



Transition Region Brightening in a Moss Region and Their Relation with Lower Atmospheric Dynamics

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Abstract

Small-scale brightenings (SBs) are commonly observed in the transition region (TR) that separates the solar chromosphere from the corona. These brightenings, omnipresent in active region patches known as “moss” regions, could potentially contribute to the heating of active region plasma. In this study, we investigate the properties of SB events in a moss region and their associated chromospheric dynamics, which could provide insights into the underlying generation mechanisms of the SBs. We analyzed the data sets obtained by coordinated observations using the Interface Region Imaging Spectrograph and the Goode Solar Telescope at Big Bear Solar Observatory. We studied 131 SB events in our region of interest and found that 100 showed spatial and temporal matches with the dynamics observed in the chromospheric $H\alpha$ images. Among these SBs, 98 of them were associated with spicules that are observed in $H\alpha$ images. Furthermore, detailed analysis revealed that one intense SB event corresponded to an Ellerman bomb (EB), while another SB event consisted of several recurring brightenings caused by a stream of falling plasma. We observed that $H\alpha$ far wings often showed flashes of strong brightening caused by the falling plasma, creating an $H\alpha$ spectral profile similar to an EB. However, 31 of the 131 investigated SB events showed no noticeable spatial and temporal matches with any apparent features in $H\alpha$ images. Our analysis indicated that the predominant TR SB events in moss regions are associated with chromospheric phenomena primarily caused by spicules. Most of these spicules display properties akin to dynamic fibrils.

Unified Astronomy Thesaurus concepts: Solar active regions (1974); Solar spicules (1525); Solar fibrils (532); Solar chromosphere (1479); Solar transition region (1532)

Materials only available in the [online version of record](#): animations

1. Introduction

The solar atmosphere is dominated by intense magnetic activities where strong heating events occur throughout the different layers of the atmosphere. Small-scale brightenings (SBs), ranging in size from subarcseconds to a few arcseconds, are common features in the transition region (TR) images of the Sun (H. Tian et al. 2014b; J. Martínez-Sykora et al. 2015; H. Skogsrud et al. 2015) and could be potential sources of the heating of the upper solar atmosphere. Among these dynamic SBs, some appear roundish in shape, while others are thin elongated jet-like features. Although SBs with a variety of sizes and intensities are observed in different regions of the Sun, the mechanisms responsible for generating these SBs remain unclear.

In quiet-Sun and coronal hole regions, chromospheric spicules (J. M. Beckers 1968; R. G. Athay & T. E. Holzer 1982; A. C. Sterling 2000; B. De Pontieu et al. 2011; G. Tsiropoulou et al. 2012) are suggested to be a dominant source of fast-moving jet-like SBs in the TR network, often called network jets (H. Tian et al. 2014a; J. Martínez-Sykora et al. 2015; N. Narang

et al. 2016). Rapid blueshifted excursions (RBEs) are chromospheric jets observed as absorption features in far-blue wings of $H\alpha$ and Ca II IR lines due to their high line-of-sight velocities (Ø. Langangen et al. 2008; L. Rouppe van der Voort et al. 2009). They are believed to be on-disk counterparts of the high-speed chromospheric jets, so-called type II spicules (B. de Pontieu et al. 2007), and are often related to TR network jets (L. Rouppe van der Voort et al. 2015). These spicules are suggested to be generated by the reconnection process and often cause heating of more than TR temperatures (T. M. D. Pereira et al. 2014; L. Rouppe van der Voort et al. 2015; T. Samanta et al. 2019), mostly appearing as SB in the TR. SB events are also abundantly seen in internetwork regions of the quiet Sun and coronal holes. J. Martínez-Sykora et al. (2015) suggested that acoustic shocks could generate these SB events in the internetwork regions. They proposed that upward-moving acoustic waves in a nonmagnetic environment get converted into shocks due to a steep decrease in the density, resulting in heating.

The TR SB events are pervasive in active regions, including above sunspots. Several studies investigated SB events above sunspots (H. Tian et al. 2014b; G. J. M. Vissers et al. 2015; N. Deng et al. 2016; T. Samanta et al. 2017). Most of these TR brightening events within sunspots are short-lived and small in size. Coordinated observations of chromospheric and TR above sunspots (H. Tian et al. 2014b; S. K. Tiwari et al. 2016; T. Samanta et al. 2017) suggest that some of the TR SBs in



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penumbrae are heating signatures of chromospheric penumbral microjets (Y. Katsukawa et al. 2007). It is suggested that the TR SBs could be generated by magnetic reconnection (S. K. Tiwari et al. 2016) or due to falling plasma (T. Samanta et al. 2017).

Ellerman bombs (EBs), UV bursts, and explosive events often produce SBs around active regions; however, they generally create “intense” SB events in the TR (K. P. Dere et al. 1989; M. D. Ding et al. 1998; H. Peter et al. 2014; H. Tian et al. 2016; L. P. Chitta et al. 2017; P. R. Young et al. 2018; H. Tian et al. 2018; Y. Chen et al. 2019b). These events are produced by magnetic reconnection at different heights in the atmosphere. EBs are transient brightening events that appear in the wings of $H\alpha$ and $\text{Ca II } 8542 \text{ \AA}$ (F. Ellerman 1917; M. K. Georgoulis et al. 2002; R. J. Rutten et al. 2013). EBs generally show a significant intensity enhancement in the $H\alpha$ wings without any noticeable intensity enhancement in the core, which creates a mustache-like $H\alpha$ spectral profile. They are believed to be caused by magnetic reconnection between oppositely oriented field lines of U-shaped loops (H. Watanabe et al. 2008; E. Pariat et al. 2012). Recently, it was suggested that UV bursts (H. Peter et al. 2014) can also be generated by a similar reconnection process as EBs; however, the reconnection height may be higher in the solar atmosphere (Y. Chen et al. 2019c). Explosive events result from busty reconnection at TR heights, producing strong flows within the TR, which generally affects the TR spectral line profiles as well as producing localized intensity enhancement (Y. Chen et al. 2019a).

Interface Region Imaging Spectrograph (IRIS; B. De Pontieu et al. 2014) observations of active regions revealed short-lived, bright, grain-like features of size $\sim 0''.5$ and lifetime ~ 3 minutes at TR temperatures (H. Skogsrud et al. 2016). The study suggested that TR SBs (bright grains in their paper) in plage regions result from acoustic shocks that drive dynamic fibrils. Dynamic fibrils are fine, thread-like structures mostly observed near the $H\alpha$ line core in active region plagues (B. De Pontieu et al. 2007). They usually follow parabolic paths with a symmetrical ascending and descending phase, typically having speeds of around $15\text{--}20 \text{ km s}^{-1}$. These structures are believed to form due to evanescent photospheric p -modes leaking into the atmosphere through inclined magnetic fields, leading to shock formation, and are commonly known as type I spicules. Recently, coronal signatures of these chromospheric dynamic fibrils have also been identified by using high-resolution observations from the Extreme Ultraviolet Imager (P. Rochus et al. 2020) on board the Solar Orbiter (S. Mandal et al. 2023a, 2023b). On the other hand, recent active region observations by N. Vilangot Nhalil et al. (2022) showed that many RBEs are related to TR SB events in the plage regions. Note the presence of RBEs and their association with TR emission are generally investigated in quiet-Sun and coronal hole regions (T. M. D. Pereira et al. 2014; L. Rouppe van der Voort et al. 2015) and rarely reported in active regions. These investigations clearly show that the exact generation mechanism of SB events is not settled. Therefore, it is crucial to delve into more high-resolution data and perform statistical analyses to gain a better understanding of the relationship between SBs and chromospheric dynamics.

