background events from reactor neutrinos, we set the mass range for our DM search to be higher than 15 MeV.

4.2 DSNB

The DSNB is a cumulative neutrino emission generated from core-collapse supernovae with all flavors of neutrinos and anti-neutrinos in the observable universe. The isotropic DSNB flux is given by [36]:

$$\frac{d\phi_{\text{DSNB}}}{dE_\nu} = \int_0^{z_{\text{max}}} R_{\text{SN}}(z) \frac{dN_\nu(E'_\nu)}{dE'_\nu} (1+z) \cdot c \left| \frac{dt}{dz} \right| dz, \quad (4.1)$$

with z_{max} being the maximal redshift to be covered, and c the speed of light. The first term, $R_{\text{SN}}(z)$, is the core-collapse supernova (CCSN) rate, which is related to the star formation rate and the initial mass function of the forming stars. The second term, $dN_\nu(E'_\nu)/dE'_\nu$, is the averaged energy spectrum of the emitted neutrinos per supernova explosion. The energy E'_ν at the source is linked with E_ν observed on the Earth through the redshift relation, $E'_\nu = (1+z)E_\nu$. The last term represents the assumed cosmological model, which relates the redshift z to the cosmic time t according to $|dt/dz| = 1/(H_0(1+z)\sqrt{\Omega_m(1+z)^3 + \Omega_\Lambda})$ with Ω_m being the present-day density parameter of matter, Ω_Λ the fraction of the energy density provided by the dark energy, and H_0 the Hubble constant. We take the standard Λ CDM cosmology with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

flux model	f_{BH}	$\langle E_\nu \rangle$ [MeV]	$R_{\text{SN}}(z)$ [$\text{yr}^{-1}\text{Mpc}^{-3}$]	N_{evts} in (12–100) MeV [147.7 kt · yr]
low	0%	12	0.5×10^{-4}	4.2
nominal	27%	15	1.0×10^{-4}	2.3×10^1
high	40%	18	2.0×10^{-4}	6.6×10^1
SuperK				1.9×10^2

Table 3. Summary of DSNB flux models used in the background analysis.

The major uncertainty of the DSNB flux arises from the cosmological SN rate $R_{\text{SN}}(z)$ and the average energy spectrum of SN neutrinos $dN_\nu(E'_\nu)/dE'_\nu$. We adopt the nominal SN rate as $R_{\text{SN}}(0) = 1.0 \times 10^{-4} \text{Mpc}^{-3}\text{yr}^{-1}$ [37] while the range for the SN rate is taken to be $0.5 \times 10^{-4} \text{Mpc}^{-3}\text{yr}^{-1} \leq R_{\text{SN}}(0) \leq 2.0 \times 10^{-4} \text{Mpc}^{-3}\text{yr}^{-1}$. The average energy spectrum of SN neutrinos varies with astrophysical parameters such as the explodability of the progenitor, the maximum baryonic neutron star mass, and the cosmic CCSN rate. Therefore, we consider $12 \leq \langle E_\nu \rangle / \text{MeV} \leq 18$ for the average energy of SN neutrinos [38]. The above-mentioned parameters determine the fraction of failed SN, f_{BH} , which in turn determines $dN_\nu(E'_\nu)/dE'_\nu$. We adopt $f_{\text{BH}} = 27\%$ as the fiducial model [38] and the range for f_{BH} is taken as $0 \leq f_{\text{BH}} \leq 0.4$ in our study.

The IBD event spectrum induced by the DSNB is given by

$$\frac{dN_{\text{DSNB}}(E_{\bar{\nu}_e})}{dE_{\bar{\nu}_e}} = \sigma_{\text{IBD}}(E_{\bar{\nu}_e}) \cdot \frac{d\phi_{\text{DSNB}}}{dE_{\bar{\nu}_e}} \cdot N_{\text{target}} \cdot t \cdot \epsilon_{\text{DSNB}}, \quad (4.2)$$

where the total efficiency for the DSNB, ϵ_{DSNB} , is 47.3% based on IBD signal selection, muon veto and PSD applied for DM search. The low efficiency is largely due to the implementation of PSD, which will be discussed in section 5. Furthermore, we consider the upper bound of the DSNB flux obtained from the SuperK search [39] as the largest possible DSNB flux. As a result, we obtain the DSNB energy spectrum with an event number ranging from $N_{\text{evts}} = 4.2$ given by the low flux model to $N_{\text{evts}} = 193.5$ given by the SuperK flux upper bound for 147.7 kt · year exposure in the visible energy range of (12–100) MeV. The relevant parameters for the three flux models and the corresponding DSNB event rates are presented in table 3.

4.3 Atmospheric ν charged current background

We study the IBD events induced by the atmospheric neutrino charged current interactions in JUNO. We adopt the atmospheric neutrino flux calculated by M. Honda et al. [40], which considers all three flavors of neutrinos and their anti-neutrinos, the effect of Earth’s magnetic field on the flux, and neutrino oscillation effects at the JUNO site. The atmospheric neutrino flux below 100 MeV is from FLUKA [41] simulation results, which is normalized to match the flux in [40] in the overlapping energy region from 100 MeV to 944 MeV. Atmospheric ν_τ and $\bar{\nu}_\tau$ are neglected in our study due to their low fluxes (both less than 1% of the total atmospheric neutrino flux), while the remaining ν_μ ($\bar{\nu}_\mu$) and ν_e ($\bar{\nu}_e$) fluxes are comparable. These two flavors of atmospheric neutrino fluxes can be separated with an accuracy better than 99.9% in the JUNO detector for the energy range of (12–100) MeV [24]. The background due to atmospheric ν_μ and $\bar{\nu}_\mu$ can be suppressed through a T_{res} cut that is applied on the root mean square of the time residual profile as mentioned in section 3. The signal selection efficiency with the T_{res} cut remains at 99.98%. Therefore we focus on backgrounds due to atmospheric ν_e and $\bar{\nu}_e$. We have simulated $\bar{\nu}_e + p$ and $\nu_e(\bar{\nu}_e) + {}^{12}\text{C}$ interactions with GENIE (2.12.0) [42]

isotope	decay mode	Q in MeV	half-life $T_{1/2}$	Rate per 10 years ($E > 10$ MeV)
^9Li	$e^- + n$	11.9	178 ms	2.8×10^4
^{11}Li	$e^- + n$	20.6	8.75 ms	9.3×10^1
^{12}Be	$e^- + n$	11.7	21.5 ms	1.2
^{14}B	$e^- + n$	20.6	12.6 ms	2.4

Table 4. The estimated rates of cosmogenic isotopes in JUNO for energies above 10 MeV.

and the JUNO offline framework, which produces results consistent with previous studies on neutrino-nucleus interactions [43]. The IBD channel, $\bar{\nu}_e + p \rightarrow e^+ + n$, gives the dominant event rate, 30.5 ± 7.6 for $147.7 \text{ kt} \cdot \text{year}$ for the visible energy range of (12–100) MeV. The second background channel, $\bar{\nu}_e + ^{12}\text{C} \rightarrow ^{12}\text{B}^* + e^+$ with the secondary decay $^{12}\text{B}^* \rightarrow ^{11}\text{B} + n$, gives the next-to-leading event rate for the same exposure and energy range. This channel can mimic IBD events with a prompt e^+ and a delayed neutron coming from the secondary $^{12}\text{B}^*$ decay. Due to comparable event rates with the signal, the two channels, $\bar{\nu}_e + p \rightarrow e^+ + n$ and $\bar{\nu}_e + ^{12}\text{C} \rightarrow e^+ + n + ^{11}\text{B}$, will be included in the following pulse shape and sensitivity analyses.

4.4 Cosmogenic isotopes

The β -n decay from isotopes could mimic IBD events by emitting a β -particle (e^- or e^+) and a neutron. This effect has been measured by both KamLAND [44] and Borexino [45]. The isotopes could be produced in the JUNO site through interactions between energetic cosmic muons and ^{12}C . We only focus on isotopes giving rise to β -n decays with significant event rates and Q values higher than 10 MeV. Using FLUKA, the rates of ^9Li , ^{11}Li , ^{12}Be and ^{14}B are presented in table 4 [21]. The background from ^9Li and ^{12}Be can be neglected due to a 12 MeV cut on the prompt energy, which is the lower energy limit of our DM search. The half-lives of ^{11}Li and ^{14}B are 8.75 ms and 12.6 ms, respectively. They both contribute to β -n decays with branching fractions of 83% and 6.1%, respectively. For the total event rate of ^{11}Li and ^{14}B , we assume a 10% uncertainty based on the Poisson error. A flat spectrum is assumed for the event rate estimation. As a result, the total event rate of the β -n decays is 57.0 ± 5.7 for $147.7 \text{ kt} \cdot \text{year}$ for the visible energy above 12 MeV.

Muon veto strategies in JUNO have been studied in [22]. To suppress ^{11}Li and ^{14}B background events, we choose 0.2 s as the deadtime interval, i.e., after each muon event a 0.2 s of exposure time is ignored. Due to this veto, the exposure efficiency in JUNO becomes 97.5% while more than 99% of ^{11}Li and ^{14}B events are vetoed. Therefore, after applying the muon veto and the 12 MeV energy cut, ^9Li , ^{11}Li , ^{12}Be and ^{14}B events are negligible.

4.5 Muon induced fast neutrons

The above-mentioned muon veto tags muons passing through the LS and the water buffer with excellent efficiencies of 100% and 99.8%, respectively [22]. However, untagged muons that are either corner clipping the detector or passing through the rocks surrounding the detector could produce energetic neutrons. Any of these neutrons may enter in the LS and produce the prompt proton-recoil signal before being captured by the hydrogen, which is called a fast neutron (FN) event. These signatures mimic IBD events. The FN simulation has been performed based on the JUNO official simulation framework. The simulated atmospheric muon event sample corresponds to around 1200 days of data taking. A flat FN energy spectrum is obtained and the event rate is 1340 ± 270 for $200 \text{ kt} \cdot \text{year}$ of exposure within the visible

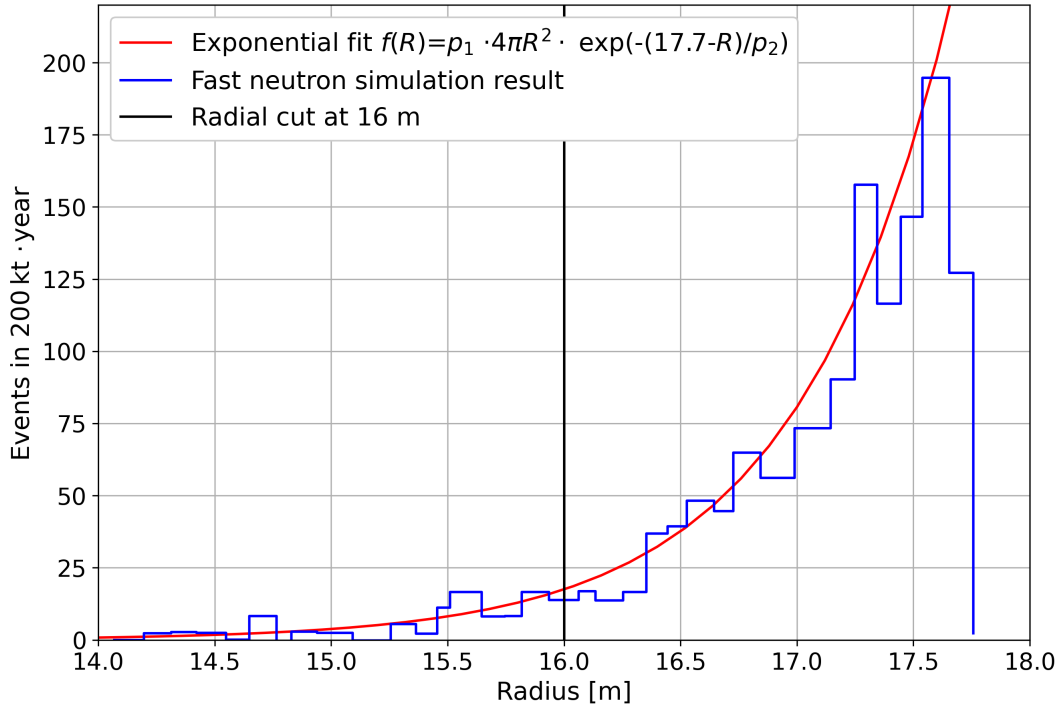


Figure 3. Histogram: radial distribution of muon-induced FN events. Red line: fit to the histogram with the function $f(R) = p_1 \cdot 4\pi R^2 \cdot \exp(-(17.7 - R)/p_2)$ with $p_1 = 0.75 \pm 0.04$, and $p_2 = 0.71 \pm 0.04$. Black line: radial cut at 16 m.

energy range of (12–100) MeV of the prompt signal. The spatial vertex distributions of FN events are shown in figure 3. The FN event rate decreases from the edge to the center of the detector, which results from the neutron attenuation.

We perform the fiducial volume cut at the radius $R = 16$ m, which significantly reduces the FN event rate. Using an exponential fit to the simulated FN data, we estimate the fiducial volume cut efficiency to be 7.1%. In contrast, the fiducial volume cut efficiency for the uniformly distributed events, such as DM, DSNB, and atmospheric neutrino events is 73.9%. The FN events with $R < 16$ m can be suppressed by the PSD method, which will be discussed in the next section. Eventually, these FN backgrounds are negligible.

4.6 Atmospheric ν neutral current background

A detailed study on the neutral-current interaction between the atmospheric neutrino and ^{12}C has been carried out in [46]. We have reproduced the above simulations and applied them to the DM study in the JUNO offline framework. We have neglected the channel $\nu + p \rightarrow \nu + p$ since the fraction of neutron-less IBD-like events from atmospheric neutrino NC interactions is only 0.99%.

Using the atmospheric neutrino flux calculated by M. Honda et al. [40], we estimate the Atm- ν NC background. The uncertainty of this flux calculation is less than 10% in the energy range of (1–10) GeV while it varies between (10–30)% outside this energy range due to the lack of observational results. We use the neutrino generator GENIE (2.12.0) [42] to model neutral current interactions between atmospheric neutrinos and ^{12}C . We adopt the default setting in GENIE, where the axial mass M_A in the parametrization of the nucleon

axial-vector form factor is taken as 0.99 GeV, the relativistic Fermi gas (RFG) model is adopted for nuclear structures, and the Intranuclear Cascade (INC) model is applied for final-state interactions [47]. The de-excitation of the final-state nuclei is simulated with the package TALYS (1.8) [48]. A statistical configuration model [49–51] is applied for providing the de-excitation probability before the TALYS simulation.

We conclude that the most important NC interaction channel is $\nu + {}^{12}\text{C} \rightarrow \nu + {}^{11}\text{C} + n$, and the total event rate before IBD signal selection is estimated to be $49.0 \text{ year}^{-1} \text{ kt}^{-1}$, which is consistent with the previous study [46]. Applying IBD event selection criteria, the efficiency for Atm- ν NC event becomes $(7.3 \pm 0.5)\%$. A conservative 15% total uncertainty is taken from the study in [52]. Hence the event rate of IBD-like atmospheric neutrino NC events is 670 ± 100 for $147.7 \text{ kt} \cdot \text{year}$ in the visible energy range (12–100) MeV.

5 Pulse Shape Discrimination

Different types of particles show distinct photon emission time profiles that result from LS excitation induced by the deposited energy. Pulse shape discrimination (PSD) is a powerful way to separate Atm- ν NC and FN events from IBD signal events by analyzing the pulse shapes of their prompt signals. The tail-to-total ratio (TTR) method [53] is adopted here, where the ratio between the charge in a specific time window corresponding to the tail of the pulse and the total charge of the pulse is the parameter to distinguish between different event types. Although multivariate analysis (TMVA) [54] and machine learning methods [55] show superior efficiencies, the TTR method is sufficient for suppressing Atm- ν NC events to the same order as the neutrino floor in our DM study.

We perform a full simulation to produce pulse shapes of different event types, which is based on the official JUNO offline simulation and analysis frameworks. To do this, we have employed the DM flux model discussed in section 3, the DSNB flux model discussed in section 4.2, the Atm- ν CC and NC events from the GENIE interaction models and the FN background events. In our analysis, we select all events that pass the IBD signal selection and muon veto cut. We scan through the tail settings with 50 ns steps on the start and end times of the tail. The optimized tail window is from 200 ns to 600 ns (see discussions later). The TTR ratio versus the visible energy and the reconstructed position are presented via scatter plots in figure 4. There are also contours corresponding to different event types. Each contour marks the region that contains 90% of a given type of events. DM signal and DSNB events occurring through the IBD process result into an identical TTR ratio distribution. This is a relatively stable distribution as reflected by its smooth 90% event boundary. FN events follow from neutron elastic scattering with proton or ${}^{12}\text{C}$ and producing gammas, which give rise to a more complex TTR ratio distribution. Finally, Atm- ν NC events result from neutral-current interactions between atmospheric neutrinos and ${}^{12}\text{C}$, which produces the most complicated TTR ratio distribution due to the variety of interaction channels characterized by different prompt-signal spectral shapes. The black solid curve on each panel represents the event selection criterion such that those events situated below this curve are classified as signals. It is seen that the IBD contour region is entirely below this curve on the left panel, while a small part of the IBD contour region is above the event selection curve for $E_{\text{vis}} \leq 30 \text{ MeV}$. This implies that the DM signal efficiency after the PSD cut remains higher than 90% for $E_{\text{vis}} \geq 30 \text{ MeV}$. On the other hand, for $E_{\text{vis}} < 30 \text{ MeV}$, a significant fraction of IBD events is also removed by the PSD cut. As will be discussed later, such an event reduction is seen for DM signals illustrated by figure 5 and DSNB events illustrated by figures 6 and 7.

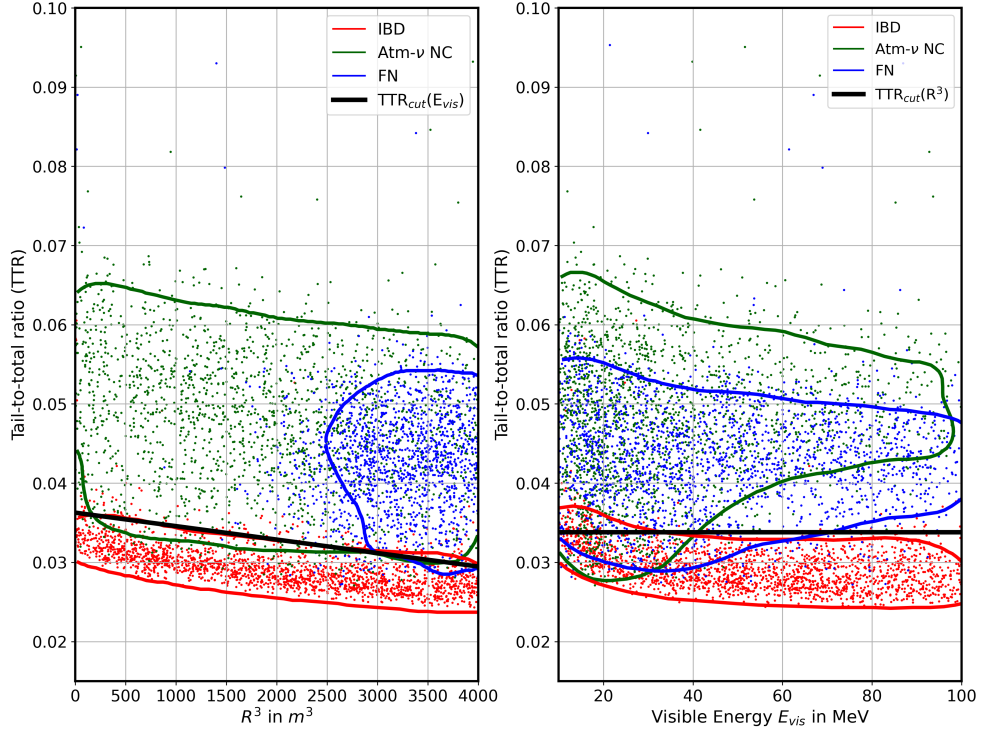


Figure 4. Tail-to-total ratios (TTR) versus visible energy and the reconstructed position for IBD events, Atm- ν NC and FN events that pass the IBD selection criteria. Different event types are represented by scatter plots with different colors. Each contour marks the region that contains 90% of a given type of events. The black solid curve on each panel is the signal selection criterion such that those events situated below this curve are classified as signals. On the left panel the curve is parametrized by $\text{TTR}_{\text{cut}}(R) = 0.036 - 1.7 \cdot 10^{-6} \text{ m}^{-3} \cdot R^3$ while the curve on the right panel is parametrized by $\text{TTR}_{\text{cut}}(E_{\text{vis}}) = 0.034$. The PSD cut selects only those events which are situated below both curves.

We evaluate the PSD performance based on the signal to background ratio $N_S/\sqrt{N_S + N_B}$ averaged over the DM mass range of (15–100) MeV (with a 5 MeV step size), where N_S is the number of signal events from DM annihilation for a specific DM mass and N_B is the sum of all backgrounds. The current best limit on the thermally averaged DM annihilation rate set by SuperK [17] is adopted here to represent the highest allowed signal to background ratio under the latest constraints. The PSD efficiencies of IBD events giving rise to the best signal to background ratio are analyzed while keeping the PSD efficiency of NC background events fixed at 2%, 3%, ..., 6%, and 7%, respectively. In table 5, we present the tail window and the resulting PSD efficiency of IBD events which optimize the signal-to-background ratio for a fixed NC-background PSD efficiency. Figure 5 illustrates PSD efficiencies as functions of the prompt energy in two different cases, $\epsilon_{\text{PSD,NC}} = 2\%$ and 4%, respectively. We adopt the setting with 4% PSD efficiency for NC background events since it gives rise to a better signal-to-background ratio.

As shown in figures 4 and 5, PSD is a powerful tool to discriminate IBD events against Atm- ν NC events, especially for visible energies above 30 MeV. Below 30 MeV, the discrimination power decreases because of the lower photon statistics. Furthermore, there is a large fraction of γ 's emitted in the de-excitation processes of Atm- ν NC events in this energy range.

$\epsilon_{\text{PSD,NC}}$	$\epsilon_{\text{PSD,IBD}}$	$\epsilon_{\text{PSD,FN}}$	Tail window (ns)	mean $N_S/\sqrt{N_S + N_B}$
2.0%	81.7%	0.5%	[200,600]	4.092
3.0%	87.5%	1.2%	[250,600]	4.140
4.0%	90.5%	2.0%	[200,600]	4.151
5.0%	92.3%	3.9%	[200,600]	4.120
6.0%	93.6%	7.0%	[150,600]	4.071
7.0%	95.5%	8.2%	[200,600]	4.054

Table 5. PSD cut efficiencies for Atm- ν NC, IBD and FN events, corresponding tail windows and the resulting best mean signal-to-background ratios.