We note that several detailed studies of “intense” brightening events have been conducted in recent years in active regions (P. R. Young et al. 2018). However, using high-resolution data, only a few studies investigated the most abundant and faint SBs

in detail within active region “moss.” Moss regions are finely textured and low-lying EUV emission patches above the active region plage (P. C. H. Martens et al. 2000; D. Tripathi et al. 2010). This low-lying emission was suggested to be coming from the locations of the footpoints of coronal loops in the dense cores of active regions (T. E. Berger et al. 1999; S. K. Antiochos et al. 2003; D. Tripathi et al. 2010) where many SBs are observed in the TR. We can now resolve incredibly small brightening features in the lower solar atmosphere and study their contributions to the heating of the TR and corona with the progress and developments in observational solar physics such as the 1.6 m Goode Solar Telescope (GST; W. Cao et al. 2010) at Big Bear Solar Observatory (BBSO), IRIS, and the Atmospheric Imaging Assembly (AIA; J. R. Lemen et al. 2012) on board the Solar Dynamics Observatory. In this article, we explored the properties of TR SB events in a moss region and their chromospheric connection in detail using high-resolution IRIS and GST data.

2. Observations

We analyzed the data recorded by coordinated observations between the 1.6 m GST at BBSO and the IRIS on 2017 May 27. IRIS performed a 16-step raster of NOAA AR 12659 at a step size of $2''$, covering $120''$ along the slit. It provided 2832, 2796, and 1330 \AA slit-jaw images (SJIs) with cadences of 83, 21, and 21 s, respectively, with a plate scale of $\sim 0''.165 \text{ pix}^{-1}$. The emission in the IRIS 2796 \AA filter is dominated by the emission from the upper chromosphere ($\sim 10^4 \text{ K}$), whereas the IRIS 1330 \AA filter is sensitive to lower TR temperatures ($\sim 10^{4.4} \text{ K}$). The Visible Imaging Spectrometer of GST recorded narrowband images of the $H\alpha$ line core and $H\alpha$ wings at ± 1 , ± 0.8 , ± 0.6 , ± 0.4 , and $\pm 0.2 \text{ \AA}$, alternately, with a spatial pixel size of $\sim 0''.029$ and a cadence of 53 s approximately. The $H\alpha$ line core generally forms at the chromospheric heights, whereas the wings form at lower atmospheric heights (J. Leenaarts et al. 2006, 2012). We used 2832 \AA SJIs to coalign the IRIS SJIs with the GST data set by comparing common photospheric features with $H\alpha - 1 \text{ \AA}$ wing images for each frame. Note that the mentioned data sets were previously utilized by Y. Chen et al. (2019c), and the detailed data reduction techniques have already been outlined there. We analyzed the recorded data during 17:50–19:24 UT, and our region of interest (ROI) ranged from $720''$ to $745''$ solar X and from $245''$ to $263''$ solar Y. It is worth mentioning that the active region we are studying exhibits intense events like large jets and flares. Therefore, we selected the ROI based on the availability of GST data and to avoid strong magnetic activities (e.g., flare, large brightenings, and larger jets), which allowed us to concentrate on studying the SB events at active region moss. Furthermore, note that AIA and the Helioseismic and Magnetic Imager (HMI; J. Schou et al. 2012) data on board the SDO satellite are used for reference purposes and clearly identifying the active region plage (see Figure 1).

3. Data Analysis and Results

This study focuses on examining TR SB events in a moss region and their chromospheric connections (Figure 1). The moss regions around the sunspots are clearly visible in the AIA 1600 and 1700 \AA passbands and SJI 1330 \AA . Figure 1 and the associated animation show many dynamic TR SB events