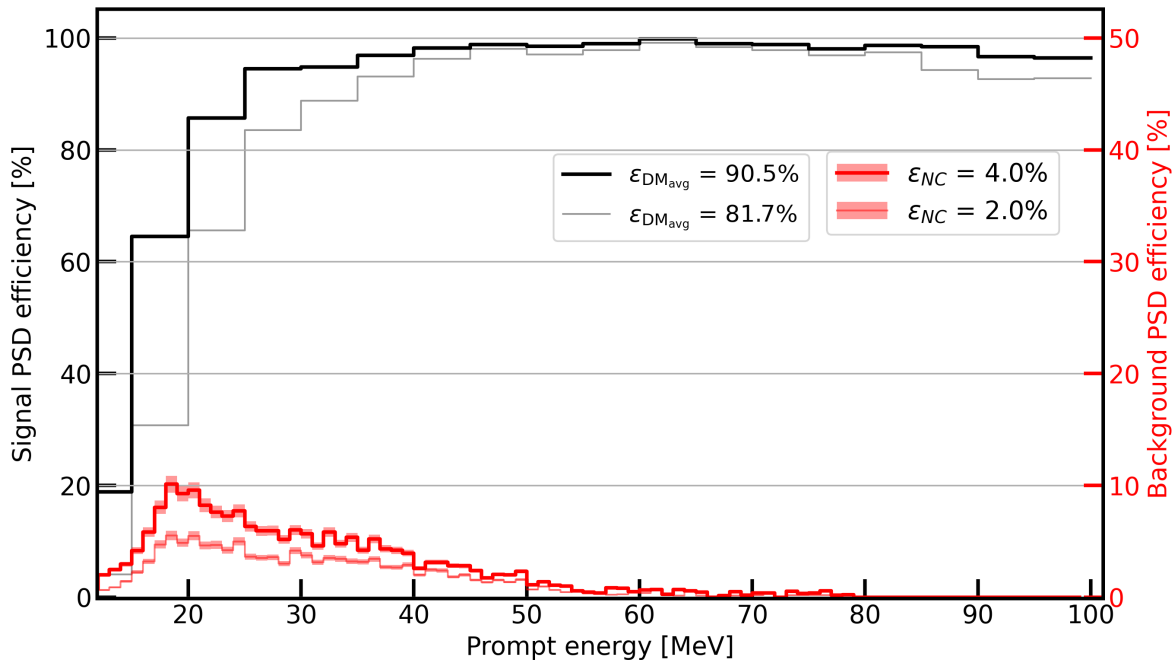


Figure 5. PSD cut efficiencies as functions of the visible energy of the prompt signal of Atm- ν NC background events (in red) and IBD signal events (in black). Two examples, $\epsilon_{\text{PSD,NC}} = 2\%$ with $\epsilon_{\text{PSD,IBD}} = 81.7\%$ (light curves), and $\epsilon_{\text{PSD,NC}} = 4\%$ with $\epsilon_{\text{PSD,IBD}} = 90.5\%$ (bold curves), are presented.

This also reduces the PSD efficiency by realizing that γ 's from prompt signals of Atm- ν NC events and positrons from prompt signals of IBD signal events give rise to almost identical photon emission time profiles, which makes the PSD method ineffective. Finally, we note that the PSD efficiency might further worsen at higher energies due to energy dependencies of pulse shapes, which require careful studies.

To determine the uncertainty of the PSD cut, we note that spallation neutrons have been proposed as ideal sources for such a study since they induce prompt and delayed pairs with energies similar to those of Atm- ν NC events. Specifically, around 180 days of muon simulation data was used for generating the aforementioned neutron sample [56], so that the number of spallation neutrons in the signal energy window can be estimated for 10 years of data taking. Taking an average PSD efficiency for Atm- ν NC events as 4%, which applies to

the spallation neutron sample as well, the statistical uncertainty for the selected spallation neutrons are determined by the number of selected neutrons, which varies with the signal energy window determined by m_χ (see figure 2). It is 16%, 10%, 15%, 24%, and 43% with 10 years of data taking for $m_\chi = 15, 20, 30, 40$, and 50 MeV, respectively. We stress that the large uncertainty in the high energy range does not cause a huge impact on the DM sensitivity because Atm- ν NC events do not dominate the total spectrum over 40 MeV after applying the PSD cut. Last but not least, several calibration sources and techniques, such as AmBe and Michel electrons, were also proposed for further constraining the PSD systematic in the future [25].

In conclusion, we obtain an average PSD efficiency of 90.5% for IBD signal events from DM annihilation and an average PSD survival probability of 4.0% for Atm- ν NC events, as well as approximately 50%, 18%, 97% and 2.0% for DSNB events from model predictions, DSNB events corresponding to the SuperK upper bound, Atm- ν CC, and FN events, respectively.

6 Sensitivity

The JUNO sensitivity to the detection of $\bar{\nu}_e$ from DM annihilation in the galactic halo will be discussed in this section. First, we summarize the DM signal and backgrounds with the corresponding veto methods applied. Two different approaches, Poisson-type log-likelihood ratio test and Bayesian analysis, are employed to calculate the DM detection sensitivities and verify the consistency of the two approaches.

6.1 Total spectrum in JUNO

We adopt 10 years of exposure time as a reasonable time scale. Canonical value, $J_{\text{avg}} = 5$ [23], is used in our analysis. The thermal relic DM annihilation rate, $\langle\sigma v\rangle = 3 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}$, is assumed, which results in few signal events for an exposure of 147.7 kt · year. Figures 6 and 7 show the final visible energy spectra of the signal and backgrounds before (upper panel) and after (lower panel) the PSD cut. DSNB and DM fluxes in figure 6 are taken from SuperK flux upper bounds [17, 39] while those in figure 7 are given by theoretical predictions. The signal spectrum is shown for an assumed DM mass of 50 MeV. It is obtained with IBD signal selection cuts, the muon veto cut, and the PSD cut applied (lower panels). With the PSD cut, non-IBD backgrounds (Atm- ν NC and FN) can be suppressed to the similar order of magnitude as IBD backgrounds (atmospheric CC, reactor, and DSNB). The IBD backgrounds and Atm- ν NC events build up the background floor and dominate the JUNO sensitivity of indirect DM search. In table 6 we summarized the event numbers with IBD signal selection, muon veto, and PSD cut. Despite the low event rate, the mono-energetic $\bar{\nu}_e$ flux yields a sharp visible energy spectrum for the signal, which can be identified against the background spectra.

6.2 Two approaches to the sensitivity

In the following we apply a Likelihood-ratio test and a Bayesian analysis to estimate the 90% confidence level upper limit on $\langle\sigma v\rangle$ expected by JUNO for the DM mass range of (15–100) MeV.

Likelihood-ratio test method. In the likelihood-ratio test method, we define

$$\chi^2 = -2 \ln \lambda = 2 \sum_{i=1}^N \left(n_i \ln \frac{n_i}{v_i} + v_i - n_i \right) + \sum_{i=1}^N \left(\frac{v_i - \bar{v}_i}{\sigma_i} \right)^2, \quad (6.1)$$

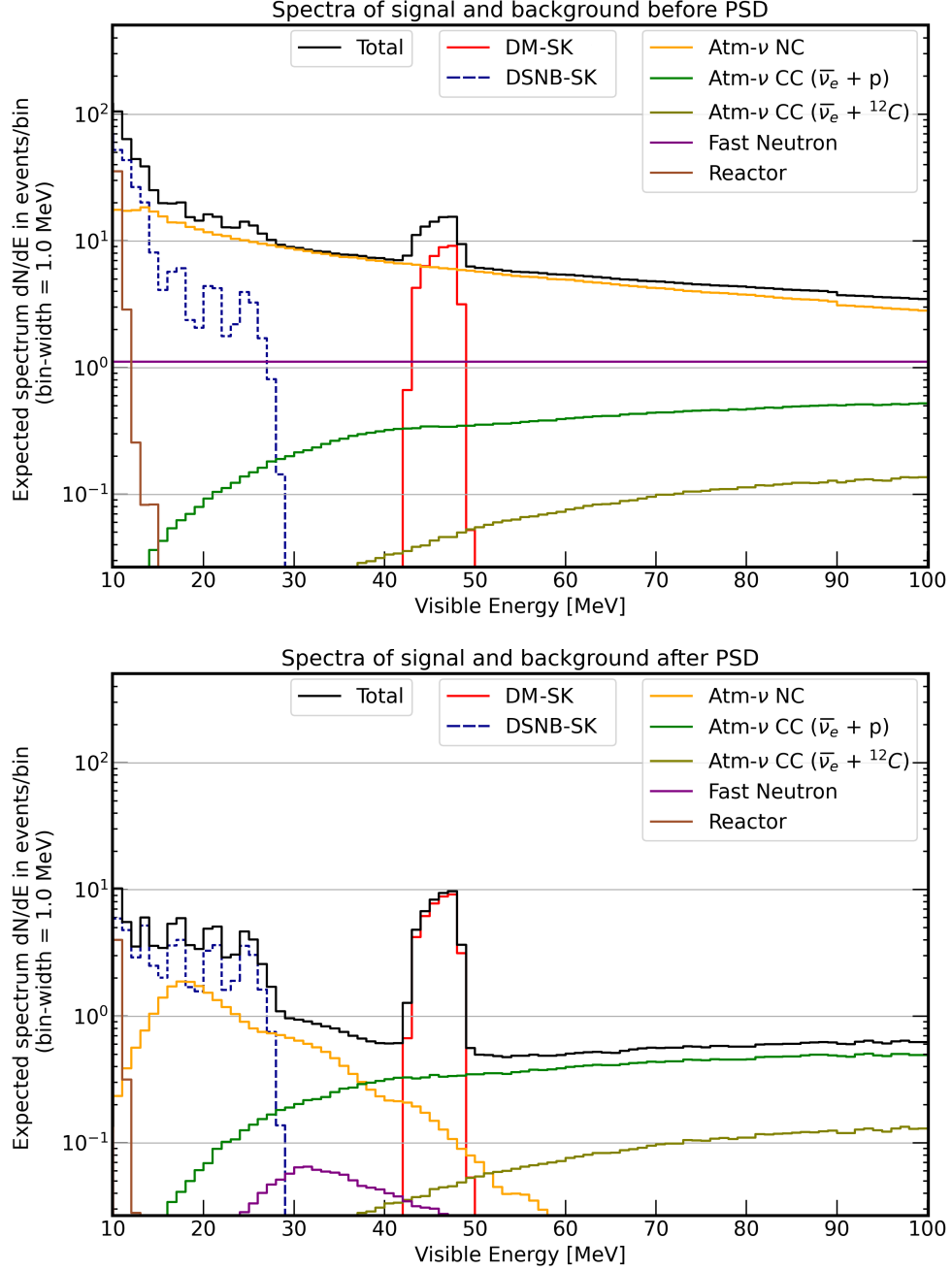


Figure 6. Total visible energy spectra of DM-IBD signals ($m_\chi = 50$ MeV) and backgrounds in JUNO before (upper panel) and after PSD (lower panel). Here DSNB [39] and DM [17] event spectra are taken from experimental upper bounds by SuperK.

where λ is the likelihood ratio, n_i represents observed (signal plus background) events per bin, v_i represents background events per bin with its central value denoted by $\bar{v}_i$, and σ_i represents the uncertainty of the total backgrounds for the energy bin i . Here we take $\bar{v}_i$ as part of the observed event number per bin included in n_i . We note that the total background uncertainty varies with DM mass since a different targeted m_χ leads to a different selection of the visible