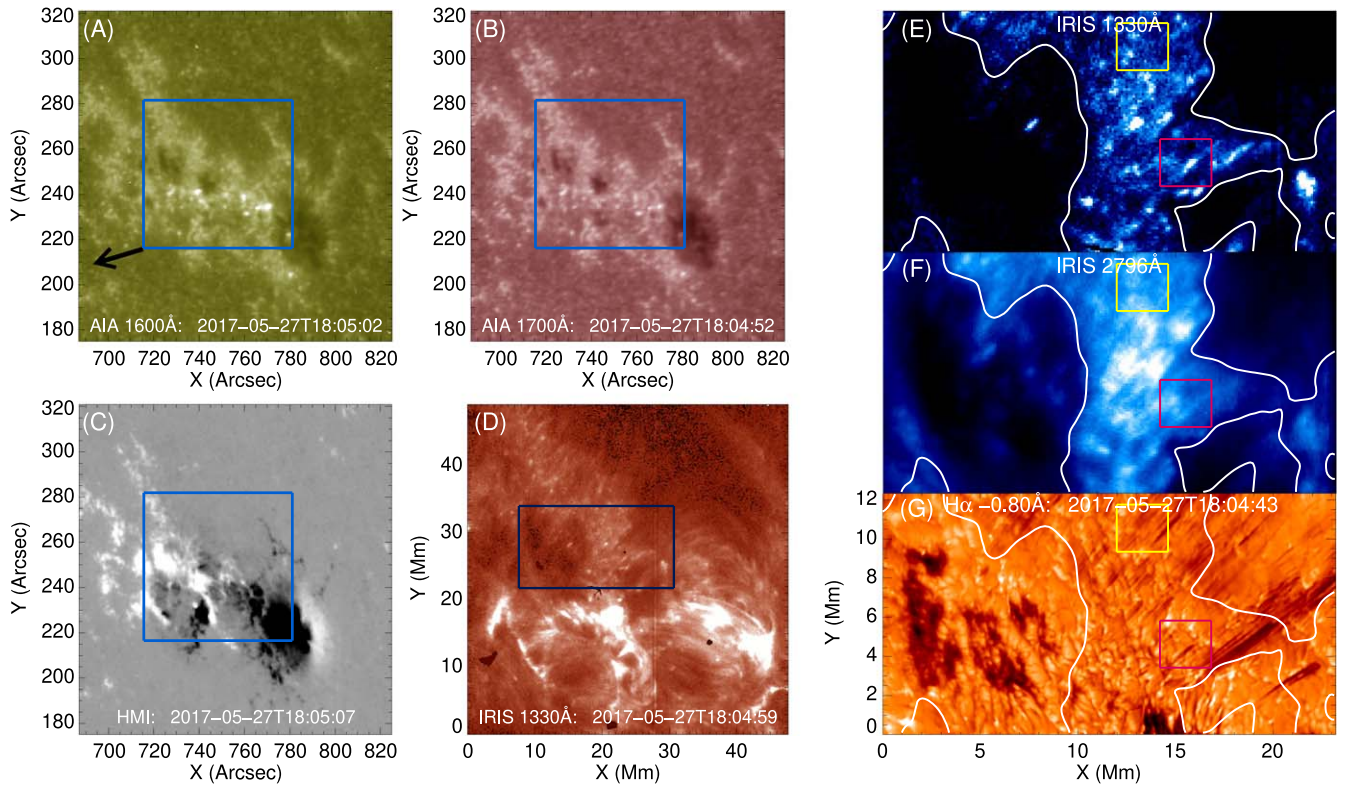


Figure 1. (A)–(B) The plage (moss) region in AIA 1600 and 1700 Å filters, respectively, at 18:05:02 UT on 2017 May 27. (C) The HMI line-of-sight magnetic field map of the same region. The black arrow in panel (A) points toward the disk center. The pale blue rectangle in panels (A)–(C) shows the field of view (FOV) of SJI 1330 Å in panel (D). Many TR SB events are clearly observed in the SJI 1330 Å (panel (D)). The dark blue rectangle marks our ROI and is shown in panels (E)–(G) in different filters as marked. The white contours, derived from a smoothed SJI 1330 Å, represent the boundary of the moss region in SJI 1330 Å. The red square represents the FOV shown in Figure 2, where we present examples of SBs in SJI 1330 Å that clearly match with spicules in the wings of H α . The yellow square corresponds to the FOV of Figure 4, which shows an example of brightening in 1330 Å images that do not match with any apparent dynamic features in H α images. Note that H α images in all the figures are displayed in the logarithmic scale. An animation of this figure presents the entire FOV as seen by the IRIS 1330 Å filter throughout the complete observational period. The animation includes the ROI highlighted in panels (E), (F), and (G), providing an overview of the ROI. (An animation of this figure is available in the [online article](#).)

visible in the IRIS SJI 1330 Å within the moss regions. We have mostly focused on the small part of the moss region marked by the white contours where most of the SB events appear. Initially, we visually investigated these TR SB events and noted that spicules (dark collimated jets observed in H α wings) are a common occurrence in this region. We selected those SBs that are observed in three or more consecutive 1330 Å images. Afterward, we searched for any corresponding chromospheric counterparts of these SBs. We found that a few SBs were observed at locations where numerous chromospheric features were present simultaneously, making it challenging to identify the relevant ones for the respective SBs accurately. Consequently, we excluded those SBs from our study. Based on these criteria, 131 SB events were identified in our ROI and selected for further analysis. Finally, we focused on exploring the chromospheric counterparts of these SBs in detail.

We first started investigating whether spicules could be a counterpart of some of the SBs. We define a good spatiotemporal match between an SB and a spicule when a spicule appears within a radius of $\sim 0''.5$ of the structure of an SB during its lifetime. This spatial offset between SBs and spicules can be caused by the projection effect since the ROI is located close to the limb. Figure 2 shows an example of a clear spatiotemporal match between SBs and spicules. The location of this observation is marked in Figure 1 (E)–(G) by a red-

colored rectangle. The time evolution of the GST H α wing images and IRIS SJIs (panels (A)–(D)) shows the two SB events that occur at the top of spicules. We further produced spacetime plots (panels (E)–(H)) to demonstrate this connection clearly. Note that the spacetime plots of IRIS SJIs in Figure 2 (also in Figures 3 and 4) are derived by taking the average intensity of all the pixels along the width of the slit. Additionally, the spacetime plots of H α images are derived by taking the minimum intensity of all the pixels along the width of the slit. The spatiotemporal match between spicules and SB events in SJI 1330 Å images is clearly visible in these spacetime plots. Detailed analysis of all the events reveals that 98 SB events in SJI 1330 Å images are related to spicules. These SBs show good spatiotemporal matches with spicules and typically follow their tips. We further notice that many of these SBs also appear in 2796 Å images; however, most of them map the middle part of these spicules (see bottom panels of Figure 2).

We conducted a more detailed investigation to determine whether these TR SB-related spicules can be classified as either dynamic fibrils or RBEs. Figure 3 provides insights into the distinct characteristics of the spicules that are linked to SB events. Three rows in the figure depict three prevalent types of observational features that we identified through our analysis. In order to comprehend their dynamics better, we present

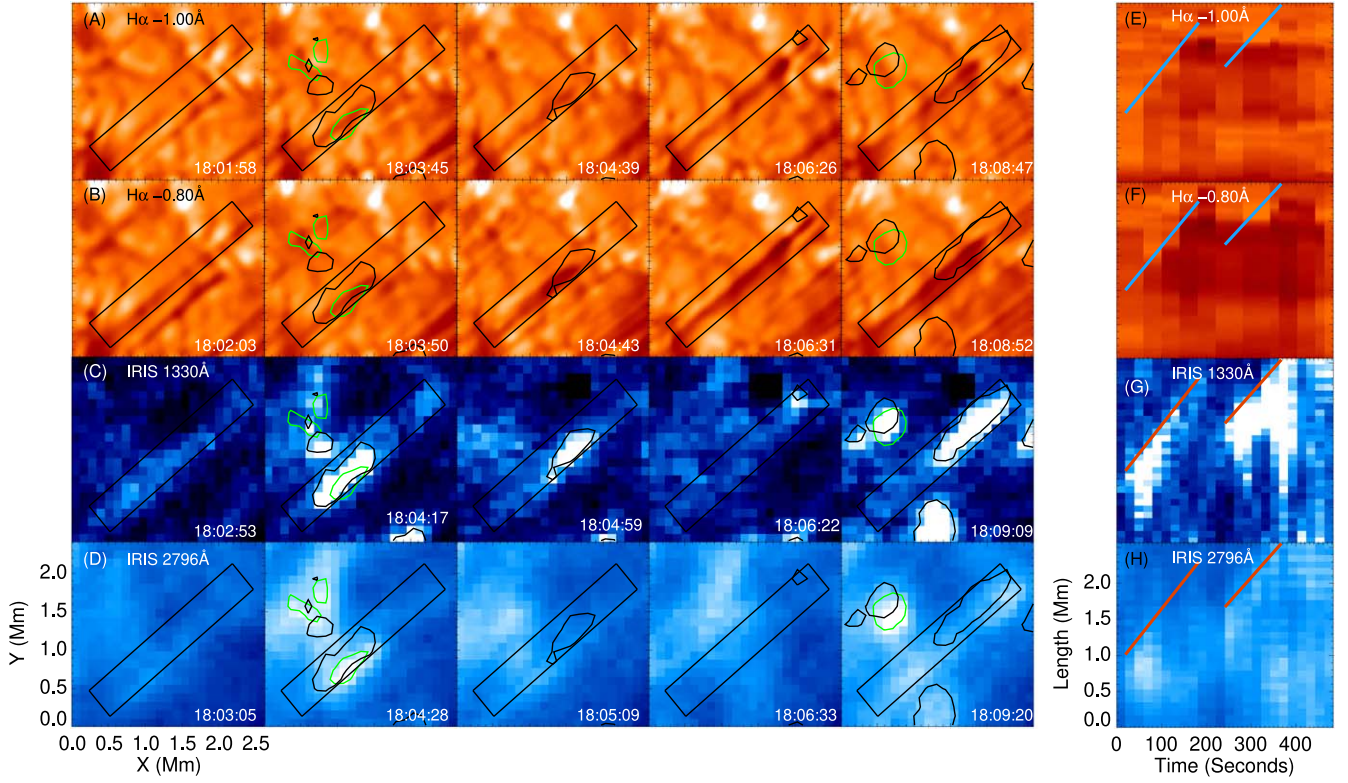


Figure 2. Time evolution of TR SB events that clearly match with the evolution of two spicules. (A)–(D) The FOV of the red square in panels (E)–(G) of Figure 1. (A) and (B) Time evolution of two spicules in $H\alpha$ -1.00 and -0.80 Å. (C) and (D) Time evolution of two SB events in SJI 1330 and 2796 Å. The black and green contours correspond to the SBs seen in SJI 1330 and 2796 Å images, respectively. The black rectangle corresponds to a $0''.8$ -wide artificial slit taken to track the spicules and their counterparts in IRIS SJI. The SBs (contoured) in SJI 1330 Å can clearly be seen following the ends of the spicules in $H\alpha$ as they move upward. SBs in SJI 2796 Å only appear for a single frame, but faint brightening can be seen mapping the body of the spicules throughout its lifetime. (E)–(H) Spacetime plots for the artificial slit shown as a black rectangle in panels (A)–(D). Inclined lines mark the motion of two SB events corresponding to two spicules.

spacetime diagrams in panels (D)–(F) and showcase their corresponding $H\alpha$ spectral profiles in panels (G)–(I).

The top panels illustrate an example of a spicule that follows a parabolic trajectory in the spacetime plots, resembling dynamic fibrils (e.g., B. De Pontieu et al. 2007). The $H\alpha$ spectral profile exhibits a slight broadening compared to the average profile. Notably, only a small fraction of these spicules (9 out of 98) display a clear parabolic trajectory in the $H\alpha$ spacetime diagrams. The middle panel presents another example where the collimated jets fade away after reaching their apex, without a distinct indication of falling back. The spectral profile shows a slight broadening of the line without noticeable asymmetry. Although the abrupt vanishing of the jets may suggest characteristics of an RBE, the absence of blue/red wing enhancement in the $H\alpha$ line profile (as seen in the red curve in panel (H)) prevents us from categorizing these jets as RBEs. Conversely, the spicule in the bottom row exhibits an ambiguous trajectory in the spacetime diagrams, along with a spectral profile similar to the previous two types with slight broadening. These predominant spicules either fade upon reaching their peak heights or display trajectories that are challenging to classify. We perform detailed calculations on the kinematics of these spicules, including the deceleration of 7 spicules with parabolic trajectories and the speeds of 15 spicules that fade after reaching their peak heights. These calculations are performed by tracing their trajectories in their spacetime plots. The deceleration values presented here represent the average apparent deceleration shown by these spicules during the rising phase in $H\alpha$ core spacetime plots,

while the apparent speeds are the average speeds from the point they appear to when they fade away in $H\alpha$ -0.80 Å wing spacetime plots. The deceleration values for these spicules range from 40 to 211 m s^{-2} , while the apparent rise velocities vary between 5 and 25 km s^{-1} . These values are on the lower end of the range observed by B. de Pontieu et al. (2007). In summary, this analysis indicates that these spicules generally demonstrate broadened profiles, decelerations, and speeds akin to those caused by dynamic fibrils.

Detailed analysis also revealed that several SB events do not have spatiotemporal correspondence with any dynamical features in $H\alpha$ images. Note that, consistent with our previous analysis, we searched for any noticeable spatiotemporal dynamics in $H\alpha$ images corresponding to an SB that may occur within approximately $0''.5$ of the structure the SB's radius during its lifetime. Figure 4 highlights an example where the SB event does not show an apparent spatiotemporal match with a spicule or any other feature in $H\alpha$ wings. Note that the location of this event is also marked in Figure 1 by a yellow-colored rectangle. Figure 4 presents the time evolution of such an SB event in SJI 1330 and 2796 Å images (see panels (C) and (D)). Panels (A) and (B) clearly reveal that there is no apparent signature of the corresponding chromospheric features inside the black rectangle in the wings of $H\alpha$ (panels (A)–(B)). Panels (E)–(F) show the spacetime plots of $H\alpha$ -0.40 and -0.80 Å wings, respectively, for the slit marked by the black rectangles in panels (A)–(D). Meanwhile, panels (G)–(H) are the spacetime plots of SJI 1330 and SJI 2796 Å for the same slit. The spacetime maps also make it evident that there is no apparent