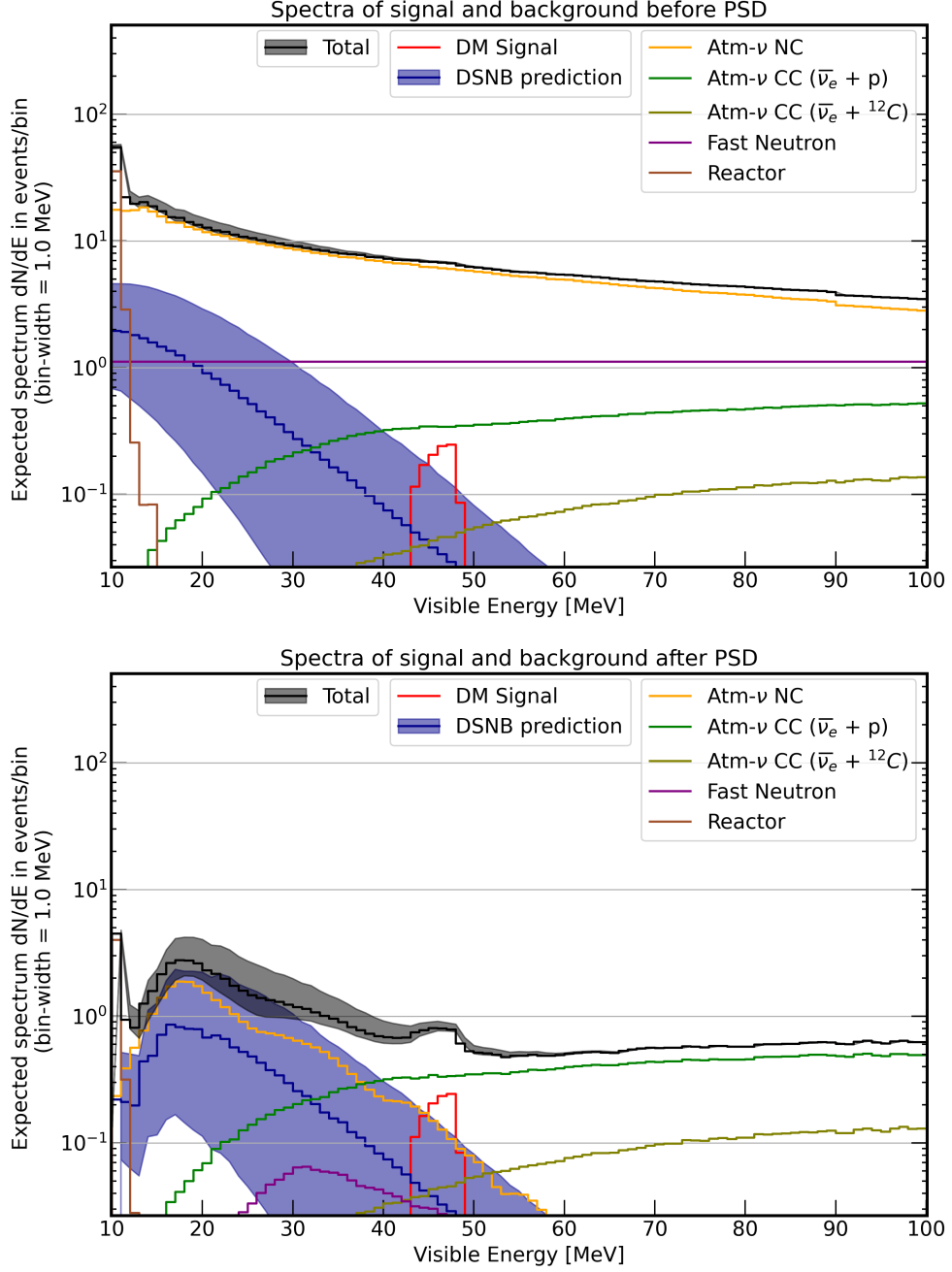


Figure 7. Same as figure 6 with DSNB event spectra given by theoretical predictions and the DM signal event spectrum based upon thermal relic DM annihilation rate.

energy window for the analysis. The uncertainty is 19%, 16%, 20%, and 24% for $m_\chi = 15, 20, 30$, and 40 MeV, respectively, while it is 25% for $50 \leq m_\chi/\text{MeV} \leq 100$. The number of energy bins is determined by the energy resolution $3\%/\sqrt{E[\text{MeV}]}$ of JUNO. The degrees of freedom correspond to the number of energy bins. We note that χ^2 equals to zero for $n_i = v_i$ and $v_i = \bar{v}_i$. Denoting χ^2 in this case as $\chi^2_{\min}$, we then look for $\Delta\chi^2 = (\chi^2 - \chi^2_{\min}) = (1.645)^2$ to obtain the 90% C.L. sensitivity limit for $\langle\sigma v\rangle$. The sensitivity limit depends on the DSNB

$\text{DM}, \langle\sigma v\rangle=3 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}, J_{\text{avg}} = 5$	Event number, IBD signal selection	muon veto [22]	muon veto [22]+PSD cut	Signal to Background Ratio ¹
15 MeV	1.9	1.8	0.6	0.32
20 MeV	1.5	1.4	0.8	0.37
90 MeV	0.9	0.9	0.9	0.27
100 MeV	0.8	0.8	0.8	0.26
Backgrounds				
Cosmogenic Isotopes				
Fast neutrons	$(0.6 \pm 0.1) \cdot 10^2$	negligible	negligible	
Atmos NC	$(1.0 \pm 0.2) \cdot 10^2$	$(0.9 \pm 0.2) \cdot 10^2$	$(0.2 \pm 0.1) \cdot 10^1$	
Atmos CC	$(6.7 \pm 1.0) \cdot 10^2$	$(6.5 \pm 1.0) \cdot 10^2$	$(2.6 \pm 0.7) \cdot 10^1$	
DSNB [SK Max Nom. Min]	$(0.3 \pm 0.1) \cdot 10^2$	$(0.3 \pm 0.1) \cdot 10^2$	$(2.9 \pm 0.8) \cdot 10^1$	
	$1.9 \cdot 10^2 \mid 0.7 \cdot 10^2 \mid 0.2 \cdot 10^2 \mid 4.2$	$1.9 \cdot 10^2 \mid 0.6 \cdot 10^2 \mid 0.2 \cdot 10^2 \mid 4.1$	$3.4 \cdot 10^1 \mid 3.5 \cdot 10^1 \mid 1.1 \cdot 10^1 \mid 0.1 \cdot 10^1$	
Total background	$10.9 \cdot 10^2 \mid 9.7 \cdot 10^2 \mid 9.2 \cdot 10^2 \mid 9.0 \cdot 10^2$	$10.1 \cdot 10^2 \mid 8.8 \cdot 10^2 \mid 8.4 \cdot 10^2 \mid 8.2 \cdot 10^2$	$9.3 \cdot 10^1 \mid 9.4 \cdot 10^1 \mid 7.0 \cdot 10^1 \mid 6.0 \cdot 10^1$	
Total background uncertainty	$\pm 1.0 \cdot 10^2$	$\pm 1.0 \cdot 10^2$	$\pm 1.1 \cdot 10^1$	

¹ The nominal DSNB flux is adopted in the computation of signal-to-background ratio.

Table 6. Summary of DM signal and background event numbers with $147.7 \text{ kt} \cdot \text{year}$ of exposure in JUNO.

flux we adopt as part of the total background. Table 3 summarizes four DSNB flux models adopted in this analysis. Hence, a specific sensitivity curve corresponds to a specific choice of DSNB flux as shown in figure 8.

Bayesian analysis with Markov Chain Monte Carlo sampling. The second approach is based on Bayesian analysis. A likelihood function is defined with the observed (n_i) and expected (λ_i) number of events in the i -th bin of the spectrum by

$$p(\text{spec}|S, B_{\text{DSNB}}, \dots, B_{\text{atmoNC}}) = \prod_{i=1}^N \frac{\lambda_i(S, B_{\text{DSNB}}, \dots, B_{\text{atmoNC}})^{n_i}}{n_i!} e^{-\lambda_i(S, B_{\text{DSNB}}, \dots, B_{\text{atmoNC}})}. \quad (6.2)$$

With Bayes' theorem, the posterior probability is $p(S, \dots, B_{\text{atmoNC}}|\text{spec}) \propto p(\text{spec}|S, \dots, B_{\text{atmoNC}}) \cdot p_0(S) \cdots p_0(B_{\text{atmoNC}})$ with eq. (6.2) and the prior probabilities for signal and background contributions p_0 where spec means the total spectrum. It represents the probability that the observed spectrum can be explained by the set of parameters S and B and is marginalized with respect to the background contributions

$$p(S|\text{spec}) = \int p(S, B_{\text{DSNB}}, \dots, B_{\text{atmoNC}}|\text{spec}) dB_{\text{DSNB}} \cdots dB_{\text{atmoNC}}. \quad (6.3)$$

The 90% probability upper limit S_{90} on the number of signal events can therefore be calculated by equating the integral of eq. (6.3) with a probability of 90%: $\int_0^{S_{90}} p(S|\text{spec}) dS = 0.90$ [57]. Data sets representing the expected number of events λ_i are generated from the background-only spectrum following Poisson distribution and are analyzed for different mass settings of the DM signal. We set a flat prior probability of the signal contribution, $p_0(S) = 1/S_{\text{max}}$ (with $S_{\text{max}} = 60$ according to the current limits of [17]). The prior probabilities of backgrounds are set by a Gaussian distribution with the mean value μ_B and width σ_B :

$$p_0(B) = \exp\left(-\frac{(B - \mu_B)^2}{2\sigma_B^2}\right) / \int_0^\infty \exp\left(-\frac{(B - \mu_B)^2}{2\sigma_B^2}\right) dB, \quad (6.4)$$

when $B \geq 0$. The marginalization of the posterior probability function is performed with Markov Chain Monte Carlo sampling using the Python package of [58], which is based on [59]. The 90% upper limit S_{90} of the number of signal events are calculated for each data set and the distribution of the values of S_{90} can be interpreted as a probability density to determine the mean 90% upper limit on the number of signal events as well as the upper limits on the anti-neutrino flux from DM self-annihilation and on the DM self-annihilation cross-section, respectively.

Two different statistical approaches obtain consistent sensitivities to $\langle\sigma v\rangle$. In figure 8, we present JUNO's expected 90% C.L. upper limit on $\langle\sigma v\rangle$ together with upper limits obtained or expected from other neutrino detectors. We analyze the total event spectrum for a 90.5% PSD cut efficiency for IBD events and a 4.0% PSD cut efficiency for Atm- ν NC background events with Bayesian analysis and a likelihood-ratio test. The spectral behaviors of the total background events and the DM signal could explain why the best search sensitivity occurs around $m_\chi = 40$ MeV. Figures 6 and 7 show that the total number of background events gradually decreases with increasing visible energy until around 40 MeV. On the other hand, figure 2 shows that the DM signal spectrum becomes broader with an increasing m_χ . The best sensitivity of JUNO therefore occurs at $m_\chi \simeq 40$ MeV, which results in a 90% upper limit on the thermally averaged DM annihilation rate, $\langle\sigma v\rangle(m_\chi = 40 \text{ MeV}) = 1.1 \times 10^{-25} \text{ cm}^3 \text{ s}^{-1}$

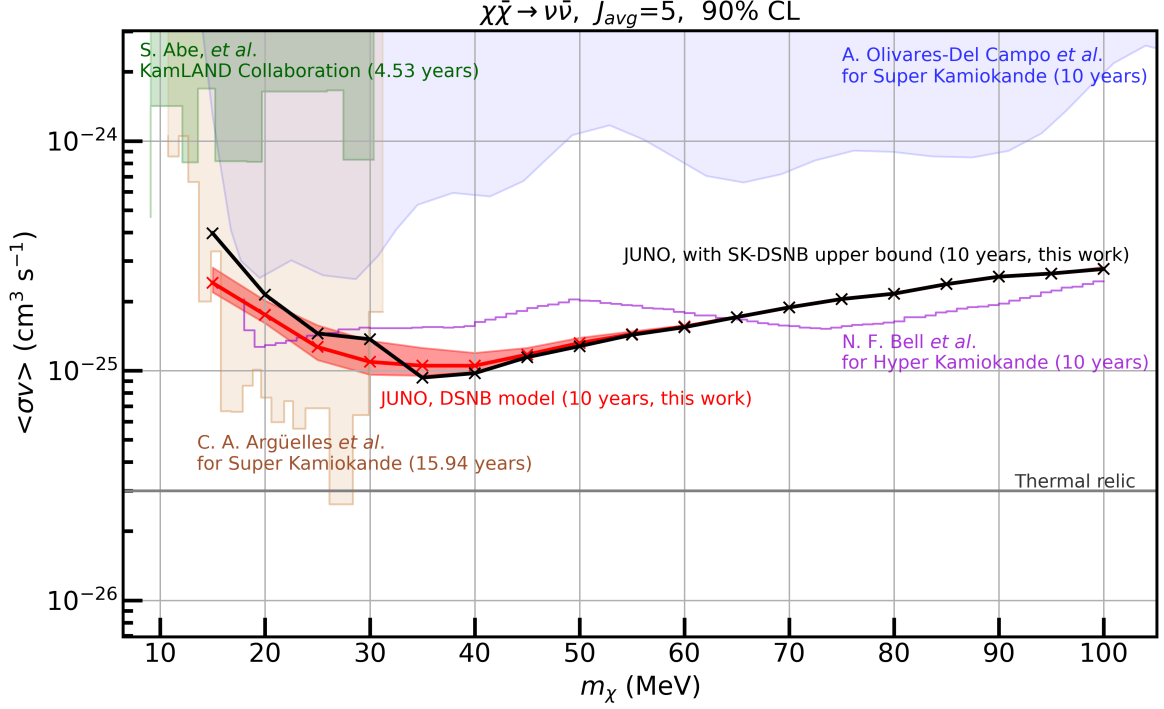


Figure 8. 90% C.L. upper limits on $\langle\sigma v\rangle$ with 10 years of data taking in JUNO. Black and red lines represent upper limits with the theoretical model and experimental limit of DSNB, respectively. Due to consistent results between the Bayesian analysis and Likelihood-ratio test, one single upper limit curve for each DSNB flux setting is presented here. Colored areas are the excluded parameter regions from KamLAND [12] and SuperK [17, 18] observations. HyperK expectation by N. F. Bell et al. [20] is also shown for comparison. The thermal relic value $\langle\sigma v\rangle = 3 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}$ is also shown.

for both Likelihood and Bayesian analysis over 10 years of data taking. The sensitivity curves in figure 8, except of those by JUNO, are scaled to the convention $J_{\text{avg}} = 5$. We compare our result with those of KamLAND [12], SuperK [17, 18] and the HyperK expectation [20]. We note that results by [17] and [18] differ due to differences in the analyzed dataset and the background modeling as pointed out in [18]. JUNO will probe into $\langle\sigma v\rangle$'s that are up to an order of magnitude smaller than the ones SuperK obtained for the DM mass range of (30–100) MeV. Such a sensitivity will be comparable to the one expected by HyperK in the DM mass range of (15–100) MeV.