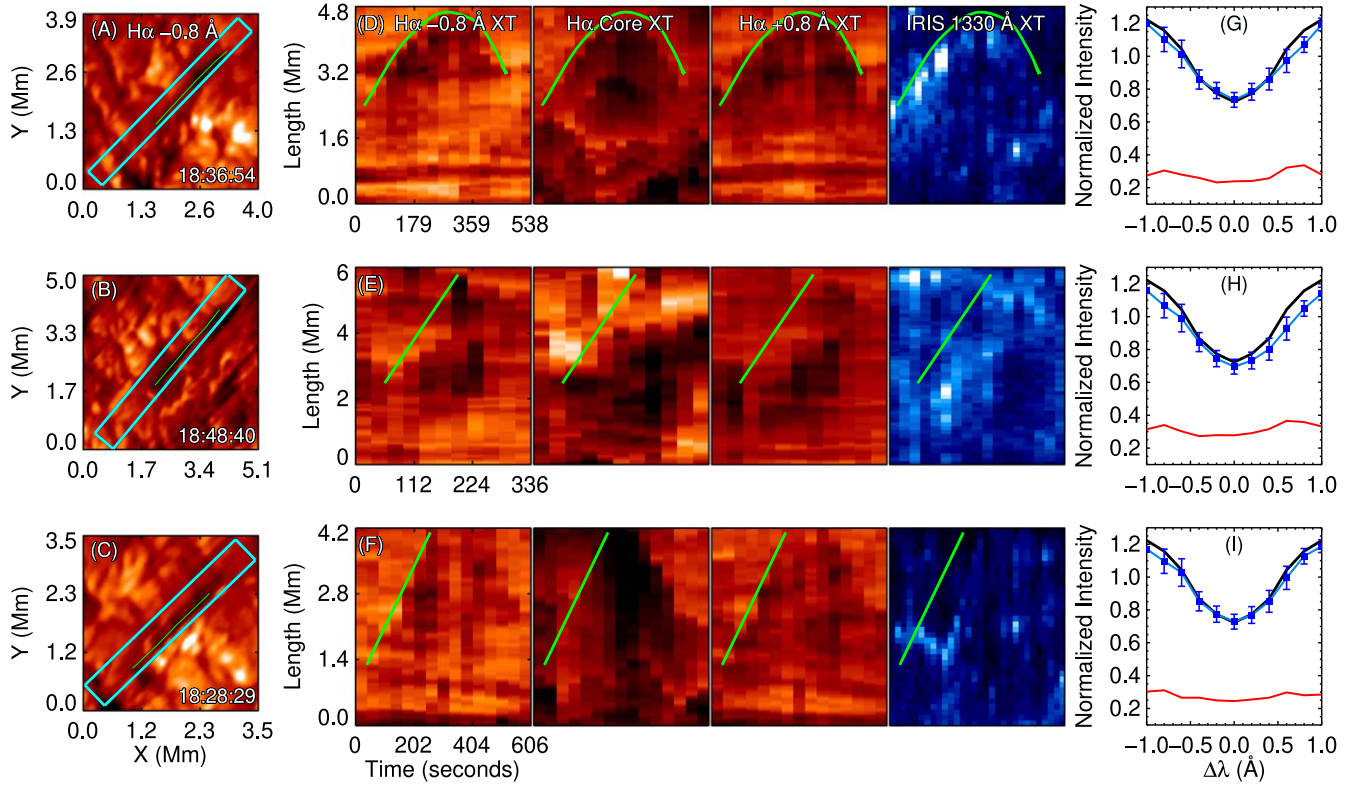


Figure 3. Examples of three spicules with different spectral profiles that have a correspondence with TR SB events. (A)–(C) The spicules as seen in $H\alpha$ -0.80 Å wing. (D)–(F) Their spacetime diagrams for the cyan-colored artificial slits in panels (A)–(C). From left to right in each panel, we display the $H\alpha$ -0.80 Å, line core, $+0.80$ Å wing, and IRIS 1330 Å, respectively. The spicule shown in panel (A) follows a parabolic trajectory, which is marked by a green curve. The green line in panels (E) and (F) also marks the trajectory followed by the spicules, as seen in the $H\alpha$ -0.80 Å wing. (G)–(I) The black curve shows the average $H\alpha$ spectral profile. The blue curve corresponds to the $H\alpha$ spectral profiles of the respective spicules calculated by averaging the intensity of the 3 pixel-wide slit along the jet shown as a green line in panels (A)–(C). The error bars represent 1σ uncertainties in the measurements. The red curve shows the difference between the average spectral profile and the spectral profile of the spicule.

sign of any interrelated features. We observed 31 such events (out of 131 investigated events) where SBs do not show any apparent spatial and temporal match with any visible features in $H\alpha$ wings, which suggests that either the corresponding features have a width below the resolution limit of our observational data or these SB events are entirely a TR phenomenon.

Our investigation reveals long-lasting intense recurring brightenings in SJI 1330 and 2796 Å in the neighborhood of the brightenings in $H\alpha$ far wings (panels (C)–(D) of Figure 5). For further analysis, we calculate $H\alpha$ spectral profiles of these bright features by taking the intensity of the brightest pixel inside the small squares (boxes of 7×7 pixels) at which the arrows are pointing in panels (C), (E), and (G). The thick black line in panel (B) is the reference profile derived by taking the average intensity of a large region for the entire observational period. The $H\alpha$ spectral profiles (drawn in green and blue in panel B) corresponding to the SB event in panels (C) and (E) look like those of an EB. These profiles show a large intensity enhancement at the wings and remain faint in the core throughout the observational period. Additionally, these brightenings in the $H\alpha$ $+1.00$ Å wing images exhibit an intensity enhancement exceeding 3σ above the mean. Using the same data, Y. Chen et al. (2019c) also reported many similar events in this active region. They suggested that the brightenings in SJI 1330 Å near EBs can occur if the reconnection current sheet extends well above the temperature minimum region. When the reconnection near the temperature minimum region generates

EBs, the same reconnection current sheet at chromospheric heights can produce UV bursts. The spatial offset between SBs and EBs can be a consequence of this height difference and the projection effect since the ROI is close to the limb. Therefore, the temporal match between these bright features becomes crucial to identify their corresponding features. The EB events, often occurring in fragmented bursts, persisted throughout the entire duration, reappearing intermittently in the same vicinity. Sometimes, the SBs appear right above the EB. We measured the distance between the centers of the EBs and the SBs, finding offsets primarily in the radially outward direction, with SBs appearing above the EBs. The example in panels (C)–(D) shows an offset of around 900 km between the EB and SB, while the EB and SB in panels (E)–(F) exhibit almost no offset. The distances during different occurrences varied from 0 to 1000 km, with a median of around 600 km. This range is consistent with the findings of Y. Chen et al. (2019c), who reported similar offsets in their study of EB and UV bursts.

Our analysis also reveals recurring transient brightenings in SJI 1330 Å, far wings of $H\alpha$, and sometimes in SJI 2796 Å near the footpoints of a stream of falling plasma (see the animation associated with Figure 5). These prominent $H\alpha$ wing brightenings appear to be similar to that reported by J. Jing et al. (2016), where the falling plasma along the flare loops causes strong brightenings at their footpoint. We investigated the $H\alpha$ spectral profiles of these events in detail. An example of the falling-plasma-related brightening event is presented in panels (G) and (H) of Figure 5. The $H\alpha$ spectral profile at the region of

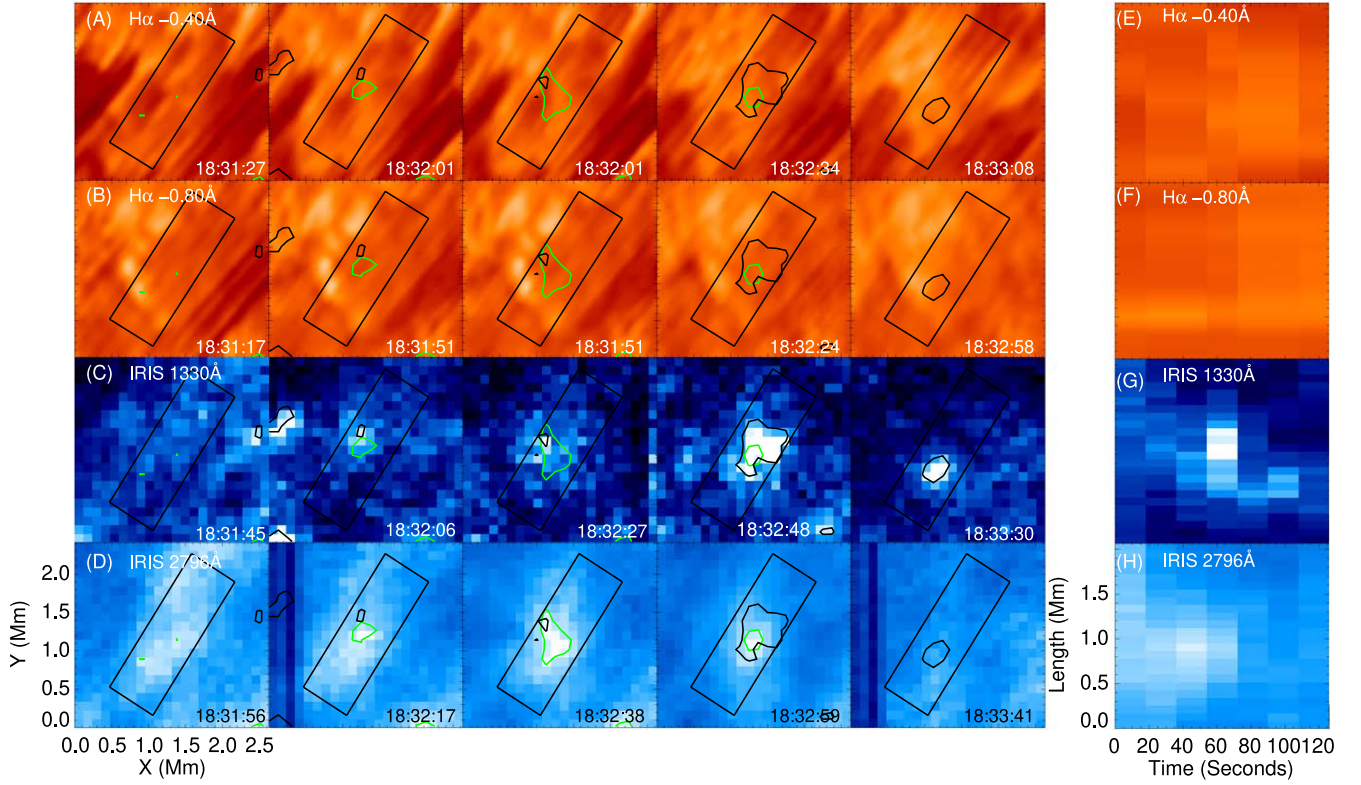


Figure 4. A typical example of an SB event that does not correspond with any apparent feature in $H\alpha$. The panels are the same as in Figure 2. The SBs in SJI 1330 and 2796 Å (panels in (A)–(B)) do not show any counterparts. The spacetime plots ((E)–(H)) also clearly show no signature of the IRIS SBs in $H\alpha$ wings.

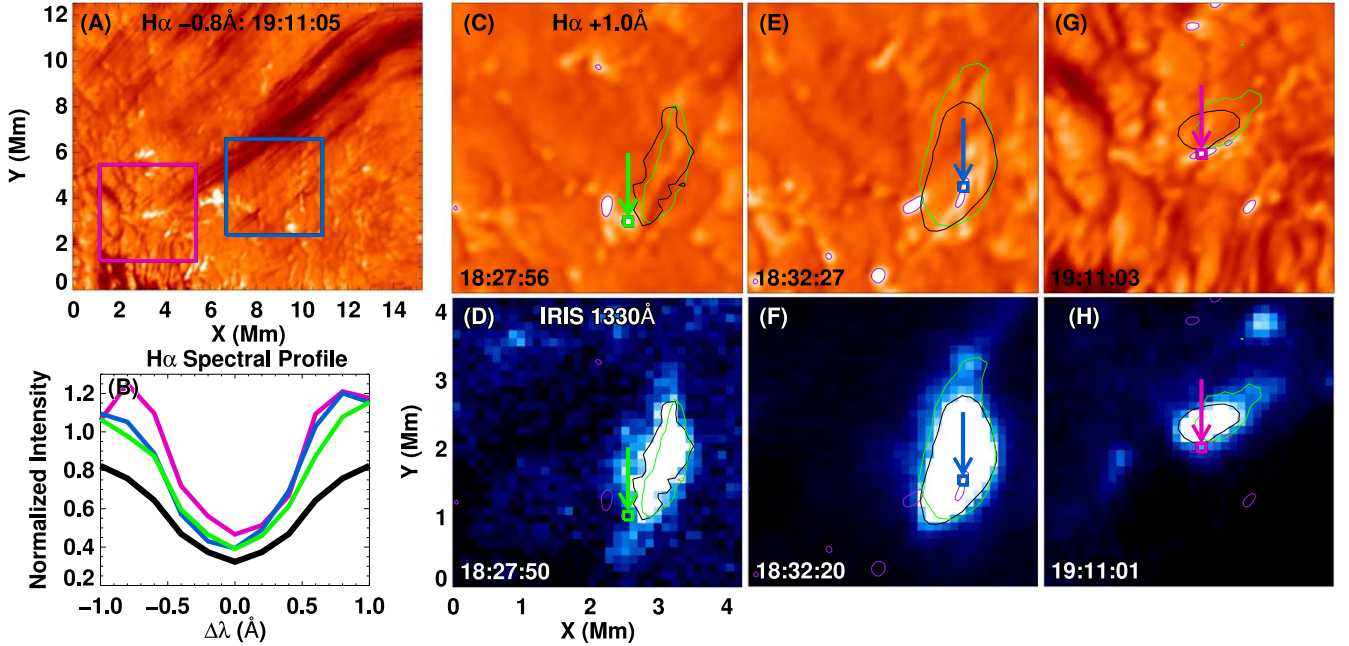


Figure 5. Relation of SBs with EB and falling plasma. (A) Large FOV covering the falling plasma and an EB in $H\alpha$ -0.8 Å at 19:11:05 UT on May 27. The blue square corresponds to the FOV of panels (C) and (E), and the pink square outlines the FOV of panel (G). (B) $H\alpha$ spectral profile: the thick black line shows the average spectral profile, whereas the green, blue, and pink lines correspond to the spectral profiles of the events shown in panels (C), (E), and (G), respectively. (C) and (E) The EB in $H\alpha$ $+1.00$ Å. (D) and (F) The same FOV as panels (C) and (E), covering profound brightening near the EB in SJI 1330 Å. Purple contours indicate 3σ intensity enhancement above the mean in $H\alpha$ $+1.00$ Å wing images. Black and green contours map the brightening in SJI 1330 and 2796 Å images, respectively. (G) The brightening in $H\alpha$ $+1.00$ Å images related to falling plasma. (H) SJI 1330 Å, which shows SBs near the brightening in $H\alpha$ $+1.00$ Å. The green, blue, and pink arrows point to the location of the brightenings in $H\alpha$ $+1.00$ Å. The green and blue arrows point to the brightening of an EB, whereas the pink arrow points to the falling-plasma-related brightening. An animation for this figure includes panels (A) and (C)–(H), where the downflows and brightenings are clearly visible.

(An animation of this figure is available in the [online article](#).)