Besides presenting JUNO sensitivities to $\langle\sigma v\rangle$ for 10 years of data taking, we also present in figure 9 the progress of sensitivities over the running time for benchmark DM masses of 20, 40, and 100 MeV, respectively. It is seen that, for $m_\chi = 40$ MeV, the sensitivity to $\langle\sigma v\rangle$ can reach $6 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}$ over 20 years of data taking.

7 Conclusion

In this study, the JUNO sensitivity to the detection of neutrinos from DM self-annihilation in the Milky Way is investigated for the DM mass range of (15–100) MeV. The expected $\bar{\nu}_e$ signal spectrum from the annihilation $\chi\chi \rightarrow \nu\bar{\nu}$ is evaluated with the JUNO offline simulation framework. Moreover, all possible background contributions to the above indirect DM search are investigated, which include the reactor $\bar{\nu}_e$ background, DSNB, atmospheric neutrino

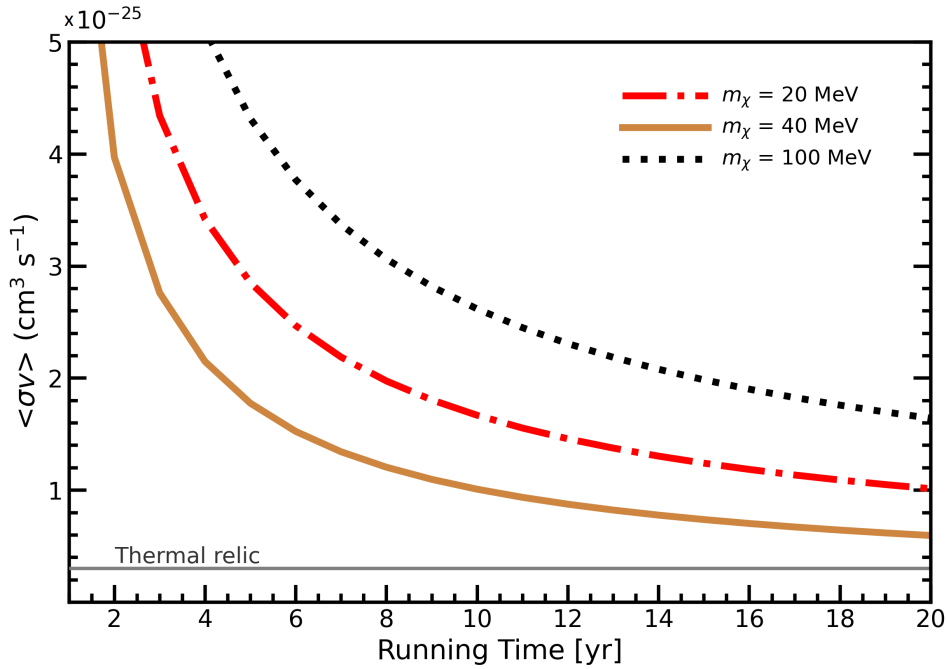


Figure 9. The progress of JUNO 90% C.L. sensitivities to $\langle\sigma v\rangle$ over the running time up to 20 years for DM masses of 20, 40, and 100 MeV, respectively. The thermal relic value $\langle\sigma v\rangle = 3 \times 10^{-26} \text{ cm}^3 \text{s}^{-1}$ is also shown for comparison. The nominal DSNB flux and the convention $J_{\text{avg}} = 5$ are adopted in the sensitivity calculations.

backgrounds from CC and NC interactions as well as muon-induced backgrounds. To further reduce the non-IBD Atm- ν NC and FN background events, a PSD method based on the tail-to-total ratio is applied. We apply a customized PSD cut and IBD signal selection based on each DM mass, which results in a different event spectrum and sensitivity from the DSNB study [56]. The sensitivity of JUNO, i.e. the 90% C.L. upper limit on the DM annihilation rate, $\langle\sigma v\rangle$, is analyzed based on both the Bayesian analysis and the Likelihood ratio test. JUNO could reach the sensitivity of $\langle\sigma v\rangle = 1.1 \times 10^{-25} \text{ cm}^3 \text{s}^{-1}$ with both Likelihood and Bayesian analysis for $m_\chi = 40 \text{ MeV}$ over 10 years of exposure. It corresponds to an improvement by a factor of 2 to 10 for the DM mass range of (15–100) MeV compared to the present-day best limits obtained by SuperK and would be comparable to that expected by HyperK.

Acknowledgments

We are grateful for the ongoing cooperation from the China General Nuclear Power Group. This work was supported by the Chinese Academy of Sciences, the National Key R&D Program of China, the CAS Center for Excellence in Particle Physics, Wuyi University, and the Tsung-Dao Lee Institute of Shanghai Jiao Tong University in China, the Institut National de Physique Nucléaire et de Physique de Particules (IN2P3) in France, the Istituto Nazionale di Fisica Nucleare (INFN) in Italy, the Italian-Chinese collaborative research program MAECI-NSFC, the Fonds de la Recherche Scientifique (F.R.S-FNRS) and FWO under the “Excellence of Science — EOS” in Belgium, the Conselho Nacional de Desenvolvimento Científico e Tecnológico in Brazil, the Agencia Nacional de Investigación y Desarrollo in Chile, the Charles University Research Centre and the Ministry of Education, Youth, and Sports in Czech

Republic, the Deutsche Forschungsgemeinschaft (DFG), the Helmholtz Association, and the Cluster of Excellence PRISMA+ in Germany, the Joint Institute of Nuclear Research (JINR) and Lomonosov Moscow State University in Russia, the joint Russian Science Foundation (RSF) and National Natural Science Foundation of China (NSFC) research program, the NSTC and MOE in Taiwan, the Chulalongkorn University and Suranaree University of Technology in Thailand, and the University of California at Irvine in U.S.A.

References

- [1] V.C. Rubin and W.K. Ford Jr., *Rotation of the Andromeda Nebula from a Spectroscopic Survey of Emission Regions*, *Astrophys. J.* **159** (1970) 379 [[INSPIRE](#)].
- [2] G. Jungman, M. Kamionkowski and K. Griest, *Supersymmetric dark matter*, *Phys. Rept.* **267** (1996) 195 [[hep-ph/9506380](#)] [[INSPIRE](#)].
- [3] SUPER-KAMIOKANDE collaboration, *Search for dark matter WIMPs using upward through-going muons in Super-Kamiokande*, *Phys. Rev. D* **70** (2004) 083523 [Erratum *ibid.* **70** (2004) 109901] [[hep-ex/0404025](#)] [[INSPIRE](#)].
- [4] ICECUBE collaboration, *The Design and Performance of IceCube DeepCore*, *Astropart. Phys.* **35** (2012) 615 [[arXiv:1109.6096](#)] [[INSPIRE](#)].
- [5] FERMI-LAT collaboration, *Searching for Dark Matter Annihilation from Milky Way Dwarf Spheroidal Galaxies with Six Years of Fermi Large Area Telescope Data*, *Phys. Rev. Lett.* **115** (2015) 231301 [[arXiv:1503.02641](#)] [[INSPIRE](#)].
- [6] AMS collaboration, *Observation of Complex Time Structures in the Cosmic-Ray Electron and Positron Fluxes with the Alpha Magnetic Spectrometer on the International Space Station*, *Phys. Rev. Lett.* **121** (2018) 051102 [[INSPIRE](#)].
- [7] HESS collaboration, *Search for γ -Ray Line Signals from Dark Matter Annihilations in the Inner Galactic Halo from 10 Years of Observations with H.E.S.S.*, *Phys. Rev. Lett.* **120** (2018) 201101 [[arXiv:1805.05741](#)] [[INSPIRE](#)].
- [8] G. Barenboim, P.B. Denton and I.M. Oldengott, *Constraints on inflation with an extended neutrino sector*, *Phys. Rev. D* **99** (2019) 083515 [[arXiv:1903.02036](#)] [[INSPIRE](#)].
- [9] C.A. Argüelles et al., *New opportunities at the next-generation neutrino experiments I: BSM neutrino physics and dark matter*, *Rept. Prog. Phys.* **83** (2020) 124201 [[arXiv:1907.08311](#)] [[INSPIRE](#)].
- [10] R. Primulando and P. Uttayarat, *Dark Matter-Neutrino Interaction in Light of Collider and Neutrino Telescope Data*, *JHEP* **06** (2018) 026 [[arXiv:1710.08567](#)] [[INSPIRE](#)].
- [11] F. Capozzi, I.M. Shoemaker and L. Vecchi, *Solar Neutrinos as a Probe of Dark Matter-Neutrino Interactions*, *JCAP* **07** (2017) 021 [[arXiv:1702.08464](#)] [[INSPIRE](#)].
- [12] KAMLAND collaboration, *A study of extraterrestrial antineutrino sources with the KamLAND detector*, *Astrophys. J.* **745** (2012) 193 [[arXiv:1105.3516](#)] [[INSPIRE](#)].
- [13] KAMLAND collaboration, *Limits on Astrophysical Antineutrinos with the KamLAND Experiment*, *Astrophys. J.* **925** (2022) 14 [[arXiv:2108.08527](#)] [[INSPIRE](#)].
- [14] SUPER-KAMIOKANDE collaboration, *Indirect search for dark matter from the Galactic Center and halo with the Super-Kamiokande detector*, *Phys. Rev. D* **102** (2020) 072002 [[arXiv:2005.05109](#)] [[INSPIRE](#)].
- [15] ICECUBE collaboration, *Search for GeV-scale dark matter annihilation in the Sun with IceCube DeepCore*, *Phys. Rev. D* **105** (2022) 062004 [[arXiv:2111.09970](#)] [[INSPIRE](#)].
- [16] ANTARES and ICECUBE collaborations, *Combined search for neutrinos from dark matter self-annihilation in the Galactic Center with ANTARES and IceCube*, *Phys. Rev. D* **102** (2020) 082002 [[arXiv:2003.06614](#)] [[INSPIRE](#)].