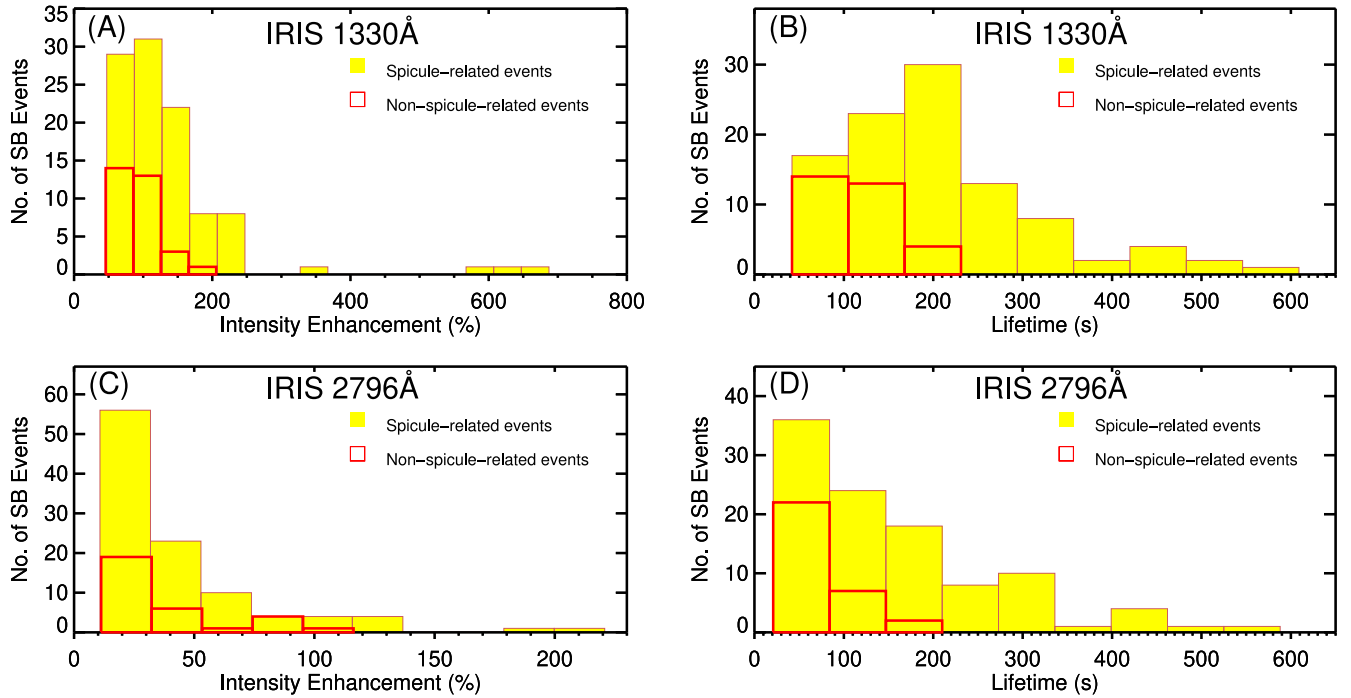


Figure 6. (A)–(B) Distribution of intensity enhancements and lifetimes of SB events in SJI 1330 Å, respectively. (C)–(D) Distribution of intensity enhancements and lifetimes of SB events in SJI 2796 Å, respectively. The solid yellow bins represent those SB events that match with the spicules, while the hollow red bins correspond to those SB events that do not match with any features in the $H\alpha$ images.

this brightening is shown in pink in panel (B). We note that the $H\alpha$ spectral profile looks similar to an EB. However, sometimes, the line core also shows intensity enhancement along with the wings. The SBs in IRIS passbands appear cotemporally near the falling-plasma regions. These SBs tend to move erratically along with their corresponding $H\alpha$ brightenings. SBs corresponding to these falling-plasma-related brightening events have different shapes and sizes. The evolution of these events is extremely difficult to track due to the temporal resolution of the GST data set and the small sample size. Some of these events only appear for a few frames (~ 1 – 4) in $H\alpha$ images, while others last throughout the entire observational period. Furthermore, some of the falling-plasma-related brightening events merge with each other, while some of them split into two or more, making it difficult to track these individual events. Additionally, we estimated the speeds of the falling plasma by producing spacetime plots along the falling streams and found that the falling-plasma speeds range from 15 to 20 km s^{-1} .

Our analysis showed that the majority of the SBs accompany spicules; however, there are many SBs where no associated dynamic changes are observed in the chromosphere. We compared the properties of these two different types of SB events, primarily focusing on their intensity enhancements relative to background intensity and their lifetime. To estimate these parameters, we utilized the slits in panels (C)–(D) in Figures 2 and 4. By examining spacetime diagrams of SB events (panels (G)–(H) in Figures 2 and 4), we determined their lifetime by setting an intensity threshold for each IRIS passband. The pixels exceeding this threshold in the frame were isolated, and their maximum intensities were recorded. Additionally, we calculated the average intensity below this threshold. The pixels exceeding this threshold in the frame were isolated, and their maximum intensities were recorded. Additionally, we calculated the minimum average intensity of

the isolated pixels throughout the lifetime of an SB event, which served as the background intensity. The ratio of the maximum intensity to the average background intensity yielded the percentage of intensity enhancement attributed to each SB event.

Figure 6 illustrates the distribution of intensity enhancements and lifetimes of SB events for both the 1330 and 2796 Å passbands. The histograms clearly demonstrate that SB events exhibiting a cospatial and cotemporal correspondence with spicules (depicted in yellow) are generally brighter and have longer lifetimes compared to those lacking spatial-temporal dynamics in the chromosphere (depicted in red). However, it is important to highlight that the characteristics of nonspicule-related SB events fall at the lower end of the spectrum compared to SB events caused by spicules. It is plausible that these SB events represent a scaled-down version of spicule-related SB events, with spicules themselves potentially going undetected in $H\alpha$ images due to current observational limitations. Additionally, it is worth noting that SBs associated with EBs and falling-plasma-related events exhibit significantly higher brightness and longer durations than the commonly observed SBs in this moss region. Specifically, the maximum intensity enhancement attributed to SBs due to falling plasma reached 522%, while the intensity enhancement for the SB event associated with EB was as high as 1700% in SJI 1330 Å images.

4. Discussion

Higher-resolution observations of the TR have revealed a plethora of brightening events originating from different shapes of magnetic reconnection at different heights, e.g., through shocks (H. Tian et al. 2014a; H. Skogsrud et al. 2015). Using coordinated observations of GST and IRIS, we have examined a part of an active region plage/moss to understand the

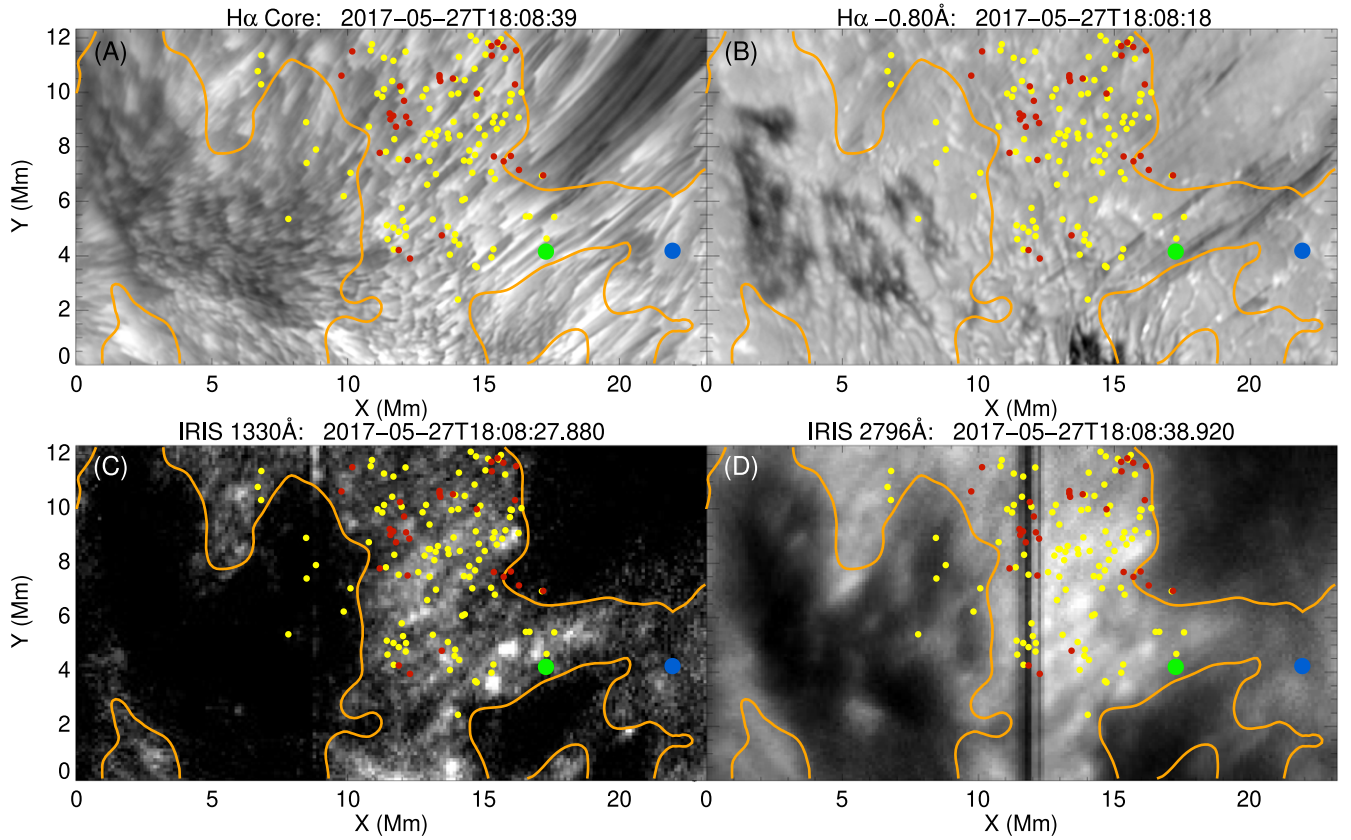


Figure 7. Locations of all the studied SB events in the ROI. (A) and (B) $H\alpha$ core and $H\alpha -0.80 \text{ \AA}$ images, respectively. (C) and (D) SJI 1330 and 2796 \AA images for the same instance. The orange contour, derived from the smoothed SJI 1330 \AA , represents the boundary of the moss region in SJI 1330 \AA . Small yellow dots show the locations of the SB events seen in SJI 1330 \AA , which are related to the spicules in the chromosphere. Small red dots show the locations of the SB events that do not match with any discernible feature in $H\alpha$. The large green dot is the location of the SB event related to the falling plasma. The large blue dot represents the location of the SB event caused by the EB.

generation mechanisms of these brightening events. We have investigated 131 TR SB events in the region and analyzed their dynamical and multiwavelength behavior. Figure 7 illustrates the locations of these 131 TR SB events. Our analysis reveals that 98 SB events show spatial and temporal correspondence with spicules in $H\alpha$ images; their locations are shown as yellow dots in this figure. Most of these SBs in SJI 1330 \AA appear at the tops of the spicules, while most of the SBs in 2796 \AA images appear to map the body of the spicules. These observations show that the majority of SB events in the plage regions accompany spicules. These spicules typically exhibit broadened $H\alpha$ spectral profiles, and some of them (9 events) demonstrate parabolic trajectories in the spacetime diagrams reminiscent of dynamic fibrils. H. Skogsrud et al. (2016) observed that a substantial number of SB events in the plage regions are driven by shocks in the chromosphere, producing dynamic fibrils that appear as jets near the $H\alpha$ line core. In contrast, our statistical investigation reveals that most spicules associated with SB events do not show parabolic trajectories in the spacetime diagrams. N. Vilangot Nhalil et al. (2022) indicated a spatiotemporal correspondence between RBEs and SB events in the active regions. However, our detailed analysis shows that the spicules we observe typically do not follow RBE-like $H\alpha$ spectral profiles.

We also observe an event where an EB in $H\alpha$ images appears very close to the TR SBs and shows a temporal match with them in SJI 1330 and 2796 \AA images. The location of this event is shown as a large blue dot in

Figure 7. Y. Chen et al. (2019c) suggested that if the magnetic reconnection happens well above the temperature minimum region, an EB and a nearby UV burst can have a similar underlying generation mechanism.

Furthermore, our analysis also reveals a correspondence between the stream of falling plasma and subsequent brightenings near its footpoint with EB-like $H\alpha$ spectral profiles. The approximate location of these falling-plasma-related brightening events is shown as a large green dot in Figure 7. The appearance of EB-like $H\alpha$ profiles near the SBs and their temporal match suggest that these have a common origin. Therefore, we report that EB-like $H\alpha$ profiles can be created due to a falling stream of plasma. Our analysis shows that the speed of the downflows is around $15\text{--}20 \text{ km s}^{-1}$. Please note that these speeds are measured in the plane of the sky, and the actual speed can be much higher. Furthermore, these downflows are clearly seen in the $H\alpha +1 \text{ \AA}$ line position, which is equivalent to around 45 km s^{-1} of Doppler shift. The falling plasma with these high speeds is enough to create termination shocks in the chromosphere, which might generate transient brightenings in TR SJIs and $H\alpha$ far wings due to shock heating. The strong $H\alpha$ wing brightenings associated with falling plasma are also reported by J. Jing et al. (2016), where a stream of falling plasma along the flare loops creates intense brightenings in the $H\alpha$ at the ending of the stream. They suggested that when siphon-like plasma flows reach supersonic velocities, they could form a shock wave in the atmosphere and

cause brightenings. We need further investigations to uncover their true generation mechanism.

The red dots in Figure 7 represent 31 SB events with no discernible chromospheric counterparts. Some of these events are localized intensity enhancements, which might correspond to explosive events. This can be confirmed using spectroscopic analysis, which is out of the scope of this study. However, only a slight difference is observed in the physical properties (see Figure 6), which may indicate that they might also correspond to even thinner or less energetic spicules that are not detected in the $H\alpha$ chromospheric images due to observational limits. It is also possible that these SBs may be connected to other higher atmospheric phenomena (e.g., coronal rain) that are not studied in this article. It is worth mentioning that some spicules in the chromosphere do not show heating, and as a result, they do not have spatial and temporal correspondence with any SB events in TR SJI. Some of these spicules could form in regions where the plasma density is relatively high, and the energy is not sufficient to cause heating up to TR temperatures.

5. Conclusions



Our analysis showed that TR SBs observed in IRIS passbands can be attributed to various chromospheric phenomena, including spicules, EBs, and falling-plasma streams. Specifically, we found that a significant portion of regular SBs (98 of 131) accompany spicules. These spicules exhibit properties similar to dynamic fibrils and are likely induced by chromospheric shocks, which impact the TR. Furthermore, we identify SB events associated with EBs, which exhibit significantly higher intensity and longer lifetimes compared to typical SBs in the studied moss region. EBs are believed to stem from magnetic reconnection, and the extended current sheet from this reconnection can trigger brightenings in the TR images. Our investigation also reveals that in certain areas, SBs may arise from falling-plasma streams, producing compact and intense events likely due to termination shocks. The $H\alpha$ line profiles in these regions resemble those of EBs, indicating that falling-plasma streams can also generate mustache-like $H\alpha$ profiles. Additionally, we also found several brightenings (31 of 131) that do not have a clear chromospheric counterpart. We speculate that these events might be generated independently in the TR or may be linked to other higher atmospheric phenomena that are beyond the scope of this study. It is also possible that these brightenings are smaller versions of spicule-related events, with spicules themselves potentially going unnoticed in $H\alpha$ images due to current observational constraints. These findings suggest a diverse range of mechanisms contributing to brightening events occurring in the TR, highlighting the complexity of solar atmospheric dynamics.

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