- [17] A. Olivares-Del Campo, C. Boehm, S. Palomares-Ruiz and S. Pascoli, *Dark matter-neutrino interactions through the lens of their cosmological implications*, *Phys. Rev. D* **97** (2018) 075039 [[arXiv:1711.05283](#)] [[INSPIRE](#)].
- [18] C.A. Argüelles et al., *Dark matter annihilation to neutrinos*, *Rev. Mod. Phys.* **93** (2021) 035007 [[arXiv:1912.09486](#)] [[INSPIRE](#)].
- [19] C.A. Argüelles, A. Kheirandish and A.C. Vincent, *Imaging Galactic Dark Matter with High-Energy Cosmic Neutrinos*, *Phys. Rev. Lett.* **119** (2017) 201801 [[arXiv:1703.00451](#)] [[INSPIRE](#)].
- [20] N.F. Bell, M.J. Dolan and S. Robles, *Searching for Sub-GeV Dark Matter in the Galactic Centre using Hyper-Kamiokande*, *JCAP* **09** (2020) 019 [[arXiv:2005.01950](#)] [[INSPIRE](#)].
- [21] JUNO collaboration, *Neutrino Physics with JUNO*, *J. Phys. G* **43** (2016) 030401 [[arXiv:1507.05613](#)] [[INSPIRE](#)].
- [22] JUNO collaboration, *JUNO physics and detector*, *Prog. Part. Nucl. Phys.* **123** (2022) 103927 [[arXiv:2104.02565](#)] [[INSPIRE](#)].
- [23] S. Palomares-Ruiz and S. Pascoli, *Testing MeV dark matter with neutrino detectors*, *Phys. Rev. D* **77** (2008) 025025 [[arXiv:0710.5420](#)] [[INSPIRE](#)].
- [24] JUNO collaboration, *JUNO sensitivity to low energy atmospheric neutrino spectra*, *Eur. Phys. J. C* **81** (2021) 10 [[arXiv:2103.09908](#)] [[INSPIRE](#)].
- [25] JUNO collaboration, *Calibration Strategy of the JUNO Experiment*, *JHEP* **03** (2021) 004 [[arXiv:2011.06405](#)] [[INSPIRE](#)].
- [26] JUNO collaboration, *The JUNO experiment Top Tracker*, [arXiv:2303.05172](#) [[INSPIRE](#)].
- [27] OPERA collaboration, *Evidence for $\nu_\mu \rightarrow \nu_\tau$ appearance in the CNGS neutrino beam with the OPERA experiment*, *Phys. Rev. D* **89** (2014) 051102 [[arXiv:1401.2079](#)] [[INSPIRE](#)].
- [28] T. Adam et al., *The OPERA experiment target tracker*, *Nucl. Instrum. Meth. A* **577** (2007) 523 [[physics/0701153](#)] [[INSPIRE](#)].
- [29] J.F. Navarro, C.S. Frenk and S.D.M. White, *The Structure of cold dark matter halos*, *Astrophys. J.* **462** (1996) 563 [[astro-ph/9508025](#)] [[INSPIRE](#)].
- [30] A.V. Kravtsov, A.A. Klypin, J.S. Bullock and J.R. Primack, *The Cores of dark matter dominated galaxies: Theory versus observations*, *Astrophys. J.* **502** (1998) 48 [[astro-ph/9708176](#)] [[INSPIRE](#)].
- [31] B. Moore et al., *Cold collapse and the core catastrophe*, *Mon. Not. Roy. Astron. Soc.* **310** (1999) 1147 [[astro-ph/9903164](#)] [[INSPIRE](#)].
- [32] A. Strumia and F. Vissani, *Precise quasielastic neutrino/nucleon cross-section*, *Phys. Lett. B* **564** (2003) 42 [[astro-ph/0302055](#)] [[INSPIRE](#)].
- [33] P. Vogel and J.F. Beacom, *Angular distribution of neutron inverse beta decay, $\bar{\nu}_e + p \rightarrow e^+ + n$* , *Phys. Rev. D* **60** (1999) 053003 [[hep-ph/9903554](#)] [[INSPIRE](#)].
- [34] T.A. Mueller et al., *Improved Predictions of Reactor Antineutrino Spectra*, *Phys. Rev. C* **83** (2011) 054615 [[arXiv:1101.2663](#)] [[INSPIRE](#)].
- [35] P. Huber, *On the determination of anti-neutrino spectra from nuclear reactors*, *Phys. Rev. C* **84** (2011) 024617 [Erratum *ibid.* **85** (2012) 029901] [[arXiv:1106.0687](#)] [[INSPIRE](#)].
- [36] S. Ando and K. Sato, *Relic neutrino background from cosmological supernovae*, *New J. Phys* **6** (2004) 179.
- [37] S. Horiuchi, J.F. Beacom and E. Dwek, *The Diffuse Supernova Neutrino Background is detectable in Super-Kamiokande*, *Phys. Rev. D* **79** (2009) 083013 [[arXiv:0812.3157](#)] [[INSPIRE](#)].
- [38] A. Priya and C. Lunardini, *Diffuse neutrinos from luminous and dark supernovae: prospects for upcoming detectors at the $O(10)$ kt scale*, *JCAP* **11** (2017) 031 [[arXiv:1705.02122](#)] [[INSPIRE](#)].

- [39] SUPER-KAMIOKANDE collaboration, *Diffuse supernova neutrino background search at Super-Kamiokande*, *Phys. Rev. D* **104** (2021) 122002 [[arXiv:2109.11174](#)] [[INSPIRE](#)].
- [40] M. Honda et al., *Atmospheric neutrino flux calculation using the NRLMSISE-00 atmospheric model*, *Phys. Rev. D* **92** (2015) 023004 [[arXiv:1502.03916](#)] [[INSPIRE](#)].
- [41] G. Battistoni et al., *Overview of the FLUKA code*, *Annals Nucl. Energy* **82** (2015) 10 [[INSPIRE](#)].
- [42] C.Andreopoulos et al., *The GENIE Neutrino MC Generator*, *Nucl. Instrum. Meth. A* **87** (2010) 614.
- [43] T. Yoshida et al., *Neutrino-Nucleus Reaction Cross Sections for Light Element Synthesis in Supernova Explosions*, *Astrophys. J.* **686** (2008) 448 [[arXiv:0807.2723](#)] [[INSPIRE](#)].
- [44] KAMLAND collaboration, *Production of Radioactive Isotopes through Cosmic Muon Spallation in KamLAND*, *Phys. Rev. C* **81** (2010) 025807 [[arXiv:0907.0066](#)] [[INSPIRE](#)].
- [45] BOREXINO collaboration, *Cosmogenic Backgrounds in Borexino at 3800 m water-equivalent depth*, *JCAP* **08** (2013) 049 [[arXiv:1304.7381](#)] [[INSPIRE](#)].
- [46] J. Cheng, Y.-F. Li, L.-J. Wen and S. Zhou, *Neutral-current background induced by atmospheric neutrinos at large liquid-scintillator detectors: I. model predictions*, *Phys. Rev. D* **103** (2021) 053001 [[arXiv:2008.04633](#)] [[INSPIRE](#)].
- [47] S. Dytman, *Final state interactions in neutrino-nucleus experiments*, *Acta Phys. Polon. B* **40** (2009) 2445 [[INSPIRE](#)].
- [48] A.J. Koning, S. Hilaire and M.C. Duijvestijn, *TALYS: Comprehensive Nuclear Reaction Modeling*, *AIP Conf. Proc.* **769** (2005) 1154 [[INSPIRE](#)].
- [49] Y.A. Kamyshkov and E. Kolbe, *Signatures of nucleon disappearance in large underground detectors*, *Phys. Rev. D* **67** (2003) 076007 [[nucl-th/0206030](#)] [[INSPIRE](#)].
- [50] N. Auerbach, N. Van Giai and O.K. Vorov, *Neutrino scattering from C-12 and O-16*, *Phys. Rev. C* **56** (1997) R2368 [[nucl-th/9705003](#)] [[INSPIRE](#)].
- [51] E. Kolbe, K. Langanke and P. Vogel, *Weak reactions on C-12 within the continuum random phase approximation with partial occupancies*, *Nucl. Phys. A* **652** (1999) 91 [[nucl-th/9903022](#)] [[INSPIRE](#)].
- [52] J. Cheng, Y.-F. Li, H.-Q. Lu and L.-J. Wen, *Neutral-current background induced by atmospheric neutrinos at large liquid-scintillator detectors. II. Methodology for insitu measurements*, *Phys. Rev. D* **103** (2021) 053002 [[arXiv:2009.04085](#)] [[INSPIRE](#)].
- [53] R. Möllenberget al., *Detecting the Diffuse Supernova Neutrino Background with LENA*, *Phys. Rev. D* **91** (2015) 032005 [[arXiv:1409.2240](#)] [[INSPIRE](#)].
- [54] TMVA collaboration, *TMVA — Toolkit for Multivariate Data Analysis*, [physics/0703039](#) [[INSPIRE](#)].
- [55] F. Pedregosa et al., *Scikit-learn: Machine Learning in Python*, *J. Machine Learning Res.* **12** (2011) 2825 [[arXiv:1201.0490](#)] [[INSPIRE](#)].
- [56] JUNO collaboration, *Prospects for detecting the diffuse supernova neutrino background with JUNO*, *JCAP* **10** (2022) 033 [[arXiv:2205.08830](#)] [[INSPIRE](#)].
- [57] A. Caldwell and K. Kröninger, *Signal discovery in sparse spectra: A Bayesian analysis*, *Phys. Rev. D* **74** (2006) 092003 [[physics/0608249](#)] [[INSPIRE](#)].
- [58] D. Foreman-Mackey, D.W. Hogg, D. Lang and J. Goodman, *emcee: The MCMC Hammer*, *Publ. Astron. Soc. Pac.* **125** (2013) 306 [[arXiv:1202.3665](#)] [[INSPIRE](#)].
- [59] J. Goodman and J. Weare, *Ensemble samplers with affine invariance*, *Commun. Appl. Math. Comput. Sci.* **5** (2010) 65.

The JUNO collaboration

Angel Abusleme^{5,41}, Thomas Adam⁴⁶, Shakeel Ahmad⁶⁷, Rizwan Ahmed⁶⁷, Sebastiano Aiello⁵⁶, Muhammad Akram⁶⁷, Abid Aleem⁶⁷, Tsagkarakis Alexandros⁴⁹, Fengpeng An²⁹, Qi An²², Giuseppe Andronico⁵⁶, Nikolay Anfimov⁶⁸, Vito Antonelli⁵⁸, Tatiana Antoshkina⁶⁸, Burin Asavapibhop⁷², João Pedro Athayde Marcondes de André⁴⁶, Didier Auguste⁴⁴, Weidong Bai²⁰, Nikita Balashov⁶⁸, Wander Baldini⁵⁷, Andrea Barresi⁵⁹, Davide Basilico⁵⁸, Eric Baussan⁴⁶, Marco Bellato⁶¹, Antonio Bergnoli⁶¹, Thilo Birkenfeld⁴⁹, Sylvie Blin⁴⁴, David Blum⁵⁵, Simon Blyth¹⁰, Anastasia Bolshakova⁶⁸, Mathieu Bongrand⁴⁸, Clément Bordereau^{45,40}, Dominique Breton⁴⁴, Augusto Brigatti⁵⁸, Riccardo Brugnera⁶², Riccardo Bruno⁵⁶, Antonio Budano⁶⁵, Jose Busto⁴⁷, Anatael Cabrera⁴⁴, Barbara Caccianiga⁵⁸, Hao Cai³⁴, Xiao Cai¹⁰, Yanke Cai¹⁰, Zhiyan Cai¹⁰, Riccardo Callegari⁶², Antonio Cammi⁶⁰, Agustín Campeny⁵, Chuanya Cao¹⁰, Guofu Cao¹⁰, Jun Cao¹⁰, Rossella Caruso⁵⁶, Cédric Cerna⁴⁵, Chi Chan³⁸, Jinfan Chang¹⁰, Yun Chang³⁹, Guoming Chen²⁸, Pingping Chen¹⁸, Po-An Chen⁴⁰, Shaomin Chen¹³, Xurong Chen²⁶, Yixue Chen¹¹, Yu Chen²⁰, Zhiyuan Chen¹⁰, Zikang Chen²⁰, Jie Cheng¹¹, Yaping Cheng⁷, Yu Chin Cheng⁴⁰, Alexey Chetverikov⁶⁸, Davide Chiesa⁵⁹, Pietro Chimenti³, Ziliang Chu¹⁰, Artem Chukanov⁶⁸, Gérard Claverie⁴⁵, Catia Clementi⁶³, Barbara Clerbaux², Selma Conforti Di Lorenzo⁴⁵, Daniele Corti⁶¹, Flavio Dal Corso⁶¹, Olivia Dalager⁷⁵, Christophe De La Taille⁴⁵, Zhi Deng¹³, Ziyang Deng¹⁰, Wilfried Depnering⁵², Marco Diaz⁵, Xuefeng Ding⁵⁸, Yayun Ding¹⁰, Bayu Dirgantara⁷⁴, Sergey Dmitrievsky⁶⁸, Tadeas Dohnal⁴², Dmitry Dolzhikov⁶⁸, Georgy Donchenko⁷⁰, Jianmeng Dong¹³, Evgeny Doroshkevich⁶⁹, Wei Dou¹³, Marcos Dracos⁴⁶, Frédéric Druillolle⁴⁵, Ran Du¹⁰, Shuxian Du³⁷, Stefano Dusini⁶¹, Martin Dvorak⁴², Jessica Eck⁵⁵, Timo Enqvist⁴³, Andrea Fabbri⁶⁵, Ulrike Fahrendholz⁵³, Donghua Fan²⁴, Lei Fan¹⁰, Jian Fang¹⁰, Wenxing Fang¹⁰, Marco Fargetta⁵⁶, Dmitry Fedoseev⁶⁸, Zhengyong Fei¹⁰, Li-Cheng Feng³⁸, Qichun Feng²¹, Federico Ferraro⁵⁸, Richard Ford⁵⁸, Amélie Fournier⁴⁵, Haonan Gan³², Feng Gao⁴⁹, Alberto Garfagnini⁶², Arsenii Gavrikov⁶⁸, Marco Giammarchi⁵⁸, Nunzio Giudice⁵⁶, Maxim Gonchar⁶⁸, Guanghua Gong¹³, Hui Gong¹³, Yuri Gornushkin⁶⁸, Alexandre Göttel^{51,49}, Marco Grassi⁶², Vasily Gromov⁶⁸, Minghao Gu¹⁰, Xiaofei Gu³⁷, Yu Gu¹⁹, Mengyun Guan¹⁰, Yuduo Guan¹⁰, Nunzio Guardone⁵⁶, Cong Guo¹⁰, Jingyuan Guo²⁰, Wanlei Guo¹⁰, Xinheng Guo⁸, Yuhang Guo³⁵, Caren Hagner⁵⁰, Ran Han⁷, Yang Han²⁰, Miao He¹⁰, Wei He¹⁰, Tobias Heinz⁵⁵, Patrick Hellmuth⁴⁵, Yuekun Heng¹⁰, Rafael Herrera⁵, YuenKeung Hor²⁰, Shaojing Hou¹⁰, Yee Hsiung⁴⁰, Bei-Zhen Hu⁴⁰, Hang Hu²⁰, Jianrun Hu¹⁰, Jun Hu¹⁰, Shouyang Hu⁹, Tao Hu¹⁰, Yuxiang Hu¹⁰, Zhuojun Hu²⁰, Guihong Huang²⁴, Hanxiong Huang⁹, Kaixuan Huang²⁰, Wenhao Huang²⁵, Xin Huang¹⁰, Xingtao Huang²⁵, Yongbo Huang²⁸, Jiaqi Hui³⁰, Lei Huo²¹, Wenju Huo²², Cédric Huss⁴⁵, Safeer Hussain⁶⁷, Ara Ioannisian¹, Roberto Isocrate⁶¹, Beatrice Jelmini⁶², Ignacio Jeria⁵, Xiaolu Ji¹⁰, Huihui Jia³³, Junji Jia³⁴, Siyu Jian⁹, Di Jiang²², Wei Jiang¹⁰, Xiaoshan Jiang¹⁰, Xiaoping Jing¹⁰, Cécile Jollet⁴⁵, Jari Joutsenvaara⁴³, Leonidas Kalousis⁴⁶, Philipp Kampmann^{54,51}, Li Kang¹⁸, Rebin Karaparambil⁴⁸, Narine Kazarian¹, Amina Khatun⁷¹, Khanchai Khosonthongkee⁷⁴, Denis Korabely⁶⁸, Konstantin Kouzakov⁷⁰, Alexey Krasnoperov⁶⁸, Nikolay Kutovskiy⁶⁸, Pasi Kuusiniemi⁴³, Tobias Lachenmaier⁵⁵, Cecilia Landini⁵⁸, Sébastien Leblanc⁴⁵, Victor Lebrin⁴⁸, Frederic Lefevre⁴⁸, Ruiting Lei¹⁸, Rupert Leitner⁴², Jason Leung³⁸, Daozheng Li¹⁰, Demin Li³⁷, Fei Li¹⁰, Fule Li¹³, Gaosong Li¹⁰, Huiling Li¹⁰, Mengzhao Li¹⁰, Min Li¹⁰, Nan Li¹⁰, Nan Li¹⁶, Qingjiang Li¹⁶, Ruhui Li¹⁰, Rui Li³⁰, Shanfeng Li¹⁸, Tao Li²⁰, Teng Li²⁵, Weidong Li^{10,14}, Weiguo Li¹⁰, Xiaomei Li⁹, Xiaonan Li¹⁰, Xinglong Li⁹, Yi Li¹⁸, Yichen Li¹⁰, Yufeng Li¹⁰, Zepeng Li¹⁰, Zhaoan Li¹⁰, Zhibing Li²⁰, Ziyuan Li²⁰, Zonghai Li³⁴, Hao Liang⁹, Hao Liang²², Jiajun Liao²⁰, Ayut Limphirat⁷⁴, Guey-Lin Lin³⁸, Shengxin Lin¹⁸, Tao Lin¹⁰, Jiajie Ling²⁰, Ivano Lippi⁶¹, Fang Liu¹¹, Haidong Liu³⁷, Haotian Liu³⁴, Hongbang Liu²⁸, Hongjuan Liu²³, Hongtao Liu²⁰, Hui Liu¹⁹, Jianglai Liu^{30,31}, Jinchang Liu¹⁰,

Min Liu²³, Qian Liu¹⁴, Qin Liu²², Runxuan Liu^{51,49}, Shubin Liu²², Shulin Liu¹⁰, Xiaowei Liu²⁰, Xiwen Liu²⁸, Yan Liu¹⁰, Yunzhe Liu¹⁰, Alexey Lokhov^{70,69}, Paolo Lombardi⁵⁸, Claudio Lombardo⁵⁶, Kai Loo⁵², Chuan Lu³², Haoqi Lu¹⁰, Jingbin Lu¹⁵, Junguang Lu¹⁰, Peizhi Lu²⁰, Shuxiang Lu³⁷, Bayarto Lubsandorzhiev⁶⁹, Sultim Lubsandorzhiev⁶⁹, Livia Ludhova^{51,49}, Arslan Lukanov⁶⁹, Daibin Luo¹⁰, Fengjiao Luo²³, Guang Luo²⁰, Shu Luo³⁶, Wuming Luo¹⁰, Xiaojie Luo¹⁰, Vladimir Lyashuk⁶⁹, Bangzheng Ma²⁵, Bing Ma³⁷, Qiumei Ma¹⁰, Si Ma¹⁰, Xiaoyan Ma¹⁰, Xubo Ma¹¹, Jihane Maalmi⁴⁴, Marco Magoni⁵⁸, Jingyu Mai²⁰, Yury Malyshev⁵⁴, Roberto Carlos Mandujano⁷⁵, Fabio Mantovani⁵⁷, Xin Mao⁷, Yajun Mao¹², Stefano M. Mari⁶⁵, Filippo Marini⁶², Cristina Martellini⁶⁵, Gisele Martin-Chassard⁴⁴, Agnese Martini⁶⁴, Matthias Mayer⁵³, Davit Mayilyan¹, Ints Mednieks⁶⁶, Artur Meinusch⁵², Yue Meng³⁰, Anselmo Meregaglia⁴⁵, Emanuela Meroni⁵⁸, David Meyhöfer⁵⁰, Mauro Mezzetto⁶¹, Jonathan Miller⁶, Lino Miramonti⁵⁸, Marta Colomer Molla², Paolo Montini⁶⁵, Michele Montuschi⁵⁷, Axel Müller⁵⁵, Massimiliano Nastasi⁵⁹, Dmitry V. Naumov⁶⁸, Elena Naumova⁶⁸, Diana Navas-Nicolas⁴⁴, Igor Nemchenok⁶⁸, Minh Thuan Nguyen Thi³⁸, Feipeng Ning¹⁰, Zhe Ning¹⁰, Hiroshi Nunokawa⁴, Lothar Oberauer⁵³, Juan Pedro Ochoa-Ricoux^{75,5,41}, Alexander Olshevskiy⁶⁸, Domizia Orestano⁶⁵, Fausto Ortica⁶³, Rainer Othegegraven⁵², Alessandro Paoloni⁶⁴, Sergio Parmeggiano⁵⁸, Yatian Pei¹⁰, Luca Pelicci^{51,49}, Nicomede Pelliccia⁶³, Anguo Peng²³, Haiping Peng²², Yu Peng¹⁰, Zhaoyuan Peng¹⁰, Frédéric Perrot⁴⁵, Pierre-Alexandre Petitjean², Fabrizio Petrucci⁶⁵, Oliver Pilarczyk⁵², Luis Felipe Piñeres Rico⁴⁶, Artyom Popov⁷⁰, Pascal Poussot⁴⁶, Ezio Previtali⁵⁹, Fazhi Qi¹⁰, Ming Qi²⁷, Sen Qian¹⁰, Xiaohui Qian¹⁰, Zhen Qian²⁰, Hao Qiao¹², Zhonghua Qin¹⁰, Shoukang Qiu²³, Gioacchino Ranucci⁵⁸, Neill Raper²⁰, Reem Rasheed⁴⁵, Alessandra Re⁵⁸, Henning Rebber⁵⁰, Abdel Rebi⁴⁵, Mariia Redchuk^{62,61}, Bin Ren¹⁸, Jie Ren⁹, Barbara Ricci⁵⁷, Mariam Rifai^{51,49}, Mathieu Roche⁴⁵, Narongkiat Rodphai⁷², Aldo Romani⁶³, Bedřich Roskovec⁴², Xichao Ruan⁹, Arseniy Rybnikov⁶⁸, Andrey Sadovsky⁶⁸, Paolo Saggese⁵⁸, Simone Sanfilippo⁶⁵, Anut Sangka⁷³, Utane Sawangwit⁷³, Julia Sawatzki⁵³, Michaela Schever^{51,49}, Cédric Schwab⁴⁶, Konstantin Schweizer⁵³, Alexandr Selyunin⁶⁸, Andrea Serafini⁵⁷, Giulio Settanta^{51,a}, Mariangela Settimo⁴⁸, Zhuang Shao³⁵, Vladislav Sharov⁶⁸, Arina Shaydurova⁶⁸, Jingyan Shi¹⁰, Yanan Shi¹⁰, Vitaly Shutov⁶⁸, Andrey Sidorenkov⁶⁹, Fedor Šimkovic⁷¹, Chiara Sirignano⁶², Jaruchit Siripak⁷⁴, Monica Sisti⁵⁹, Maciej Slupecki⁴³, Mikhail Smirnov²⁰, Oleg Smirnov⁶⁸, Thiago Sogo-Bezerra⁴⁸, Sergey Sokolov⁶⁸, Julanan Songwadhana⁷⁴, Boonrucksar Soonthornthum⁷³, Albert Sotnikov⁶⁸, Ondřej Šrámek⁴², Warintorn Sreethawong⁷⁴, Achim Stahl⁴⁹, Luca Stanco⁶¹, Konstantin Stankevich⁷⁰, Dušan Štefánek⁷¹, Hans Steiger^{52,53}, Jochen Steinmann⁴⁹, Tobias Sterr⁵⁵, Matthias Raphael Stock⁵³, Virginia Strati⁵⁷, Alexander Studenikin⁷⁰, Jun Su²⁰, Shifeng Sun¹¹, Xilei Sun¹⁰, Yongjie Sun²², Yongzhao Sun¹⁰, Zhengyang Sun³⁰, Narumon Suwonjandee⁷², Michal Szelezniak⁴⁶, Jian Tang²⁰, Qiang Tang²⁰, Quan Tang²³, Xiao Tang¹⁰, Vidhya Thara Hariharan⁵⁰, Eric Theisen⁵², Alexander Tietzsch⁵⁵, Igor Tkachev⁶⁹, Tomas Tmej⁴², Marco Danilo Claudio Torri⁵⁸, Konstantin Treskov⁶⁸, Andrea Triossi⁴⁶, Giancarlo Troni⁵, Wladyslaw Trzaska⁴³, Cristina Tuve⁵⁶, Nikita Ushakov⁶⁹, Vadim Vedin⁶⁶, Giuseppe Verde⁵⁶, Maxim Vialkov⁷⁰, Benoit Viaud⁴⁸, Cornelius Moritz Vollbrecht^{51,49}, Cristina Volpe⁴⁴, Katharina Von Sturm⁶², Vit Vorobel⁴², Dmitriy Voronin⁶⁹, Lucia Votano⁶⁴, Pablo Walker^{5,41}, Caishen Wang¹⁸, Chung-Hsiang Wang³⁹, En Wang³⁷, Guoli Wang²¹, Jian Wang²², Jun Wang²⁰, Lu Wang¹⁰, Meifen Wang¹⁰, Meng Wang²³, Meng Wang²⁵, Ruiguang Wang¹⁰, Siguang Wang¹², Wei Wang²⁷, Wei Wang²⁰, Wenshuai Wang¹⁰, Xi Wang¹⁶, Xiangyue Wang²⁰, Yangfu Wang¹⁰, Yaoguang Wang¹⁰, Yi Wang¹⁰, Yi Wang¹³, Yi Wang²⁴, Yifang Wang¹⁰, Yuanqing Wang¹³, Yuman Wang²⁷, Zhe Wang¹³, Zheng Wang¹⁰, Zhimin Wang¹⁰, Zongyi Wang¹³, Apimook Watcharangkool⁷³, Wei Wei¹⁰, Wei Wei²⁵, Wenlu Wei¹⁰, Yadong Wei¹⁸, Kaile Wen¹⁰, Liangjian Wen¹⁰, Jun Weng¹³, Christopher Wiebusch⁴⁹, Steven Chan-Fai Wong²⁰, Bjoern Wonsak⁵⁰, Diru Wu¹⁰, Qun Wu²⁵, Zhi Wu¹⁰, Michael Wurm⁵²,

Jacques Wurtz⁴⁶, Christian Wysotzki⁴⁹, Yufei Xi³², Dongmei Xia¹⁷, Xiang Xiao²⁰, Xiaochuan Xie²⁸, Yuguang Xie¹⁰, Zhangquan Xie¹⁰, Zhao Xin¹⁰, Zhizhong Xing¹⁰, Benda Xu¹³, Cheng Xu²³, Donglian Xu^{31,30}, Fanrong Xu¹⁹, Hangkun Xu¹⁰, Jilei Xu¹⁰, Jing Xu⁸, Meihang Xu¹⁰, Yin Xu³³, Baojun Yan¹⁰, Qiyu Yan¹⁴, Taylor Yan⁷⁴, Wenqi Yan¹⁰, Xiongbo Yan¹⁰, Yupeng Yan⁷⁴, Changgen Yang¹⁰, Chengfeng Yang²⁸, Huan Yang¹⁰, Jie Yang³⁷, Lei Yang¹⁸, Xiaoyu Yang¹⁰, Yifan Yang¹⁰, Yifan Yang², Haifeng Yao¹⁰, Jiakuan Ye¹⁰, Mei Ye¹⁰, Ziping Ye³¹, Frédéric Yermia⁴⁸, Zhengyun You²⁰, Boxiang Yu¹⁰, Chiye Yu¹⁸, Chunxu Yu³³, Hongzhao Yu²⁰, Miao Yu³⁴, Xianghui Yu³³, Zeyuan Yu¹⁰, Zezhong Yu¹⁰, Cenxi Yuan²⁰, Chengzhuo Yuan¹⁰, Ying Yuan¹², Zhenxiong Yuan¹³, Baobiao Yue²⁰, Noman Zafar⁶⁷, Vitalii Zavadskyi⁶⁸, Shan Zeng¹⁰, Tingxuan Zeng¹⁰, Yuda Zeng²⁰, Liang Zhan¹⁰, Aiqiang Zhang¹³, Bin Zhang³⁷, Binting Zhang¹⁰, Feiyang Zhang³⁰, Guoqing Zhang¹⁰, Honghao Zhang²⁰, Jialiang Zhang²⁷, Jiawen Zhang¹⁰, Jie Zhang¹⁰, Jin Zhang²⁸, Jingbo Zhang²¹, Jinnan Zhang¹⁰, Mohan Zhang¹⁰, Peng Zhang¹⁰, Qingmin Zhang³⁵, Shiqi Zhang²⁰, Shu Zhang²⁰, Tao Zhang³⁰, Xiaomei Zhang¹⁰, Xin Zhang¹⁰, Xuanton Zhang¹⁰, Yinhong Zhang¹⁰, Yiyu Zhang¹⁰, Yongpeng Zhang¹⁰, Yu Zhang¹⁰, Yuanyuan Zhang³⁰, Yumei Zhang²⁰, Zhenyu Zhang³⁴, Zhijian Zhang¹⁸, Fengyi Zhao²⁶, Jie Zhao¹⁰, Rong Zhao²⁰, Runze Zhao¹⁰, Shujun Zhao³⁷, Dongqin Zheng¹⁹, Hua Zheng¹⁸, Yangheng Zheng¹⁴, Weirong Zhong¹⁹, Jing Zhou⁹, Li Zhou¹⁰, Nan Zhou²², Shun Zhou¹⁰, Tong Zhou¹⁰, Xiang Zhou³⁴, Jiang Zhu²⁰, Jingsen Zhu²⁹, Kangfu Zhu³⁵, Kejun Zhu¹⁰, Zhihang Zhu¹⁰, Bo Zhuang¹⁰, Honglin Zhuang¹⁰, Liang Zong¹³, Jiaheng Zou¹⁰, Sebastian Zwickel⁵³

¹ Yerevan Physics Institute, Yerevan, Armenia

² Université Libre de Bruxelles, Brussels, Belgium

³ Universidade Estadual de Londrina, Londrina, Brazil

⁴ Pontificia Universidade Católica do Rio de Janeiro, Rio de Janeiro, Brazil

⁵ Pontificia Universidad Católica de Chile, Santiago, Chile

⁶ Universidad Técnica Federico Santa María, Valparaíso, Chile

⁷ Beijing Institute of Spacecraft Environment Engineering, Beijing, China

⁸ Beijing Normal University, Beijing, China

⁹ China Institute of Atomic Energy, Beijing, China

¹⁰ Institute of High Energy Physics, Beijing, China

¹¹ North China Electric Power University, Beijing, China

¹² School of Physics, Peking University, Beijing, China

¹³ Tsinghua University, Beijing, China

¹⁴ University of Chinese Academy of Sciences, Beijing, China

¹⁵ Jilin University, Changchun, China

¹⁶ College of Electronic Science and Engineering, National University of Defense Technology, Changsha, China

¹⁷ Chongqing University, Chongqing, China

¹⁸ Dongguan University of Technology, Dongguan, China

¹⁹ Jinan University, Guangzhou, China

²⁰ Sun Yat-Sen University, Guangzhou, China

²¹ Harbin Institute of Technology, Harbin, China

²² University of Science and Technology of China, Hefei, China

²³ The Radiochemistry and Nuclear Chemistry Group in University of South China, Hengyang, China

²⁴ Wuyi University, Jiangmen, China

²⁵ Shandong University, Jinan, China, and Key Laboratory of Particle Physics and Particle Irradiation of Ministry of Education, Shandong University, Qingdao, China

²⁶ Institute of Modern Physics, Chinese Academy of Sciences, Lanzhou, China

²⁷ Nanjing University, Nanjing, China

²⁸ Guangxi University, Nanning, China

²⁹ East China University of Science and Technology, Shanghai, China

³⁰ School of Physics and Astronomy, Shanghai Jiao Tong University, Shanghai, China

³¹ Tsung-Dao Lee Institute, Shanghai Jiao Tong University, Shanghai, China

- ³² *Institute of Hydrogeology and Environmental Geology, Chinese Academy of Geological Sciences, Shijiazhuang, China*
- ³³ *Nankai University, Tianjin, China*
- ³⁴ *Wuhan University, Wuhan, China*
- ³⁵ *Xi'an Jiaotong University, Xi'an, China*
- ³⁶ *Xiamen University, Xiamen, China*
- ³⁷ *School of Physics and Microelectronics, Zhengzhou University, Zhengzhou, China*
- ³⁸ *Institute of Physics, National Yang Ming Chiao Tung University, Hsinchu, Taiwan*
- ³⁹ *National United University, Miao-Li, Taiwan*
- ⁴⁰ *Department of Physics, National Taiwan University, Taipei, Taiwan*
- ⁴¹ *Millennium Institute for SubAtomic Physics at the High-energy Frontier (SAPHIR), ICN2019_044, ANID, Chile*
- ⁴² *Charles University, Faculty of Mathematics and Physics, Prague, Czech Republic*
- ⁴³ *University of Jyväskylä, Department of Physics, Jyväskylä, Finland*
- ⁴⁴ *IJCLab, Université Paris-Saclay, CNRS/IN2P3, 91405 Orsay, France*
- ⁴⁵ *Univ. Bordeaux, CNRS, LP2i, UMR 5797, F-33170 Gradignan, France*
- ⁴⁶ *IPHC, Université de Strasbourg, CNRS/IN2P3, F-67037 Strasbourg, France*
- ⁴⁷ *Aix-Marseille Univ., CNRS/IN2P3, CPPM, Marseille, France*
- ⁴⁸ *SUBATECH, Université de Nantes, IMT Atlantique, CNRS-IN2P3, Nantes, France*
- ⁴⁹ *III. Physikalisches Institut B, RWTH Aachen University, Aachen, Germany*
- ⁵⁰ *Institute of Experimental Physics, University of Hamburg, Hamburg, Germany*
- ⁵¹ *Forschungszentrum Jülich GmbH, Nuclear Physics Institute IKP-2, Jülich, Germany*
- ⁵² *Institute of Physics and EC PRISMA⁺, Johannes Gutenberg Universität Mainz, Mainz, Germany*
- ⁵³ *Technische Universität München, München, Germany*
- ⁵⁴ *Helmholtzzentrum für Schwerionenforschung, Planckstrasse 1, D-64291 Darmstadt, Germany*
- ⁵⁵ *Eberhard Karls Universität Tübingen, Physikalisches Institut, Tübingen, Germany*
- ⁵⁶ *INFN Catania and Dipartimento di Fisica e Astronomia dell'Università di Catania, Catania, Italy*
- ⁵⁷ *Department of Physics and Earth Science, University of Ferrara and INFN Sezione di Ferrara, Ferrara, Italy*
- ⁵⁸ *INFN Sezione di Milano and Dipartimento di Fisica dell'Università di Milano, Milano, Italy*
- ⁵⁹ *INFN Milano Bicocca and University of Milano Bicocca, Milano, Italy*
- ⁶⁰ *INFN Milano Bicocca and Politecnico di Milano, Milano, Italy*
- ⁶¹ *INFN Sezione di Padova, Padova, Italy*
- ⁶² *Dipartimento di Fisica e Astronomia dell'Università di Padova and INFN Sezione di Padova, Padova, Italy*
- ⁶³ *INFN Sezione di Perugia and Dipartimento di Chimica, Biologia e Biotecnologie dell'Università di Perugia, Perugia, Italy*
- ⁶⁴ *Laboratori Nazionali di Frascati dell'INFN, Roma, Italy*
- ⁶⁵ *University of Roma Tre and INFN Sezione Roma Tre, Roma, Italy*
- ⁶⁶ *Institute of Electronics and Computer Science, Riga, Latvia*
- ⁶⁷ *Pakistan Institute of Nuclear Science and Technology, Islamabad, Pakistan*
- ⁶⁸ *Joint Institute for Nuclear Research, Dubna, Russia*
- ⁶⁹ *Institute for Nuclear Research of the Russian Academy of Sciences, Moscow, Russia*
- ⁷⁰ *Lomonosov Moscow State University, Moscow, Russia*
- ⁷¹ *Comenius University Bratislava, Faculty of Mathematics, Physics and Informatics, Bratislava, Slovakia*
- ⁷² *Department of Physics, Faculty of Science, Chulalongkorn University, Bangkok, Thailand*
- ⁷³ *National Astronomical Research Institute of Thailand, Chiang Mai, Thailand*
- ⁷⁴ *Suranaree University of Technology, Nakhon Ratchasima, Thailand*
- ⁷⁵ *Department of Physics and Astronomy, University of California, Irvine, California, U.S.A.*

^a *Present address: Istituto Superiore per la Protezione e la Ricerca Ambientale, 00144 Rome, Italy*